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Superresolution trends in the ALMA Taurus survey: structured inner discs and compact discs

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ABSTRACT

The 1.33-mm survey of protoplanetary discs in the Taurus molecular cloud found annular gaps and rings to be common in extended sources ($\gtrsim 55$ AU), when their 1D visibility distributions were fit parametrically. We first demonstrate the advantages and limitations of *non-parametric* visibility fits for data at the survey's 0.12-arcsec resolution. Then we use the non-parametric model in Frankenstein (FRANK) to identify new substructure in three compact and seven extended sources. Among the new features, we identify three trends: a higher occurrence rate of substructure in the survey's compact discs than previously seen, underresolved (potentially azimuthally asymmetric) substructure in the innermost disc of extended sources, and a 'shoulder' on the trailing edge of a ring in discs with strong depletion at small radii. Noting the shoulder morphology is present in multiple discs observed at higher resolution, we postulate it is tracing a common physical mechanism. We further demonstrate how a superresolution FRANK brightness profile is useful in motivating an accurate parametric model, using the highly structured source DL Tau in which FRANK finds two new rings. Finally, we show that sparse (u, v) plane sampling may be masking the presence of substructure in several additional compact survey sources.

Key words: techniques: high angular resolution – techniques: image processing – techniques: interferometric – submillimetre: general – submillimetre: planetary systems – accretion, accretion discs.

1 BACKGROUND

Numerous physical mechanisms are capable of producing axisymmetric (and in many cases also asymmetric) substructures in protoplanetary discs, with the list of candidates growing. Categories include forming and newly formed planets (e.g. Goldreich & Tremaine 1979; Lin & Papaloizou 1986; Kley & Nelson 2012); opacity effects due to ice sublimation fronts (e.g. Okuzumi et al. 2016; Zhang et al. 2016; Hu et al. 2019); gas-dust coupling effects, including preferential dust growth in localized regions (e.g. Pinilla et al. 2012; Dullemond et al. 2018; Sierra et al. 2019), gravitational instability (e.g. Dipierro et al. 2015; Dong, Najita & Brittain 2018b; Hall et al. 2018), dynamical effects of a central binary (e.g. Ragusa et al. 2017; Price et al. 2018; Longarini et al. 2021), and internal photoevaporation (e.g. Clarke, Gendrin & Sotomayor 2001; Alexander, Clarke & Pringle 2006; Ercolano, Clarke & Drake 2009); and magnetic field effects including dead-zone boundaries (e.g. Varnière & Tagger 2006; Flock et al. 2015; Pinilla et al. 2016), magnetic flux concentration and zonal flows (e.g. Johansen, Youdin & Klahr 2009; Bai & Stone 2014; Cui & Bai 2021), and the vertical shear instability (e.g. Flock et al. 2017; Manger & Klahr 2018; Pfeil & Klahr 2021).

Determining which of these mechanisms dominate in observed systems requires both in-depth studies of individual sources and a

large ensemble of discs with characterized substructure. Interferometric observations offer the highest spatial resolution to characterize disc features, and numerous works at the best resolutions achieved to-date in the (sub-)mm, ≈ 25 – 75 mas (≈ 1 – 10 AU), have confirmed (along with many critical works at moderate resolution) that the \approx mm dust distribution in these discs is commonly structured. At these high resolutions, the DSHARP (Andrews et al. 2018; Huang et al. 2018) and ODISEA (Cieza et al. 2021) surveys, as well as several high-resolution case studies of individual systems (e.g. ALMA Partnership et al. 2015; Andrews et al. 2016; Clarke et al. 2018; Dong et al. 2018a; Kudo et al. 2018; Sheehan & Eisner 2018; Keppler et al. 2019; Pérez et al. 2019; Pinte et al. 2019; Tsukagoshi et al. 2019; Huang et al. 2020; Benisty et al. 2021; Casassus et al. 2021; Hashimoto et al. 2021; Macías et al. 2021), have identified a ubiquity of annular gaps and rings, as well as multiple instances of asymmetric arcs (crescents) and spiral arms, in the continuum emission.

When applied to such high-resolution observations, superresolution techniques that fit the observed visibilities directly, such as GALARTO (Tazzari, Beaujean & Testi 2018) and FRANK (Jennings et al. 2020), have found a yet greater occurrence rate of disc substructure. This includes identification of previously unseen features across the DSHARP (Andrews et al. 2021; Jennings et al. 2022) and ODISEA surveys (Cieza et al. 2021); in compact sources, including those that appear featureless in a CLEAN image (Kurtovic et al. 2021; Pinilla et al. 2021); and for observations at the highest available ALMA resolutions, such as in PDS 70 (Benisty et al. 2021).

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The next question is whether superresolution techniques are also able to identify more substructure in moderate resolution observations. This would be particularly valuable for a statistical approach to substructure characterization over a large sample of discs, enabling a fuller investigation of demographic trends by exploiting the large archive of data sets at $\approx 100\text{--}300$ mas. This archive includes many discs that are not particularly large or bright, which current models predict should also contain substructure in order to counteract radial drift and retain reasonable dust disc sizes on few Myr time-scales (Toci et al. 2021). We can ask for example whether compact discs that routinely appear smooth in CLEAN images are intrinsically featureless, or if this tends to be an artefact of observational or model resolution.

Long et al. (2018, 2019) demonstrated at the survey level that parametric visibility fits can identify more substructure in moderate resolution (120 mas, ≈ 16 AU) observations than the CLEAN images alone. Here we will push superresolution visibility fits to still higher resolution, using the *non-parametric* approach in FRANK to fit the observed visibilities yet more accurately. This will allow us to investigate how much more substructure in the Taurus survey data can be recovered from the observed visibilities – including in compact sources – and whether the identified features suggest new trends.

In this work, we characterize new substructure in 10 of the Taurus survey discs using the 1D code FRANK, which reconstructs a disc’s brightness profile at superresolution scales by non-parametrically fitting the azimuthally averaged visibility distribution.¹ Section 2 summarizes the FRANK modeling approach and its limitations. Section 3 more closely examines the major advantages (Section 3.1) and limitations (Sections 3.2–3.3) of non-parametric visibility fitting for data sets at the Taurus survey resolution, exploring how they affect substructure inference in FRANK fits to these observations. In Section 4, we present fits for the 10 sources, grouping substructure findings into trends in compact discs (Section 4.1) and extended discs (Section 4.2). We further divide the extended sources into those with an inner and outer disc (Section 4.2.1) and those with an inner cavity (Section 4.2.2). Section 5 summarizes our findings and briefly places them in the context of superresolution substructure found in data sets outside the survey.

2 MODEL

A full description of the FRANK model framework and its limitations is in Jennings et al. (2020). In short, FRANK reconstructs the 1D (axisymmetric) brightness profile of a source as a function of disc radius by directly fitting the real component of the deprojected, unbinned visibilities as a function of baseline. The brightness profile is determined non-parametrically by fitting the visibilities with a Fourier–Bessel series, which is linked to the real space profile by a discrete Hankel transform. A Gaussian process regularizes the fit, with the covariance matrix non-parametrically learned from the visibilities under the assumption that this matrix is diagonal in Fourier space. The free parameters (diagonal elements) of the matrix correspond to the power spectrum of the reconstructed brightness profile. The fitting procedure takes $\lesssim 1$ min on a standard laptop for each data set shown here.

To obtain the results shown in this work, we vary three of the five FRANK model hyperparameters across data sets: R_{\max} , N , and α . The hyperparameters R_{\max} and N simply set the maximum radius of the fit and number of brightness points in the fit, which we increase for larger discs. α controls the prior on the Gaussian process, effectively

determining the signal-to-noise ratio (SNR) threshold at which the model no longer attempts to fit the data. By varying α , we can thus account for the unique visibility distribution and noise properties of each data set, with higher α values imposing a stronger constraint that in practice causes the model to stop fitting the data at shorter maximum baseline. Most of the Taurus survey data sets become noise-dominated at their longest baselines, as (u, v) plane sampling becomes increasingly sparse. We will thus choose α such that we fit the data sets out to long baselines, but stop before fitting clearly noise-dominated data (using $\alpha \in [1.01, 1.10]$). Pushing a fit out to long baselines to extract higher resolution information does none the less come at the cost of fitting some noise. The noise imprints on the brightness profile as short period, low amplitude oscillations; we will note non-trivial instances.

There are three major limitations in the current version of FRANK:

(i) A FRANK fit drives to a visibility amplitude of zero once it stops fitting the data. This is intentional, given the difficulty of generically extrapolating a fit beyond the edge of the observed visibilities, but we often expect the true visibility distribution would continue oscillating beyond the longest observed baselines if the disc is sufficiently structured. We will motivate how this affects substructure inference in data set characteristic of the Taurus survey in Section 3.2. Ultimately, this issue stems from the ill-posed nature of reconstructing the sky brightness from Fourier data, and it is also why the uncertainty on a FRANK brightness profile is easily underestimated, particularly for deep gaps (CLEAN brightness profiles can similarly exhibit underestimated uncertainties for this reason). We thus will not show uncertainties in the FRANK profiles in this work.

(ii) The 1D approach in FRANK fits for the azimuthal average of the visibility data at each baseline. While this is an accurate representation of the azimuthally averaged brightness profile, in the presence of azimuthal asymmetries the brightness profile should be interpreted with caution, as (particularly superresolution) asymmetries can be misidentified as annular features. We will demonstrate this in Section 3.3.

(iii) The FRANK real space model is not positive definite and so can exhibit regions of small amplitude, negative brightness. When this unphysical behaviour occurs we can enforce positivity by finding the most probable brightness profile for a given set of power spectrum parameters and the constraint that the brightness be non-negative, using a non-negative least-squares solver. This sometimes alters features across the disc (i.e. not just in regions of negative brightness) because the enforced positivity condition affects the visibility fit at long baselines. We will remove this limitation in a forthcoming work and version of the code by fitting in logarithmic brightness space, but for the current analysis we will show non-negative fits for those FRANK models that would otherwise exhibit regions of negative brightness; we will note which fits include this correction.

2.1 Data reduction

In this work, we reanalyse the ALMA Taurus survey published by Long et al. (2018, 2019), to which we refer for details on the observational setup and calibration procedure. To apply FRANK to the data sets, we first apply channel averaging (one channel per spectral window) and time averaging (60 s) to all spectral windows in the self-calibrated measurement set, then extract the unflagged visibilities. We then use the disc geometries and phase centres in Long et al. (2019) to deproject the visibilities in FRANK prior to fitting their 1D distribution. After deprojection, we re-estimate the weights by a constant factor of order unity to approximate the relation $w =$

¹The code is available at <https://discsim.github.io/FRANK>.

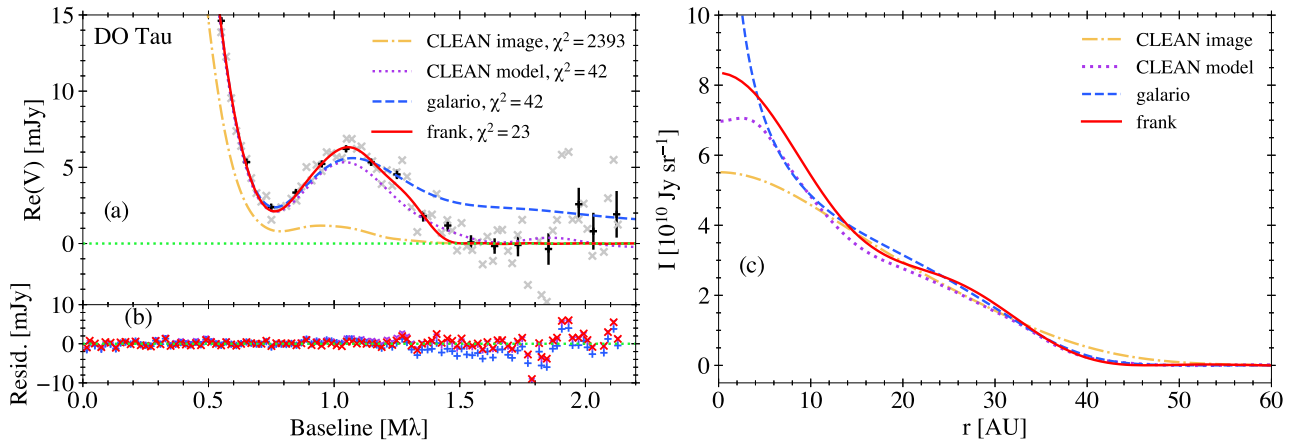


Figure 1. Improved visibility fit accuracy better resolves disc structure. Panel (a): a zoom on the Taurus survey visibilities for DO Tau (20- and 100-k λ bins, with 1σ uncertainties shown for the 100-k λ points); the parametric fit from Long et al. (2019); the non-parametric FRANK fit; and the Fourier transform of brightness profiles extracted from the CLEAN image and CLEAN model. Panel (b): residuals of the parametric and FRANK visibility fits and the CLEAN model transform (20-k λ bins). Panel (c): brightness profiles for DO Tau corresponding to the visibility fits in panel (a).

$1/\sigma^2$, where w is the weight of a visibility point and σ^2 is the variance of its real and imaginary components.

3 METHODOLOGIES – ADVANTAGES AND LIMITATIONS OF A 1D, NON-PARAMETRIC VISIBILITY FIT

Here we examine the benefits and drawbacks of 1D, non-parametric visibility fits (both generally and specific to FRANK) for brightness profile reconstruction at resolutions typical of the Taurus survey, ≈ 120 mas.

3.1 Advantages – a highly accurate fit to the observed data

Recovering superresolution structure in a brightness profile with a 1D visibility model is a matter of fit accuracy; even a modest improvement in accuracy can correspond to new or more highly resolved profile features. To demonstrate how a non-parametric visibility fit’s improved accuracy can better constrain superresolution structure in Taurus survey data, Fig. 1 compares the parametric visibility fit from Long et al. (2019) for the compact disc DO Tau with the non-parametric FRANK fit.² Long et al. (2019) inferred that structure in the visibility distribution for this source indicates a sharp outer edge in the brightness profile, and so they modeled the profile parametrically as an exponentially tapered power law. The resulting visibility fit in Fig. 1(a) is more accurate than the Fourier transform of a brightness profile extracted from the CLEAN image.³ The parametric visibility fit’s improved accuracy in turn corresponds to superresolution structure recovery in the brightness profile; this structure is also apparent in a profile extracted from the

CLEAN model.⁴ While the parametric visibility fit is accurate at short and intermediate baselines, its residuals in Fig. 1(b) show non-trivial error at long baselines. By comparison, the FRANK visibility model in Fig. 1(a) and its residuals in (b) demonstrate a yet higher accuracy across intermediate and long baselines.

We can quantify an improvement in fit accuracy with the χ^2 statistic,

$$\chi^2 = \sum_{k=1}^N w_k [Re(V_{k, \text{obs}}) - Re(V_{k, \text{fit}})]^2,$$

where we neglect the imaginary component of the visibilities because FRANK only fits the real component. As given in the legend of Fig. 1(a), both the parametric fit and the Fourier transform of a brightness profile extracted from the CLEAN model exhibit a smaller χ^2 than the transform of a profile extracted from the CLEAN image by a factor of 57.0 for this source, while the FRANK fit yields a further reduction of the χ^2 value by a factor of 1.8. This comparatively small improvement in fit accuracy with FRANK corresponds to a clear change in the disc morphology in the FRANK brightness profile in panel (c), with the bump at 28 AU in the FRANK profile not seen in the parametric profile and only hinted at in the CLEAN model profile. (Note that while the CLEAN model profile has lower integrated flux than the FRANK profile – because there is visibility information left in the residuals during the CLEAN process – reducing the `tclean` threshold value also results in fitting more noise.) Thus even in a data set with a simple visibility distribution and relatively featureless brightness profile, a fairly small improvement to the accuracy of a visibility fit can non-trivially inform the scale and location of superresolution structure in the recovered profile. This is the main advantage of a non-parametric fit, and it motivates why, for the sources in Section 4 that all exhibit more structured visibility distributions than DO Tau, a more accurate visibility fit with FRANK yields new brightness profile

²All visibility fits from Long et al. (2018) and Long et al. (2019) shown in this work are obtained by taking the 1D Fourier transform of their bestfit GALARIO brightness profiles.

³This difference is primarily due to the resolution loss induced by CLEAN beam convolution, which results in the transform of the CLEAN image poorly representing the observed visibilities. While we should thus not expect the transform of a CLEAN image profile to be accurate at long baselines, we will include this visibility profile in comparisons throughout this work because the CLEAN .image is the most common imaging product on which analysis is conducted in this field.

⁴All CLEAN brightness profiles for Taurus survey data in this work are extracted from CLEAN model images (the .model output of `tclean`) and convolved images (the .image output of `tclean`) generated using `tclean` in CASA 5.6.1-8 with the multiscale deconvolver (pixel size of 30 mas and scales of 1, 2, 4, 6 pixels); a threshold of 3σ , where σ is the rms noise measured in a region of the image far from the source; and Briggs weighting with a robust value of 0.5.

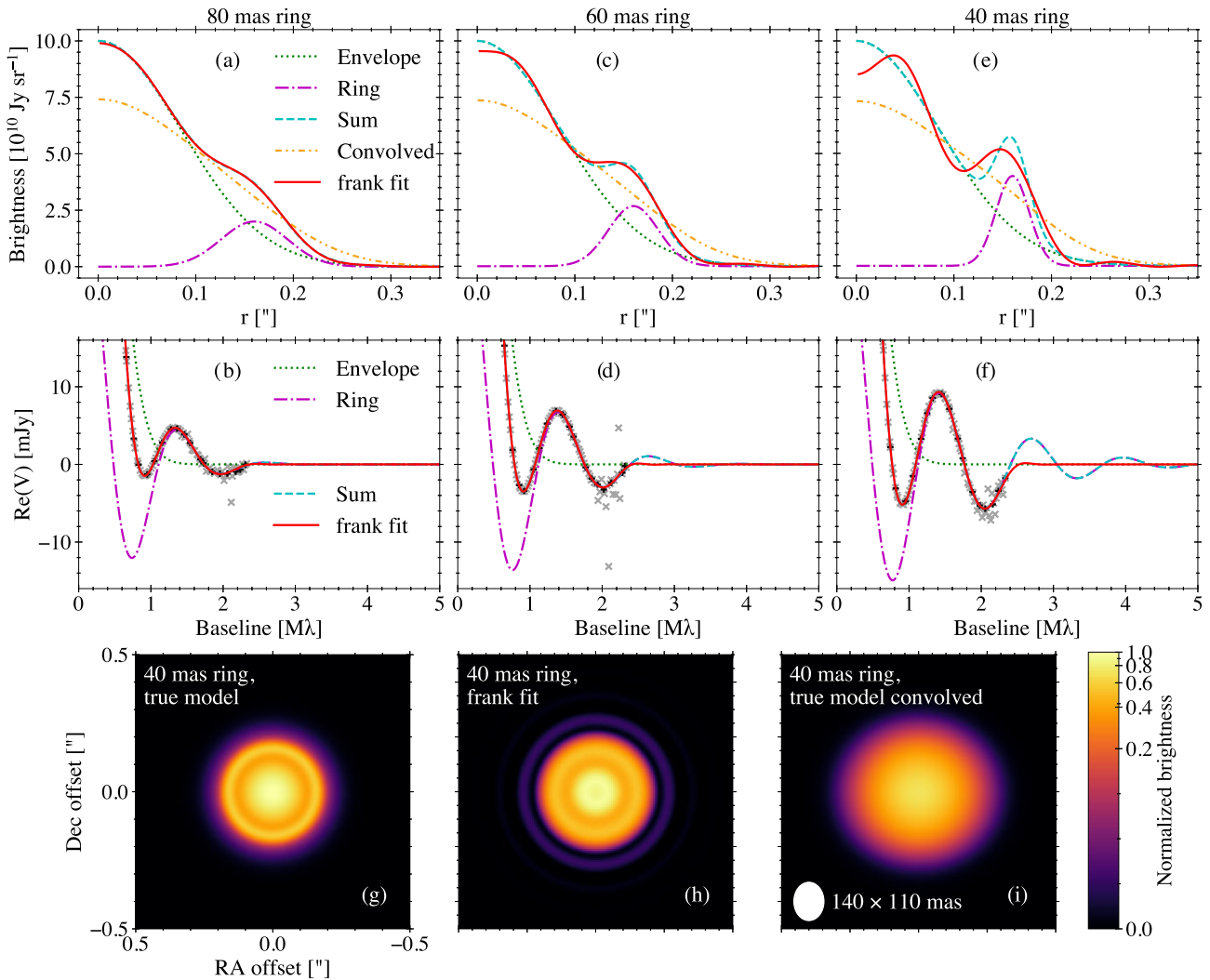


Figure 2. *FRANK* brightness profile accuracy decreases as a ring becomes increasingly superresolution. Panel (a): mock brightness profile of a compact disc with a shallow Gaussian ring; the profile is the sum of a Gaussian envelope and Gaussian ring, with each component shown. The ring’s FWHM is given in the plot title. The *FRANK* recovery of the summed profile is also shown, as is the summed profile convolved with a 140×110 mas beam. Panel (b): the real component of the 1D Fourier transform of the Gaussian envelope, ring, and their sum (the summed profile peaks at ≈ 150 mJy). Also shown are noisy mock observations of the summed profile (20- and 100-k λ bins), and the *FRANK* fit to these mock data. Panels (c)–(d) and (e)–(f): as in panels (a)–(b), but with the Gaussian ring’s FWHM successively decreased and amplitude correspondingly increased to conserve the disc’s total (2D) flux. Panels (g)–(i): the noiseless, true model image of the disc in panel (e), the 1D *FRANK* fitted profile swept over 360° in azimuth, and the true model convolved with the 140×110 mas beam. The images use an arcsinh stretch ($I_{\text{stretch}} = \text{arcsinh}(I/a) / \text{arcsinh}(1/a)$, $a = 0.02$) and the same absolute brightness normalization.

features (as well as more highly resolved known features) relative to the parametric fits and the *CLEAN* models.

3.2 Limitations – extrapolating the fit to unobserved baselines

A fair question to then ask is how much we trust the morphology of features in a superresolution profile. An important consideration is that even superresolution fits can still be expected to underresolve most disc features (even broad ones, albeit to a lesser extent), as is evident when comparing fits to lower and higher resolution observations of the same source (differing in resolution by a factor of say 3). True features in a disc that are highly superresolution (very roughly, a factor $\gtrsim 3$ narrower than the nominal spatial resolution) tend to be inaccurately recovered in a *FRANK* fit, and in some cases they can induce erroneous oscillations in the brightness profile, as we will now show.

Fig. 2 demonstrates the accuracy of a *FRANK* fit to a disc with superresolution features using mock data. In Figs 2(a)–(b), we first consider a simple disc – the sum of a Gaussian envelope and a shallow Gaussian ring whose 80-mas full width at half-maximum (*FWHM*) is superresolution relative to the ≈ 120 -mas full width at half-maximum of the mock observations by a factor of ≈ 1.5 .⁵ While the profile convolved with a 140×110 mas beam (typical of the Taurus survey) in panel (a) shows no clear indication of the superresolution ring, fitting the visibilities in panel (b) with *FRANK* gives an accurate recovery of the true brightness profile. But if we then narrow the ring to 60 mas (and increase its surface brightness to conserve total flux) in Figs 2(c)–(d), it is now superresolution by a factor of ≈ 2 , and the *FRANK* recovered profile

⁵The mock data set is generated with a baseline distribution and noise properties that emulate the Taurus survey observations of DR Tau.

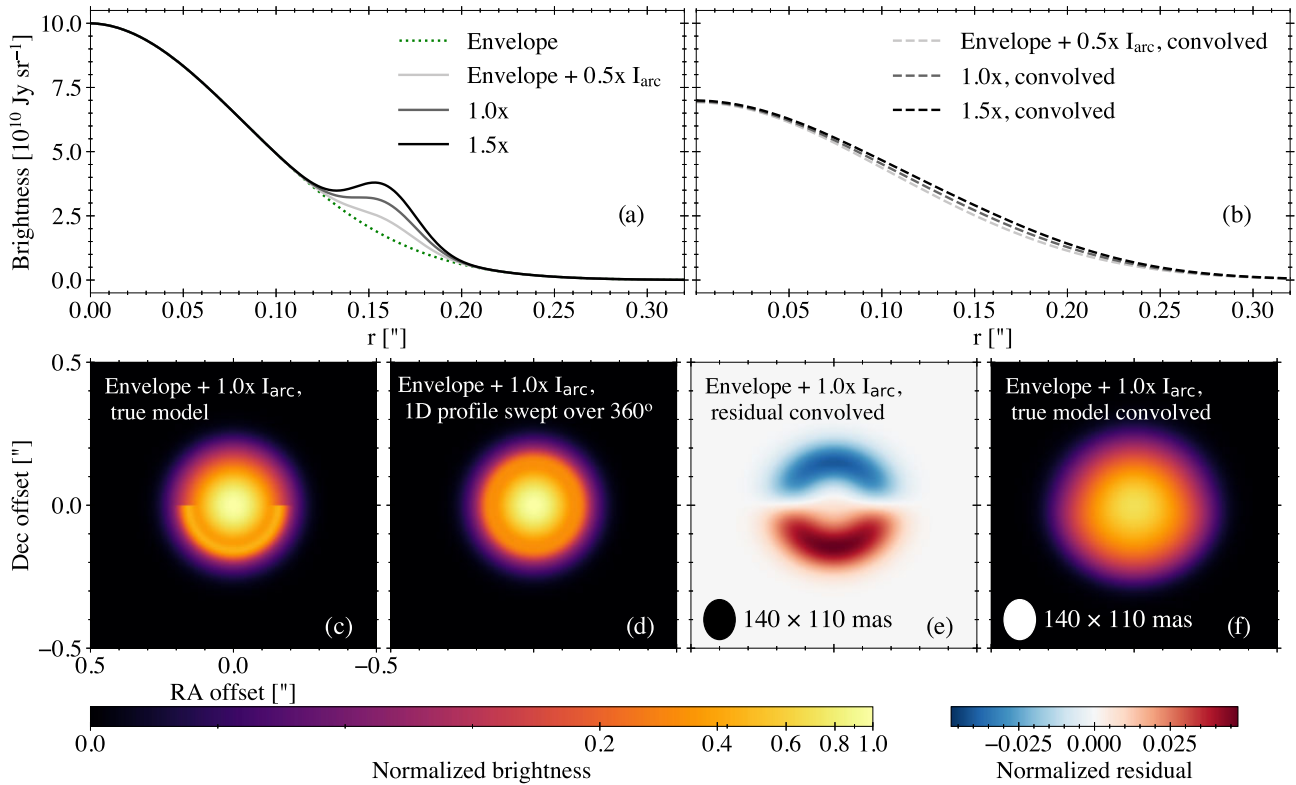


Figure 3. Low-contrast, asymmetric substructure emulates an underresolved annular ring in a brightness profile. Panel (a): the brightness profile of a Gaussian envelope, as well as the azimuthally averaged profile of the envelope summed with a 40-mas arc that spans 180° in azimuth. The summed profile is shown for various amplitudes of the arc (0.5, 1.0, and 1.5 times an arbitrary value of $I_{\text{arc}} = 2.67 \times 10^{10}$ Jy sr^{-1}). Panel (b): brightness profiles obtained from the 2D images of the envelope + arc, when convolved with a Gaussian beam whose 140×110 mas size is typical of the Taurus survey. Panels (c)–(f): for the $1.0x I_{\text{arc}}$ case, the noiseless, true model image; the 1D brightness profile extracted from this image (which is shown in panel (a)) swept over 360° ; the residual between these two images, convolved with the beam; and the true model image convolved with the beam (corresponding to the brightness profile in panel (b)). The disc images use an arcsinh stretch ($I_{\text{stretch}} = \text{arcsinh}(I/a) / \text{arcsinh}(1/a)$, $a = 0.02$) and the same absolute brightness normalization; the residual image uses a linear stretch symmetric about zero.

begins to show some clear inaccuracy. It exhibits a plateau around 0.16 arcsec, underresolving the true gap/ring pair. This is due to an inaccurate extrapolation of the FRANK visibility fit beyond the mock observation’s longest baselines, where the true profile’s visibility distribution continues to oscillate. A further consequence of the fit’s underestimated visibility amplitudes at unsampled baselines is the underestimated peak brightness in the FRANK brightness profile.

Narrowing and brightening the ring even further so that it has a 40-mas FWHM (superresolution by a factor of ≈ 3) in Figs 2(e)–(f), the convolved profile in panel (e) (and the 2D image of this profile swept over 2π in panel (i)) still shows no hint of the ring. The FRANK profile in panel (e) identifies the gap/ring pair, but underresolves the feature amplitudes and misidentifies their centroids. The FRANK profile also underestimates the peak brightness more severely, showing an erroneous turnover near $r = 0$. This turnover is a consequence of the narrower ring in the true profile increasing the absolute visibility amplitudes at all baselines; accurately fitting the higher amplitude features in the visibilities introduces higher contrast structure into the brightness profile. Because the FRANK fit has a visibility amplitude of ≈ 0 beyond the edge of the data, while the true visibility distribution has non-trivial amplitude there, these higher contrast structures are not well constrained. This effect also introduces the erroneous, shallow bump into the FRANK brightness profile between 0.2 and 0.3 arcsec, appearing in the 2D image of the swept FRANK brightness profile in Fig. 2(h) as a faint but fake ring (compare the true 2D image in panel (g)). It is thus possible

for highly superresolution features in a true brightness profile to introduce erroneous oscillations into a FRANK brightness profile.

For some data sets in the Taurus survey such as DO Tau in Fig. 1(a), this is not much of a concern, as the observed visibilities appear to plateau at zero at the longest baselines. But for other data sets, it is less clear whether higher resolution and/or deeper observations would show the visibilities to continue oscillating beyond the baselines at which the current data become noise-dominated. While any extrapolation of a fit beyond the data’s longest baselines is highly uncertain, it can be useful to compare a FRANK brightness profile to that obtained with a parametric visibility fit, where the parametric profile’s functional form is motivated for example by the FRANK fit or by structure in the observed visibilities (as Long et al. 2018, 2019 have done). We will perform an in-depth comparison in Section 4.2.1 for the most structured disc in our results, DL Tau.

3.3 Limitations – distinguishing azimuthally symmetric from asymmetric substructure

If we have found a superresolution feature in a disc, the next question is whether it is an annular ring (gap) or an azimuthally asymmetric brightness excess (depletion). Because a 1D brightness profile averages the flux in a given annulus over 2π in azimuth, a low – moderate contrast asymmetric feature within that annulus can mimic an underresolved (or shallow) ring in the profile. Fig. 3 demonstrates this with mock data, using a Gaussian disc with an

additional brightness ‘arc’ that is produced by sweeping a Gaussian ring only over 180° in azimuth in panel (c). The arc emulates a brightness excess on top of the background envelope, and a 1D profile in Fig. 3(a) extracted from the image in panel (c) shows a slight bump at the arc’s radial location. From the brightness profile alone this could be misidentified as an annular feature, and because the arc is superresolution by a factor of ≈ 3 , the true model image convolved with a 140×110 mas beam in panel (f) – and the corresponding convolved brightness profile in panel (b) – show no clear indication of it. When we increase the asymmetry’s brightness by 50 per cent, it emulates a shallow gap/ring pair in (a), while the convolved profile in panel (b) is effectively unchanged.

How then can we distinguish superresolution asymmetries from annular features? We do not have an unambiguous method for this, so jointly consider three metrics: contouring the CLEAN image, identifying structure in the imaginary component of the visibilities, and imaging a FRANK fit’s residual visibilities. The first of these, contouring the CLEAN image at levels of the rms noise, can be useful in identifying the convolved representation of superresolution asymmetries. A limitation is that low contrast or sufficiently narrow features are often not identifiable. Second, while an asymmetric feature is represented in the real component of the 1D visibilities exactly as an annular feature at the same location and that has the same width and total surface brightness (as integrated over 360° in azimuth),⁶ structure in the imaginary component of the visibilities indicates scales at which there is asymmetry with respect to the phase centre. A limitation here is that without a robust model to fit $\text{Im}(V)$, interpretation of its structure can be complicated by the comparatively low amplitude (and thus low binned SNR) relative to $\text{Re}(V)$, and by the typical uncertainty in the disc phase centre of $\lesssim 3$ mas. Third, imaging the FRANK residual visibilities effectively isolates azimuthal asymmetries in the image by subtracting out the (fitted) average brightness at each radius. A limitation is that there is typically ambiguity in interpreting structure in imaged residuals, due to potential artefacts of an incorrect disc geometry and/or phase centre, imaging artefacts, and loss in resolution by convolving the residuals with the CLEAN beam (as demonstrated in Fig. 3e). While each of these three approaches is thus imperfect, together they can aid in distinguishing superresolution asymmetries from annular features.

4 RESULTS AND ANALYSIS

Of the 32 sources in the Taurus survey, our analysis focused only on the 24 single-disc systems; among these, here we show the 10 for which we obtain a brightness profile with prominent new substructure. The remaining 14 comprise four extended discs where the FRANK visibility fit is highly similar to the parametric fit in Long et al. (2018) and 10 compact discs whose fitted brightness profiles lack substructure. In Section 4.1 we will discuss the general potential for substructure in these compact sources based on the observed visibility distributions. All FRANK fits in this work are available at <https://zenodo.org/record/6686456>.

The FRANK fit hyperparameters for the 10 data sets where we find new substructure are summarized in Table 1. We divide our analysis into compact and extended discs. The compact discs – BP Tau, DR Tau, and FT Tau – have an effective radius

⁶We can intuit this by recalling that the Fourier transform is a linear operation; the Fourier transform of a feature is equal to the sum of its ‘components’ Fourier transforms. Thus the transform of a ring is equal to the sum of the transforms of its azimuthal segments.

Table 1. For each Taurus survey disc in Section 3 and 4, the distance to the source (using *Gaia* DR2 measurements from Bailer-Jones et al. 2018), and the values for the five FRANK hyperparameters: SNR criterion α , strength of smoothing w_{smooth} applied to the reconstructed power spectrum, outer radius of the fit R_{out} , number of radial (and spatial frequency) points N , and floor value p_0 for the reconstructed power spectral mode amplitudes.

Disc	d [pc]	α	$\log_{10} w_{\text{smooth}}$	R_{out} [arcsec]	N	p_0 [Jy^2]
Compact discs						
BP Tau ^a	129	1.01	−4	1.0	200	10^{-15}
DO Tau	139	‘	‘	‘	‘	‘
DR Tau	195	‘	‘	‘	‘	‘
FT Tau	127	‘	‘	‘	‘	‘
Extended discs						
UZ Tau E	131	1.01	−4	1.0	200	10^{-15}
CIDA 9 A ^a	171	1.05	‘	‘	‘	‘
DS Tau ^a	159	1.05	‘	‘	‘	‘
RY Tau	128	1.10	‘	‘	‘	‘
DL Tau ^a	159	1.01	‘	1.5	300	‘
GO Tau	144	1.00	‘	1.5	300	‘
MWC 480 ^a	161	1.01	‘	1.5	300	‘

Notes. Section 2 gives a fuller explanation of α . All FRANK fits in this work are available at <https://zenodo.org/record/6686456>.

^aThese fits have enforced brightness profile positivity (see Section 2).

$R_{\text{eff, 90 per cent}} \leq 50$ AU, where the integrated flux $f(R_{\text{eff, 90 per cent}}) = 0.9 \times 2\pi \int_0^\infty I(r) r dr$. The extended discs – CIDA 9 A, DL Tau, DS Tau, GO Tau, MWC 480, RY Tau, and UZ Tau E – have $R_{\text{eff, 90 per cent}} > 60$ AU. Long et al. (2019) fit BP Tau and DR Tau parametrically with GALARIO, using an exponentially tapered power law to model the brightness profile, motivated by structure in the observed visibility distributions. Long et al. (2018) fit the remaining eight discs shown here with a parametric form in GALARIO comprised of a sum of Gaussians (for CIDA 9 A, DS Tau, RY Tau, and UZ Tau) or an exponentially tapered power law summed with Gaussians (for DL Tau, FT Tau, GO Tau, and MWC 480). Their choice of the number of Gaussians for each source is motivated by a brightness profile extracted (along the disc’s major axis) from the CLEAN image.

For each of these 10 sources, we compare the FRANK visibility fit to the parametric fit, as well as the Fourier transform of the CLEAN brightness profile, in Fig. 4 for the compact discs (the corresponding brightness profiles, discussed below, are in Fig. 6) and Fig. 5 for the extended discs (brightness profiles in Figs 8 and 11). In every case the parametric model matches the data more accurately than the Fourier transform of the CLEAN image profile, the transform of the CLEAN model profile is generally comparable to or in some cases slightly more accurate than the parametric model, and the residuals and χ^2 values demonstrate that the FRANK fit is more accurate than each of the CLEAN image, CLEAN model, and parametric model visibility profiles. The FRANK fits that have enforced brightness profile positivity (BP Tau, CIDA 9 A, DL Tau, DS Tau, MWC 480) underestimate data amplitudes at long baselines. None the less, we recall from Section 3.1 that even modest improvements in visibility fit accuracy can yield appreciably more highly resolved brightness profile features and can in some cases identify new features.

In order to examine whether new features may be non-axisymmetric, Figs 4 and 5 also show the imaginary component of the observed visibility distributions (which FRANK treats as zero at all baselines). We will discuss $\text{Im}(V)$ in relation to disc asymmetries in the following subsections.

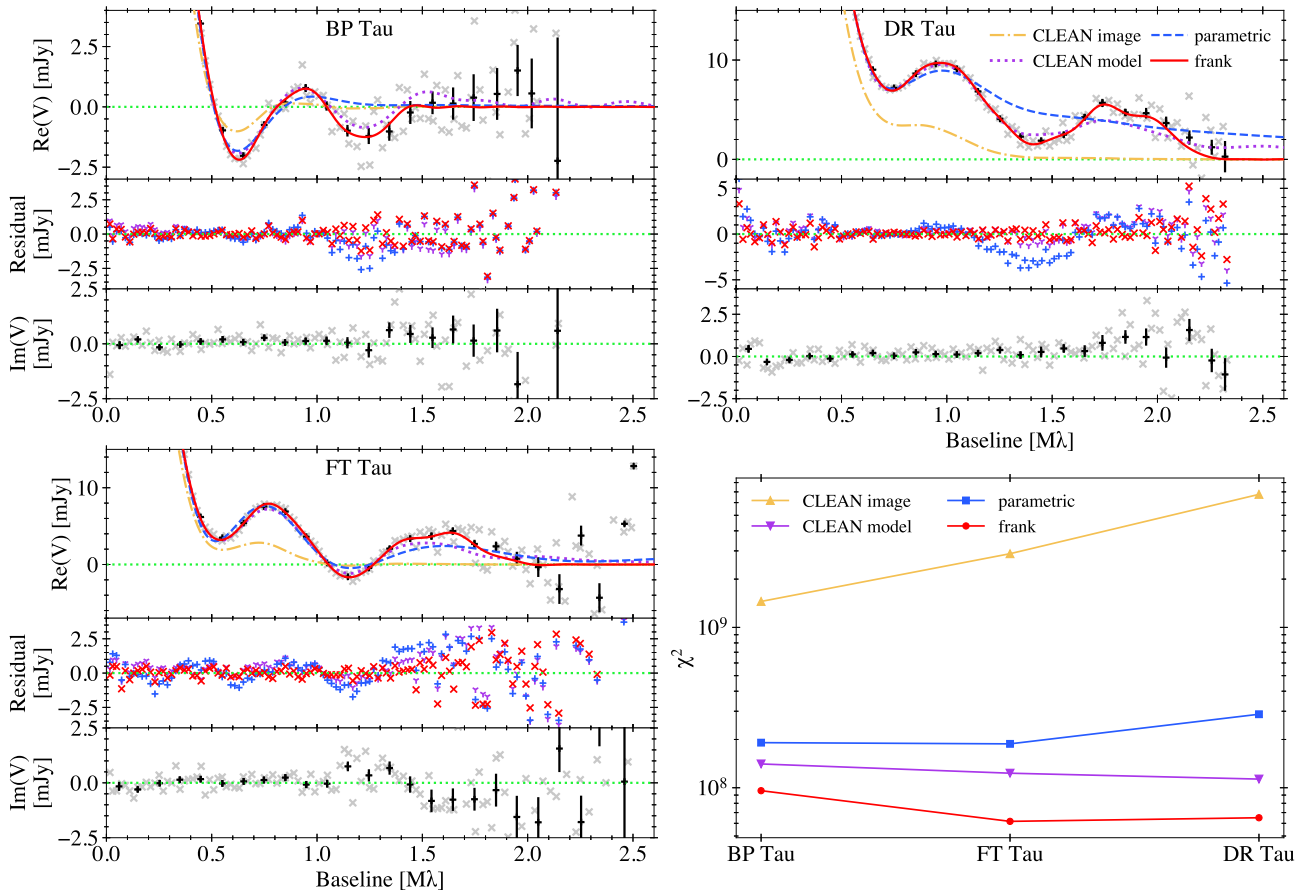


Figure 4. Improved visibility model accuracy in *FRANK* fits to compact discs. For the compact discs in Section 4.1, a zoom on the visibilities ($>0.30 \text{ M}\lambda$; 20- and 100-k λ bins, with 1σ uncertainties shown for the 100-k λ points). The parametric visibility fit from either Long et al. (2018) or Long et al. (2019), the *FRANK* fit, and the Fourier transforms of the *CLEAN* image and model brightness profiles are shown. Also shown are residuals for the parametric and *FRANK* fits and the *CLEAN* model transform (20-k λ bins; larger amplitude residuals at the longest baselines are beyond the y-range in some panels), as well as the imaginary component of the observed visibilities. Discs are arranged from the top to bottom and then left to right in increasing *FRANK* fit resolution. The bottom right-hand panel shows the χ^2 statistic for each fit.

4.1 New substructure in compact discs

Across three compact sources in the survey – BP Tau, DR Tau, and FT Tau – the *FRANK* fits in Fig. 6 find new substructure. Additionally we note that for the highly compact ($R_{\text{eff, 90 per cent}} < 25 \text{ AU}$) disc T Tau N, in which Yamaguchi et al. (2021) recently found a gap/ring pair with their 2D, superresolution modeling framework *PRIISM*, the *FRANK* fit (not shown here) demonstrates agreement in the gap/ring pair’s location and approximate amplitude.

BP Tau. The *FRANK* fit to BP Tau identifies a new turnover in the inner disc that is not seen in the Long et al. (2019) parametric profile because the corresponding visibility fit does not recover the negative peak in the data at $1.25 \text{ M}\lambda$. The parametric model instead finds an almost flat inner disc (power law index of 0.1), resulting in a quasi-linear region of the brightness profile between ≈ 8 and 17 AU ; this can be understood as a result of underresolving the turnover (which may itself be an underresolved ring). Representation of an underresolved brightness excess as a quasi-linear region in a brightness profile is demonstrated with mock data in Fig. 3(b). We can further motivate the turnover by the observed visibilities; their amplitudes are preferentially negative between ≈ 0.5 and $1.5 \text{ M}\lambda$, which is an indication of a wide Gaussian in the brightness profile that is not centred at zero radius. Notice how the visibility distributions in Fig. 5 for the three discs with an apparent inner cavity (and thus, a

Gaussian ring not centred at zero) – CIDA 9 A, RY Tau, and UZ Tau E – also each exhibit preferentially negative visibility amplitudes at intermediate baselines. The *FRANK* profile also better localizes the structure beyond 20 AU in the disc than the parametric profile, with the *CLEAN* model profile showing rough agreement with *FRANK* here.

To further assess the *FRANK* profile features (using BP Tau as an example for the analysis we will more succinctly cover in subsequent discs), we can consider how the model limitations in Sections 3.2 and 3.3 may affect the fit. Given the demonstration in Section 3.2 of the difficulty in accurately extrapolating a fit to unobserved (and noise-dominated) baselines – and how this can introduce fake oscillations into a profile when the underlying disc has highly superresolution features – we emphasize that a more accurate visibility fit or higher resolution/deeper observations can find the features in a *FRANK* brightness profile to become either more or less prominent. This is particularly true in the innermost disc, where substructure can routinely be highly underresolved. The turnover in the *FRANK* profile for BP Tau, for example, may resolve into a ring or something more complicated, as may be indicated by the inner 15 AU of the *CLEAN* model profile in Fig. 6. We do expect the turnover to be indicating the presence of real substructure, given the data set’s preferentially negative visibility amplitudes at intermediate baselines

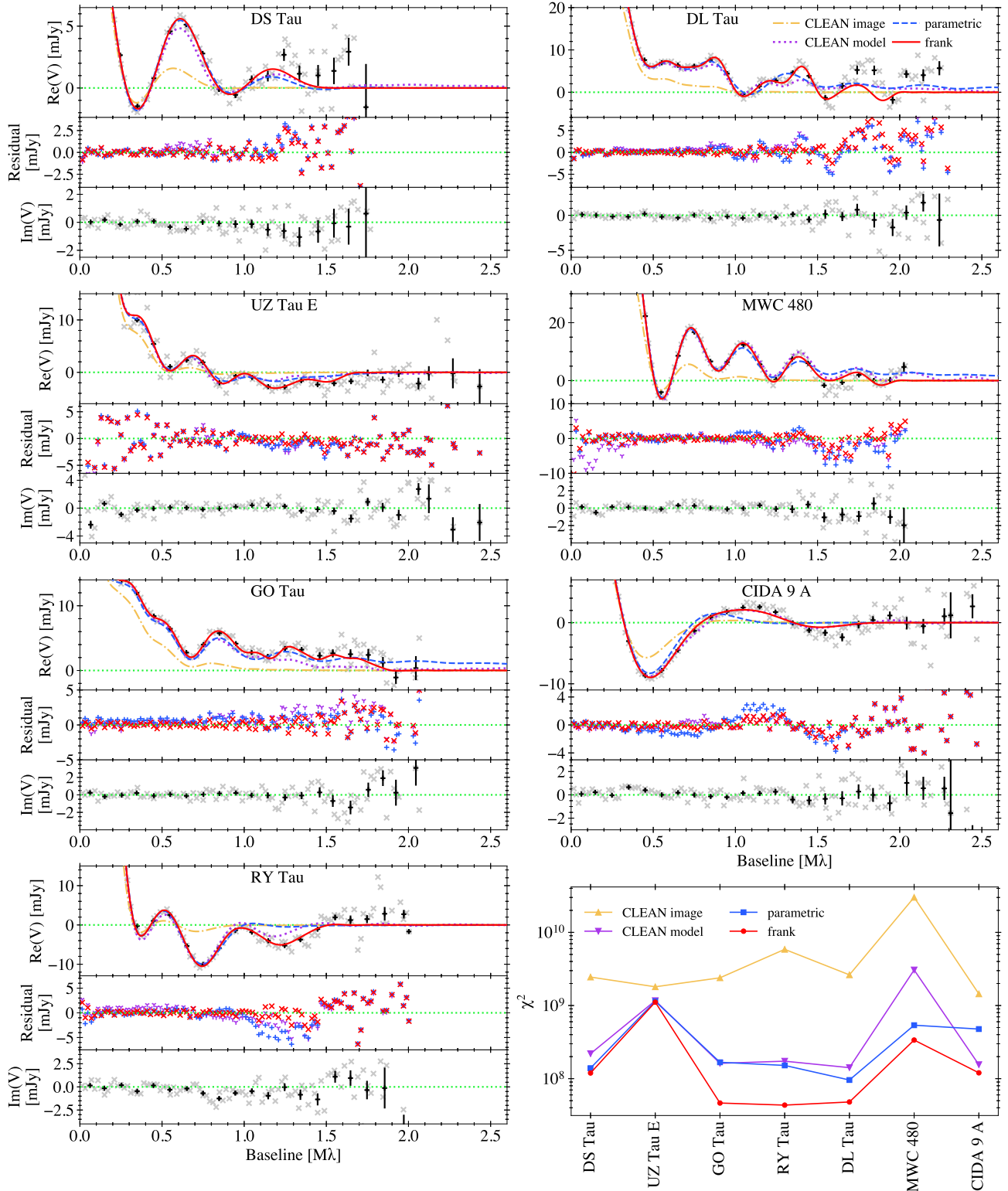


Figure 5. Improved visibility model accuracy in *FRANK* fits to extended discs. As in Fig. 4, but for the extended discs in Section 4.2.

as discussed above. In the outer disc, the broad, shallow feature in the *FRANK* brightness profile between ≈ 52 and 65 AU is at least partly due to noise (influenced by the visibility fit's extrapolation of zero amplitude beyond ≈ 1.5 $M\lambda$, analogous to Fig. 2e), but it may also have contributions from real, diffuse emission.

In light of the discussion in Section 3.3 on how non-axisymmetric features can mimic the appearance of a partially resolved ring, we can also use the *CLEAN* image, imaginary component of the visibilities and imaged *FRANK* residuals to examine whether any superresolution features in the *FRANK* profile may be artefacts of

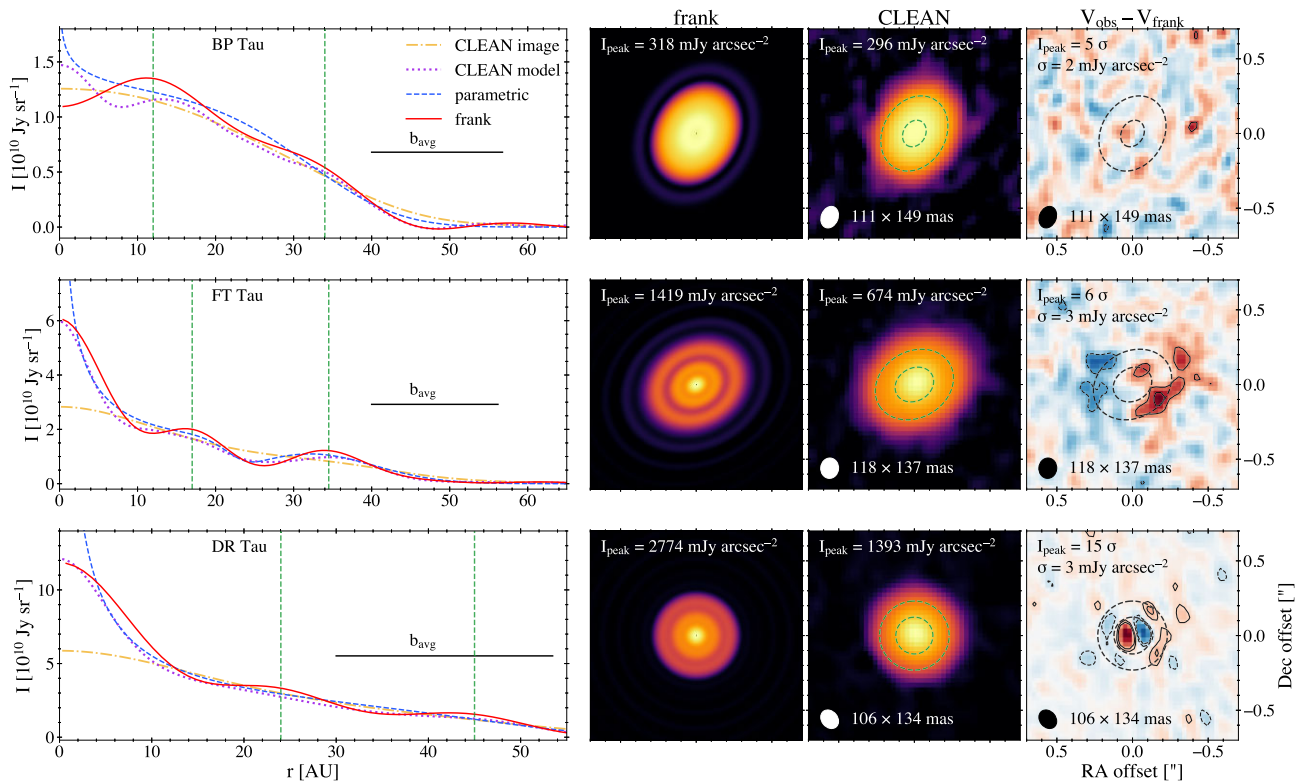


Figure 6. New substructure in `FRANK` fits to compact discs. Brightness profiles extracted from the `CLEAN` image and model, the parametric profile, and `FRANK` profile for three of the compact ($R_{\text{eff}, 90 \text{ per cent}} \leq 50$ AU) systems in the Taurus survey, corresponding to the visibility fits in Fig. 4 (the parametric fits are from Long et al. 2018 and Long et al. 2019; b_{avg} shows the mean of the `CLEAN` beamwidth along its major and minor axes). Also shown are an image of the `FRANK` profile swept over 2π and reprojected, the `CLEAN` image, and the imaged `FRANK` residual visibilities (zero `CLEAN` iterations; contours at -5σ , -3σ , $+3\sigma$, $+5\sigma$). Vertical lines in the brightness profile plots denote features that are shown as ellipses in the `CLEAN` image and imaged `FRANK` residuals for reference. The `FRANK` and `CLEAN` images use an arcsinh stretch ($I_{\text{stretch}} = \text{arcsinh}(I/a) / \text{arcsinh}(1/a)$, $a = 0.02$), but different brightness normalization (indicated by the given peak brightness). The imaged `FRANK` residuals use a linear stretch symmetric about zero. We use image brightness units of $[\text{mJy arcsec}^{-2}]$ to facilitate comparison between data sets of different beam size.

azimuthally asymmetric emission. Contouring the `CLEAN` image of the source shows no clear signs of an asymmetry; the imaginary component of the visibility distribution in Fig. 4 does not exhibit prominent structure, indicating that asymmetries in the image must be particularly faint and/or small-scale; and the imaged `FRANK` residuals do not have clear features within the disc (the small, 5σ blob in the west of the imaged residuals that is also in the `CLEAN` image). We thus infer that the profile’s features are likely annular.

FT Tau. The `FRANK` fit in Fig. 6 finds a new gap/ring pair around 11–17 AU, underresolved in the parametric brightness profile as the quasi-linear region (and hinted at in the `CLEAN` model profile). The `FRANK` profile also determines the gap at 26 AU identified in Long et al. (2018) to be wider, with a brighter adjacent ring. While the difference between the parametric and `FRANK` visibility fits for FT Tau in Fig. 4 may not look dramatic enough to correspond to a new gap/ring pair, it is important first that `FRANK` exhibits an improved fit accuracy over a large span in baseline (≈ 1.0 – 1.7 $M\lambda$). Secondly, while the `FRANK` fit converges on zero visibility amplitude at ≈ 2.0 $M\lambda$, the parametric fit remains positive and continues to slowly oscillate out to the longest baselines and beyond. The data instead appear by eye to indicate that the true visibility distribution becomes negative beyond 2.0 $M\lambda$ [denser (u , v) plane sampling at these baselines would be needed to confirm].

Considering disc asymmetries, the imaginary component of the visibilities for FT Tau in Fig. 4 have clear structure on scales between ≈ 1.2 and 1.7 $M\lambda$, and this structure has amplitude comparable to

the difference between the `FRANK` and parametric fit residuals for $\text{Re}(V)$. The imaged `FRANK` residuals also have $\leq 5\sigma$ features within and beyond the gap at ≈ 26 AU. Together this suggests that there may be some faint asymmetric structure in the disc, particularly in the gap centred at 26 AU, where the residual amplitude is largest.

DR Tau. We find two new gaps relative to Long et al. 2019 (the `FRANK` fit to DR Tau was previously shown in Jennings et al. 2020) in Fig. 6. This can be motivated by the significant difference in visibility fit accuracy between the parametric and `FRANK` models in Fig. 4; the `CLEAN` model visibility profile is also more accurate than the parametric model, with the `CLEAN` model brightness profile having a hint of the outer ring found in the `FRANK` brightness profile. The qualitative similarity in structure between the observed visibility distributions for FT Tau and DR Tau also motivates why the `FRANK` fit shows two gaps in both discs. The visibilities for DR Tau do not exhibit a zero-crossing, indicating the data contain underresolved structure at small spatial scales; this seems most likely to be an indication of a partially resolved inner disc.

Considering the inner disc, while the imaginary component of the visibilities for DR Tau do not show clear structure, the imaged `FRANK` residuals in Fig. 6 do have strong features in the innermost radii ($\leq 15\sigma$, or $\lesssim 5$ per cent of the background brightness in the `CLEAN` image). This is likely affecting the morphology of the inner gap in the `FRANK` profile to some extent. We find that the inner disc residuals in DR Tau and other discs discussed below are not attributable solely

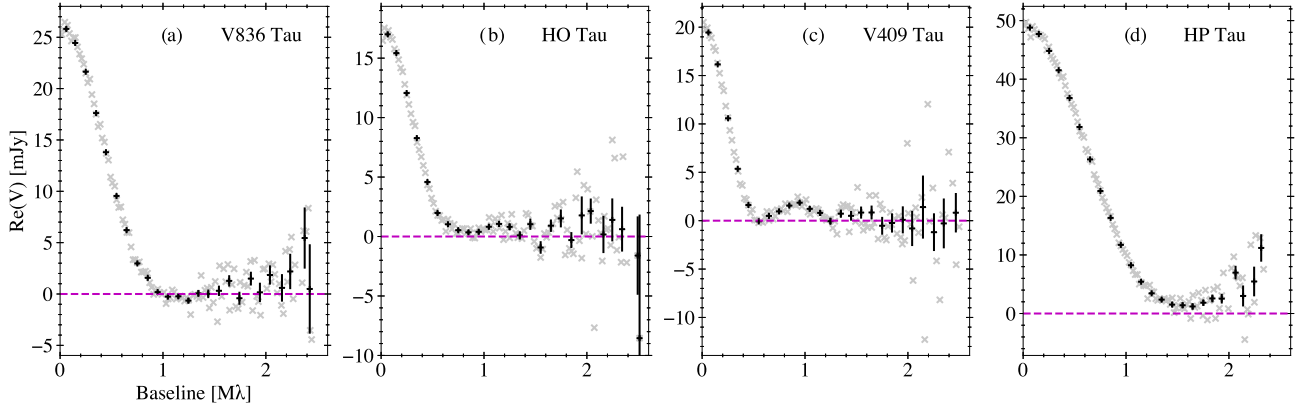


Figure 7. Noisy visibility distributions for compact discs. Visibility distributions for four compact ($R_{\text{eff, 90 per cent}} \leq 50$ AU) sources in the Taurus survey, whose structure at long baselines is unclear due to (u, v) plane sparsity.

to an incorrect determination of the disc phase centre (assessed by varying the applied phase centre in Appendix A).

4.1.1 Occurrence rate of substructure in compact discs

The FRANK fits to BP Tau and DR Tau raise the number of compact, single-disc systems with substructure from two – FT Tau and the cavity disc IP Tau (shown in Long et al. 2018) – to four, out of a total of 14 in the survey. Among the 14 compact discs, these four are neither the largest nor brightest, which prompts the question of whether more of the survey’s compact objects may be structured. To partially address this, we can consider whether the survey data strongly exclude the presence of substructure in the remaining 10 compact sources. The visibility distributions for four additional compact discs in Fig. 7 each show tentative or clear indications of structure at intermediate baselines and become highly noisy at longer baselines due to sparse (u, v) plane sampling. Whether this structure at intermediate baselines corresponds in each case only to the brightness profile becoming steeper in the outer disc (i.e. no gap/ring substructure), or instead to substructure, is not clear from these data. The visibility distributions in Figs 7(b)–(d) also do not exhibit a zero-crossing at short baseline, characteristic of an underresolved inner disc and/or highly superresolution substructure. It is thus possible that higher resolution and/or deeper observations would identify substructure in a larger subset of the survey’s compact sources.

That the current data do show substructure in four of the survey’s compact discs – BP Tau, DR Tau, FT Tau, and IP Tau – is in line with multiple features detected in the FRANK fits to the DSHARP observations of the compact sources SR 4, DoAr 33, and WSB 52 (Jennings et al. 2022), as well as substructure recovered in the parametric visibility fits to the compact discs CIDA 1, MHO 6, and J0433 (Kurtovic et al. 2021; Pinilla et al. 2021). Collectively these results demonstrate that many compact discs are not intrinsically featureless; their lack of apparent substructure is instead in some cases an artefact of either the resolving power of the model applied to the data or of the data itself. This may be a tentative indication that a non-trivial fraction of compact dust discs follow the same evolutionary pathway as extended discs, which tend to be structured.

4.2 New substructure in extended discs

For each of seven extended discs – CIDA 9 A, DL Tau, DS Tau, GO Tau, MWC 480, RY Tau, and UZ Tau E – the FRANK brightness

profiles identify new features and more highly resolve those found in Long et al. (2018). We divide our analysis here into sources with a deep gap separating the inner and outer disc (Section 4.2.1) and those with an apparent inner cavity (Section 4.2.2).

4.2.1 Sources with an inner and outer disc

In four sources with an inner disc separated from one or more outer rings by a deep gap – DS Tau, MWC 480, DL Tau, and GO Tau – we find new substructure as shown in Fig. 8.

DS Tau. The FRANK fit finds a new feature, a broad plateau, in the gap separating the inner and outer disc (at 30 AU). This arises from the small improvement in visibility fit accuracy in Fig. 5, and it may be underresolving smaller scale substructure. The CLEAN model profile also exhibits this plateau. The improved fit accuracy with FRANK additionally yields slightly steeper gap walls. The feature in the gap may be informed to some extent by non-axisymmetric emission, given structure in the imaginary component of the visibilities and in the imaged FRANK residuals of Fig. 8 in the inner disc.

GO Tau. The FRANK profile in Fig. 8 finds the quasi-linear region between ≈ 21 and 45 AU in the parametric profile to resolve into two rings. This may seem surprising when comparing the fairly similar FRANK and parametric visibility fits for GO Tau in Fig. 5, but it can be understood by the FRANK visibility fit exhibiting regions of comparatively steeper slope beyond ≈ 1.1 M λ as it more closely traces the data. The inner disc features in the FRANK profile can none the less be expected to evolve with longer baseline data that more strongly condition structure on small scales. In Fig. 8, we also again see a bimodal pattern in the imaged FRANK residuals of the innermost disc, with $\geq 5\sigma$ and $\leq -5\sigma$ features interior to the inner ring. We can expect that the inner disc features may evolve considerably with higher resolution observations.

The deep gap separating inner from outer disc (at 55 AU) in the FRANK fit exhibits a slight bump (see the inset in Fig. 8), suggesting it may not be empty. This is reminiscent of structure in the deep gap between inner and outer disc in the FRANK fits to the ≈ 35 -mas resolution DSHARP observations of AS 209, Elias 24, HD 163296, and SR 4 (see fig. 12 in Jennings et al. 2022); it may be indicative of a common gap forming mechanism. The fractional uncertainty in a FRANK profile is largest at faint brightness though, and the rms noise level in the CLEAN image of GO Tau, 0.01×10^{10} Jy sr $^{-1}$, is of comparable amplitude to the bump, so inference on structure within the deep gap is limited. In the outer disc, the FRANK profile

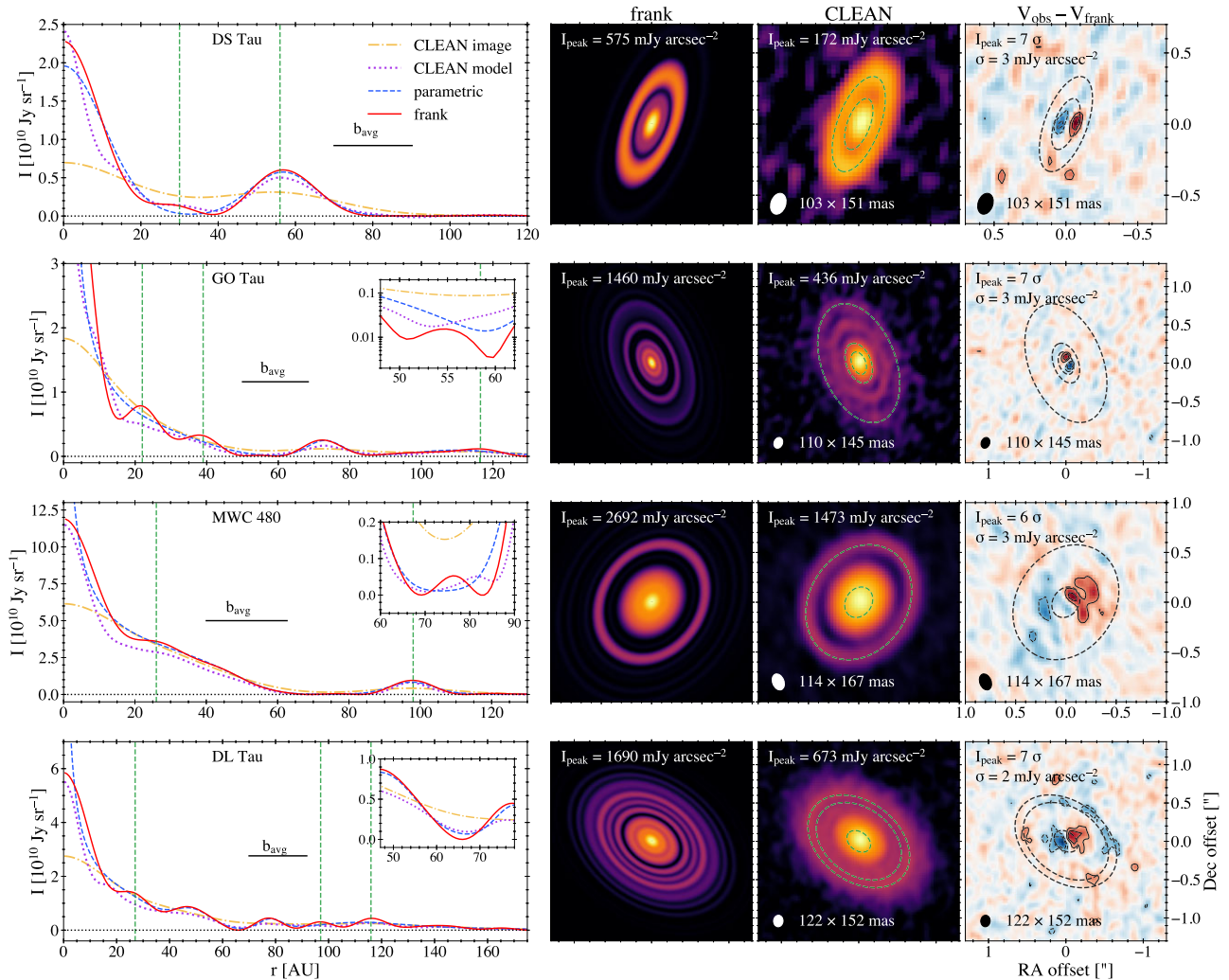


Figure 8. New substructure in *FRANK* fits to extended discs with outer rings. As in Fig. 6, but for the four extended ($R_{\text{eff, 90 per cent}} > 50$ AU) systems in the Taurus survey that exhibit an inner disc and one or more outer rings, discussed in Section 4.2.1. Parametric profiles are from Long et al. (2018); the visibilities and fits for these discs are in Fig. 5. The inset panels zoom in on deep gaps in the brightness profiles. The *FRANK* fit to GO Tau peaks at $6.2 \times 10^{10} \text{ Jy sr}^{-1}$.

more strongly localizes the location of the outermost ring and better resolves its faint amplitude. As a note, the two rings in the outer disc are clearly visible in the *CLEAN* image due to the colourscale, but are relatively faint, and imaging artefacts are likely introducing the apparent diffuse emission into the gap between the rings; these rings are thus shallow and broad in the *CLEAN* image profile (though visible in the *CLEAN* model profile).

MWC 480. The *FRANK* fit finds the inner disc for this source as well to be structured, with a new plateau between ≈ 20 and 27 AU. The profile’s broad, shallow, quasi-linear region between ≈ 30 and 50 AU may be a further indication of underresolved inner disc substructure. Additionally, the imaged *FRANK* residuals show $\geq 3\sigma$ asymmetries across the inner disc; from all of this we may again expect the inner disc morphology to evolve with higher resolution observations. As in GO Tau, the deep gap separating inner from outer disc (at 76 AU) has a bump in the inset in Fig. 8. Again, at low surface brightness the relative model uncertainty is higher, although the *CLEAN* model profile does also suggest there may be structure in this gap.

DL Tau. Like DS Tau, GO Tau and MWC 480, DL Tau has a deep gap that separates inner from outer disc (at 66 AU). Yet the gap in DL Tau is narrower and lacks the flat bottom morphology.

The outer disc in DL Tau is distinct as well; while in the other three discs there is one prominent ring exterior to the gap, in DL Tau we find three (in addition to broad bumps at 144 and 165 AU that either trace faint rings, diffuse emission, or potentially artefacts of the visibility model’s extrapolation). The rings at 97 and 116 AU in the *FRANK* profile are averaged over as a single, broad feature in the parametric profile from Long et al. (2018). Unique also to DL Tau is prominent asymmetry in the outer disc. The imaged *FRANK* residuals have $\geq 3\sigma$ and $\leq -3\sigma$ regions that lie roughly in the gap between the outer two prominent rings. The asymmetries have an orientation consistent with a generally brighter east side of the outer disc as identified by contouring the *CLEAN* image in Fig. 9(c). Collectively, these differences in morphology for DL Tau could indicate that the gaps in this disc are produced by a different physical process or a lower mass planet than in DS Tau, GO Tau and MWC 480. The one strong similarity between DL Tau and these other sources is a new plateau in DL Tau between ≈ 19 and 27 AU that corresponds to an asymmetry in the imaged fit residuals (note a plateau is also seen in the *CLEAN* model profile), suggesting the underresolved inner disc substructure may not be purely annular.

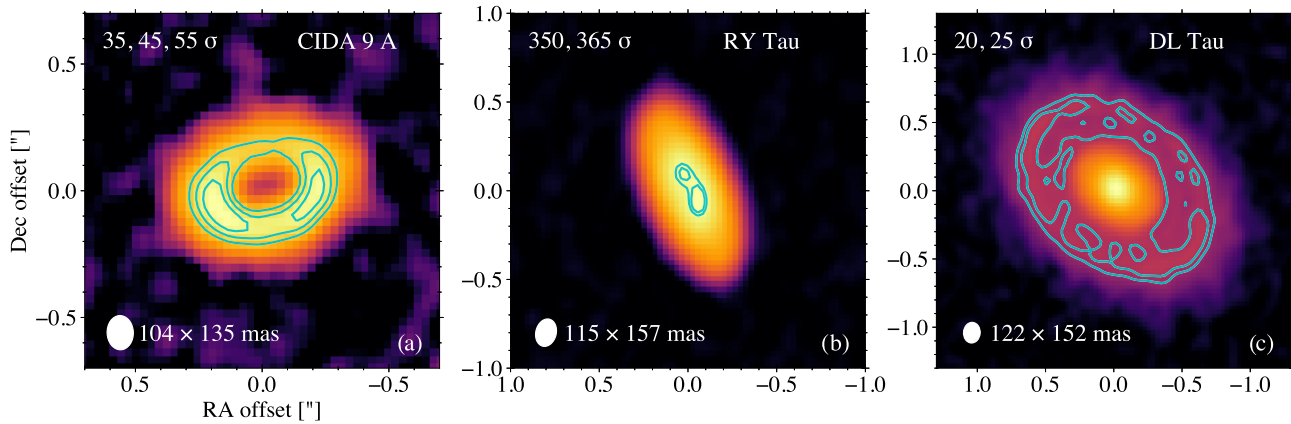


Figure 9. CLEAN image asymmetries. CLEAN images for three of the Taurus survey’s extended discs (see Section 4.2), with contours chosen to highlight asymmetries. The images are identical to those in Figs 8 and 11.

The abundance of substructure in the FRANK brightness profile for DL Tau (two new rings in addition to the three rings identified in Long et al. 2018) makes this a good disc for comparing the FRANK fit to a parametric model whose functional form is motivated by the FRANK profile. Such a comparison gives a sense of how similar we can expect non-parametric and parametric fits to be for a highly structured source. This is of particular interest in the inner disc, where FRANK fits tend to find new substructure; that is, an independent parametric model can test the recovery of the features in the FRANK profile. This comparison also demonstrates the benefit of using a rapid, superresolution FRANK brightness profile (as compared to the profile extracted from a CLEAN image or even from a CLEAN model) to motivate a parametric model that uses expensive Markov Chain Monte Carlo (MCMC).

Fig. 10 shows this comparison for DL Tau, between a 10 Gaussian parametric model and the FRANK fit. The parametric modelling approach and results, including the corner plot and analysis of sampling convergence, are more fully presented in Appendix B. The 10 Gaussian parametric form is composed of 2 Gaussians based on the plateau and ring in the inner disc of the FRANK fit, 3 Gaussians for the three prominent rings in the outer disc of the FRANK fit, 1 Gaussian for the broad bump at 144 AU in the FRANK fit, 2 additional Gaussians to describe the disc interior to 25 AU, and 2 more Gaussians to account for the brightness profile’s small offset from zero brightness out to large radii. Fig. 10(b) shows the median of the posterior samples for each of these 10 Gaussians, as well as the spread in randomly drawn samples for each.

The median brightness profile for the parametric model is in general agreement with the FRANK profile for DL Tau in Fig. 10(a), with the FRANK profile lying within the 2σ confidence interval of the parametric model at almost all radii. Both models find the outer disc between 65 and 130 AU to resolve into three rings, and both prefer a (likely underresolved) deviation from the smooth Gaussian envelope in the inner disc, between 15 and 30 AU. Relative to the FRANK profile, the parametric median profile exhibits narrower and brighter rings in the outer disc (and thus more flat-bottomed gaps between these rings), as well as a slight turnover near $r = 0$. These differences arise from the different extrapolation of the parametric median visibility fit and the FRANK fit beyond the end of the data in Fig. 10(e). The true visibility distribution likely continues to oscillate beyond the longest baselines sampled, but the observations of course provide no constraint on visibility amplitudes at unsampled baselines (apart from flux conservation). Since the differences between the

parametric and FRANK profiles in visibility space are essentially limited to noisy or unsampled baselines, the precise ring widths, flatness of the gap bottoms, and turnover near $r = 0$ in the parametric brightness profile should thus be considered uncertain. Overall though, the general agreement between the parametric and FRANK profiles provides further evidence that DL Tau is densely structured, and the comparison illustrates the benefit of using a FRANK profile to initialize a parametric visibility fit, particularly for a disc with a large number of features.

4.2.2 Sources with an apparent inner cavity

For each of the three discs with an apparent inner cavity identified in the Taurus survey – CIDA 9 A, RY Tau, and UZ Tau E – the FRANK fit in Fig. 11 finds one or more new features.

RY Tau. The FRANK fit finds the cavity hinted at in the parametric fit to be almost fully cleared, with a steep outer wall. The adjacent ring in the parametric profile resolves into a narrower/brighter ring and an emission excess, a ‘shoulder’, in the FRANK profile (the shoulder is also hinted at in the CLEAN model profile). The contoured CLEAN image in Fig. 9(b) shows asymmetry in the innermost disc, and the imaged FRANK residuals in Fig. 11 have a strong asymmetric pattern at small radii (roughly interior to the shoulder) that is $\leq 17\sigma$, or ≈ 5 per cent of the peak brightness in the CLEAN image. This is a smaller contrast by a factor of a few than the shoulder in the FRANK profile, suggesting that feature is not purely due to an asymmetry. The residual structure could be dominated by an elevated/flared emission surface, as $\geq 5\sigma$ and $\leq -5\sigma$ residuals span most of the disc, and the source has a large fitted inclination of $\approx 65^\circ$. A cleared inner cavity and inner disc asymmetry are seen in higher resolution observations (20×40 mas beam) of this source (Francis & van der Marel 2020). In the outer disc, the plateau in the parametric profile between ≈ 40 and 50 AU becomes a gap/ring pair in the FRANK profile (and to a lesser extent in the CLEAN model profile), as may be expected from a higher resolution fit; note how the parametric fit in Fig. 5 misses the trough in the visibilities centred at 1.25 M λ that the FRANK fit recovers and the CLEAN model visibility profile partially recovers.

UZ Tau E. As in RY Tau, the FRANK profile finds the cavity to be more devoid of material than previously seen, with a steeper edge and brighter adjacent ring, and a shoulder on the ring’s trailing edge. The broad region of quasi-linear slope in both the parametric and FRANK brightness profiles (between ≈ 40 and 70 AU in the latter) is

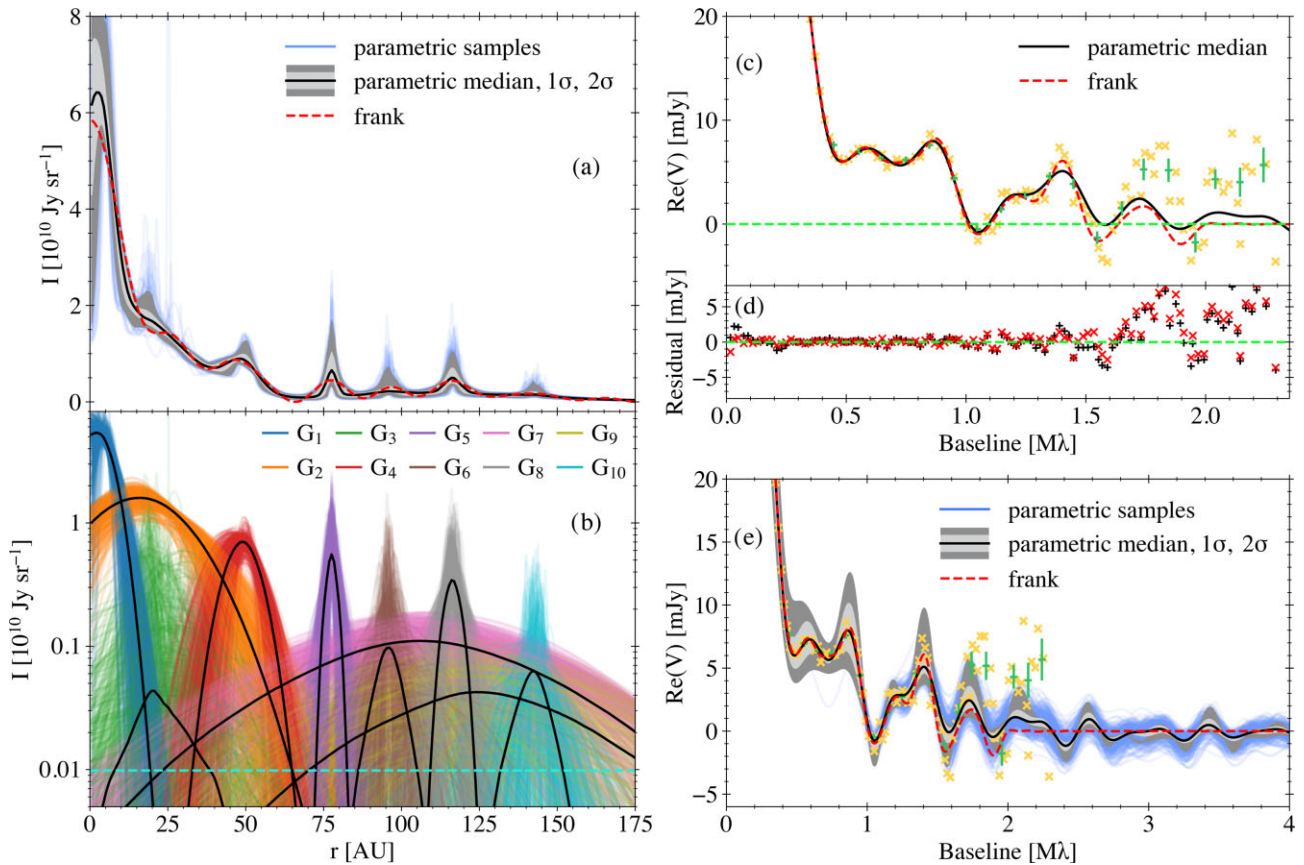


Figure 10. Ten Gaussian parametric fit to DL Tau. Panel (a): posterior median, 1σ and 2σ confidence intervals of the 10 Gaussian fit to DL Tau, and 500 randomly drawn posterior samples. Also shown is the `FRANK` fit from Fig. 8. Panel (b): posterior median (black lines) for each of the 10 Gaussians in the fit, and the same 500 samples. The dashed horizontal line is at the `CLEAN` image rms noise level. Panel (c): a zoom on the observed visibilities ($>0.30 M\lambda$; 20- and 100-k λ bins), with 1σ uncertainties shown for the 100-k λ points, and the parametric median and `FRANK` visibility fits. Panel (d): residuals for the visibility fits (20-k λ bins). Panel (e): as in panel (c), but with the 1σ and 2σ confidence intervals and the 500 posterior samples included. Longer baselines are shown to demonstrate the difference in fit extrapolations at unsampled scales.

potentially suggestive of underresolved substructure at these radii. In the outer disc, the `FRANK` fit finds the ring at 82 AU to be narrower and brighter. The imaginary component of the visibilities for UZ Tau do show structure at the shortest baselines, but this is due to the disc-bearing binary system UZ Tau Wa and Wb in the field of view.

CIDA 9 A. As in RY Tau and UZ Tau E, the `FRANK` profile finds the cavity wall to be steeper, with a brighter adjacent ring and an accompanying shoulder that is also apparent in the `CLEAN` model profile. The imaginary component of the visibilities in Fig. 5 show structure across a wide range of baselines, and the contoured `CLEAN` image of the source in Fig. 9(a) correspondingly traces brightness excesses in the south-east and south-west of the disc. These roughly coincide with the ring’s peak location in the `FRANK` profile and the strong structure in the imaged `FRANK` residuals. The residual features have brightness up to 18 per cent of the peak brightness in the `CLEAN` image; such a high contrast entails they are affecting the `FRANK` profile in the bright ring’s vicinity. The profile also indicates an additional, faint ring within the cavity (at 9 AU).

4.2.2.1 The shoulder morphology as a trend A shoulder is present on the trailing edge of the bright ring in all three Taurus survey discs with an apparent inner cavity, suggesting a trend. The shoulder morphology is also seen in several discs beyond the survey that have an inner cavity or deep gap. These shoulders have been

identified using a variety of fitting techniques, and over a range of observational resolutions and wavelengths. Like the Taurus discs, the shoulder’s contrast varies across discs observed at similar resolution and wavelength. And like the Taurus discs there are often brightness asymmetries in the vicinity of the ring and shoulder, identified in either a `CLEAN` image or imaged fit residuals.

In some discs, a brightness arc in an otherwise empty annulus seen in the `CLEAN` image manifests in the `CLEAN` brightness profile as a shoulder. Examples include the arc exterior to a ring outside a deep gap in the 1.3-mm DSHARP observations of HD 143006 (Huang et al. 2018; Pérez et al. 2018), as well as the arc exterior to a ring that surrounds an inner cavity in the 0.9-mm observations of V1247 Ori and HD 135344 B (van der Marel et al. 2019; the shoulder in HD 135344 B is also seen in the `FRANK` profile in Norfolk et al. 2021). In other cases, similar to the Taurus survey discs, the shoulder morphology is present not as the result of a clearly isolated arc, but within an annulus that in the `CLEAN` image appears to contain emission across all azimuthal angles. The 2.1-mm `CLEAN` brightness profile of GM Aur shows such a shoulder on the trailing edge of a bright ring exterior to a cavity (Huang et al. 2020), with the `CLEAN` image showing hints of a brightness asymmetry in the radial region of the gap and shoulder; lower resolution observations of the same source at 0.93 and 7 mm (Macías et al. 2018) also find a shoulder.

`FRANK` fits to four of the six DSHARP sources that have a bright ring in the inner disc – AS 209, HD 142666, HD 163296, and Sz 129

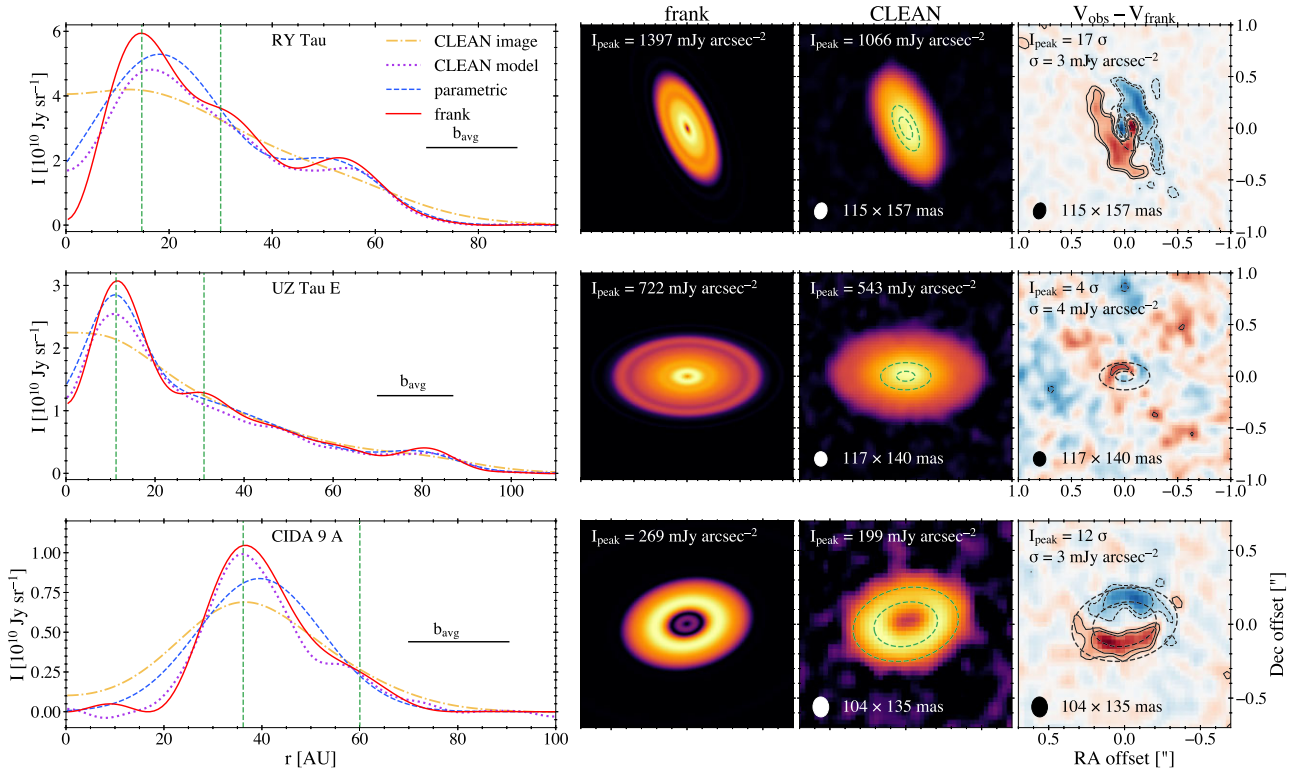


Figure 11. New substructure in `FRANK` fits to extended discs with cavities. As in Fig. 6, but for three of the extended ($R_{\text{eff}, 90 \text{ per cent}} > 50 \text{ AU}$) systems in the Taurus survey that exhibit an inner cavity, discussed in Section 4.2.2. Parametric profiles are from Long et al. 2018; the visibilities and fits for these discs are in Fig. 5. Vertical lines in the brightness profile plots denote the radial location of a ring and its shoulder; ellipses in the `CLEAN` image and imaged `FRANK` residuals correspond to these radii.

– show a shoulder on the ring’s trailing edge (Jennings et al. 2022). In Sz 129 the ring is exterior to an inner cavity, while in AS 209, HD 142666, and HD 163296 it is exterior to a deep gap in the inner disc. These fits are reproduced in Fig. 12, with brightness asymmetries consistently present in the imaged `FRANK` residuals interior to and/or at the radial location of the ring. Asymmetries are also identified at these radii in the `CLEAN` image for HD 142666, HD 163296, and Sz 129 (Huang et al. 2018). The shoulder’s contrast varies across the `FRANK` brightness profiles, from a faint, wide bump in Sz 129 to an apparent ring in HD 142666. Fig. 12 also shows a `FRANK` fit to the 40-mas observations of CI Tau from Clarke et al. (2018), where the broad ring in the parametric profile at 27 AU resolves into an inner narrow ring and an outer, fainter ring (the shoulder) in the `FRANK` fit. The `FRANK` profile also finds the deep gap interior to the rings to be structured.

We suspect this shoulder morphology (regardless of whether a given shoulder is underresolving a ring) is tracing some common physical mechanism whose relative effect varies between sources. Perhaps the most viable candidates are ones that can produce azimuthal brightness asymmetries in a disc with a cavity or deep gap, such as those discussed in section 3.3 of Long et al. (2018): planet-induced dust traps (Ataiee et al. 2013; van der Marel et al. 2013) and eccentric cavities in a circumbinary disc (Ragusa et al. 2017); or migrating planets (Meru et al. 2018; Nazari et al. 2019).

5 CONCLUSIONS

We used `FRANK` to identify new features and more highly resolve known features in 10 Taurus survey discs observed at ≈ 120 -mas

resolution.⁷ Relative to the parametric visibility fits in Long et al. (2018) and Long et al. (2019) and the `CLEAN` model brightness profiles, which both yielded substantially more disc substructure than the `CLEAN` image brightness profiles, we demonstrated how further improvements to visibility fit accuracy with the non-parametric approach in `FRANK` could find yet more features. The most notable example was DL Tau, in which the `FRANK` fit recovered two new rings in a disc with three previously identified rings. We also used this source to show how a superresolution `FRANK` profile is advantageous for motivating a parametric form that can be modelled with tools such as `GALARIO`, and how this parametric fit provided further confidence in the `FRANK` profile features. Among the substructures characterized across the 10 discs, we identified three main trends:

(i) Increased substructure in compact discs: Of the survey’s 14 discs with radii $\lesssim 55 \text{ AU}$, we found two previously smooth discs (BP Tau, DR Tau) to exhibit substructure and identified a new gap in the inner disc of another (FT Tau). These discs were not systematically larger or brighter than the compact sources without detected substructure, and we motivated how sparse (u, v) plane sampling at long baselines in many of the latter does not exclude the presence of substructure at the observed spatial scales.

(ii) Increased inner disc substructure: Across the compact and extended sources considered, we found evidence of underresolved substructure at small ($\lesssim 30 \text{ AU}$) radii, in many cases coinciding with azimuthally asymmetric fit residuals.

⁷All `FRANK` fits in this work are available at <https://zenodo.org/record/668645>.

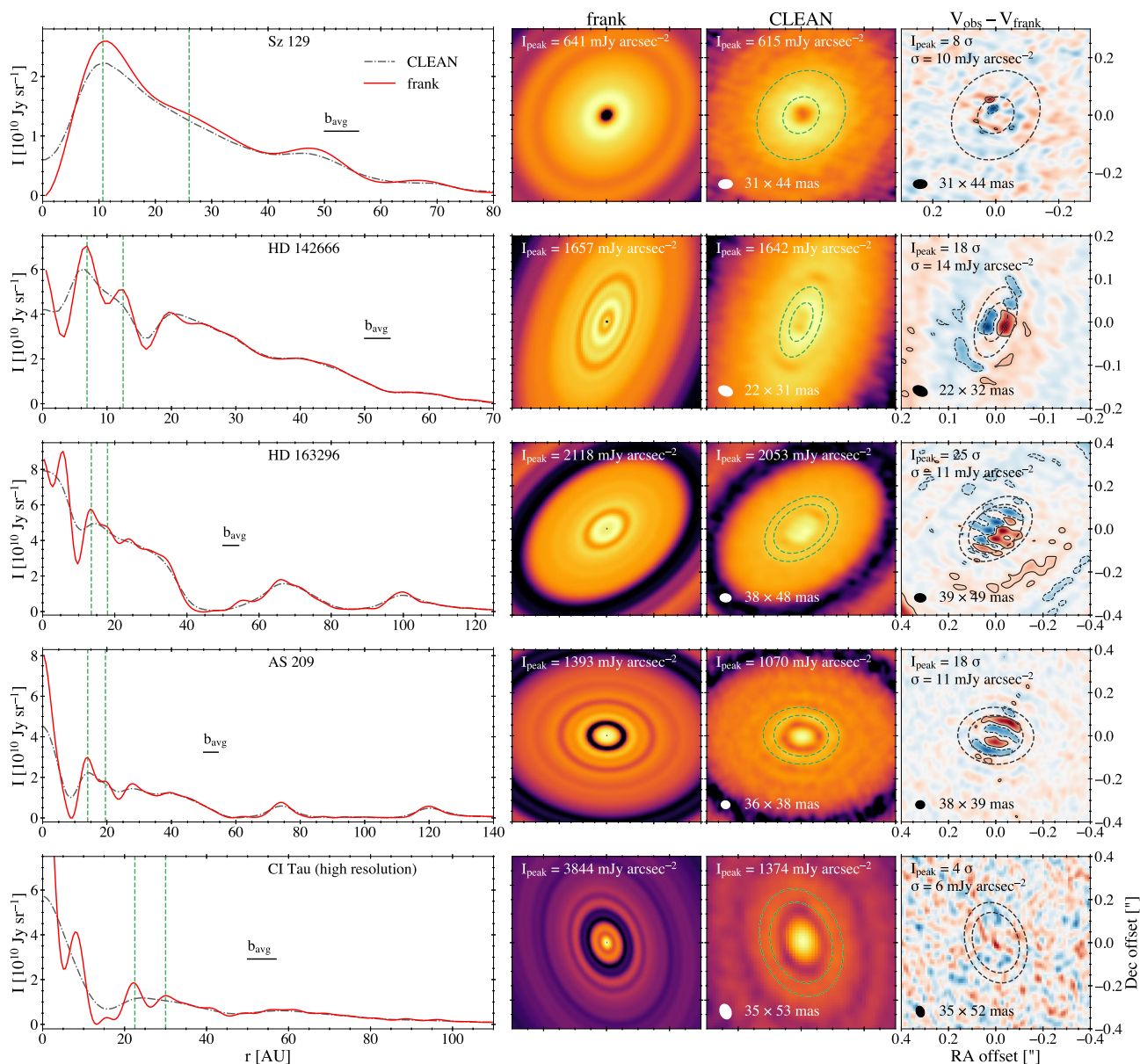


Figure 12. Shoulder morphology in high-resolution observations. The brightness profile extracted from the CLEAN image and the FRANK brightness profile for four discs in the ≈ 35 -mas resolution DSHARP survey (Andrews et al. 2018; Huang et al. 2018) and the ≈ 40 -mas resolution observations of CI Tau (Clarke et al. 2018; the y-scale zooms on lower brightness). Vertical lines in the brightness profile plots denote the radial location of a ring and its shoulder (Section 4.2); ellipses in the CLEAN image and imaged FRANK residuals correspond to these radii. Images zoom on the inner disc of each source. The imaged FRANK residuals have contours at -5σ , $+5\sigma$. FRANK fits for the DSHARP discs are from Jennings et al. (2022); the FRANK fit for CI Tau is previously unpublished (Appendix C shows the visibility fit).

(iii) A ring/shoulder morphology in inner discs: The three survey sources with an apparent inner cavity (CIDA 9 A, RY Tau, UZ Tau E) showed a shoulder on the trailing edge of the disc’s bright ring. We noted numerous instances of this same morphology exterior to a cavity or deep gap in discs outside the survey, positing it may trace a common physical mechanism.

Identification of new substructure in Taurus survey discs complements recent applications of FRANK to the DSHARP survey (Jennings et al. 2022) and ODISEA survey (Cieza et al. 2021). Along with superresolution fits obtained using other methods such as GALARIO in Long et al. (2018, 2019), these results contribute to the growing evidence that it is not only bright, large discs that

exhibit substructure. Instead a lack of substructure in a disc may often be an artefact of a data set’s or model’s resolution. This underscores the utility of superresolution methods across a range of observational resolutions to better constrain substructure occurrence rates and discern morphological trends. Ultimately, a large ensemble of sources characterized at superresolution scales will help to discriminate between candidate physical mechanisms producing disc features.

Softwares: ASTROPY (Astropy Collaboration 2013, 2018), CASA (McMullin et al. 2007), CORNER.PY (Foreman-Mackey 2016), EMCEE (Foreman-Mackey et al. 2013), FRANK (Jennings et al. 2020), GALARIO (Tazzari et al. 2018), MATPLOTLIB (Hunter 2007), NUMPY (Walt, Colbert & Varoquaux 2011), SCIPY (Virtanen et al. 2020).

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DATA AVAILABILITY

The data used in this work are available on the ALMA archive at <https://almascience.eso.org/asax/>, project code 2016.1.01164.S.

All FRANK fits in this work are available at <https://zenodo.org/record/6686456>.

REFERENCES

- Alexander R. D., Clarke C. J., Pringle J. E., 2006, *MNRAS*, 369, 229
 ALMA Partnership T. et al., 2015, *ApJ*, 808, L3
 Andrews S. M. et al., 2016, *ApJ*, 820, L40
 Andrews S. M. et al., 2018, *ApJ*, 869, L41
 Andrews S. M. et al., 2021, *ApJ*, 916, 51
 Astropy Collaboration, 2013, *A&A*, 558, A33
 Astropy Collaboration, 2018, *AJ*, 156, 123
 Ataiee S., Pinilla P., Zsom A., Dullemond C. P., Dominik C., Ghanbari J., 2013, *A&A*, 553, L3
 Bai X.-N., Stone J. M., 2014, *ApJ*, 796, 31
 Bailer-Jones C. A. L., Rybizki J., Fousneau M., Mantelet G., Andrae R., 2018, *AJ*, 156, 58
 Benisty M. et al., 2021, *ApJ*, 916, L2
 Casassus S. et al., 2021, *MNRAS*, 507, 3789
 Cieza L. A. et al., 2021, *MNRAS*, 501, 2934
 Clarke C. J., Gendrin A., Sotomayor M., 2001, *MNRAS*, 328, 485
 Clarke C. J. et al., 2018, *ApJ*, 866, L6
 Cui C., Bai X.-N., 2021, *MNRAS*, 507, 1106
 Dipierro G., Pinilla P., Lodato G., Testi L., 2015, *MNRAS*, 451, 974
 Dong R. et al., 2018a, *ApJ*, 860, 124
 Dong R., Najita J. R., Brittain S., 2018b, *ApJ*, 862, 103
 Dullemond C. P. et al., 2018, *ApJ*, 869, L46
 Ercolano B., Clarke C. J., Drake J. J., 2009, *ApJ*, 699, 1639
 Flock M., Ruge J. P., Dzyurkevich N., Henning T., Klahr H., Wolf S., 2015, *A&A*, 574, A68
 Flock M., Nelson R. P., Turner N. J., Bertrang G. H.-M., Carrasco-González C., Henning T., Lyra W., Teague R., 2017, *ApJ*, 850, 131
 Foreman-Mackey D., 2016, *J. Open Source Softw.*, 1, 24
 Foreman-Mackey D., Hogg D. W., Lang D., Goodman J., 2013, *PASP*, 125, 306
 Francis L., van der Marel N., 2020, *ApJ*, 892, 111
 Goldreich P., Tremaine S., 1979, *ApJ*, 233, 857
 Hall C., Rice K., Dipierro G., Forgan D., Harries T., Alexander R., 2018, *MNRAS*, 477, 1004
 Hashimoto J., Muto T., Dong R., Liu H. B., van der Marel N., Francis L., Hasegawa Y., Tsukagoshi T., 2021, *ApJ*, 911, 5
 Hu X., Zhu Z., Okuzumi S., Bai X.-N., Wang L., Tomida K., Stone J. M., 2019, *ApJ*, 885, 36
 Huang J. et al., 2018, *ApJ*, 869, L42
 Huang J. et al., 2020, *ApJ*, 891, 48
 Hunter J. D., 2007, *Comput. Sci. Eng.*, 9, 90
 Jennings J., Booth R. A., Tazzari M., Rosotti G. P., Clarke C. J., 2020, *MNRAS*, 495, 3209
 Jennings J., Booth R. A., Tazzari M., Clarke C. J., Rosotti G. P., 2022, *MNRAS*, 509, 2780
 Johansen A., Youdin A., Klahr H., 2009, *ApJ*, 697, 1269
 Keppler M. et al., 2019, *A&A*, 625, A118
 Kley W., Nelson R. P., 2012, *ARA&A*, 50, 211
 Kudo T., Hashimoto J., Muto T., Liu H. B., Dong R., Hasegawa Y., Tsukagoshi T., Konishi M., 2018, *ApJ*, 868, L5
 Kurtovic N. T. et al., 2021, *A&A*, 645, A139
 Lin D. N. C., Papaloizou J., 1986, *ApJ*, 309, 846
 Long F. et al., 2018, *ApJ*, 869, 17
 Long F. et al., 2019, *ApJ*, 882, 49
 Longarini C., Lodato G., Toci C., Aly H., 2021, *MNRAS*, 503, 4930
 McMullin J. P., Waters B., Schiebel D., Young W., Golap K., 2007, in Shaw R. A., Hill F., Bell D. J., eds, *ASP Conf. Ser. Vol. 376, Astronomical Data Analysis Software and Systems XVI*. Astron. Soc. Pac., San Francisco, p. 127
 Macías E. et al., 2018, *ApJ*, 865, 37
 Macías E., Guerra-Alvarado O., Carrasco-González C., Ribas Á., Espaillat C. C., Huang J., Andrews S. M., 2021, *A&A*, 648, A33
 Manger N., Klahr H., 2018, *MNRAS*, 480, 2125
 Meru F., Rosotti G. P., Booth R. A., Nazari P., Clarke C. J., 2018, *MNRAS*, 482, 3678
 Nazari P., Booth R. A., Clarke C. J., Rosotti G. P., Tazzari M., Juhasz A., Meru F., 2019, *MNRAS*, 485, 5914
 Norfolk B. J. et al., 2021, *MNRAS*, 502, 5779
 Okuzumi S., Momose M., Sirono S.-i., Kobayashi H., Tanaka H., 2016, *ApJ*, 821, 82
 Pérez L. M. et al., 2018, *ApJ*, 869, L50
 Pérez S., Casassus S., Baruteau C., Dong R., Hales A., Cieza L., 2019, *AJ*, 158, 15
 Pfeil T., Klahr H., 2021, *ApJ*, 915, 130
 Pinilla P., Birnstiel T., Ricci L., Dullemond C. P., Uribe A. L., Testi L., Natta A., 2012, *A&A*, 538, A114
 Pinilla P., Flock M., Ovelar M. d. J., Birnstiel T., 2016, *A&A*, 596, A81
 Pinilla P. et al., 2021, *A&A*, 649, A122
 Pinte C. et al., 2019, *Nat. Astron.*, 3, 1109
 Price D. J. et al., 2018, *MNRAS*, 477, 1270
 Ragusa E., Dipierro G., Lodato G., Laibe G., Price D. J., 2017, *MNRAS*, 464, 1449
 Sheehan P. D., Eisner J. A., 2018, *ApJ*, 857, 18
 Sierra A., Lizano S., Macías E., Carrasco-González C., Osorio M., Flock M., 2019, *ApJ*, 876, 7
 Tazzari M., Beaujean F., Testi L., 2018, *MNRAS*, 476, 4527
 Toci C., Rosotti G., Lodato G., Testi L., Trapman L., 2021, *MNRAS*, 507, 818
 Tsukagoshi T. et al., 2019, *ApJ*, 878, L8
 van der Marel N. et al., 2013, *Science*, 340, 1199
 van der Marel N., Dong R., di Francesco J., Williams J. P., Tobin J., 2019, *ApJ*, 872, 112
 Varnière P., Tagger M., 2006, *A&A*, 446, L13
 Virtanen P. et al., 2020, *Nat. Methods*, 17, 261

- Walt S. v. d., Colbert S. C., Varoquaux G., 2011, *Comput. Sci. Eng.*, 13, 22
- Yamaguchi M., Tsukagoshi T., Muto T., Nomura H., Nakazato T., Ikeda S., Tamura M., Kawabe R., 2021, *ApJ*, 923, 121
- Zhang K., Bergin E. A., Blake G. A., Cleeves L. I., Hogerheijde M., Salinas V., Schwarz K. R., 2016, *ApJ*, 818, L16

APPENDIX A: EFFECT OF PHASE CENTER UNCERTAINTY ON IMAGED FRANK RESIDUALS

To assess the robustness of features in the imaged FRANK residuals shown in the main text, here we consider how the fitted phase centre alters their morphology and brightness. We focus on the phase centre, rather than the fitted inclination or position angle, or out-of-plane effects, because several of the imaged FRANK residuals in the main text show a bimodal asymmetry in the inner disc that may reasonably be expected as an artefact of the applied phase centre (see appendix A in Andrews et al. 2021 for a good demonstration). We seek to determine whether they may instead be indications of real asymmetries (see also appendix B in Jennings et al. 2022).

Long et al. (2018) found the typical 1σ uncertainties in fitted Right Ascension and Declination offsets for a source (relative to the centre of the field of view) to be <1 mas, while we have found

uncertainties with mock and real data to commonly be 1–3 mas. While <1 mas shifts in phase centre typically have a trivial effect on residual visibility amplitudes, shifts of 1–3 mas can induce visible differences in imaged FRANK residuals. To test whether these shifts can remove high residual brightness in the inner disc, for each source in the main text we have applied a phase centre differing from the published value by 1, 2, or 3 mas – with the shift at $\pi/4$ intervals over the full 2π in azimuth – then fit the shifted visibilities and compared the fit to that with the published phase centre (this is the same test described in appendix B of Jennings et al. 2022). The effects of a phase shift of 1–3 mas on the visibilities and thus the FRANK brightness profile are largely imperceptible, but differences are evident in the imaged FRANK residuals.

As an example, the imaged FRANK residuals for DR Tau using the published phase centre contain $>5\sigma$, bimodal features in the inner disc. Fig. A1 shows the results of the above test for a 2-mas phase centre shift at each of the $\pi/4$ azimuthal angles (the imaged residuals for shifts of 1 and 3 mas are qualitatively similar). The phase shifts do result in a variation in the peak residual brightness by a factor of ≤ 2 , and in the orientation of the bimodal pattern. But the pattern persists in all cases, and phase shifts that reduce the pattern’s brightness (which we may at first interpret as the applied phase

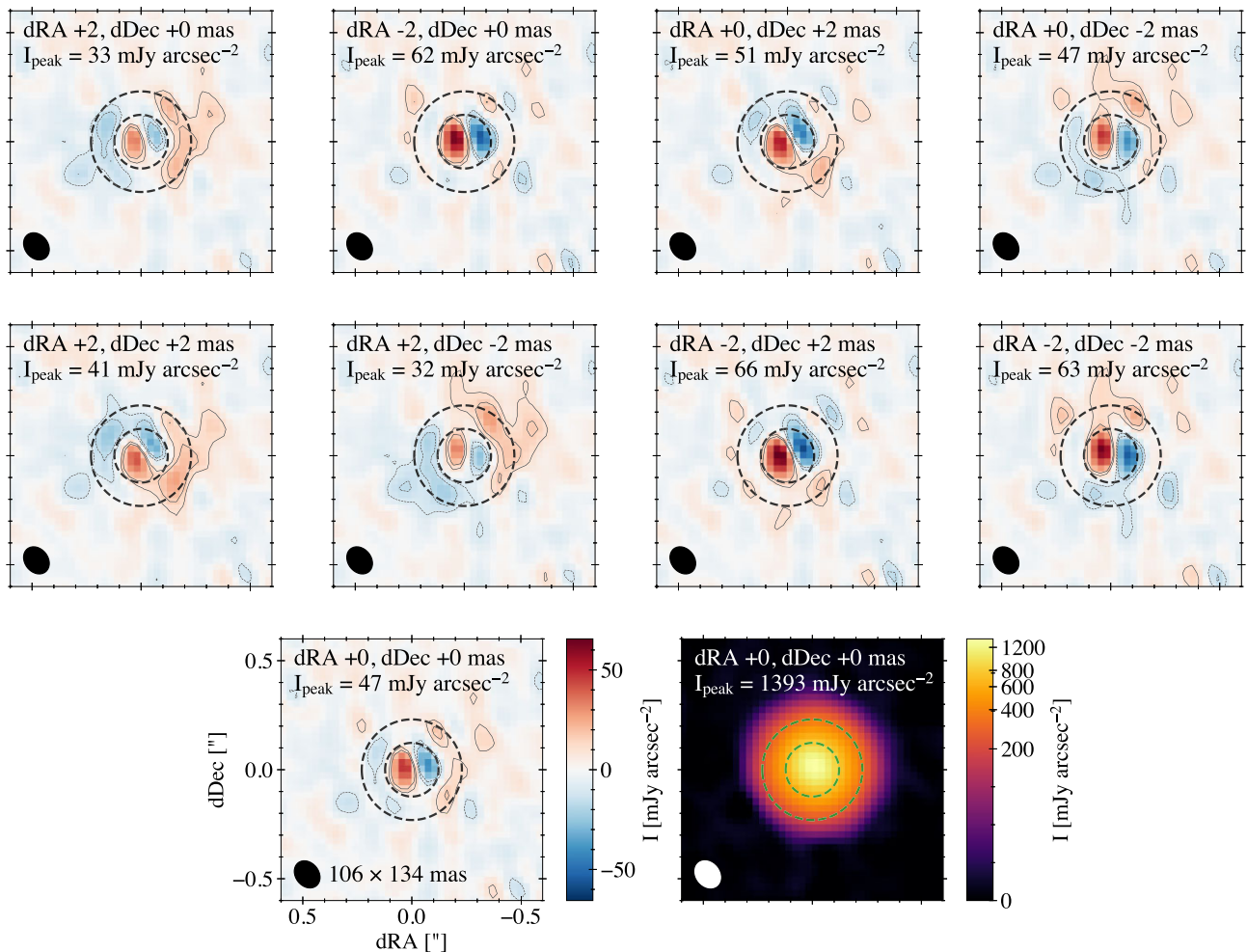


Figure A1. Effect of fitted phase centre on imaged residual visibilities. Bottom row: the imaged FRANK residual visibilities for the fit in the main text to DR Tau (zero CLEAN iterations; contours at -5σ , -3σ , $+3\sigma$, $+5\sigma$, with $\sigma = 3 \text{ mJy arcsec}^{-2}$), and the CLEAN image. Top and centre rows: the imaged FRANK residual visibilities when the fitted dRA and/or dDec is varied by ± 2 mas (as listed in each panel). The imaged FRANK residual panels all use the same absolute linear stretch shown in the colorbar.

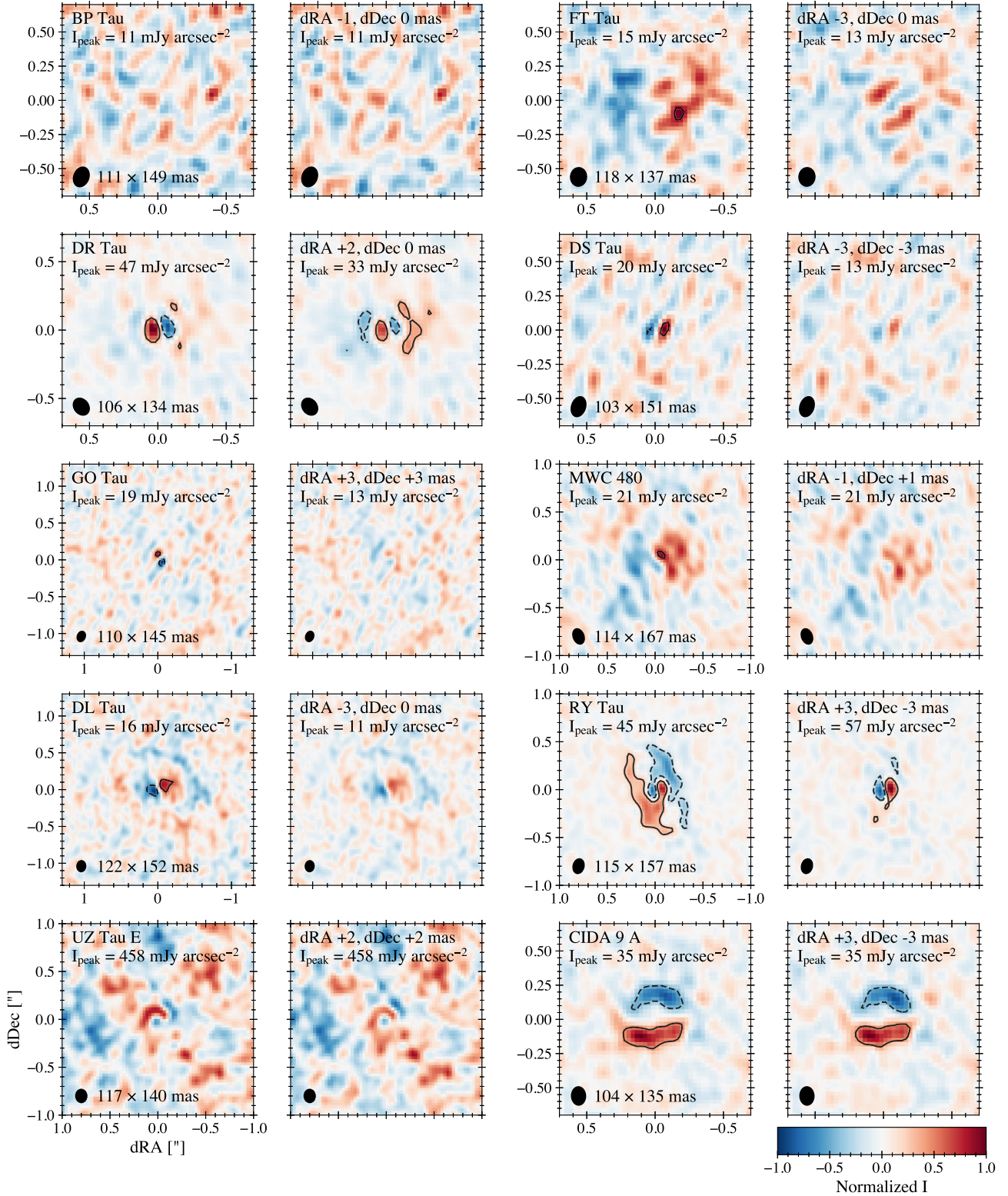


Figure A2. Varying the phase centre to minimize the imaged residual visibilities. The imaged FRANK residual visibilities from the main text for each source in Section 4, alongside the imaged visibilities obtained by varying the fitted dRA and/or dDec to minimize the absolute image brightness. The images are produced with zero CLEAN iterations; contours are at -5σ , $+5\sigma$. The peak brightness is given, as is the phase shift applied to minimize the absolute brightness.

centre being more accurate) also increase the residual amplitude at larger disc radii. We could expect that this is due to a more complex combination of an incorrect phase centre, incorrect inclination and/or position angle, and out-of-plane effects; however, this disc is nearly face-on (fitted inclination of $5^\circ.4$). The persistence of $>10\sigma$ features in the residuals thus suggests there is real inner disc structure that the `FRANK` fit to these data is not resolving.

To consider the full set of 10 sources in Section 4, Fig. A2 compares the imaged `FRANK` residuals from the main text for each disc with the residuals produced when we shift the phase centre to minimize the absolute brightness in the image. In some cases, a bimodal asymmetry in the inner disc is weakened, while in others it persists. This suggests these inner disc residual features are not (always) purely an artefact of an incorrect visibility deprojection.

APPENDIX B: PARAMETRIC FIT TO DL TAU

Section 4.2.1 shows a 10 Gaussian parametric fit to DL Tau using `GALARIO`. Here we present the fit in more detail. The model contains 34 free parameters: a centroid, amplitude, and standard deviation for each of the 10 Gaussians, as well as the disc geometry (inclination, position angle, and the Right Ascension and Declination offsets). We perform an initial maximum likelihood estimate using the BFGS solver in `scipy.optimize.minimize`, then initialize the

MCMC walkers in a Gaussian ball around this estimate (by adding to each parameter value a draw from the standard normal distribution multiplied by 10^{-4}). We run the MCMC with `EMCEE` (Foreman-Mackey et al. 2013), using 160 walkers (≈ 5 per parameter) and a uniform prior on each parameter in the brightness profile Gaussians, as well as a Gaussian prior on the disc geometry parameters (centred on the published geometry), as listed in Table B1. We run the MCMC for 3×10^5 steps and then estimate the autocorrelation time τ for each chain at various points in the run. We do not reach convergence across all chains during the run, with the estimate of the autocorrelation time averaged over all dimensions $\hat{\tau}$ continually increasing as a power law in Fig. B1 rather than plateauing. This demonstrates how the high dimensionality of the parameter space would require a significantly larger number of steps to reach sampling convergence.

From the full set of remove a burn-in of $2 \times \max(\tau) \approx 6 \times 10^4$ steps, with τ estimated at the last step in the chains. Using the resulting samples, Table B1 gives the posterior 16th, 50th, and 84th percentiles for each parameter; unsurprisingly the faintest Gaussians (G_3 , G_6 , G_9 , G_{10}) have the highest uncertainty on their width and amplitude. Fig. B1 shows the corner plot using `CORNER.PY` (Foreman-Mackey 2016), with few instances of strong covariance in the 2D distributions, but also non-Gaussianity in the 1D distributions for the centroid and the standard deviation of some of the brightness profile Gaussians.

Table B1. Priors and posterior 16th, 50th, and 84th percentiles for each parameter in the 10 Gaussian parametric fit to DL Tau.

Parameter (unit)	Priors	Prior
r_i in $G_i(r_i, \sigma_i, I_i)$ [arcsec]	Uniform:	$(0.00, 0.08)$, $i = 1$
		$(0.08, 0.10)$, $i = 2$
		$(0.10, 0.20)$, $i = 3$
		$(0.20, 0.40)$, $i = 4$
		$(0.40, 0.55)$, $i = 5$
		$(0.55, 0.63)$, $i = 6$
		$(0.63, 0.65)$, $i = 7$
		$(0.65, 0.70)$, $i = 8$
		$(0.70, 0.80)$, $i = 9$
		$(0.80, 0.95)$, $i = 10$
σ_i in $G_i(r_i, \sigma_i, I_i)$ [arcsec]	Uniform: $(0.00, 0.30)$ for $i \in [1...10]$	
I_i in $G_i(r_i, \sigma_i, I_i)$ [$\log_{10}(\text{Jy sr}^{-1})$]	Uniform: $(8, 12)$ for $i \in [1...10]$	
inc [$^\circ$]	$G(x_0 = 44.95, \sigma_x = 5.0)$	
PA [$^\circ$]	$G(x_0 = 52.14, \sigma_x = 5.0)$	
dRA [mas]	$G(x_0 = 240, \sigma_x = 5)$	
dDec [mas]	$G(x_0 = -60, \sigma_x = 5)$	
	Posteriors	
	Brightness profile Gaussians	Disc geometry
$G_1(r = 0.01^{+0.01}_{-0.01}, \sigma = 0.03^{+0.01}_{-0.01}, I = 10.76^{+0.08}_{-0.06})$		inc = $45.10^{+0.32}_{-0.30}$ $^\circ$
$G_2(r = 0.10^{+0.02}_{-0.01}, \sigma = 0.10^{+0.02}_{-0.02}, I = 10.21^{+0.06}_{-0.07})$		PA = $51.90^{+0.45}_{-0.46}$ $^\circ$
$G_3(r = 0.14^{+0.04}_{-0.03}, \sigma = 0.05^{+0.10}_{-0.04}, I = 8.90^{+0.86}_{-0.62})$		dRA = 236^{+1}_{-1} mas
$G_4(r = 0.31^{+0.01}_{-0.01}, \sigma = 0.03^{+0.01}_{-0.01}, I = 9.86^{+0.05}_{-0.06})$		dDec = -59^{+1}_{-1} mas
$G_5(r = 0.49^{+0.00}_{-0.00}, \sigma = 0.01^{+0.01}_{-0.00}, I = 9.80^{+0.24}_{-0.22})$		
$G_6(r = 0.60^{+0.02}_{-0.02}, \sigma = 0.03^{+0.05}_{-0.02}, I = 9.09^{+0.43}_{-0.45})$		
$G_7(r = 0.67^{+0.02}_{-0.02}, \sigma = 0.24^{+0.04}_{-0.09}, I = 9.05^{+0.13}_{-0.22})$		
$G_8(r = 0.73^{+0.01}_{-0.01}, \sigma = 0.02^{+0.01}_{-0.01}, I = 9.58^{+0.23}_{-0.15})$		
$G_9(r = 0.79^{+0.04}_{-0.04}, \sigma = 0.22^{+0.05}_{-0.12}, I = 8.65^{+0.24}_{-0.38})$		
$G_{10}(r = 0.89^{+0.02}_{-0.03}, \sigma = 0.03^{+0.04}_{-0.02}, I = 8.89^{+0.34}_{-0.32})$		

Notes. $G(r, \sigma, I)$ denotes a Gaussian of radial position r , standard deviation σ , and logarithmic brightness I . The disc geometry parameters listed are inclination (inc), position angle (PA), Right Ascension offset (dRA), and Declination offset (dDec).

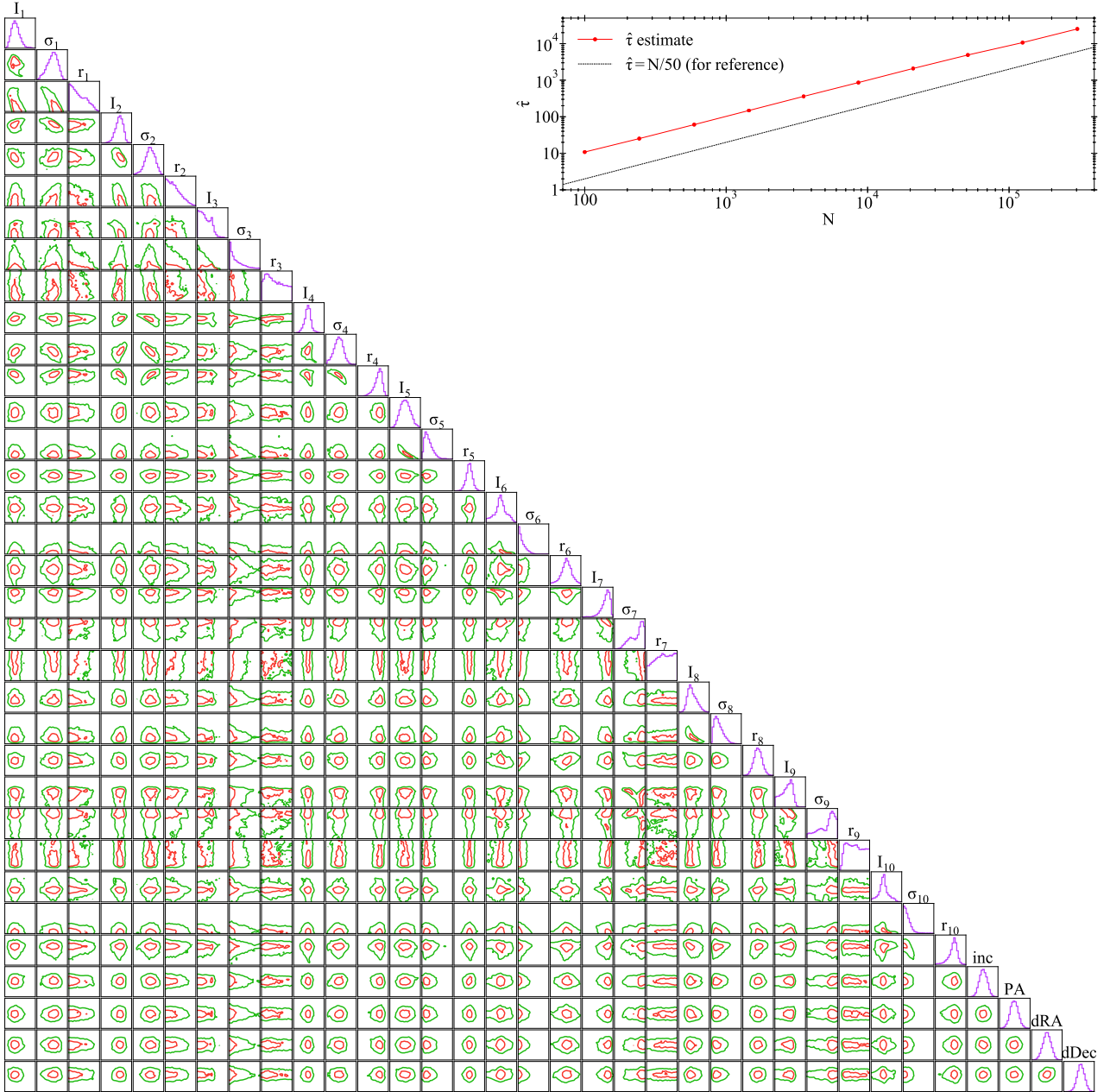


Figure B1. Corner plot for parametric fit to DL Tau. For the parametric fit to DL Tau, a corner plot showing the posterior for each fitted parameter (along the diagonal) and the covariance between parameters (red 1σ and green 2σ confidence intervals). The top right-hand panel shows the estimate for the autocorrelation time averaged over all dimensions $\hat{\tau}$ as a function of the number of samples N .

APPENDIX C: FRANK FIT TO HIGH-RESOLUTION OBSERVATIONS OF CI TAU

The `FRANK` fit to the ≈ 40 -mas observations of CI Tau⁸ in Fig. C1 finds new features in the disc’s brightness profile: a (very likely underresolved) gap/ring pair at 5 AU, structure in the deep gap at 15 AU, and a separation of the single ring at 25 AU into two rings. The parametric `GALARIO` profile from Clarke et al. (2018), also shown

⁸The `FRANK` fit uses visibilities deprojected and phase centred by $i = 47^\circ 3$, $PA = 14^\circ 1$, (dRA, dDec) = (330, -93) mas. These were determined in `FRANK` by fitting a 2D Gaussian to the visibilities. The model hyperparameters for the brightness profile fit are $\alpha = 1.05$, $w_{\text{smooth}} = 10^{-4}$, $R_{\text{out}} = 1.5$ arcsec, $N = 500$, $p_0 = 10^{-15}$ Jy².

in Fig. C1, exhibits a change in slope at the location of the 5-AU gap in the `FRANK` fit, giving further credence to this feature. The fast oscillations in the `FRANK` brightness profile are artefacts of the visibility fit. The `FRANK` fit shows a large improvement in accuracy in the visibility domain relative to the 1D Fourier transform of a brightness profile extracted from the `CLEAN` image,⁹ with a factor of ≈ 11 lower χ^2 . The `FRANK` fit to the Taurus survey observations of CI Tau (not shown) does not resolve any indication of the new features seen in the fit to the higher resolution data.

⁹The `CLEAN` image was generated using `tclean` in `CASA` 5.6.1-8 with the multiscale deconvolver (pixel size of 10 mas and scales of 1, 2, 4, 6 pixels) and Briggs weighting with a robust value of 0.5.

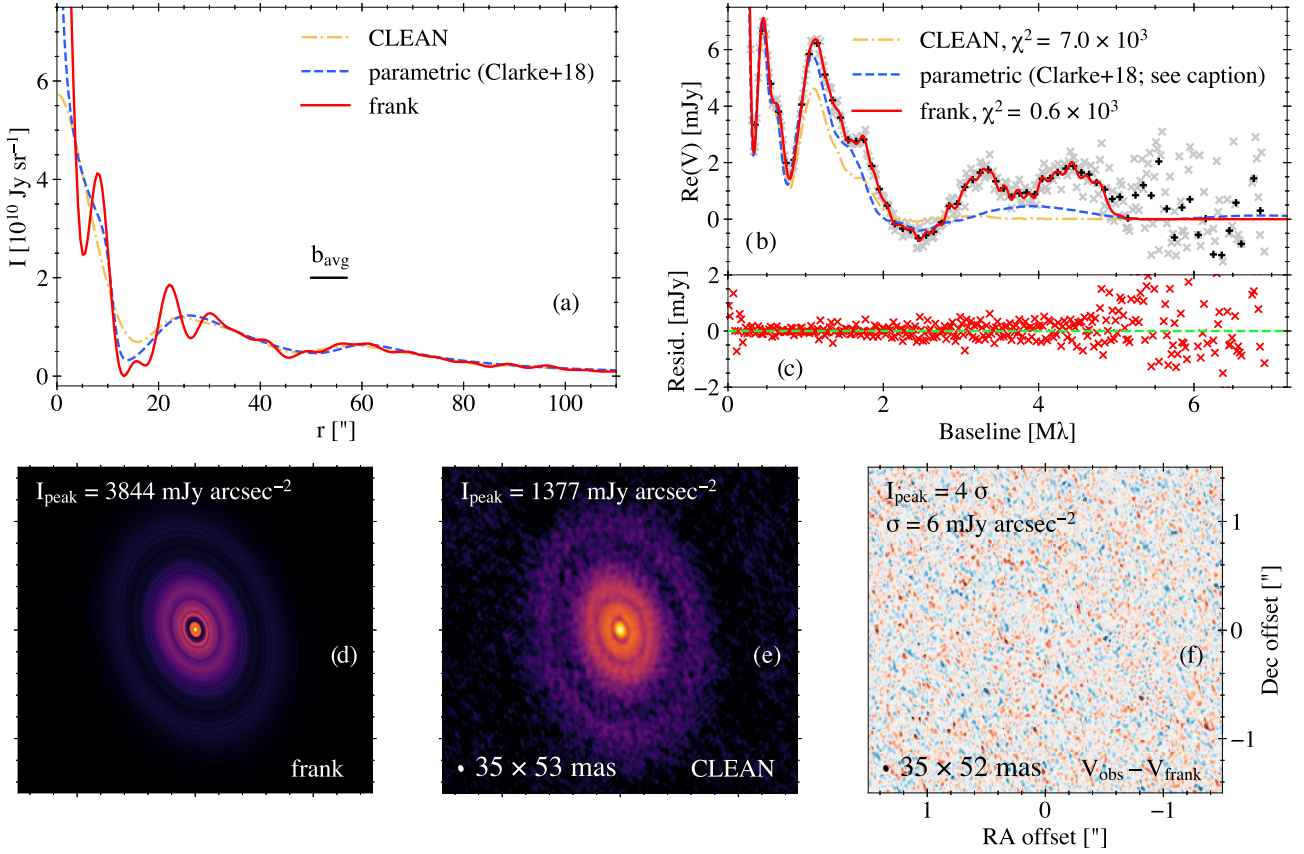


Figure C1. *FRANK* fit to high-resolution CI Tau observations. Panel (a): brightness profile fits for ≈ 40 -mas observations of CI Tau, with the parametric fit from Clarke et al. (2018), *FRANK* fit (which peaks at $16 \times 10^{10} \text{ Jy sr}^{-1}$), and the *CLEAN* image brightness profile. b_{avg} shows the mean of the *CLEAN* beamwidth along its major and minor axes. Panel (b): observed visibilities (20- and 100- $k\lambda$ bins) and fits corresponding to the brightness profiles in panel (a). The parametric fit used a different frequency normalization to convert the (u, v) distances into units of $[\lambda]$ and a different geometry to deproject the visibilities, so it is not directly comparable to the data shown here, the *FRANK* fit or the *CLEAN* fit. Hence, we do not report a χ^2 . Panel (c): residuals of the *FRANK* visibility fit (20- $k\lambda$ bins). Panels (d)–(f): an image of the *FRANK* profile swept over 2π and reprojected; the *CLEAN* image; and the imaged *FRANK* residual visibilities (zero *CLEAN* iterations; contours at -3σ , $+3\sigma$). The *FRANK* and *CLEAN* images use an arcsinh stretch ($I_{\text{stretch}} = \text{arcsinh}(I/a) / \text{arcsinh}(1/a)$, $a = 0.02$), but different brightness normalization (indicated by the given peak brightness). The imaged *FRANK* residuals use a linear stretch symmetric about zero.

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