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LETTER TO THE EDITOR

The dynamics of small-scale magnetic fields modulated by the solar cycle

M. Stangalini^{1,*}, G. Verth², V. Fedun³, D. Perrone¹, M. Berretti^{4,5}, P. Bushby⁶, S. Jafarzadeh^{7,8},
D. B. Jess^{7,9}, F. Giannattasio¹⁰, P. H. Keys⁷, R. Bruno¹¹, and F. Berrilli⁵

¹ ASI Italian Space Agency, Via del Politecnico snc, 00133 Rome, Italy

² Plasma Dynamics Group, School of Mathematical and Physical Sciences, The University of Sheffield, Hicks Building, Hounsfield Road, Sheffield S3 7RH, UK

³ Plasma Dynamics Group, School of Electrical and Electronic Engineering, The University of Sheffield, Mappin Street, Sheffield S1 3JD, UK

⁴ University of Trento, Via Calepina 14, 38122 Trento, Italy

⁵ University of Rome Tor Vergata, Department of Physics, Via della Ricerca Scientifica 3, 00133 Rome, Italy

⁶ School of Mathematics, Statistics and Physics, Newcastle University, Newcastle upon Tyne NE1 7RU, UK

⁷ Astrophysics Research Centre, School of Mathematics and Physics, Queen's University Belfast, Belfast BT7 1NN, Northern Ireland, UK

⁸ Niels Bohr International Academy, Niels Bohr Institute, Blegdamsvej 17, DK-2100 Copenhagen, Denmark

⁹ Department of Physics and Astronomy, California State University Northridge, Northridge, CA 91330, USA

¹⁰ INGV, Istituto Nazionale di Geofisica e Vulcanologia, Rome, Italy

¹¹ INAF-IAPS, Istituto Nazionale di Astrofisica, Rome, Italy

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ABSTRACT

In addition to sunspots, the most easily visualized manifestation of solar magnetism, cutting-edge observations of the solar atmosphere have uncovered a plethora of magnetic flux tubes, down to the resolving power of modern high-resolution telescopes (a few tens of kilometers), revealing how the Sun is a fully magnetized star. These magnetic elements are advected and buffeted by ambient plasma flows and turbulent convection, resulting in perturbations of the flux tubes that make them natural conduits for channeling wave energy into the upper layers of the Sun's atmosphere and significantly contributing to the acceleration of the solar wind. Data acquired by the Helioseismic and Magnetic Imager (HMI) on board NASA's Solar Dynamics Observatory (SDO) have made it possible to study the dynamics of small-scale magnetic fields over long timescales. Here, for the first time, we present the discovery of a modulation in the dynamical behavior of small-scale magnetic concentrations in the photosphere over temporal scales consistent with the solar activity cycle (i.e., 11 years), which has only been made possible by the long observing lifetime of the SDO/HMI spacecraft. Furthermore, we also find a temporal varying polarization of their perturbations on similar timescales. This demonstrates how the small-scale dynamics of magnetic fields are also affected by the global dynamo. These discoveries were realized through automated tracking of magnetic fields in the solar photosphere over 11 continuous years, resulting in the most extended statistical analysis of its kind so far, with more than 31 million magnetic concentrations examined.

Key words. Sun: magnetic fields – Sun: oscillations – Sun: photosphere

1. Introduction

The solar photosphere presents magnetic features on a wide range of scales, from tens of megameters (typical scales for sunspots) down to spatial scales close to the resolving power of modern high-resolution telescopes (50–100 km; [de Wijn et al. 2009](#); [Lagg et al. 2010](#); [López Ariste & Sainz Dalda 2012](#)). Subject to the turbulent forcing of the ambient photospheric plasma, small-scale magnetic fields are advected and diffused over the solar surface ([Abramenko et al. 2011](#); [Lepreti et al. 2012](#); [Giannattasio et al. 2013](#); [Jafarzadeh et al. 2017](#)). This forcing can also excite magnetohydrodynamic waves ([Edwin & Roberts 1983](#); [Stangalini et al. 2014](#); [Jess & Verth 2016](#); [Jess et al. 2023](#)), which can propagate upward and take part in the energization of the outer layers of the solar atmosphere and

thus in the acceleration of the solar wind ([Hasan et al. 2003](#); [Hasan & van Ballegoijen 2008](#); [Matsumoto & Shibata 2010](#); [Sakurai 2017](#); [Van Doorselaere et al. 2020](#)).

From high-resolution observations of the solar atmosphere, many authors have reported horizontal perturbations of small-scale magnetic elements with velocity amplitudes on the order of 1–2 km/s ([Keys et al. 2011](#); [Jafarzadeh et al. 2013](#); [Stangalini et al. 2013](#)), consistent with the horizontal velocity flows of the ambient photospheric plasma ([Matsumoto & Kitai 2010](#)). However, these analyses were based on data sequences limited to 1–2 hours maximum and thus could not capture a possible evolution of their dynamics over longer timescales.

Launched in 2012, and with more than 10 years of continuous and stable observations of the solar photospheric magnetic fields, the Helioseismic and Magnetic Imager (HMI) magnetograph ([Scherrer et al. 2012](#)) on board NASA's Solar

* Corresponding author; marco.stangalini@asi.it

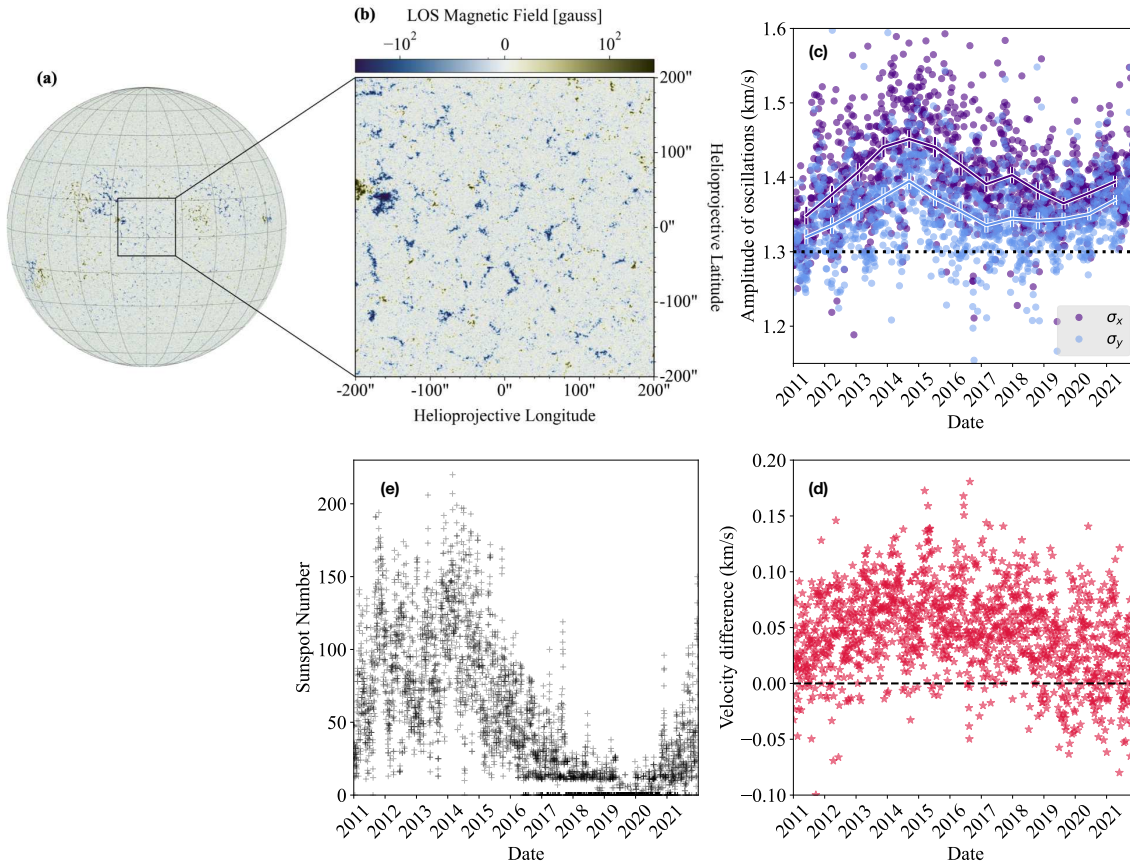


Fig. 1. Panel a: example of a full-disk magnetogram acquired by SDO/HMI. Panel b: FoV considered in the analysis. Magnetic elements were automatically detected and tracked for 2 hours every three days in the period 2011–2022. Panel c: mean amplitude of the horizontal oscillations of the magnetic elements in the east-west (x) and north-south (y) directions. The error bars represent the standard error of the mean. Panel d: difference in the amplitude in the two directions. Panel e: sunspot number in the same period, indicating the progression of the solar cycle. The dynamics of small-scale and short-lived magnetic elements in the solar atmosphere is affected by a long-term modulation on timescales that are much longer (years) than their lifetimes (several minutes). These dynamical properties cannot be explained solely by a local driver of the perturbations and indicate the presence of a global scale effect.

Dynamics Observatory (SDO) mission (Pesnell et al. 2012) has enabled the investigation of the dynamics of small-scale fields in the solar atmosphere over temporal scales typical of the solar activity cycle. This allows the investigation of the dynamics of photospheric magnetic flux tubes and their possible link to global-scale phenomena such as the dynamo. We utilized the unprecedented long-term observational capabilities of SDO/HMI to conduct a comprehensive study of small-scale magnetic concentration dynamics. Over 31 million such features were analyzed across a complete solar cycle, enabling investigation of their relationship to global solar activity.

2. Dataset

We used sequences of photospheric magnetograms obtained in the Fe I 617.3 nm spectral line and acquired by the HMI (Scherrer et al. 2012) on board the NASA SDO (Pesnell et al. 2012) with a cadence of 45 s. Specifically, magnetic flux tubes in a region of $400 \times 400 \text{ arcsec}^2$ ($290 \times 290 \text{ Mm}^2$), located at the center of the solar disk, were tracked in HMI magnetograms (Figs. 1a and 1b), and we analyzed their horizontal motions by employing an automatic tracking algorithm. For more details about the data used and the tracking, we refer the reader to Berretti et al. (2024). This analysis spans 11 years, resulting in an accumulated set of more than 31 million small-scale magnetic

concentrations. We considered data sequences of HMI magnetograms of 40 minutes every 3 days in the period 1 January 2011–31 December 2021, thus covering an entire solar activity cycle. The choice of the length of the data segments was made to allow the detection of the typical photospheric frequencies (i.e., 3 mHz) while maintaining the data volume at a reasonable level. In this regard, it is also worth considering that the average lifetime of the tracked magnetic elements is ~ 20 min. Additionally, since solar rotation introduces a characteristic timescale of ~ 27 days, a three-day sampling ensures that different solar longitudes are systematically covered, mitigating potential aliasing effects.

3. Methods and results

The tracking of photospheric magnetic elements in a small area, $400 \times 400 \text{ arcsec}^2$, at disk center was done using the SWAMIS code (DeForest et al. 2007). The code searches for magnetic “blobs” above $B = 40 \text{ G}$ that have an area of at least four pixels and are identified in at least three consecutive images. These stringent thresholds taken together ensure the reliability of the tracking information of the magnetic elements in the photosphere and resulted in more than 31 million magnetic elements tracked in total. Each subset of 40 min was co-aligned to correct for the solar rotation. However, it is worth noting that the change

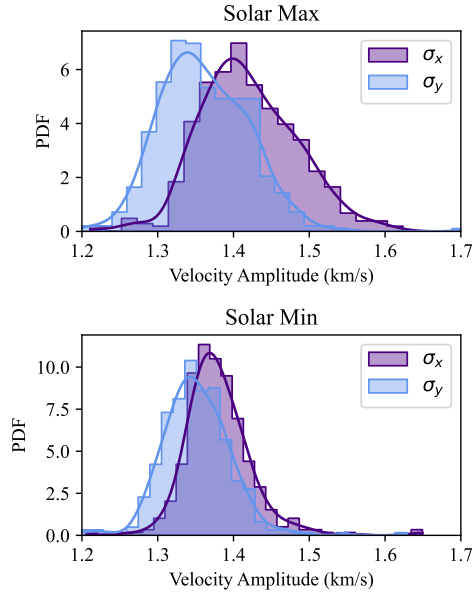


Fig. 2. Probability density function (PDF) of the horizontal oscillations of the magnetic structures in both the x and y direction at the solar maximum and minimum.

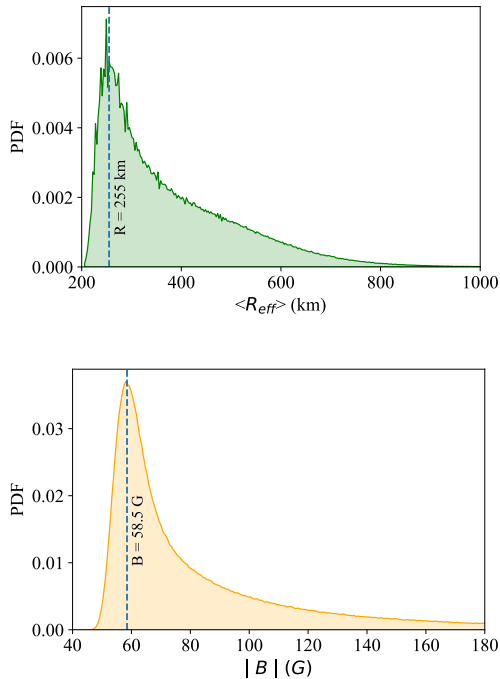


Fig. 3. Top: probability density function of the effective radius of the magnetic elements investigated in this work. The vertical dashed line indicates the most frequent value, $R = 255$ km. Bottom: probability density function of the mean magnetic field of the magnetic elements investigated in this work. The vertical line indicates the most frequent value of the magnetic field in the magnetic elements. The magnetic field values correspond to the mean values within the area of the elements.

in the degree of polarization of the horizontal perturbations seen in Fig. 1 cannot be ascribed to an effect of the tracking, as in this case it would be independent of the solar cycle. The tracking allowed the position of each magnetic feature in time to be determined, and from that we could estimate its horizontal displacement and oscillations. In this study the amplitude of the

horizontal (kink) oscillations was taken as the standard deviation of the horizontal velocity along the equatorial (east-west) and north-south axis (σ_x and σ_y , respectively; see Fig. 1c). In each temporal window of 40 minutes, the average values of σ_x and σ_y of all the magnetic elements identified in the field of view (FoV) were considered. This was done over the entire 11-year period considered, resulting in an estimate of the average amplitude of the perturbations every three days (Fig. 1c). The statistical set of tracked magnetic elements is constituted by magnetic concentrations with an effective radius peaking at about 255 km. A more in-depth characterization of the flux tubes tracked in this work and used in the analysis, as well as their statistical properties, is shown in Fig. 3. The same data were also used in Berretti et al. (2024).

We find that the velocity amplitude of the horizontal perturbations of photospheric magnetic flux concentrations is not constant, but instead undergoes a long-term variation that follows the progression of the solar activity cycle (Fig. 1c). The horizontal velocity amplitudes resulting from the buffeting action of the photospheric plasma flows are lowest at the point of solar minimum and increase as the solar cycle develops. This behavior points to a change in the driver of these perturbations on timescales consistent with that of the activity cycle. Furthermore, we also find a polarization of the direction of their perturbations, with the amplitude of the horizontal perturbations in the equatorial direction increasing faster than their counterparts aligned along the north-south direction, implying a preferential direction of the underlying driver. This can be seen in Fig. 1d, where the differences in the velocity amplitudes in the two perpendicular directions are plotted. Once again, we stress that both the amplitude and the polarization are reduced at the solar minimum. This can also be seen in Fig. 2, where we plot the probability density functions of the horizontal oscillations of the magnetic elements at both the solar maximum and minimum. Here we see that despite the distribution of the amplitude of oscillations being broader in both the x and y direction at the solar maximum, in the x direction (aligned with the equator) the peak in the distribution is shifted toward the right, while in the y direction it remains in the same position.

It is worth underlining that what we investigated is not the bulk velocity of the magnetic elements, but the oscillation amplitude with respect to it. This implies that, although the data were co-aligned before the analysis, any residual co-alignment errors resulting in a slow trend in the velocity time series associated with a magnetic element would not impact the results.

4. Discussion

Our results show that the horizontal perturbations of the magnetic elements, with typical periodicities on the order of minutes, are not only driven by local plasma processes (e.g., buffeting of the flux tubes due to granular convection) but also by a long-term process with a characteristic timescale consistent with the solar cycle (i.e., several years). In other words, in addition to the locally acting driver due to the photospheric buffeting, there exists a larger-scale driver acting on much longer timescales, those typical of the solar dynamo and activity cycles. It is worth noting that both the amplitude variation and the polarization are highest at the peak of the solar cycle and decrease at the minima. This excludes the possibility of an effect of co-alignment of the data, which would also manifest at the minima.

The observed modulation in the dynamics of these small-scale magnetic features over the solar cycle appears somewhat reminiscent of the (so-called) torsional oscillations

(Howard & Labonte 1980; Vorontsov et al. 2002; Howe et al. 2013; Getling et al. 2021), which are cyclic perturbations in the solar differential rotation profile. At low latitudes, the equatorward drift of the torsional oscillation pattern follows that of the sunspot cycle. Although the details of the driving mechanism are still not fully understood, the period of oscillation and the radial phase lag (at least at low latitudes; see, e.g., Vorontsov et al. 2002) suggest that the torsional oscillations are driven by the oscillatory large magnetic field within the solar interior (e.g., Covas et al. 2004; Bushby 2005; Guerrero et al. 2016; Pipin & Kosovichev 2019). Extending a similar analysis to higher latitudes could provide useful insights in this regard could help in preparing the polar observations of Solar Orbiter in the coming years. Our key result is that there are indications that this dynamical influence is not confined to larger scales; it also affects the smaller ones explored by HMI near the surface. In fact, our observations suggest that the large-scale magnetic fields can also influence the dynamics of much smaller-scale magnetic features. However, the precise mechanism behind this possible coupling remains unclear. One possibility is that convective flows that are perturbing these magnetic elements are subject to some weak modulation by the large-scale solar magnetic field. An arguably more likely explanation is that (at least some of) the magnetic features tracked might be relics of active region decay (Hagenaar & Shine 2005; Švanda et al. 2021), with roots deep enough to couple with the global scale magnetic field. As the underlying large-scale magnetic field varies across the solar cycle, it is plausible that we would then observe some signatures of this variation in the dynamics of these magnetic elements. In this regard, higher-resolution observations from, for example, DKIST (Daniel K. Inouye Solar Telescope; Rimmele et al. 2020) could provide important information in this context, allowing the study of the photospheric magnetic field dynamics down to very small scales (below 100 km). We note that the elements tracked in this study are mostly represented by magnetic fields at the limit of HMI resolution, and therefore they could be considered mostly as network fields or magnetic fields associated with the emergence of active regions. Whatever the underlying explanation, the solar-cycle variations in the horizontal perturbations of the magnetic elements in this study suggest that the large-scale solar magnetic field has a surprisingly important dynamical influence on the short-timescale evolution of small-scale magnetic fields, not only in the subsurface layers but also at photospheric heights and beyond.

5. Conclusions

Thanks to the availability of long-term continuous observations of the Sun's magnetic field with the SDO/HMI spacecraft, we studied the long-term modulation of the dynamical properties of small-scale magnetic structures in the photosphere. We find a clear correlation with the solar activity cycle, suggesting a possible connection between the global scales of the Sun's magnetic field cyclic regeneration, namely the solar dynamo, and the dynamics of small-scale structures in quiet-Sun regions. Interestingly, a time-dependent polarization of the velocity perturbations is also observed with timescales compatible with that of the solar dynamo, further complicating our understanding of the driving mechanism. Future efforts will be directed toward the investigation of the nontrivial coupling of different temporal and spatial scales thanks to upcoming missions and further advancements in simulations.

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