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# Why “Solar Tsunamis” Rarely Leave Their Imprints in the Chromosphere

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## Abstract

Solar coronal waves frequently appear as bright disturbances that propagate globally from the eruption center in the solar atmosphere, just like the tsunamis in the ocean on Earth. Theoretically, coronal waves can sweep over the underlying chromosphere and leave an imprint in the form of Moreton wave, due to the enhanced pressure beneath their coronal wave front. Despite the frequent observations of coronal waves, their counterparts in the chromosphere are rarely detected. Why the chromosphere rarely bears the imprints of solar tsunamis remained a mystery since their discovery three decades ago. To resolve this question, all coronal waves and associated Moreton waves in the last decade have been initially surveyed, though the detection of Moreton waves could be hampered by utilizing the low-quality  $H\alpha$  data from the Global Oscillations Network Group. Here, we present eight cases (including five in the Appendix) of the coexistence of coronal and Moreton waves in inclined eruptions where it is argued that the extreme inclination is key to providing an answer to address the question. For all these events, the lowest part of the coronal wave front near the solar surface appears very bright, and the simultaneous disturbances in the solar transition region and the chromosphere predominantly occur beneath the bright segment. Therefore, evidenced by observations, we propose a scenario for the excitation mechanism of the coronal-Moreton waves in highly inclined eruptions, in which the lowest part of a coronal wave can effectively disturb the chromosphere even for a weak (e.g., B-class) solar flare.

*Unified Astronomy Thesaurus concepts:* [Solar coronal waves \(1995\)](#); [Solar activity \(1475\)](#); [Solar corona \(1483\)](#); [Solar atmosphere \(1477\)](#)

*Supporting material:* animations

## 1. Introduction

The solar atmosphere is composed of magnetized plasma and hosts a variety of waves that are generally believed to account for many fundamental processes in solar physics (De Pontieu et al. 2007; Jess et al. 2009). One of the most spectacular waves are the solar coronal waves often detected as circular disturbances traveling outward from their eruption source (Moses et al. 1997; Thompson et al. 1998), similar to tsunamis on the Earth. Coronal waves are often associated with large-scale coronal mass ejections (CMEs), flares, or even small-scale jets. They have been interpreted differently, from e.g., the true fast-mode magnetohydrodynamic (MHD) waves (Wu et al. 2001; Ofman & Thompson 2002), pseudo waves (Attrill et al. 2007; Delannée et al. 2008), or hybrid waves (Chen & Wu 2011; Liu & Ofman 2014; Warmuth 2015). Fast-mode coronal waves, by means of solar magnetoseismology, can yield estimates of the coronal magnetic field strengths that are hard to measure directly (Nakariakov et al. 1999; Ballai 2007), and may also accelerate solar energetic particles that influence space weather (Carley et al. 2013).

Before their observational discovery, the existence of coronal waves was already conjectured by the detection of the chromospheric Moreton waves, first detected in the 1960s (Moreton & Ramsey 1960; Moreton 1960; Meyer 1968). A typical Moreton waveforms not far ( $\sim 100$  Mm) from an eruption center, and appears as an arc-shaped chromospheric disturbance. They propagate with a high speed ( $\sim 1000$  km s<sup>-1</sup>) for minutes before becoming irregular and diffusive. Moreton waves are interpreted as the ground tracks of coronal fast-mode MHD wave fronts sweeping over the chromosphere (Uchida 1968, 1974). In theory, the coronal wave front leads to local enhanced pressure that compresses the chromospheric plasma, and therefore these Moreton wave fronts appear in emission in the line core and blue wing of  $H\alpha$ , and in absorption in the red wing (Vršnak et al. 2002; Vršnak & Cliver 2008). In addition, Moreton waves are frequently associated with solar type II radio bursts, which indicates that their coronal counterparts have evolved into shocks (Mann 1995).

With the high-quality observations from the Ahead (-A) and Behind (-B) spacecraft of Solar Terrestrial Relations Observatory (Kaiser et al. 2008) (STEREO) and the Solar Dynamics Observatory (Pesnell et al. 2012; SDO), hundreds of coronal waves have been found and studied. However, only tens of Moreton waves have been observed. Since chromospheric Moreton waves are the counterparts of



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coronal waves, some fundamental questions rise naturally: Why did most coronal waves fail to ignite a detectable chromospheric disturbance? Why was no circular-shaped Moreton wave observed for all the coronal waves with circular-shaped fronts? Moreover, why are all previous Moreton waves only associated with strong M- and X-class flares? On the other hand, if a coronal wave can compress the chromosphere, the transition region (TR) as the interface between them should be compressed first, but the wave signatures in TR are also rarely reported. Hence, although it is generally believed that Moreton waves are the chromospheric counterpart of coronal waves, the bridge between two phenomena still remains elusive.

Combining observations from the dual perspectives by SDO and STEREO, we initially surveyed all coronal waves observed in the last decade from 2010 to 2021,<sup>9</sup> and the associated Moreton waves were checked by the line core of  $H\alpha$  from Global Oscillations Network Group (GONG). Note that most of the detected Moreton waves are only distinguished in animations and very faint in still images, due to the low-quality data of GONG. Though the low-quality data could decrease the detection number of total Moreton waves, the continuous observations can be related to all the coronal waves. Because the same condition of low-quality GONG data are used to detect the related Moreton waves for all coronal waves that are studied here, the data factor can be neglected when we analyze the relationship between coronal waves and Moreton waves. The fraction of coronal waves associated Moreton waves is in progress in the upcoming statistics work.

In this Letter, we picked three cases of distinct chromospheric wave signatures and coronal waves associated with inclined eruptions that are identified by the observations in dual perspectives from the sample in the quick view, to first propose a scenario of the plausible excitation mechanism for Moreton waves. In this scenario, Moreton waves are easily ignited by the enhanced pressure behind coronal waves when the eruptions are highly “inclined” and thus the nose parts of coronal waves are close to the solar surface. The strong reliance to highly inclined eruptions can reasonably explain the rarity of detectable Moreton waves.

## 2. Results

Three coronal wave events, presented and analyzed here in details, are respectively related to an X6.9-class flare on 2011 August 9 (E1), an M1.4-class flare on 2021 May 22 (E2), and a B9.5-class flare on 2011 July 3 (E3). Each event is accompanied by a CME and a type II radio burst (Supplementary Figures A1 and A2). The inclined eruptions and coronal waves are shown clearly in dual perspectives from Atmospheric Imaging Assembly (AIA; Lemen et al. 2012) on board SDO and from Extreme Ultraviolet Imager (EUVI; Howard et al. 2008) on STEREO-A (Figure 1 and Figure 1 animation). In the full-disk composite images in AIA 304, 171, and 193 Å (left panels), all the eruptions move highly nonradially (the sectors), and the inclined angles with respect to the solar radial direction are in the range of  $\sim 63^\circ$ – $76^\circ$  estimated by the reconstruction of the eruption cores or lowest part of coronal wave fronts (Supplementary Figures A3 and A4). Following the inclined eruptions, the coronal waves also propagate unilaterally. In the local difference images of AIA 193 Å and

EUVI-A 195 Å (middle and right columns in Figure 1), the wave fronts are arc-like (white arrows), and their lowest parts close to the solar surface became very bright (red arrows).

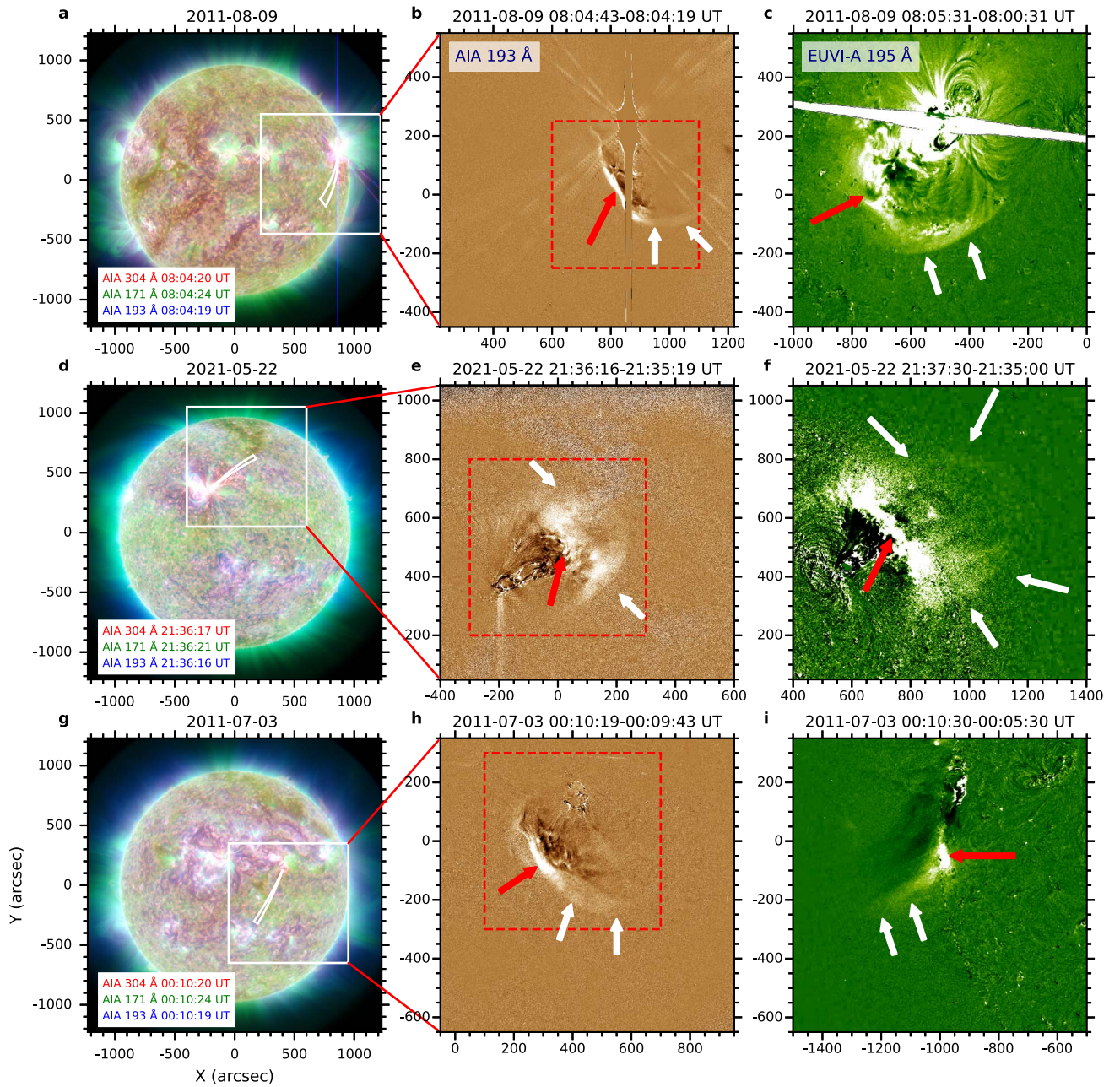
Interestingly, different from the rest of the wave front, the bright segments in AIA 193 Å had evident associated signatures in AIA 171, 304 Å and  $H\alpha$  (Figure 2 and Figure 2 animation). In AIA 171, the wave fronts are replaced by coronal dimmings, which can be interpreted by the heating of local plasma (Liu & Ofman 2014). The wave signatures in AIA 304 Å are evident, and the  $H\alpha$  responses are faint. The signatures in He II 304 Å and  $H\alpha$  can respectively represent the TR and the chromosphere counterparts of coronal waves, and both have arc shapes similar to the bright segments of 193 Å wave, indicating an intimate relation between three waves in different layers. They have similar speeds in the range of  $\sim 380$ – $580$  km s<sup>-1</sup> according to the time–distance plots (Supplementary Figure A5). Assuming a fast magnetosonic speed of  $\sim 200$ – $300$  km s<sup>-1</sup> at a low height for the bright segments (Wu et al. 2001), the coronal waves may manifest as shocks with a fast-magnetosonic Mach number of  $\sim 2$  (Warmuth et al. 2001), in line with the occurrence of the radio type II bursts in Supplementary Figure A2. Note that the disturbance in E3 is the first example of a Moreton wave associated with a weak B-class flare, though the wave signal in  $H\alpha$  is nearly invisible in the time–distance plot (Supplementary Figure A5).

For the coronal wave fronts (cyan boxes in Figure 2), we analyze the evolution of intensity and emission at different passbands (Figure 3). In the upper row, the curves represent the changes of the normalized intensities in AIA 193 (green), 171 (black), 304 Å (red), and GONG  $H\alpha$  (purple). After the wave arrival (pink solid lines), the intensity of AIA 171 Å begins to decrease suddenly, and the intensities in AIA 193 and 304 Å increase abruptly. Though the rise onset of the intensity in GONG  $H\alpha$  is uncertain due to the poor data quality, each of them achieves a distinct peak as those in AIA 193 and 304 Å (dashed lines). The peaks in AIA 193, 304 Å, and GONG  $H\alpha$  are almost simultaneous for E1 and E2. For E3, the peak time in AIA 304 Å is 36 s later than that in AIA 193 Å, but 30 s earlier than that in GONG  $H\alpha$  (panel (c)).

The bottom row of Figure 3 shows the average evolution of the differential emission measure (DEM) in the region of bright segments (cyan boxes in Figure 2) obtained by employing the sparse inversion code (Cheung et al. 2015; Su et al. 2018). After the wave passage (pink solid lines), the DEM at  $\log(T/K) \approx 6.2$  increases obviously, and the DEM at  $\log(T/K) \approx 6.0$  decreases simultaneously (white and pink arrows). It indicates that the local plasma of wave fronts is heated from  $\log(T/K) \approx 6.0$  to  $\log(T/K) \approx 6.2$ , consistent with the bright wave front in AIA 193 Å and the dimmings in AIA 171 Å. Intriguingly, the DEM at  $\log(T/K) \approx 5.7$  also increases simultaneously with that at  $\log(T/K) \approx 6.2$  (red arrows), though the increase in E3 is weak. Moreover, the continuous increase of DEM at  $\log(T/K) \approx 5.7$  seemed to be evolved from that at  $\log(T/K) \approx 5.9$ , which possibly implies their compression relationship. It indicates that the DEM increase at  $\log(T/K) \approx 5.7$  likely represents the emission enhancement for the plasma at the bottom of the corona and/or the upper TR, due to the downward expansion and compression of the bright segment of the coronal wave that is nearly horizontal and close to the solar surface.

<sup>9</sup> [https://www.lmsal.com/nitta/movies/AIA\\_Waves/oindex.html](https://www.lmsal.com/nitta/movies/AIA_Waves/oindex.html)





**Figure 1.** Coronal waves in dual perspectives. Left column: The eruptions are displayed in the full-disk composite images of AIA 304 Å (red), 171 Å (green), and 193 Å (blue). The sectors indicate the directions of eruptions and associated coronal waves, and the white boxes represent the field of view (FOV) of middle panels. Middle and right columns: Coronal waves (white arrows) and bright segments (red arrows) are shown in difference images in AIA 193 Å and EUVI-A 195 Å. The red dashed boxes indicate the FOV of panels in Figure 2. An animation of middle and right columns of this figure is available. The top row of the animation runs from approximately 08:00–08:12 on 2011 August 9. The middle row shows the 21:30–21:45 UT sequence on 2011 May 22. The bottom row of the animation gives the 00:00–00:10 UT sequence on 2011 June 4. The real-time duration of the animation is 2 s.

(An animation of this figure is available.)

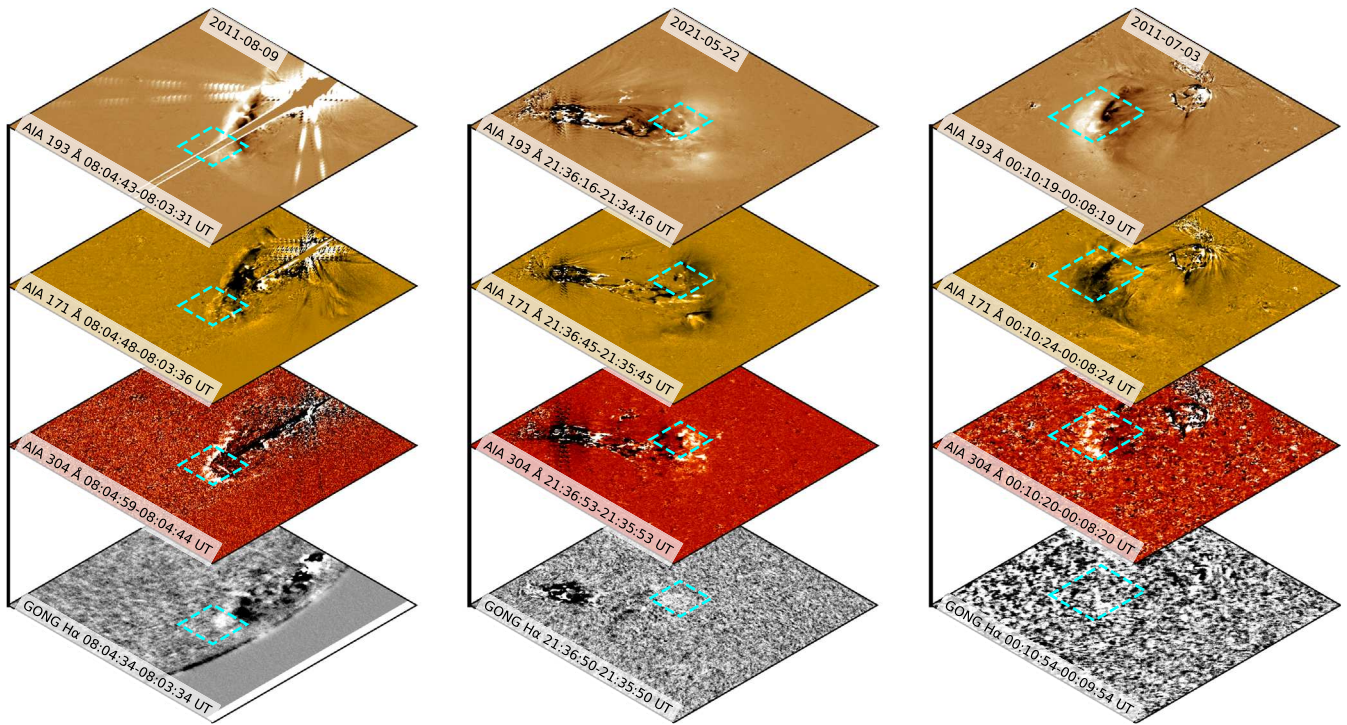
### 3. Discussion and Conclusions

Vršnak et al. (2016) conjectured that maybe an asymmetric/nonradial eruption is an additional requirement for the generation of the Moreton wave in a weak eruption, when they simulated the Moreton waves associated with impulsive eruptions. In the few cases of coronal and Moreton waves observed by SDO and ground-based observatories, all the related eruptions are nonradial (Asai et al. 2012; Shen & Liu 2012; Francile et al. 2016; Cabezas et al. 2019; Long et al. 2019; Wang et al. 2020), and the inclined eruptions could be

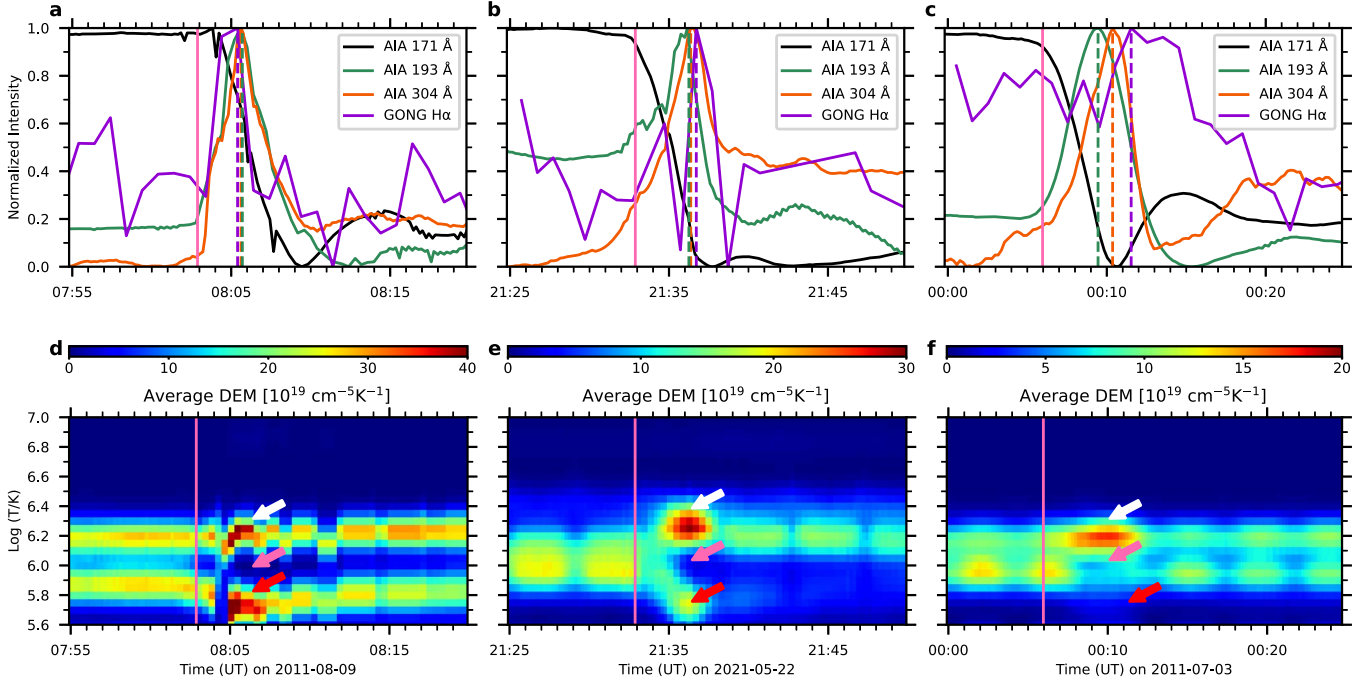
the result of filaments erupting into weaker magnetic field regions (Panasenco et al. 2013). For instance, the inclined eruption and associated Moreton wave on 2014 March 29 is shown in Supplementary Figure A6 and Figure A6 animation), which has been studied by other researchers (Francile et al. 2016; Cabezas et al. 2019; Long et al. 2019).

The AIA 304 Å channel is mainly dominated by the two He II 303.8 Å lines with a temperature of  $\log(T/K) \approx 4.7$  for coronal holes, quiet Sun (QS) regions, active regions, and flare plasma. However, in QS off-limb regions the channel can also make an important contribution by Si XI 303.33 Å line at the





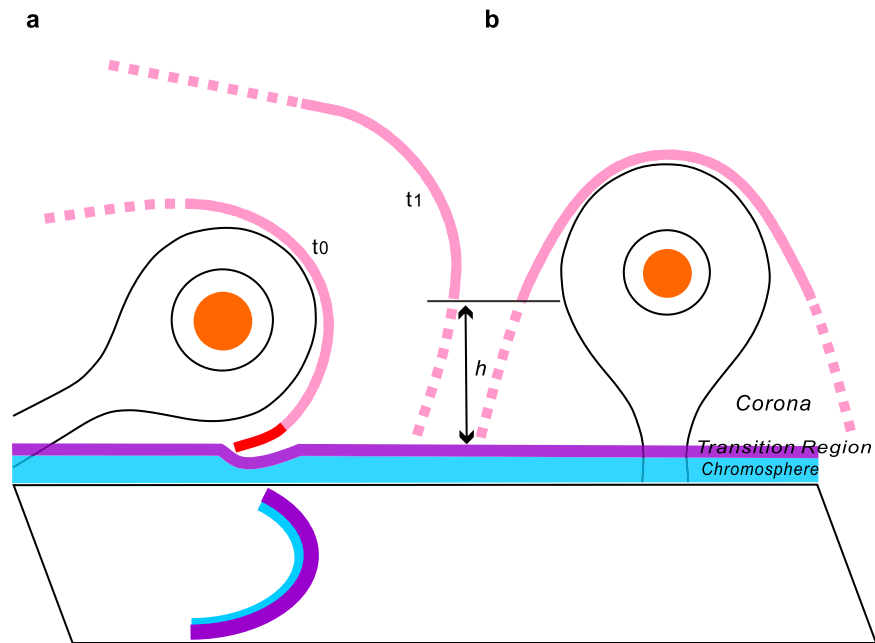
**Figure 2.** The wave signatures in different atmospheric layers. The cyan boxes in difference images outline the bright segment of coronal waves in AIA 193 Å and their responses in AIA 171, 304 Å and GONG H $\alpha$ . The cyan boxes indicate the regions for the DEM evolution in Figure 3. The left row gives the sequence from 08:00 to 08:14 UT on 2011 August 9. The middle row shows the 21:30–21:45 UT sequence on 2021 May 22. The right row of the animation gives the evolution from 00:00 to 00:14 UT on 2011 June 3. The real-time duration of the animation is 2 s. (An animation of this figure is available.)



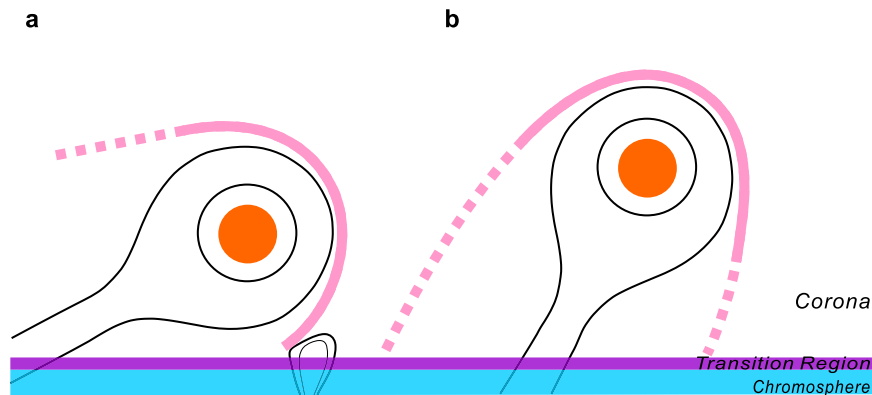
**Figure 3.** The evolution for the intensity and emission of plasma. Upper row: The light curves in the region of bright segments (cyan boxes in Figure 2) in AIA 193, 171, 304 Å and GONG H $\alpha$ . Bottom row: The average DEM evolutions in the region of bright segments (cyan boxes in Figure 2). The pink solid lines indicate the arrival times of coronal waves, and the colored dashed lines represent the peak times in the associated passbands, and the colored arrows show the DEM changes at different temperatures.

coronal temperature of  $\log(T/K) \approx 6.2$  (O’Dwyer et al. 2010). Some previous observations have shown that wave signatures in AIA 304 Å channel could be dominated in a coronal line of

Si XI 303.33 Å (Long et al. 2010), or the He II 303.8 Å lines in the TR (Long et al. 2011). In this Letter, it is very clear that the wave fronts in 304 Å were narrow, similar to wave fronts in



**Figure 4.** The scenario of exciting mechanism for Moreton waves. (a): For an inclined eruption with a core (orange), the lowest part of the associated coronal wave (the solid–dotted pink front) becomes very bright (the red patch) at the onset ( $t_0$ ), and compresses the TR (the purple layer) and chromosphere (the blue layer) below. The projected purple and blue arcs represent the disturbances in the TR and chromosphere. The nose part of the coronal wave front reaches a height of  $h$  from the coronal base after the propagation in a few minutes ( $t_1$ ). (b): For a radial eruption, the nose of the coronal waveforms at an initial height of  $h$ . Note that the coronal wave degenerates from a shock wave near the nose part (solid line) to an ordinary wave further away (dotted line).



**Figure 5.** The unsuccessful situations for Moreton waves in inclined eruptions. (a): The coronal wave encounters the low-lying coronal loops. (b): The inclination angle for the eruption is small with respect to the radial direction. The coronal wave (the solid–dotted pink curve) consists of a strong nose (solid) and weak flanks (dotted).

$H\alpha$ , and were different with wave fronts in  $193 \text{ \AA}$ . Like Moreton waves, the wave fronts in  $304 \text{ \AA}$  were narrow, and closely related to the bright parts of coronal wave fronts (red and white arrows in Figure 1 and the dashed boxes in Figure 2, similar to the case in Hou et al. (2022)). In addition, the wave fronts in  $304 \text{ \AA}$  can be shown in the disk in the perspective of SDO or STEREO. Hence, the morphological difference in the disk between the wave fronts in  $304$  and  $193 \text{ \AA}$  suggested that the wave fronts in  $304 \text{ \AA}$  in this Letter were much likely dominated by the He II line.

To further uncover the relationship between the coronal and Moreton waves, we check four more coronal waves that are associated with either inclined limb eruptions or radial eruptions. In the inclined limb eruptions (Supplementary Figure A7 and Figure A7 animation), the bright segments of coronal waves and disturbances in the TR are prominent (blue arrows), though the chromospheric wave signatures are hard to identify due to the

poor  $H\alpha$  data quality. Although the radial eruptions are associated with strong X- or M-class flares (Supplementary Figure A8 and Figure A8 animation), no wave signature is found in He II  $304 \text{ \AA}$  and  $H\alpha$ . This is then consistent with the fact that the bulk of coronal waves is obvious in AIA  $193$  and  $211 \text{ \AA}$ , while there is no signature in AIA  $304 \text{ \AA}$ . According to Uchida’s model (Uchida 1968, 1974), the TR should be compressed first if the coronal wave can disturb the chromosphere. Therefore, the waves in He II  $304 \text{ \AA}$  for the above cases likely represent the response in TR, and can be regarded as the missing link between coronal and Moreton waves. Furthermore, for the above cases of inclined eruptions, all coronal waves successfully disturb the TR and chromosphere, including the one associated with the weak B-class flare. This further implies that the inclined eruption is the key to generating chromospheric Moreton waves.

For coronal waves in the inclined eruptions presented above, they display a bright segment of wave fronts near the solar

surface, distinct from other parts (Figure 1), and the disturbances in the TR and chromosphere simultaneously occur just under the brightest segment (Figure 2). Comparing with the evolution in different regions of wave fronts, it is likely that the average DEM evolution in the region of brightest segment reveals the emission enhancement in the coronal base and the upper TR behind the coronal shocks/waves (Figure 3), consistent with the scenario of the Moreton wave (Vršnak et al. 2002). Hence, the bright segments of coronal wave fronts and the related continuous increase of DEM indicate that the TR and chromosphere can be effectively disturbed by the lowest part of the inclined coronal waves, and provide further evidence of the role of inclined eruptions in driving the Moreton waves.

Based on the above observational results, we propose a plausible excitation mechanism for Moreton waves in Figure 4. The coronal wave (the solid-dotted pink curve) consists of a shocked nose (solid part) and weak flanks (dotted parts). For the coronal wave in an inclined eruptions (panel (a)), the lowest part of the wave front is close to the nose part of the shock wave; hence it is relatively strong at the onset (at  $t_0$ ). The expanding lowest part exerts a strong compression toward the solar surface resulting in the bright segment (the red patch). The compression ignites a pressure jump that will cause a density/temperature increase not only in the TR but also in the chromosphere, in the form of wave signatures in the TR and chromosphere (the distortions and the projected arcs). When the nose part of the coronal wave front lifts higher later (at  $t_1$ ), the lowest part, being further away from the nose part of the coronal wave, becomes so weak that it hardly compresses the TR beneath, and consequently the wave signatures in the TR and chromosphere vanish. On the contrary, in the radial eruption scenario (panel (b)), the nose of coronal wave initially forms at a typical height of 50–100 Mm (Patsourakos & Vourlidas 2012) a few minutes after the early rising phase of the eruption core. The lowest rim of the coronal wave front is far from the nose part of the shock (above a height of  $h$ ) and becomes too weak to imprint the TR and the chromosphere.

In this scenario, Moreton waves are easily generated by coronal waves that are associated with inclined eruptions. However, would all coronal waves with inclined eruptions leave imprints in the chromosphere? We show two more Moreton waves associated with inclined eruptions (Supplementary Figure A9 and Figure A9 animation). The Moreton wave associated with an X1.8 flare on 2011 September 7 was hardly distinguishable, and was only confirmed by the oscillation of a low-lying filament and nearby loops (the pink and black arrows in top panels in Figure A9). The Moreton wave (blue arrows) associated with an X1.0 flare on 2021 October 28 was also weak, and the westward erupting prominence (yellow arrows) had an inclination angle of  $\sim 30^\circ$  calculated with reconstructed points of the erupting prominence in dual perspectives (bottom panels in Figure A9). Hence, the scenario can be complimented by the factors of the scale of magnetic structures beneath the inclined coronal wave fronts and the range of inclination angles. Assuming that there exists an underlying large-scale coronal loop system that would be intact during the downward compression of the coronal wave fronts (Figure 5(a)), the large-scale loop system possibly invokes the MHD wave refractions from regions of high Alfvén speed (Wang 2000). The refraction from the ambient large-scale loop systems indeed contributed to the lack of Moreton waves for some coronal waves in inclined eruptions. On the other hand, if the

inclination angle of the eruption is small to the radial direction, it should become difficult for the coronal wave to disturb the TR and the chromosphere significantly, and the Moreton wave signature will be faint or may even disappear (Figure 5(b)).

Therefore, the scenario introduced above suggests that the reliance on the highly inclined eruptions is the key to resolving the rarity puzzle of Moreton waves. Hence, the comprehensive relationship between the coronal waves and Moreton waves is likely established by the highly inclined eruptions, with the additional factors (the inclined angle and the magnetic structures beneath downward coronal wave fronts) for the scenario. We also imply that the wave in He II 304 Å may be the counterparts of Moreton waves when the Moreton waves are undetectable in low-quality  $H\alpha$  data. According to this scenario, a coronal wave related to an impulsive X-class flare need a moderately inclined eruption, and a coronal wave associated with a weak B-class flare can ignite a detectable Moreton wave when the eruption and the associated coronal wave front are extremely “inclined.” The high-quality chromospheric data from the Chinese  $H\alpha$  Solar Explorer (Li et al. 2022), launched in October of 2021, will provide more cases of Moreton waves to check the scenario. Further simulations are demanded to verify the scenario proposed in this Letter.

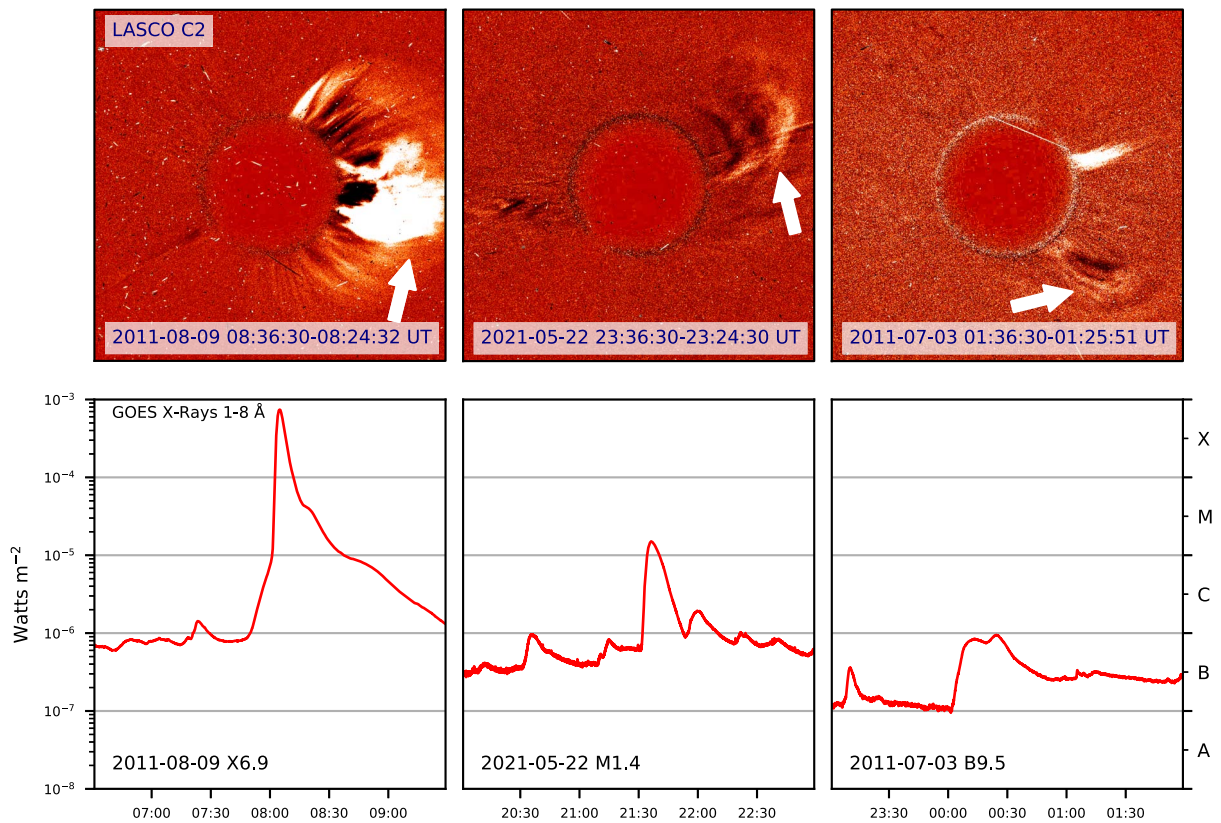
SDO is a mission of NASA’s Living With a Star Program. We gratefully acknowledge the usage of data from the SDO, STEREO, and from the ground-based GONG project. This work is supported by grants of NSFC 11790303, 12073016, 11825301, 11973031, and 12127901. R.E. is also grateful to the Science and Technology Facilities Council (STFC, grant No. ST/M000826/1) for the support received. R.E. also acknowledges support from the European Union’s Horizon 2020 research and innovation programme under the grant agreements No. 739500 (PRE-EST project), No. 824135 508 (SOLARNET project), Marie Skłodowska-Curie grant agreement No 955620, and the support from NKFIH (OTKA, grant No. K142987) Hungary for enabling this research.

## Appendix Appendix Information

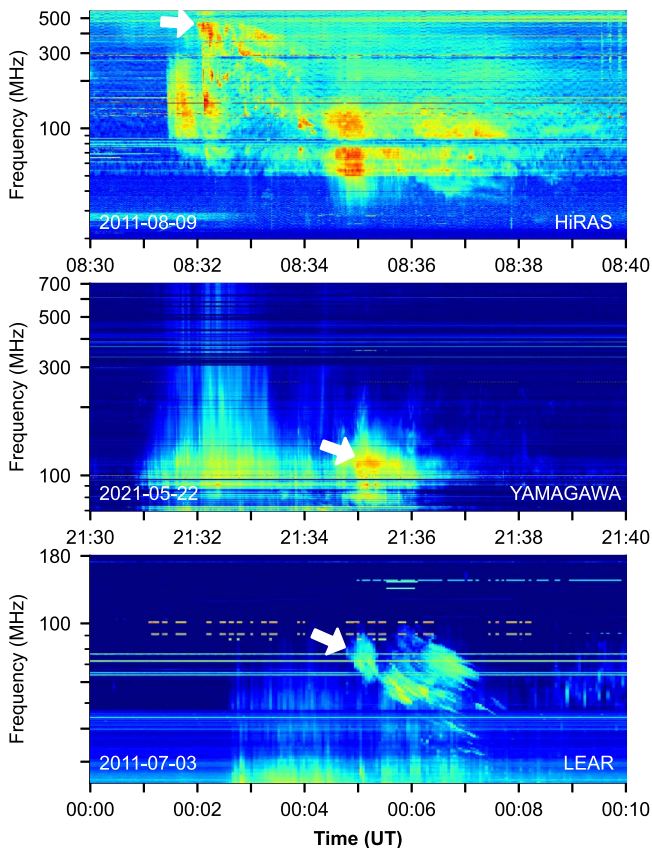
*Data.* The analysis of coronal waves is primarily based on the extreme ultraviolet (EUV) observations in dual perspectives from AIA 304, 171, and 193 Å and EUVI 304, 171, and 195 Å (Figures 1, 2, and associated animations). Each AIA image has a cadence of 12 s and a pixel size of  $0''.6$ . EUVI images have a pixel size of  $1.58''$ , and their cadences are 2.5, 5, or 10 minutes. Moreover, to study the evolution of Moreton waves, we use the  $H\alpha$  filtergrams from GONG of the National Solar Observatory. The  $H\alpha$  images have a cadence of 1 minute and a pixel size of  $\sim 1''$  (Figure 2).

The evolution of the CMEs in the high corona (white arrows) is shown by Large Angle and Spectrometric Coronagraph (Brueckner et al. 1995) on board the Solar and Heliospheric Observatory spacecraft, and the related flares are recorded by the Geostationary Operational Environment Satellite (GOES) in the form of integrated full-disk soft X-ray emission from the Sun (Supplementary Figure A1). In addition, the associated radio bursts (white arrows) were detected in the frequency range of 25–500 MHz by the metric spectrometers from Hirasó Radio Spectrograph (HiRAS) and





**Figure A1.** The associated CMEs and flares. Upper row: The related CMEs (white arrows) are shown in difference images of LACSO C2. Bottom row: The related flares are indicated by GOES X-rays 1–8 Å.

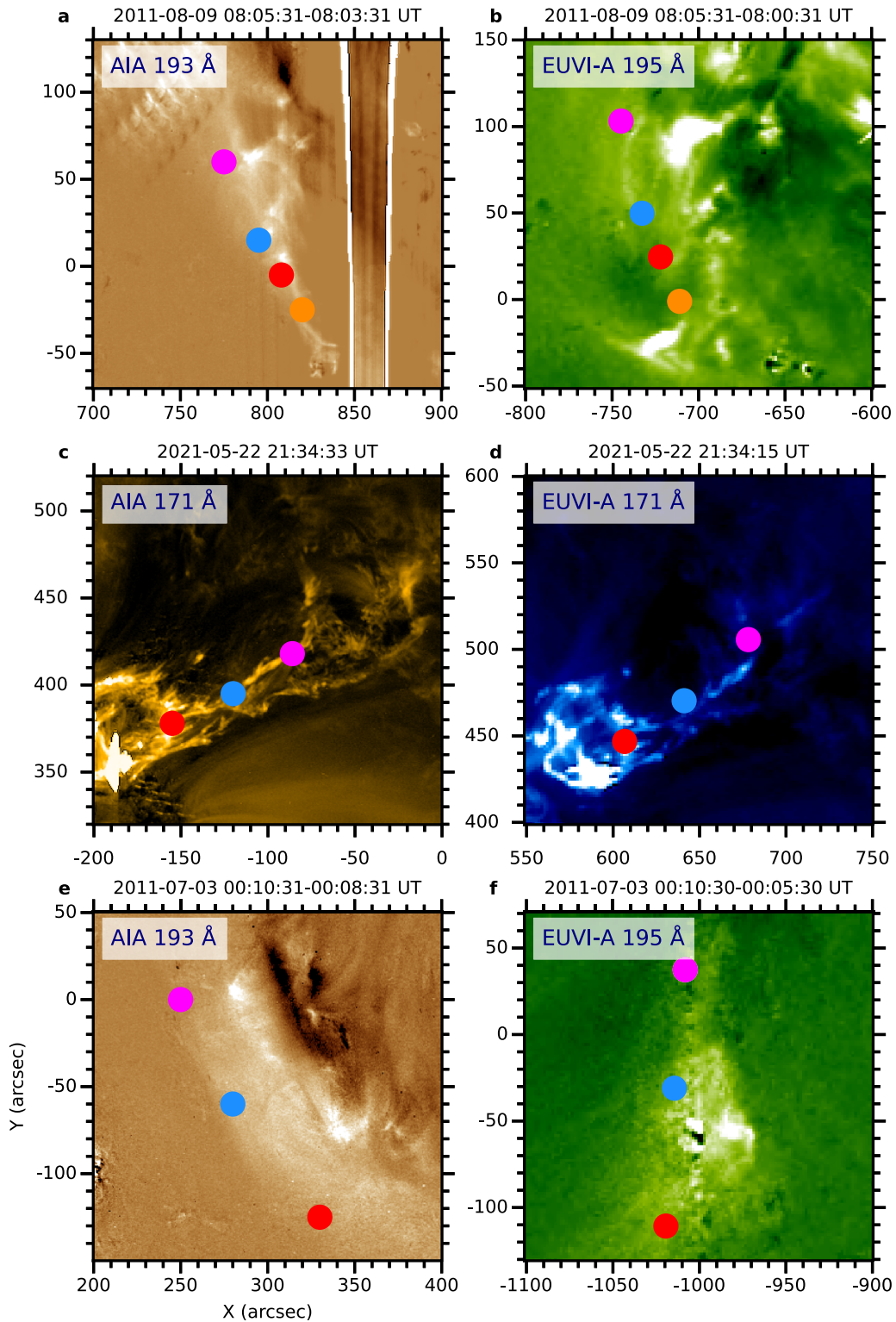


**Figure A2.** The associated radio bursts. The radio dynamic spectrums for three cases are obtained from HiRAS, YAMAGAWA, and LEAR, respectively, and the white arrows indicate the onsets of radio bursts.

YAMAGAWA of National Institute of Information and Communications Technology (Kondo et al. 1995; Iwai et al. 2017), and Learmonth (LEAR) of the radio Solar Telescope Network (Kennewell & Steward 2003) (Supplementary Figure A2).

*Methods.* To identify well the wave fronts and CMEs, the image is subtracted by one previous image, and the times of two images are attached (Figures 1, 2, Supplementary Figures A1, A3, A6, A7, A8, A9, and associated animations). The wave signatures in different layers are analyzed by the time-slice approach, in which a time–distance plot was constructed with a stack of slices along a selected sector for a set of difference images. The wave speeds and errors (blue lines) are obtained by fitting with the linear (*linfit.pro*) function in the SolarSoftWare (SSW), assuming a measurement uncertainty of 4 pixels ( $\sim 1.74$  Mm) for the selected points (Supplementary Figure A5).

Benefiting from the observations in dual perspectives (upper panels in Supplementary Figure A4), we derive the real locations of the erupting cores (the core material of erupting filaments/prominences and jets) and wave fronts by the reconstruction procedure of *scc\_measure.pro* in SSW. To estimate the eruption inclination, we reconstruct some selected points from the erupting core of E2 both that simultaneously appear in AIA 171 Å and EUVI-A 171 Å. For E1 and E3, we choose the bright segment instead of invisible erupting cores (Supplementary Figure A3), assuming that the inclination of the erupting core is approximately equal to that of the bright segment (the orange oval core and the red segment in Figure 4). Finally, we superimpose the reconstructed points on one AIA 304 Å image in which the

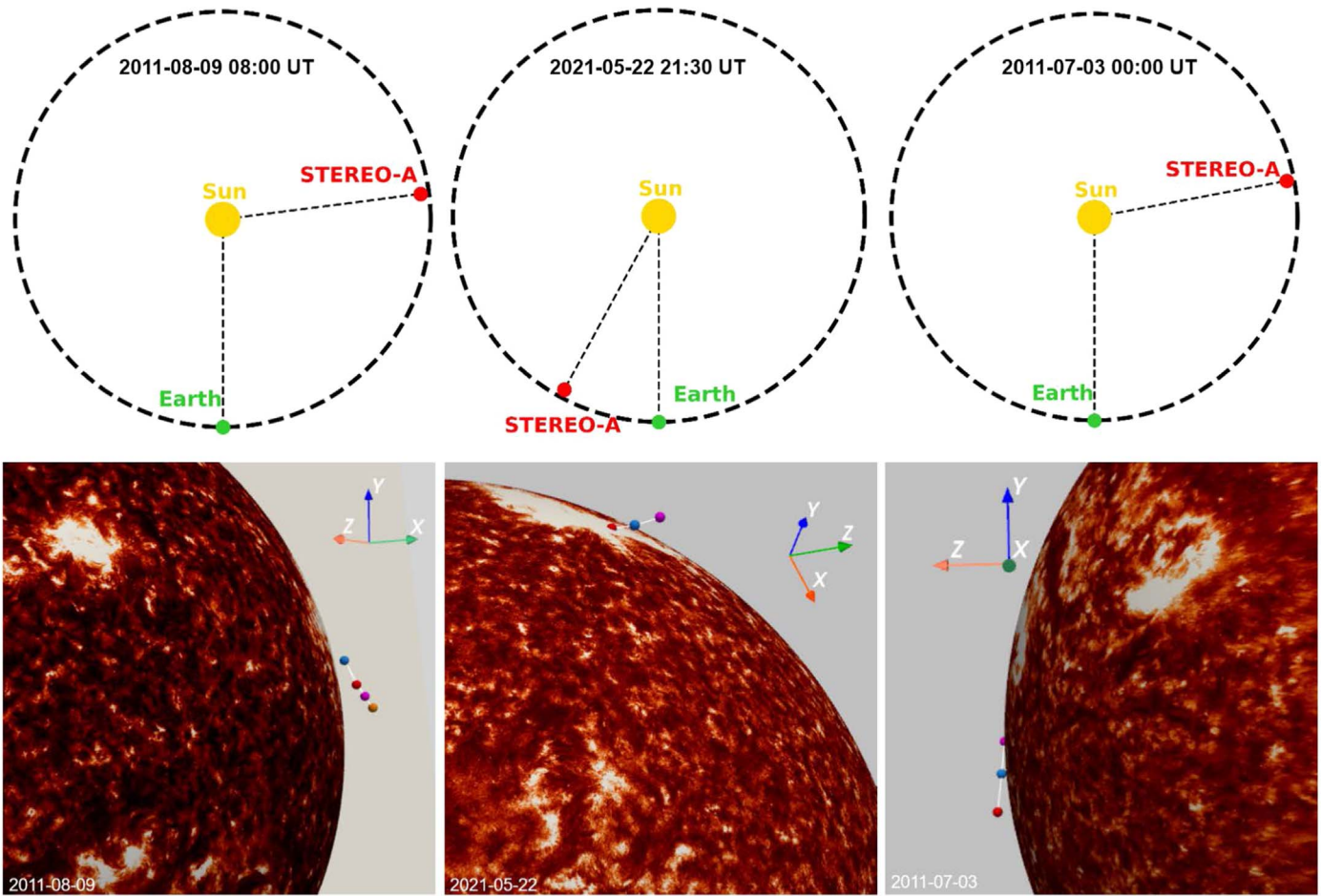


**Figure A3.** The selected points for the reconstructions. The colored points in dual perspectives from AIA and EUVI-A represent the positions for the reconstructions in Supplementary Figure A4.

eruption center is rotated to the limb (bottom panels in Supplementary Figure A4).

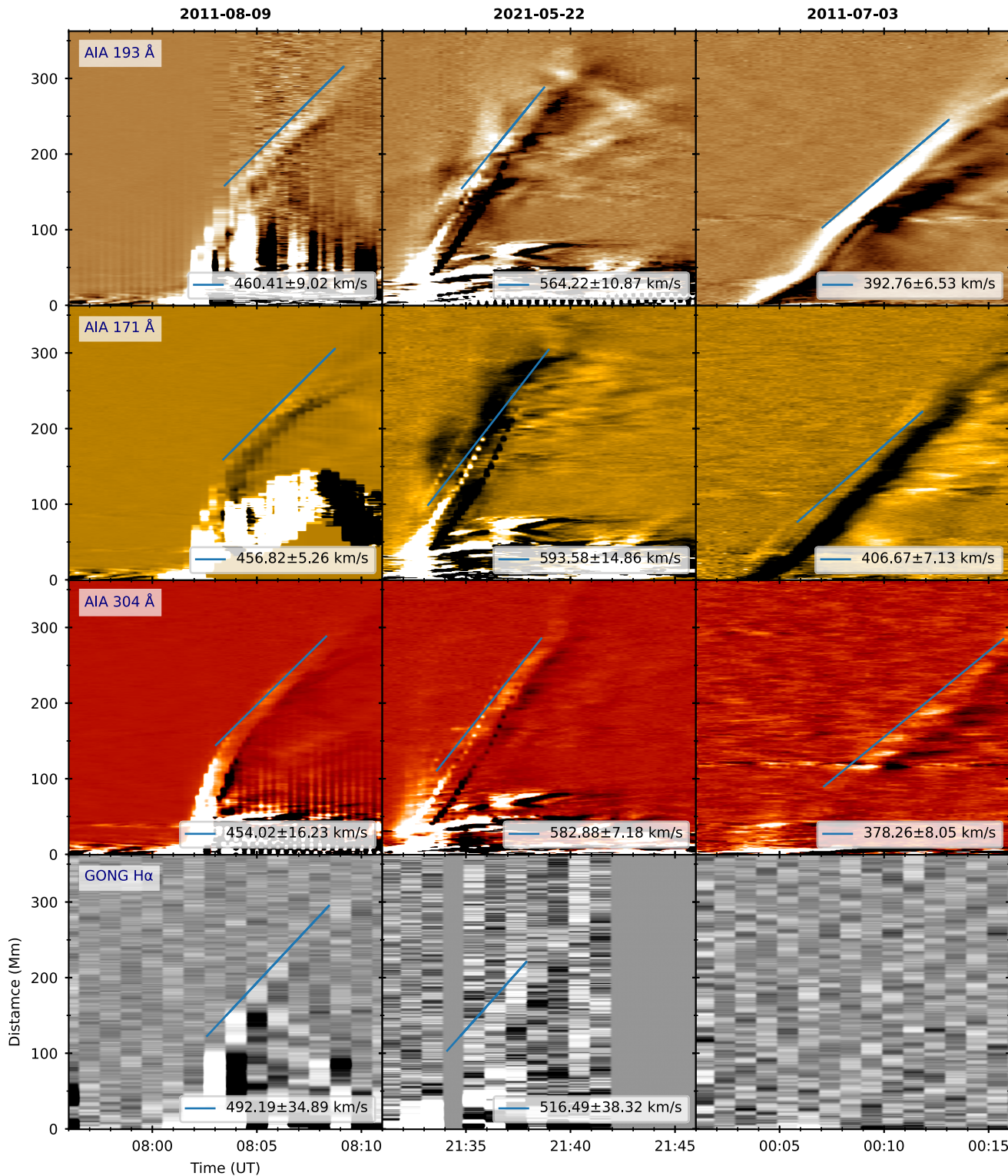
We also investigated the emission evolution in the fronts of coronal waves with the differential emission measure (DEM) method by employing the sparse inversion code (Cheung et al. 2015; Su et al. 2018). The emission at a range of temperature is obtained by a set of AIA images in

six channels, i.e., 131 Å (Fe XXI,  $\sim 10$  MK), 94 Å (Fe XVII,  $\sim 6.4$  MK), 335 Å (Fe XVI,  $\sim 2.5$  MK), 211 Å (Fe XIV,  $\sim 2.0$  MK), 193 Å (Fe XII,  $\sim 1.6$  MK), and 171 Å (Fe IX,  $\sim 0.6$  MK). The DEM evolution plots are constructed with the average value in the selected regions (cyan boxes in Figure 2) for a series of EM maps (bottom panels in Figure 3).

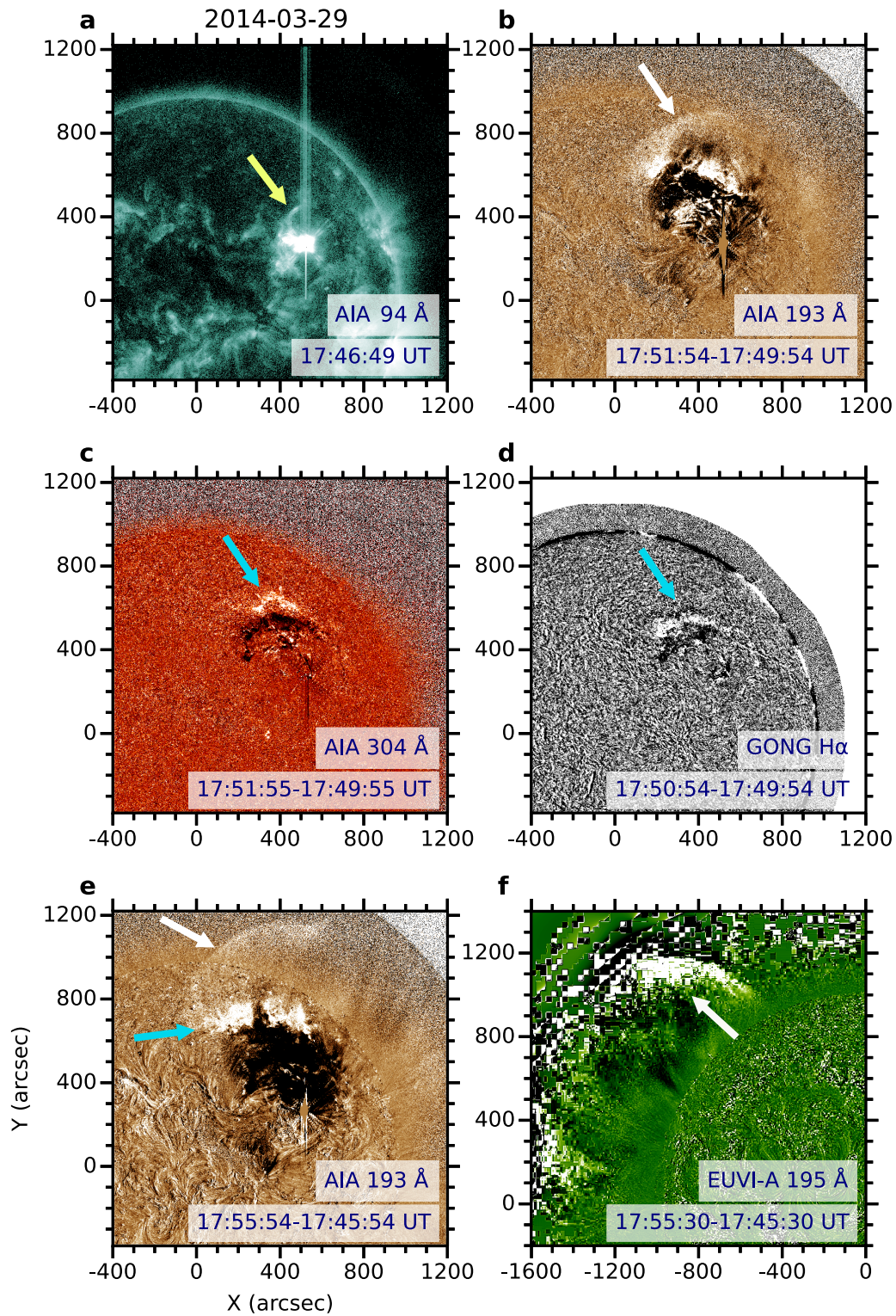


**Figure A4.** The inclination angles of the eruptions. Upper row: The positions of the Sun, the Earth, and the STEREO-A. Bottom row: The selected points in the limb of rotated AIA 304 Å images to show the the inclination angles that are  $\sim 73^\circ$ ,  $\sim 63^\circ$ , and  $\sim 76^\circ$ , respectively. The axis of  $z$  points to the Earth, and the axes of  $x$  and  $y$  define the coordinate plane of AIA images.





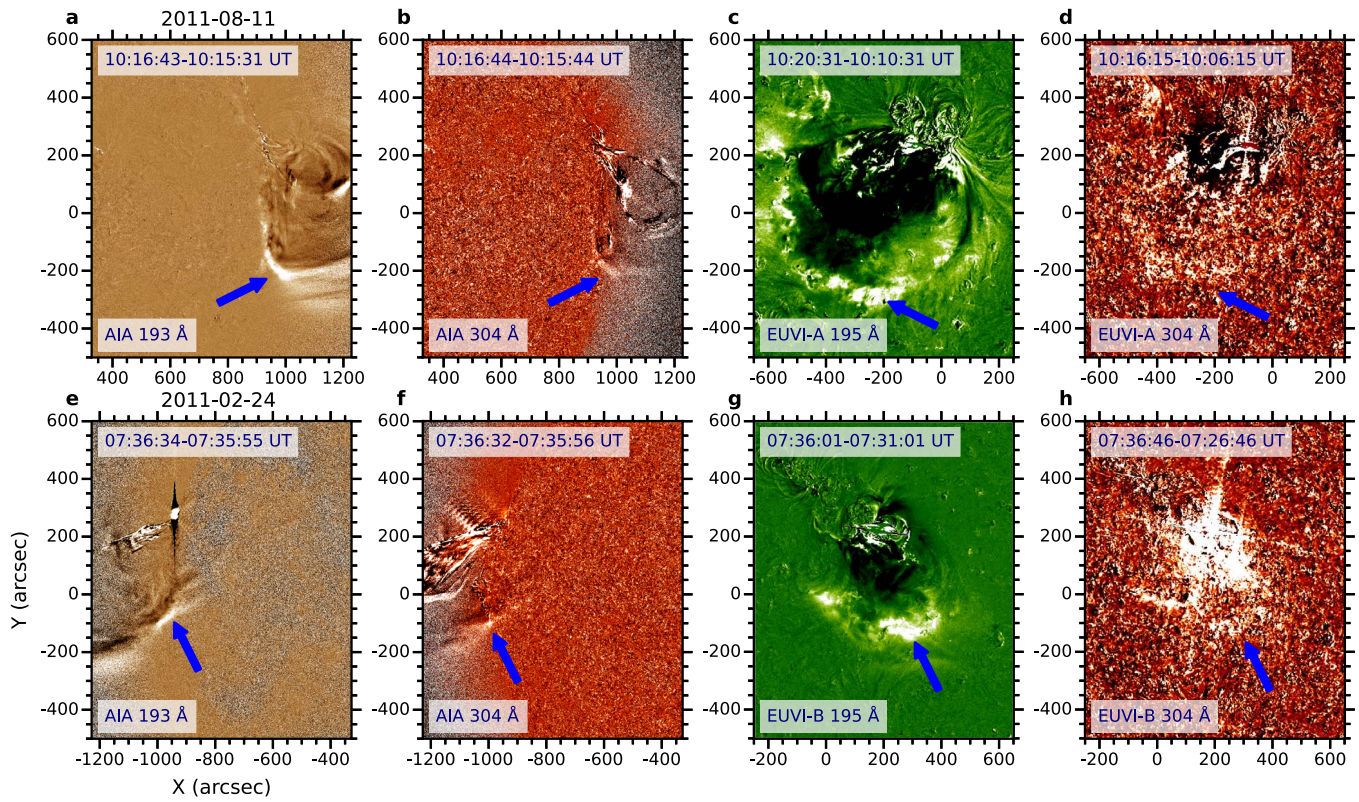
**Figure A5.** Wave velocities in different layers. The velocities in AIA 193, 171, 304 Å and GONG H $\alpha$  are calculated along the sectors in Figure 1 for three cases (left, middle, right columns), and the inclined lines indicate the measured speeds. Note that the chromospheric wave signature for E3 is hardly distinguished in the last panel.



**Figure A6.** The Moreton wave with an X1.0 flare on 2014 March 29. The inclined erupting structure (the yellow arrow) is shown in AIA 94 Å. In the difference images in dual perspectives, the blue arrows show the bright segments of coronal waves (white arrows) in AIA 193 Å and EUVI 195 Å and the simultaneous wave signatures in AIA 304 Å and GONG H $\alpha$ . The associated animation runs from approximately 17:45–18:00 UT on 2014 March 29. The real-time duration of the animation is 2 s.

(An animation of this figure is available.)

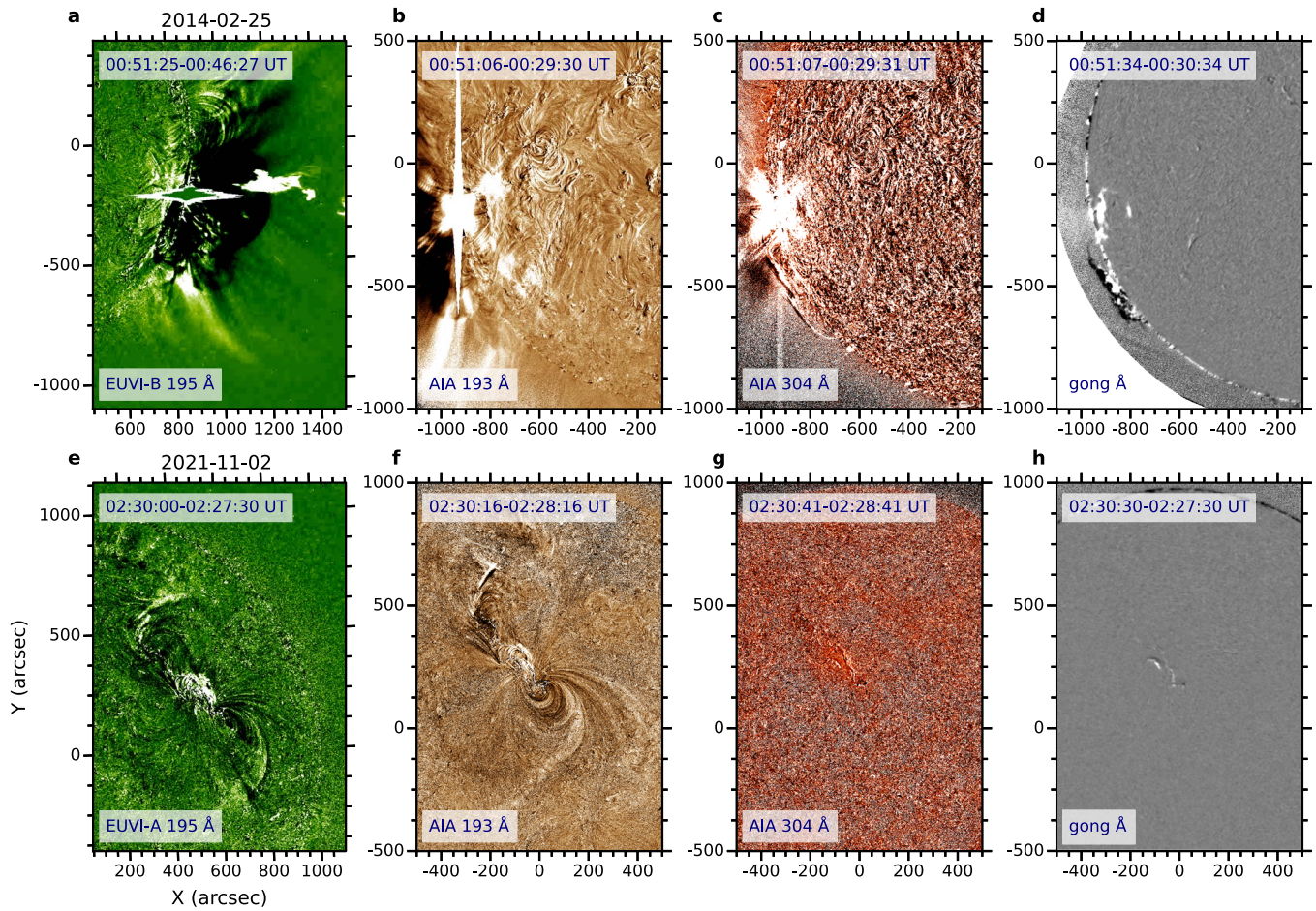




**Figure A7.** Two cases of inclined limb eruptions. Two eruptions are separately accompanied with a C6.2 flare on 2011 August 11 and an M3.5 flare on 2011 February 24. In the difference images in dual perspectives, the blue arrows show the bright segments of coronal waves in AIA 193 Å and EUVI 195 Å and the simultaneous wave signatures in He II 304 Å. The top row of the associated animation runs from approximately 10:05–10:20 UT on 2011 August 11. The bottom row of the animation gives the sequence of 07:25–07:40 UT on 2011 February 24. The real-time duration of the animation is 3 s.

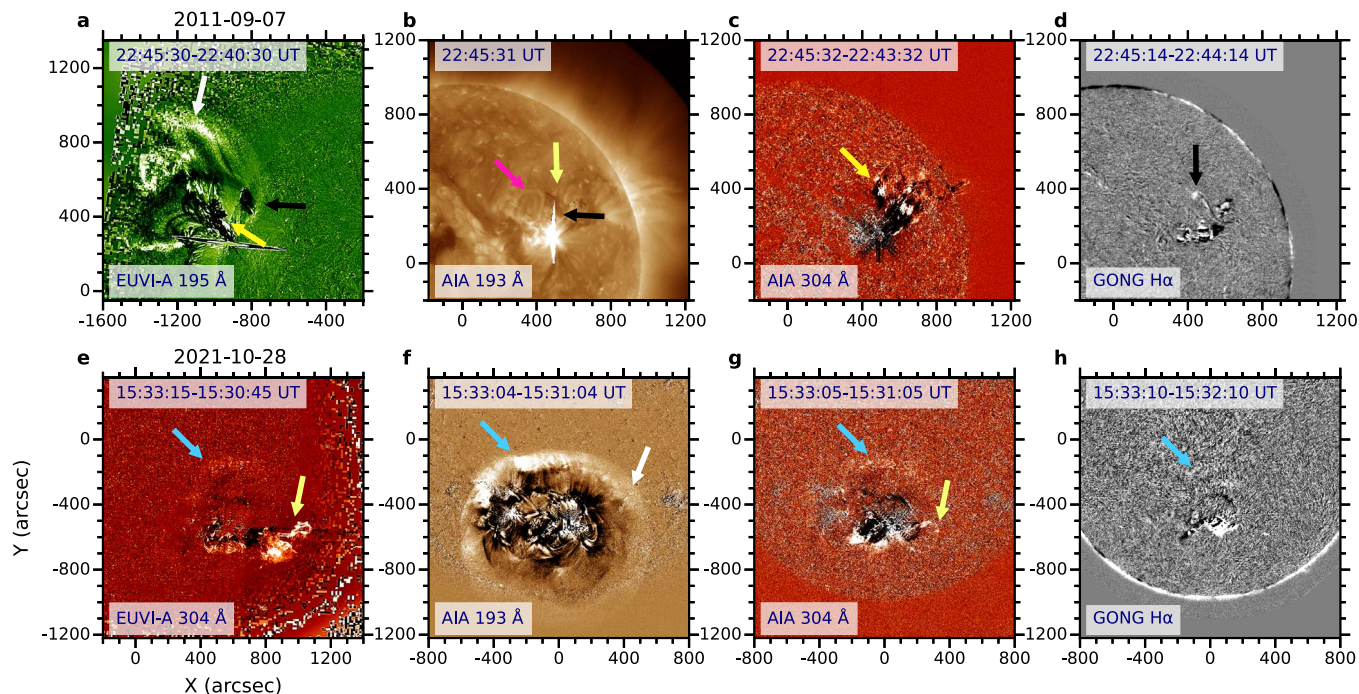
(An animation of this figure is available.)





**Figure A8.** Two cases of radial eruptions. Two eruptions are separately accompanied with an X4.9 flare on 2014 February 25 and an M1.7 flare on 2021 November 2. In the difference images in dual perspectives, the coronal waves are obvious in AIA 193 Å and EUVI 195 Å, but there is no wave signature in He II 304 Å and H $\alpha$ . The top row of the associated animation runs from approximately 00:40–00:52 UT on 2014 February 25. The bottom row of the animation gives the sequence of 02:20–02:32 UT on 2021 November 2. The real-time duration of the animation is 1 s.

(An animation of this figure is available.)



**Figure A9.** Two more cases of inclined eruptions. Two eruptions are separately accompanied with an X1.8 flare on 2011 September 7 and an X1.0 flare on 2021 October 28. In the difference images in dual perspectives, the blue arrows show the bright segments of coronal waves (white arrows) in AIA 193 Å and EUVI 195 Å and the simultaneous wave signatures in AIA 304 Å and GONG H $\alpha$ . The waves are closely associated with inclined erupting filaments (yellow arrows), and are affected by the ambient coronal loops (the pink arrow) and the underlying filament (black arrows). The top row of the associated animation runs from approximately 22:32–22:47 UT on 2011 September 7. The bottom row of the animation gives the sequence of 15:25–15:40 UT on 2021 October 28. The real-time duration of the animation is 1 s.

(An animation of this figure is available.)

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