

Dust growth and planet formation by disc fragmentation

Hans Lee ¹, Sergei Nayakshin ¹★ and Richard A. Booth ²

¹*School of Physics and Astronomy, University of Leicester, Leicester LE1 7RH, UK*

²*School of Physics and Astronomy, University of Leeds, Leeds LS2 9JT, UK*

Accepted 2025 August 29. Received 2025 August 27; in original form 2025 July 11

ABSTRACT

It is often argued that gravitational instability of realistic protoplanetary discs is only possible at distances larger than ~ 50 au from the central star, requiring high disc masses and accretion rates, and that therefore disc fragmentation results in the production of brown dwarfs rather than gas giant planets. However, the effects of dust growth on opacity can be very significant but have not been taken into account systematically in the models of fragmenting discs. We employ dust opacity that depends on both temperature and maximum grain size to evaluate analytically the properties of a critically fragmenting protoplanetary disc. We find that dust growth may promote disc fragmentation at disc radii as small as ~ 30 au. As a result, the critical disc masses and accretion rates are smaller, and the initial fragment masses are in the gas giant planet mass regime. While this suggests that formation of gas giant planets by disc fragmentation may be more likely than usually believed, we caution that numerical models of the process are needed to evaluate the effects not taken into account here, e.g. dust grain mobility and fragment evolution after disc fragmentation.

Key words: planets and satellites: formation – protoplanetary discs.

1 INTRODUCTION

Planet formation by Gravitational Instability (GI; see reviews by Helled et al. 2014; Kratter & Lodato 2016) posits that young massive protoplanetary discs fragment gravitationally into clumps (e.g. Kuiper 1951; Boss 1998), which then evolve further to form planets (Nayakshin 2017a) or brown dwarfs (Stamatellos & Whitworth 2009).

It is universally accepted that to fragment, discs must be massive and cool (Toomre 1964), and also be able to cool rapidly (Gammie 2001; Meru & Bate 2010; Deng, Mayer & Meru 2017); otherwise, a self-regulated gravitoturbulent state develops in which disc density perturbations in the disc constantly appear and dissolve away (Rice, Lodato & Armitage 2005; Cossins, Lodato & Clarke 2009; Paardekooper 2012). The disc fragmentation conditions are believed to be satisfied only at large distances, i.e. $R \gtrsim 50$ au (Rafikov 2005; Clarke & Lodato 2009). Giant planets discovered orbiting their stars on wide orbits (e.g. Marois et al. 2008, 2010; Vigan et al. 2021; Blunt et al. 2023) may have formed via GI.

However, the potential of GI to make planets rather than brown dwarfs is frequently questioned since both analytical estimates (Kratter, Murray-Clay & Youdin 2010; Forgan & Rice 2011) and most numerical simulations (Stamatellos & Whitworth 2008; Zhu et al. 2012; Xu et al. 2025) indicate initial fragment masses $M_{\text{frag}} \gtrsim 10 M_{\text{J}}$, although note that some numerical studies do suggest that sub-Jovian mass gaseous objects can be formed by GI (e.g. Boley et al. 2010; Kubli, Mayer & Deng 2023).

Turbulence in protoplanetary discs may result in very rapid grain growth (Weidenschilling 1984; Dullemond & Dominik 2005). A number of authors (e.g. Rice et al. 2004; Boley & Durisen 2010; Booth & Clarke 2016; Vorobyov & Elbakyan 2018, 2019) pointed out that dust in self-gravitating discs can grow to significant sizes, $a_{\text{max}} \sim 1\text{--}100$ μm in the first ~ 0.1 Myr (e.g. Molyarova et al. 2021). Furthermore, focused in spiral density arms (Rice et al. 2004), large grains can become self-gravitating and collapse into massive solid cores (Gibbons, Rice & Mamatsashvili 2012; Gibbons, Mamatsashvili & Rice 2014; Baehr 2023; Longarini et al. 2023a, b; Rowther et al. 2024), which could later turn into gas giant planets via gas accretion as in the core accretion scenario (Helled et al. 2014) or the pebble accretion scenario (Baehr 2023).

In this Letter, we draw attention to the complementary and potentially stronger effects of dust growth on planet formation by GI due to opacity effects. Although dust makes up approximately 1 per cent of the mass of protoplanetary discs, it contributes $\gtrsim 99.9$ per cent of the total opacity (e.g. cf. fig. 5 in Malygin et al. 2014).

Dust opacity depends significantly on the distribution of particle sizes (e.g. Draine & Lee 1984; Woitke et al. 2016). For a collection of spherical dust particles with uniform size a and fixed total mass, its total geometric cross-sectional area scales as $\propto a^{-1}$, and therefore, the Rosseland mean opacity usually decreases with grain growth (although the behaviour of frequency-dependent opacity is more nuanced, e.g. see the ‘opacity cliff’ in fig. 1 in Rosotti et al. 2019).

Fig. 1 shows several commonly used Rosseland mean opacities for dusty discs. Except for opacities of Zhu et al. (2021), all of these consider only relatively small dust grains, with maximum grain sizes $a_{\text{max}} \lesssim 10$ μm . Zhu et al. (2021), focusing on the physics of planetary atmospheres, calculated dust opacities with grain sizes up to 10 cm.

* E-mail: sergei.nayakshin@le.ac.uk

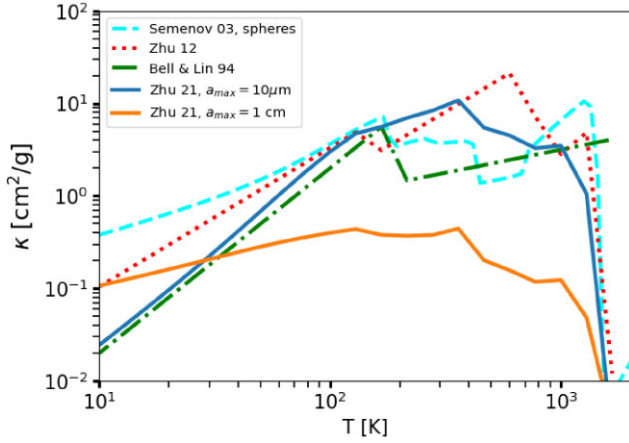


Figure 1. Comparison between dust opacities of Bell & Lin (1994), Semenov et al. (2003), Zhu et al. (2012) at $\rho = 10^{-10} \text{ g cm}^{-3}$, and Zhu et al. (2021) at $a_{\text{max}} = 10 \mu\text{m}$ and $a_{\text{max}} = 1 \text{ cm}$, respectively.

Fig. 1 shows clearly that while dust opacity depends significantly on grain physics, it is most sensitive to a_{max} , and can drop by one-two orders of magnitude in an important range of disc temperatures if grains can grow to sizes larger than $\sim 1 \text{ cm}$.

Below we present an analytic dust opacity fit for dust opacity computed for a wider range of dust sizes than in Zhu et al. (2021), and use it to re-evaluate disc fragmentation conditions and initial fragment masses.

2 DUST OPACITY

2.1 An analytic fit to opacity curves

We use the DIANA standard opacities described in Woitke et al. (2016), utilizing the OpacityTool FORTRAN package (Toon & Ackerman 1981; Dorschner et al. 1995; Zubko et al. 1996; Min, Hovenier & de Koter 2005) to produce tables of dust absorption coefficient κ for a range of wavelengths and a_{max} values. We assume that the number density of dust grains follows a power law: $n(a)da \propto a^{-p}$, where $p = 3.5$, from the minimum $a_{\text{min}} = 0.05 \mu\text{m}$ up to the maximum dust size a_{max} . These opacity tables are calculated for the wavelength range $0.05 \mu\text{m} \leq \lambda \leq 9000 \mu\text{m}$, which covers the blackbody spectrum for the temperature range we are interested in sufficiently well.

We produce 17 of these opacity tables, for values of a_{max} evenly spaced on the logarithmic scale from $1 \mu\text{m}$ to 100 m . Since our focus is on the total cooling rate via radiation in optically thick discs, we calculate the Rosseland mean opacity $\kappa_{\text{DIANA}}(T)$ for each opacity table, which are functions of temperature.

For a given a_{max} , we then fit the DIANA opacity curves to the following piecewise power-law function:

$$\kappa(a_{\text{max}}, T) = \begin{cases} \kappa_0 \left(\frac{T}{155 \text{ K}} \right)^{p_l} & T \leq 155 \text{ K} \\ \kappa_0 & 155 \text{ K} \leq T \leq 353 \text{ K} \\ \kappa_0 \left(\frac{T}{353 \text{ K}} \right)^{p_h} & T \geq 353 \text{ K}, \end{cases} \quad (1)$$

where κ_0 , p_l , and p_h vary with a_{max} . The transition temperatures are chosen empirically according to the shape of the opacity curves, and are not of particular physical significance.

The values of these parameters for each a_{max} are listed in Table 1.

Table 1. Fit parameters κ_0 , p_l , and p_h for the range of a_{max} values computed. For $a_{\text{max}} \geq 10^3 \mu\text{m}$, p_l and p_h are constant. Note that this function gives the opacity for dust only ($z = 1$), and should be multiplied by the metallicity of the system.

$a_{\text{max}} (\mu\text{m})$	κ_0	p_l	p_h
1.0	1680	1.55	1.05
3.16	2160	1.64	0.867
10.0	2490	1.64	0.516
31.6	2040	1.40	0.311
1.0×10^2	1090	0.905	0.369
3.16×10^2	564	0.640	0.432
1.0×10^3	297	0.515	0.442
3.16×10^3	173	—	—
1.0×10^4	98.5	—	—
3.16×10^4	55.8	—	—
1.0×10^5	31.4	—	—
3.16×10^5	17.7	—	—
1.0×10^6	9.92	—	—
3.16×10^6	5.56	—	—
1.0×10^7	3.12	—	—
3.16×10^7	1.75	—	—
1.0×10^8	0.982	—	—

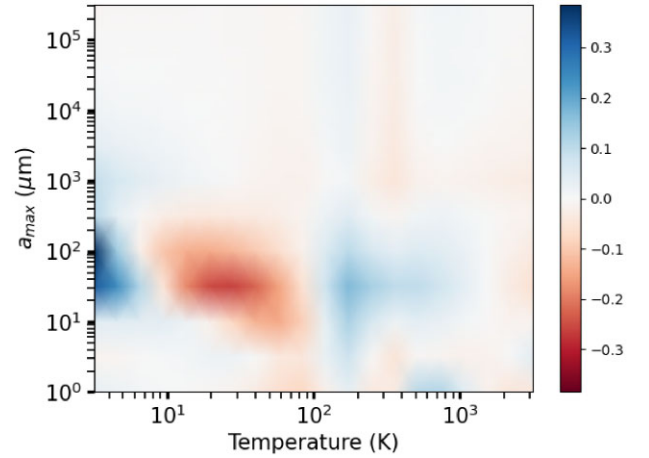


Figure 2. Logarithm of the opacity ratio between the analytic function and the DIANA opacity curves.

Using this, one can obtain intermediate values by interpolating p_l , p_h , and $\log \kappa_0$ linearly with respect to $\log a_{\text{max}}$.

For $a_{\text{max}} \geq 1 \text{ mm}$, the shape of the curve is unchanged on a log-log plot, therefore p_l and p_h are fixed. When a_{max} increases by a factor of 10, κ_0 increases by a factor of ≈ 3 .

We characterize the goodness of fit by calculating the ratio between our function $\kappa(\kappa_0, p_l, p_h)$ and the DIANA opacities κ_{DIANA} by calculating $\log_{10} \left(\frac{\kappa(\kappa_0, p_l, p_h)}{\kappa_{\text{DIANA}}} \right)$, a function of T and a_{max} , which is shown in Fig. 2.

The deviation reaches up to a factor of $10^{0.385} = 2.43$ at $a_{\text{max}} \sim 100 \mu\text{m}$ and $T \sim 3 \text{ K}$, and $10^{0.262} = 1.83$ at $a_{\text{max}} \sim 30 \mu\text{m}$ and $T \sim 25 \text{ K}$. However, Fig. 1 demonstrates that there exists a range of reasonable assumptions about dust composition and physics which results in yet larger uncertainties than the errors of the semi-analytic opacity fit. Therefore, we consider equation (1) to be a sufficiently good fit to the DIANA a_{max} -dependent opacity curves.

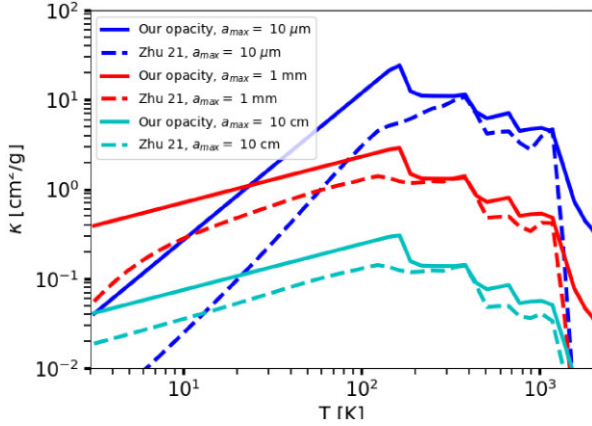


Figure 3. Comparison between Rosseland mean opacities in Zhu et al. (2021) (at $\rho = 10^{-10} \text{ g cm}^{-3}$) and this paper.

Table 2. T_i and f_i used to account for sublimation in our opacity fit.

Material	T_i (K)	f_i
Water ice	166	0.455
Organics	425	0.472
Troilite	680	0.593

2.2 Sublimation

Since the DIANA standard opacity calculation does not account for dust sublimation, we estimate how much opacity remains above each sublimation point using the results from Pollack et al. (1994), scaled according to the material composition used in DIANA (Min et al. 2011). We first consider the sublimation of three species: water ice, organics, and troilite. We represent the effects of this by multiplying the opacity function by $(f_i + (1 - f_i) \exp(\frac{T - T_i}{10 \text{ K}}))$ when $T > T_i$ for all species i . T_i is the sublimation temperature for each species considered, and f_i is the remaining opacity fraction after sublimation. The values used are shown in Table 2, and as shown in Fig. 3, there is a notable sharp decrease in opacity at each sublimation temperature.

For the evaporation of the most refractory grains, we follow Kuiper et al. (2010), multiplying the dust opacity by $0.5 - \frac{1}{\pi} \arctan(\frac{T - T_{\text{evap}}}{100 \text{ K}})$, where $T_{\text{evap}} = 2000(\frac{\rho}{1 \text{ g cm}^{-3}})^{0.0195} \text{ K}$, where ρ is the gas density.

The resulting opacity function is shown in Fig. 3, alongside the dust size-dependent opacities from Zhu et al. (2021). The difference between the curves is mainly due to the assumption we make about the dust composition. For example, the water ice mass fraction we use is about two times that of the one in Zhu et al. (2021). This can be seen in Fig. 3, where our opacity curve has a greater drop at the water ice sublimation temperature. Since our opacity is larger at all temperatures and a_{max} , we opt to use it in the rest of the paper. This approach is conservative with respect to opacity uncertainties as it leads to higher disc and fragment masses at disc fragmentation, which is the main focus of our paper.

3 DISC FRAGMENTATION WITH DUST GROWTH

In this section, we use the opacity function derived in Section 2 to calculate the properties of a critically fragmenting disc.

3.1 Critically fragmenting discs

Pseudo-viscous disc models is a convenient framework to study the disc fragmentation boundary analytically (e.g. Rafikov 2005; Levin 2007; Clarke 2009), with their results being consistent with numerical simulations within a factor of order unity (e.g. fig. 2 in Zhu et al. 2012). In this approach, a marginally stable self-gravitating protoplanetary disc has the Toomre (1964) parameter $Q = \frac{c_s \Omega}{\pi G \Sigma} \simeq 1$, which gives the sound speed $c_s = \frac{\pi G \Sigma}{\Omega}$, where Ω is the orbital angular velocity. The disc mid-plane temperature is then

$$T = \frac{2m_p c_s^2}{k_B} = \frac{2m_p}{k_B} \left(\frac{\pi G \Sigma}{\Omega} \right)^2. \quad (2)$$

Further, in marginally stable discs, heating generated by gravito-turbulence is balanced by radiative cooling (Gammie 2001). Xu et al. (2025) show that the disc's radiative cooling rate is well approximated in both optically thick and optically thin regimes by

$$F_{\text{cool}} = 8\sigma_B (T^4 - T_{\text{irr}}^4) \frac{\tau}{(1 + (0.875\tau^2)^{0.45})^{1/0.45}}, \quad (3)$$

where the optical depth of the disc $\tau = \kappa \Sigma / 2$, and the disc temperature due to stellar irradiation heating $T_{\text{irr}} = \left(\frac{0.1 L_*}{4\pi r^2 \sigma_B} \right)^{1/4}$ (Zhu et al. 2012). The constant factor of 0.1 is a simple estimate of non-local disc irradiation effects, which depend on the disc global H/R profile and properties of the dust at the height of a few H above the disc mid-plane (e.g. Chiang & Goldreich 1997). These effects, including possible disc self-shadowing (e.g. Bell 1999), cannot be computed in a local disc stability calculation. None the less, our choice of 0.1 is consistent with the flaring angle in our calculations within a factor of order unity. Supplementary materials provide examples of how our results vary depending on the choice of L_* , which is equivalent to varying the irradiation angle.

On the other hand, the heating rate due to gravitoturbulence with the corresponding viscosity parameter α is given by

$$F_{\text{heat}} = \frac{3}{8\pi} \Omega^2 \dot{M} = \frac{9}{8} \alpha \Omega c_s^2 \Sigma = \frac{9}{8} \frac{\alpha (\pi G)^2 \Sigma^3}{\Omega}. \quad (4)$$

Setting $F_{\text{cool}} = F_{\text{heat}}$ for a self-regulating disc, one obtains an equation consisting only of T , Ω , α , and physical constants. This equation can be solved numerically to find the critical mid-plane temperature T_{crit} of the disc as a function of Ω for a fixed α (Levin 2007). Since we are interested in disc fragmentation, we set $\alpha = \alpha_{\text{crit}} \simeq 0.06$, above which the disc fragments (Rice et al. 2005). Note that, in the language of self-gravitating discs in thermal balance between turbulent heating and radiative cooling, this choice of α_c corresponds to the 'β' cooling parameter of $\beta = 4/[9\gamma(\gamma - 1)\alpha_c] \approx 7$ (cf. equation 20 in Gammie 2001) for $\gamma = 5/3$ (as appropriate for the cold regions of the disc; Boley et al. 2007).

From that, we calculate the corresponding disc scale height, $H_{\text{crit}} = 0.69 \frac{c_s}{\Omega}$ (Kratter & Lodato 2016, equation 11 with $Q = 1$), surface density, Σ_{crit} (cf. equation 2), and thus arrive at the critical accretion rate above which the disc should fragment (e.g. Clarke 2009; Zhu et al. 2012), as well as the initial fragment mass, which we take to be $M_{\text{frag}} = 57 \Sigma_{\text{crit}} H_{\text{crit}}^2$ (Xu et al. 2025). While the factor in front of $\Sigma_{\text{crit}} H_{\text{crit}}^2$ differs between different studies by an order of magnitude (see Kratter & Lodato 2016), here we follow the larger value derived by Xu et al. (2025) in their radiation hydrodynamics simulations, which leads to a more conservative estimate of the initial fragment masses.

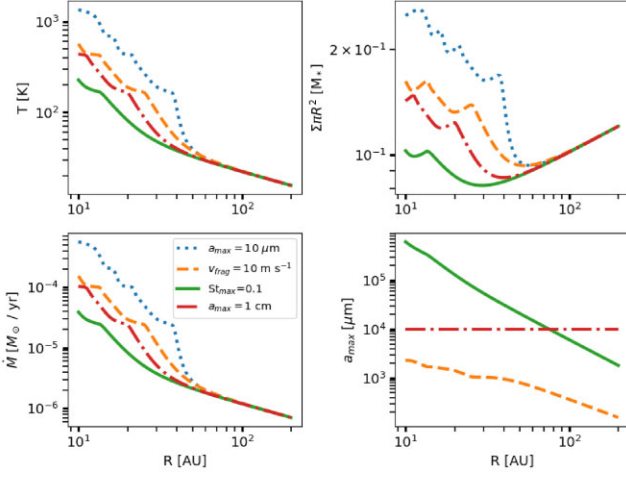


Figure 4. Plotted against R : T_{crit} (upper left), $\Sigma_{\text{crit}}\pi R^2$ (upper right), \dot{M} (lower left), a_{max} (lower right). The cases are $a_{\text{max}} = 10\ \mu\text{m}$ (dotted), $v_{\text{frag}} = 10\ \text{m s}^{-1}$ (dashed), $\text{St}_{\text{max}} = 0.1$ (solid), and $a_{\text{max}} = 1\ \text{cm}$ (dash-dotted). We choose to rescale the critical surface density plot by πR^2 to better illustrate the difference between the four cases.

The maximum grain size in a marginally stable disc depends on many parameters of the problem and is a non-local quantity due to dust dynamics being a function of grain size (e.g. Rice et al. 2004; Boley & Durisen 2010; Booth & Clarke 2016). Below, we consider four representative scenarios to evaluate the effects of dust growth on disc fragmentation. Two of the cases assume a fixed a_{max} independent of disc radius, R , $a_{\text{max}} = 10\ \mu\text{m}$ and $a_{\text{max}} = 1\ \text{cm}$, respectively, roughly indicating the ‘no growth’ and the significant grain growth limits.

For the third case, we estimate dust growth limited by the grain fragmentation velocity, which we take to be $v_{\text{frag}} = 10\ \text{m s}^{-1}$. We calculate the maximum dust size as $a_{\text{max}} = \frac{2\Sigma}{\pi\alpha\rho_s} \frac{v_{\text{frag}}^2}{c_s^2}$ (Birnstiel, Dullemond & Brauer 2009), where ρ_s is the material density of solid dust, which we set $3\ \text{g cm}^{-3}$ here. This model assumes that similarly sized grain collisions dominate both grain growth and fragmentation, and occur with relative velocity $\Delta v \sim \sqrt{\alpha}\ \text{St}\ c_s$. Other mechanisms of grain growth such as sweep-up of small grains by larger grains (Xu & Armitage 2023) may allow for growth significantly beyond this barrier, but we do not consider these effects here.

Finally, we use $\text{St} = 0.1$ as an optimistic upper limit to the grain sizes for our fourth case because previous work has shown that self-gravity produces particle motions that are correlated when $\text{St} < 1$ (e.g. Booth & Clarke 2016). As a result, collisions may occur at lower velocities than those estimated by the Birnstiel et al. (2009) formulae, unless the large-scale motions associated with the spirals can be efficiently converted into small-scale turbulence. While small-scale motion has been identified in high-resolution 3D simulations (Riols, Latter & Paardekooper 2017; Booth & Clarke 2019), its effect on the dust growth has not been quantified. Instead, we estimate the most optimistic case by assuming that the only source of collisions between dust grains is due to the differential motion created by the spiral-induced pressure gradients. Following the arguments laid out in Booth & Clarke (2016), this produces $\text{St} \lesssim 0.1$ for our parameters. The maximum dust size is then found via $a_{\text{max}} = 2\Sigma\text{St}/(\pi\rho_s)$.

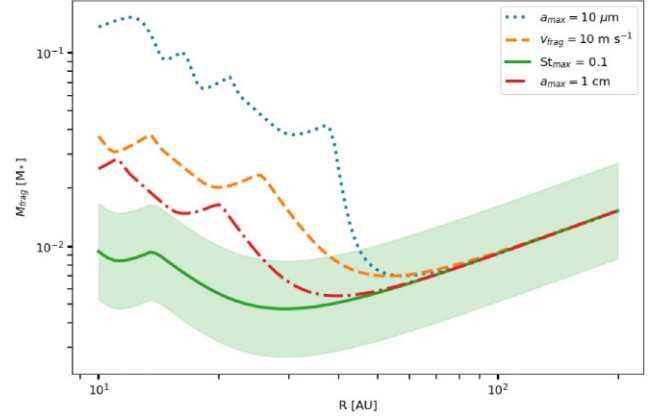


Figure 5. M_{frag} plotted against R , for $a_{\text{max}} = 10\ \mu\text{m}$ (dotted), $v_{\text{frag}} = 10\ \text{m s}^{-1}$ (dashed), $\text{St}_{\text{max}} = 0.1$ (solid), and $a_{\text{max}} = 1\ \text{cm}$ (dash-dotted). 1σ uncertainty is shown around the $\text{St}_{\text{max}} = 0.1$ case (Xu et al. 2025).

3.2 Results

Fig. 4 shows various properties of the critically fragmenting disc as functions of R , demonstrating that dust growth can lead to fragmentation in colder, less massive, and geometrically thinner discs. We observe that up to $R \sim 40\ \text{au}$, when dust grains are allowed to grow from $a_{\text{max}} = 10\ \mu\text{m}$ to $\text{St}_{\text{max}} = 0.1$, T_{crit} decreases by a factor of 5 to 10, and Σ_{crit} decreases by a factor of ≈ 3 .

In the plot of a_{max} , we see that for the radius range we consider, $\text{St}_{\text{max}} = 0.1$ is strictly above $a_{\text{max}} = 1\ \text{cm}$, which is then strictly above $v_{\text{frag}} = 10\ \text{m s}^{-1}$.

Dust growth also decreases the critical accretion rate \dot{M}_{crit} for fragmentation. From equation (4), $\dot{M} = 3\pi\alpha c_s^2 \Sigma \Omega^{-1}$, implying $\dot{M}_{\text{crit}} \propto T_{\text{crit}} \Sigma_{\text{crit}}$. Since T_{crit} and Σ_{crit} decrease when dust growth is present, \dot{M}_{crit} decreases accordingly. Hydrodynamical simulations of disc formation from protostellar cloud collapse show that accretion rates on the order of $\dot{M} \sim 10^{-5}\ \text{M}_{\odot}\ \text{yr}^{-1}$ can be expected (Vorobyov & Basu 2015). In the ISM-like case, we find that $\dot{M} > \dot{M}_{\text{crit}}$ only for $R > 40\ \text{au}$, while this condition is satisfied as close as $R \sim 20\ \text{au}$ in the $\text{St} = 0.1$ case.

Fig. 5 shows the initial fragment masses for the two cases. Xu et al. (2025) showed that clump masses forming in their simulations are distributed lognormally with the 1σ spread of $\sigma(\log M_{\text{frag}}) = 0.57$, and this spread is shown using the band around our dust growth case. We observe from the figure that the minimum M_{frag} is about a factor of two lower when dust growth up to $\text{St}_{\text{max}} = 0.1$ is allowed. While in itself this is a relatively small effect, the minimum fragment mass $M_{\text{min}} \approx 5\text{M}_J$ occurs at $R_{\text{min}} \approx 30\ \text{au}$, which is much smaller than for the ISM-like dust, when the minimum is at $R_{\text{min}} \sim 60\ \text{au}$. Since protoplanetary discs are observed to extend to $20\text{--}30\ \text{au}$ on average (Trapman et al. 2023; Guerra-Alvarado et al. 2025), the possibility of planet-mass fragments at small R is especially important.¹

In the case of $a_{\text{max}} = 1\ \text{cm}$, although the minimum is at $R_{\text{min}} \approx 40\ \text{au}$, which is larger than the average disc extent stated above, the formation of planet-mass clumps at $R = 20\text{--}30\ \text{au}$ remains possible. For $v_{\text{frag}} = 10\ \text{m s}^{-1}$, planet-mass clumps are less likely to form at

¹Although note that the epoch of disc fragmentation we are most concerned here could be limited to much earlier disc ages, $t \lesssim 0.5\ \text{Myr}$, while material infall from the dusty envelope is ongoing (e.g. Vorobyov & Basu 2006, 2015), and these discs could be more extended than older ALMA discs.

that radius range, but not impossible due to the spread in fragment masses as shown in Fig. 5 for the $\text{St}_{\text{max}} = 0.1$ case.

Since the critical β is directly correlated to our choice of α_c , we have explored the effects of varying α_c by a factor of 2 higher and lower, which correspond to $\beta \approx 3$ and $\beta \approx 15$, respectively. It is found that \dot{M} increases by a factor of ~ 2 with higher α_c , and vice versa. However, T , Σ , and a_{max} show basically no change for $R > 50$. For $R < 50$, the quantities vary by no more than 10 per cent, increasing and decreasing with α_c . The profile also shifts radially by a few au, outwards for higher α_c and vice versa. The resulting M_{frag} also varies by ~ 10 per cent with higher α_c and vice versa. To summarize, an increase in the critical β can lower the critical accretion rate, though it only weakly facilitates fragmentation otherwise.

We have also explored the effects of stellar mass, disc metallicity, and irradiation on fragmentation conditions and fragment masses (available as supplementary material). In general, fragmentation is favoured in low opacity and irradiation environments (e.g. lower mass stars, irradiation shielding, metal-poor systems), which is consistent with findings from simulations (Meru & Bate 2010).

4 SUMMARY AND DISCUSSION

We investigated the effects of dust growth on disc fragmentation conditions. We find that dust growth beyond $a_{\text{max}} \sim 1$ mm reduces the disc opacity sufficiently to allow fragmentation to occur in colder, less massive, discs, and at smaller radii, i.e. $R \sim 30$ versus $R \sim 60$ au (cf. Fig. 4) for the standard ISM dust opacity. Dust growth also reduces the initial fragment mass (Fig. 5). As a result, disc fragmentation on planetary mass clumps, $M_{\text{frag}} \leq 10 M_J$, may be more likely than usually concluded, but we note that our analytical calculations rely on a number of assumptions; numerical simulations are needed to explore the problem further.

We assumed that dust properties depend only on the radial coordinate, and that the metallicity (which encompasses both solids and vapourized materials) is constant throughout the disc volume. Numerical simulations show large inhomogeneities in the distribution of both solids (e.g. Boley & Durisen 2010; Forgan & Rice 2011; Booth & Clarke 2016; Longarini et al. 2023b; Birnstiel 2024) and volatile materials (Molyarova et al. 2021) in self-gravitating discs. These effects are multi-faceted and future dedicated numerical studies are required. For example, vertical dust settling to the disc mid-plane may reduce disc column-averaged dust opacity even further compared to our estimations. However, preferential dust accumulation in spiral density arms may increase the opacity there, making these regions less likely to fragment on to clumps.

Furthermore, simulations show that after forming, the fragments evolve significantly within the disc. The fragments gain mass by merging (Hall, Forgan & Rice 2017) or by accreting material from the disc (Zhu et al. 2012; Stamatellos & Inutsuka 2018), yet they can lose mass due to tidal effects (e.g. Nayakshin 2017b; Vorobyov & Elbakyan 2018). Therefore, simulations are required to establish the eventual outcome of disc fragmentation. We note in passing that grain growth will also affect gas accretion on to the fragments since gas can only accrete on to a planet if the cooling rate is sufficiently high (Pollack et al. 1996; Ayliffe & Bate 2009).

ACKNOWLEDGEMENTS

We thank Dimitris Stamatellos for helpful comments on the draft. HL acknowledges Science and Technology Facilities Council PhD studentship funding. RAB acknowledges support from the Royal

Society through a University Research Fellowship and Enhanced Expenses Award.

DATA AVAILABILITY

The opacity function and parameter table are available online.²

REFERENCES

- Ayliffe B. A., Bate M. R., 2009, *MNRAS*, 397, 657
 Baehr H., 2023, *MNRAS*, 523, 3348
 Bell K. R., 1999, *ApJ*, 526, 411
 Bell K. R., Lin D. N. C., 1994, *ApJ*, 427, 987
 Birnstiel T., 2024, *ARA&A*, 62, 157
 Birnstiel T., Dullemond C. P., Brauer F., 2009, *A&A*, 503, L5
 Blunt S. et al., 2023, *AJ*, 166, 257
 Boley A. C., Durisen R. H., 2010, *ApJ*, 724, 618
 Boley A. C., Hartquist T. W., Durisen R. H., Michael S., 2007, *ApJ*, 656, L89
 Boley A. C., Hayfield T., Mayer L., Durisen R. H., 2010, *Icarus*, 207, 509
 Booth R. A., Clarke C. J., 2016, *MNRAS*, 458, 2676
 Booth R. A., Clarke C. J., 2019, *MNRAS*, 483, 3718
 Boss A. P., 1998, *ApJ*, 503, 923
 Chiang E. I., Goldreich P., 1997, *ApJ*, 490, 368
 Clarke C. J., 2009, *MNRAS*, 396, 1066
 Clarke C. J., Lodato G., 2009, *MNRAS*, 398, L6
 Cossins P., Lodato G., Clarke C. J., 2009, *MNRAS*, 393, 1157
 Deng H., Mayer L., Meru F., 2017, *ApJ*, 847, 43
 Dorschner J., Begemann B., Henning T., Jaeger C., Mutschke H., 1995, *A&A*, 300, 503
 Draine B. T., Lee H. M., 1984, *ApJ*, 285, 89
 Dullemond C. P., Dominik C., 2005, *A&A*, 434, 971
 Forgan D., Rice K., 2011, *MNRAS*, 417, 1928
 Gammie C. F., 2001, *ApJ*, 553, 174
 Gibbons P. G., Rice W. K. M., Mamatsashvili G. R., 2012, *MNRAS*, 426, 1444
 Gibbons P. G., Mamatsashvili G. R., Rice W. K. M., 2014, *MNRAS*, 442, 361
 Guerra-Alvarado O. M., van der Marel N., Williams J. P., Pinilla P., Mulders G. D., Lambrechts M., Sanchez M., 2025, *A&A*, 696, A232
 Hall C., Forgan D., Rice K., 2017, *MNRAS*, 470, 2517
 Helled R. et al., 2014, *Protostars and Planets VI*. Univ. Arizona Press, Tucson, p. 643
 Kratter K. M., Lodato G., 2016, *ARAA*, 54, 271
 Kratter K. M., Murray-Clay R. A., Youdin A. N., 2010, *ApJ*, 710, 1375
 Kubli N., Mayer L., Deng H., 2023, *MNRAS*, 525, 2731
 Kuiper G. P., 1951, in Hynek J. A., ed., *Proceedings of a Topical Symposium, Commemorating the 50th Anniversary of the Yerkes Observatory and Half a Century of Progress in Astrophysics*. McGraw-Hill, New York, p. 357
 Kuiper R., Klahr H., Beuther H., Henning T., 2010, *ApJ*, 722, 1556
 Levin Y., 2007, *MNRAS*, 374, 515
 Longarini C., Lodato G., Bertin G., Armitage P. J., 2023a, *MNRAS*, 519, 2017
 Longarini C., Armitage P. J., Lodato G., Price D. J., Ceppi S., 2023b, *MNRAS*, 522, 6217
 Malygin M. G., Kuiper R., Klahr H., Dullemond C. P., Henning T., 2014, *A&A*, 568, A91
 Marois C., Macintosh B., Barman T., Zuckerman B., Song I., Patience J., Lafrenière D., Doyon R., 2008, *Science*, 322, 1348
 Marois C., Zuckerman B., Konopacky Q. M., Macintosh B., Barman T., 2010, *Nature*, 468, 1080
 Meru F., Bate M. R., 2010, *MNRAS*, 406, 2279
 Min M., Hovenier J. W., de Koter A., 2005, *A&A*, 432, 909
 Min M., Dullemond C. P., Kama M., Dominik C., 2011, *Icarus*, 212, 416

²https://github.com/hanslee6/dust_growth_opacity

- Molyarova T., Vorobyov E. I., Akimkin V., Skliarevskii A., Wiebe D., Güdel M., 2021, *ApJ*, 910, 153
- Nayakshin S., 2017a, *Publ. Astron. Soc. Aust.*, 34, e002
- Nayakshin S., 2017b, *MNRAS*, 470, 2387
- Paardekooper S.-J., 2012, *MNRAS*, 421, 3286
- Pollack J. B., Hollenbach D., Beckwith S., Simonelli D. P., Roush T., Fong W., 1994, *ApJ*, 421, 615
- Pollack J. B., Hubickyj O., Bodenheimer P., Lissauer J. J., Podolak M., Greenzweig Y., 1996, *Icarus*, 124, 62
- Rafikov R. R., 2005, *ApJ*, 621, L69
- Rice W. K. M., Lodato G., Pringle J. E., Armitage P. J., Bonnell I. A., 2004, *MNRAS*, 355, 543
- Rice W. K. M., Lodato G., Armitage P. J., 2005, *MNRAS*, 364, L56
- Riols A., Latter H., Paardekooper S. J., 2017, *MNRAS*, 471, 317
- Rosotti G. P., Tazzari M., Booth R. A., Testi L., Lodato G., Clarke C., 2019, *MNRAS*, 486, 4829
- Rowther S., Nealon R., Meru F., Wurster J., Aly H., Alexander R., Rice K., Booth R. A., 2024, *MNRAS*, 528, 2490
- Semenov D., Henning T., Helling C., Ilgner M., Sedlmayr E., 2003, *A&A*, 410, 611
- Stamatellos D., Inutsuka S.-i., 2018, *MNRAS*, 477, 3110
- Stamatellos D., Whitworth A. P., 2008, *A&A*, 480, 879
- Stamatellos D., Whitworth A. P., 2009, *MNRAS*, 392, 413
- Toomre A., 1964, *ApJ*, 139, 1217
- Toon O. B., Ackerman T. P., 1981, *Appl. Opt.*, 20, 3657
- Trapman L., Rosotti G., Zhang K., Tabone B., 2023, *ApJ*, 954, 41
- Vigan A. et al., 2021, *A&A*, 651, A72
- Vorobyov E. I., Basu S., 2006, *ApJ*, 650, 956
- Vorobyov E. I., Basu S., 2015, *ApJ*, 805, 115
- Vorobyov E. I., Elbakyan V. G., 2018, *A&A*, 618, A7
- Vorobyov E. I., Elbakyan V. G., 2019, *A&A*, 631, A1
- Weidenschilling S. J., 1984, *Icarus*, 60, 553
- Woitke P. et al., 2016, *A&A*, 586, A103
- Xu W., Armitage P. J., 2023, *ApJ*, 946, 94
- Xu W., Jiang Y.-F., Kunz M. W., Stone J. M., 2025, *ApJ*, 986, 91
- Zhu Z., Hartmann L., Nelson R. P., Gammie C. F., 2012, *ApJ*, 746, 110
- Zhu Z., Jiang Y.-F., Baehr H., Youdin A. N., Armitage P. J., Martin R. G., 2021, *MNRAS*, 508, 453
- Zubko V. G., Mennella V., Colangeli L., Bussoletti E., 1996, *MNRAS*, 282, 1321

SUPPORTING INFORMATION

Supplementary data are available at *MNRAS* online.

Dust.GI.Letter.Supp.pdf

Please note: Oxford University Press is not responsible for the content or functionality of any supporting materials supplied by the authors. Any queries (other than missing material) should be directed to the corresponding author for the article.

This paper has been typeset from a \LaTeX file prepared by the author.