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First light with HiPERCAM on the GTC

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First light with HiPERCAM on the GTC

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ABSTRACT

HiPERCAM is a quintuple-beam imager that saw first light on the 4.2 m William Herschel Telescope (WHT) in October 2017 and on the 10.4 m Gran Telescopio Canarias (GTC) in February 2018. The instrument uses re-imaging optics and 4 dichroic beamsplitters to record *ugriz* (300–1000 nm) images simultaneously on its five CCD cameras. The detectors in HiPERCAM are frame-transfer devices cooled thermo-electrically to -90°C , thereby allowing both long-exposure, deep imaging of faint targets, as well as high-speed (over 1000 windowed frames per second) imaging of rapidly varying targets. In this paper, we report on the as-built design of HiPERCAM, its first-light performance on the GTC, and some of the planned future enhancements.

Keywords: instrumentation: detectors – instrumentation: photometers – techniques: photometric

1. INTRODUCTION

The study of astrophysical objects that vary in brightness is set to be revolutionised in the coming years with the advent of major new survey facilities like Gaia, SKA, LSST and Euclid. The plethora of interesting new targets that these facilities are expected to reveal will require detailed follow-up on large-aperture telescopes. HiPERCAM has been designed to perform this follow-up role, providing CCD imaging in five optical bands (*ugriz*) simultaneously at frame rates ranging from 0.001 Hz to 1000 Hz. The scientific motivation and preliminary design of HiPERCAM have been described before,¹ and details of the HiPERCAM data acquisition system are described elsewhere at this conference.² In this paper, we report on the final design and first-light performance of the instrument when mounted on the GTC on La Palma. We also describe some of the instrument enhancements we are planning to implement in the near future.

2. FINAL DESIGN

The HiPERCAM project began in January 2014, the start date of the €3.5M European Research Council Advanced Grant that funded the instrument. First light took place just under 4 years after this date, on budget and on time.

The design of HiPERCAM is based on our successful predecessor instrument, ULTRACAM,³ but offers a very significant advance in performance, as shown in Table 1 and described in the following sections.

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Table 1. Comparison of ULTRACAM and HiPERCAM.

	ULTRACAM	HiPERCAM	Notes
Number of simultaneous bands	3 ($ug + r/i/z$)	5 ($ugriz$)	
Readout noise	$3 e^-$ at 100 kHz	$4.5 e^-$ at 263 kHz	Dummy output (HiPERCAM)
CCD operating temperature	233 K	183 K	
Dark current	$360 e^-/\text{pix}/\text{hr}$	$100 e^-/\text{pix}/\text{hr}$	
Longest exposure time	30 s	1800 s	
Highest frame rate	400 Hz	1050 Hz	24×24 pixel windows, bin 6×6
Field of view on WHT	5.1'×5.1'	10.2'×5.1'	platescale 0.3"/pixel
Field of view on GTC	–	2.8'×1.4'	0.081"/pixel (HiPERCAM)
Probability of $r = 10$ comparison	45%	78%	WHT, galactic latitude = 30°
Comparison star pick-off	No	Yes	Under development – Sect. 2.6
Dummy CCD outputs	No	Yes	
Deep depletion	No	Yes	
QE at 700/800/900/1000 nm	83/61/29/5 %	88/78/53/13 %	
Fringe suppression CCDs	No	Yes	
Fringe amplitude in z	>10%	<1%	

2.1 Optics

A ray trace through the HiPERCAM optics has been presented before.¹ Light from the telescope is first collimated and then split into five beams using four dichroic beamsplitters. The beam in each arm then passes through a re-imaging camera, which focuses the light through a bandpass filter and cryostat window onto a CCD. The HiPERCAM collimator has been designed for use on both the 4.2 m William Herschel Telescope (WHT) on La Palma and the 3.5 m New Technology Telescope (NTT) on La Silla, giving a platescale of 0.3"/pixel and 0.35"/pixel, and a field of view of 11.4' and 13.4', respectively, along the diagonal of the detector. The same collimator also gives excellent optical performance on the 10.4 m Gran Telescopio Canarias (GTC) on La Palma, giving a platescale of 0.081"/pixel and a field of view of 3.1' (diagonal).

The dichroics, lens barrels, filters and CCDs are housed in/on an aluminium hull, which forms a sealed system to light and dust – see Fig. 1. The bandpasses of the five arms are defined by a set of so-called “Super” SDSS filters (Fig. 2), which were designed specifically for HiPERCAM. These filters do not use coloured glasses, but instead rely only on multi-layer coatings to define the filter bandpasses, with the cut-on/off wavelengths designed to match the original SDSS filter set.⁴ The percentage improvements in throughput of the HiPERCAM Super SDSS filters, which we call $u_s g_s r_s i_s z_s$, over the original SDSS filters, $ugriz$, are 41/9/6/9/5 %, respectively.

2.2 Detectors

The HiPERCAM CCDs have been described in detail elsewhere.^{1,2} The specifications of the five detectors currently in use in HiPERCAM are detailed in Table 2.

The HiPERCAM detectors need to be cooled to below 187 K in order to achieve the dark current requirement of $< 360 e^-/\text{pixel}/\text{hr}$, which corresponds to 10% of the faintest sky level we can observe with HiPERCAM (set by observations in the u -band in dark time on the GTC). Cooling to below 187 K therefore ensures that dark current is always a negligible noise source in HiPERCAM. We looked at a number of cooling options¹ before deciding to use thermo-electric (peltier) coolers (TECs), which are the simplest, cheapest, lightest and most compact of coolers. Our solution uses two Marlow NL5010 five-stage TECs mounted side by side, as shown in Fig. 3. Our detector head design uses all-metal seals rather than o-rings in order to minimize leaks. We went to great lengths to avoid using any materials inside the detector heads that could outgas. So, for example, we mounted the pre-amplifier board outside the head,² and used a corrugated indium foil to make good thermal connections between the heatsink, TECs and cold plate. We also reduced outgassing by thoroughly cleaning all components prior to assembly, and then baking the assembled head whilst vacuum pumping. Even with all of these precautions, outgassing limits the vacuum hold time of the HiPERCAM CCD heads to around 1 week, due



Figure 1. The HiPERCAM hull, housing the four dichroic beamsplitters. The faces of the hull act as mounting points for the collimator and five re-imaging cameras – only three of the camera barrels are visible in this photo. The filter holders and CCD heads are mounted on the ends of the camera barrels. For scale, the large collimator lens visible in the photo has a diameter of 208 mm.

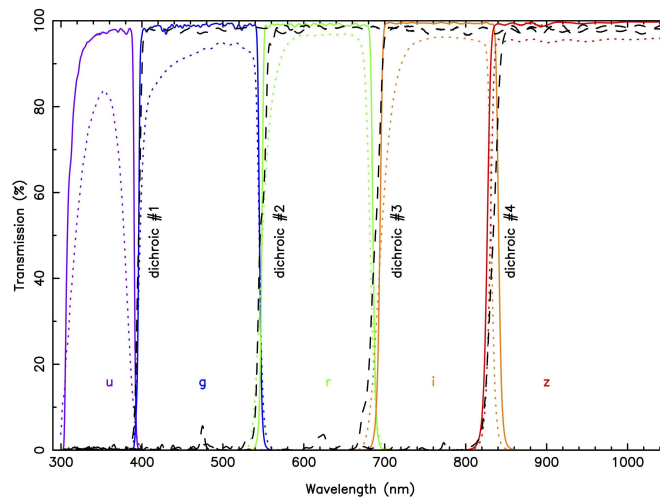


Figure 2. As-built transmission profiles of the HiPERCAM “Super” SDSS filters (solid lines), the HiPERCAM “original” SDSS filters (dotted lines), and the four HiPERCAM dichroic beamsplitters (dashed lines).

primarily to the small interior volume of the heads (~ 0.5 litres) and the lack of a sufficiently cold, large-area interior surface to give effective cryopumping. Fortunately, it only takes a few minutes to pump down the heads whilst on the telescope, thanks to the use of a 5-way vacuum manifold system permanently installed on the instrument.

The heat generated by the TECs is extracted using a 278 K water-glycol cooling circuit. To ensure that cooling fluid of the same temperature enters each of the 5 CCD heads, the heads are connected in parallel rather than series, via two 6-way manifolds (the sixth arm is for cooling the NGC controller). Each arm in this parallel

Table 2. Specifications of the as-built HiPERCAM CCDs.

CCD model	Teledyne e2v CCD231-42
Cosmetic grade	1
Architecture	Split frame transfer, back thinned, 2-phase, NIMO
Format	2048×2048 pixels
Image area	2048×1024 pixels
Storage area	2× 2048×512 pixels
Pixel size	15 μm
Outputs	4
Readout noise	3.2 e ⁻ at 200 kHz, single-ended
Gain	1.2 e ⁻ /ADU
Dark current	10 e ⁻ /pix/hr at 173 K
Deep depletion Si	i_s and z_s bands
AR coatings	Astro Broadband in u_s , Astro Multi-2 in $g_s r_s i_s z_s$
Peak QE in $u_s g_s r_s i_s z_s$ bands	76/89/88/88/78 %
Fringe suppression	i_s and z_s bands
Fringe amplitude	~ 0.1% in i_s and ~ 1% in z_s band
Full well	120 ke ⁻
Non linearity	< ±0.5%
Vertical clocking	10 μs/row
Horizontal clocking	0.1 μs/pix, split serial-register
Full frame time (bin 1×1)	3.0 s in slow, 1.25 s in fast readout mode
Full frame time (bin 2×2)	0.9 s in slow, 0.4 s in fast readout mode
Drift-mode frame time	0.0009 s with 24×24 pixel windows, bin 6×6

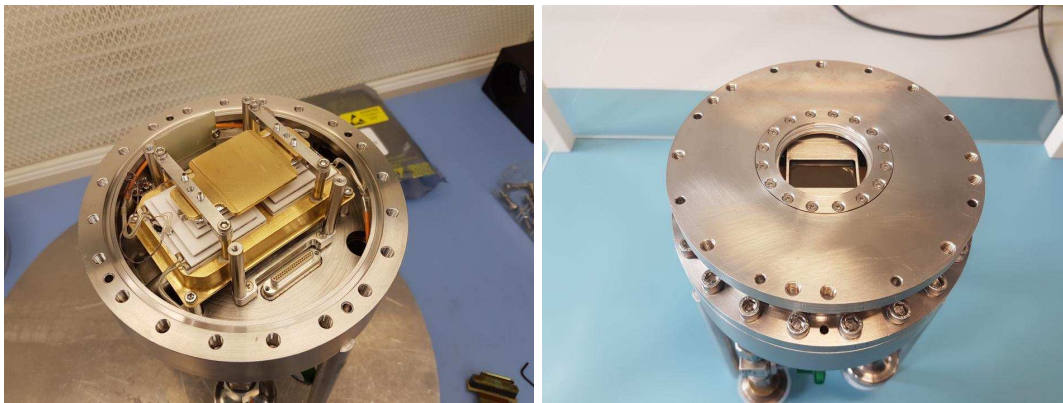


Figure 3. Left: Interior view of one of the HiPERCAM CCD heads, showing the gold-plated cold plate sitting on top of two, white 5-stage TECs, which themselves are sitting on a gold-plated heatsink through which the cooling liquid runs. Right: Exterior view of one of the HiPERCAM CCD heads. The diameter of the head is 160 mm and the weight is approximately 7 kg.

circuit is equipped with a flow sensor connected to a Honeywell Minitrend GR Data Recorder mounted in the electronics cabinet. As well as providing a display of the flow rate through each CCD head, the data recorder has relays that can switch off the power to the TEC power supplies if the flow rate in any head drops below a user-defined limit, thereby protecting the CCDs from overheating. As a backup to this system, the TEC power supplies themselves (made by Meerstetter, model LTR-1200) have a high-temperature cut-off facility: if the temperature of the heat-sink in the CCD head rises above a user-defined value, such as would occur if the coolant supply fails, the power to the TEC is automatically shut off. The TEC power supplies are able to maintain the CCD temperatures at their -90°C set points to within $0.01 - 0.1^{\circ}\text{C}$.

In case of high humidity, HiPERCAM has a 5-way manifold that enables clean, dry air from a telescope supply to be blown across each of the CCD windows at approximately 1 litre/min to prevent condensation.

2.3 Mechanics

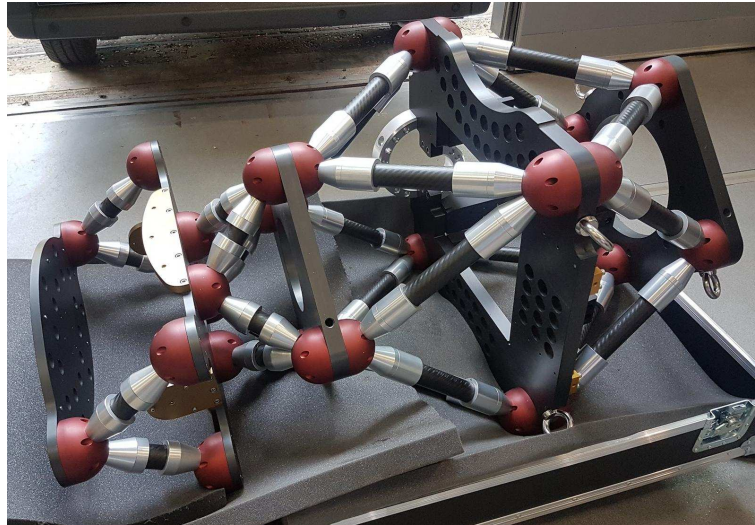


Figure 4. The empty HiPERCAM opto-mechanical chassis sitting in its packing crate. The plate at the right-hand side mounts onto the telescope. The large central plate with a square aperture in it houses the hull. The CCD controller is mounted between the two left-most plates.

The HiPERCAM opto-mechanical chassis is a triple octopod composed of 3 large aluminium plates connected by carbon fibre struts, as shown in Fig. 4. This design ensures a stiff, compact (1.25 m long), light-weight (220 kg) and open structure, which is relatively insensitive to temperature variations. These characteristics make HiPERCAM easy to maintain, transport and mount/dismount at the telescope.

The top plate of the triple octopod is used to mount the instrument onto the telescope, the middle plate is used to mount the hull on (Fig. 1), and the bottom plate is used to mount the CCD controller onto the instrument. Two different interface collars are used to attach HiPERCAM onto the rotator of the WHT and GTC, and place the instrument at the correct distance from the respective telescope focal planes (see Fig. 5).

2.4 Data acquisition system

The HiPERCAM data acquisition system (DAS) has been described in detail elsewhere.^{1,2} The DAS in HiPERCAM is detector limited, i.e. the throughput of data from the output of the CCDs to the hard disk on which it is archived is always greater than the rate at which the data comes off the CCDs. This means that the instrument is capable of running continuously all night at its maximum data rate without ever having to pause for archiving of data. All CCDs are read out simultaneously and have identical exposure start and end times. In order to change the exposure times of the CCDs with respect to each other, it is possible to skip the readout of selected CCDs using the *NSKIP* parameter. For example, if *NSKIP* is set to 3,2,1,2,3 for the *u, g, r, i, z* CCDs, and the exposure time is set to 10s, then the CCD controller will read out only the *r*-band CCD on the first readout cycle (giving *r* a 10s exposure), then the *g, r* and *i*-band CCDs on the second cycle (giving *g* and *i* a 20s exposure), then the *u, r* and *z*-band CCDs on the third cycle (giving *u* and *z* a 30s exposure), etc.

The HiPERCAM graphical user interface (GUI) is shown in Fig. 6. The top right of the GUI shows three buttons that allows the astronomer to change the readout mode: full frame, windowed or drift mode.³ Below this are the various CCD readout parameters, giving the astronomer complete control over the detector setup. The number of exposures is usually set to zero in the GUI and the green *Start* button is then pressed: the data acquisition system will then take data continuously until the red *Stop* button is pressed. All frames of a HiPERCAM run on a target are written to a single, custom-format FITS file. The HiPERCAM data-reduction

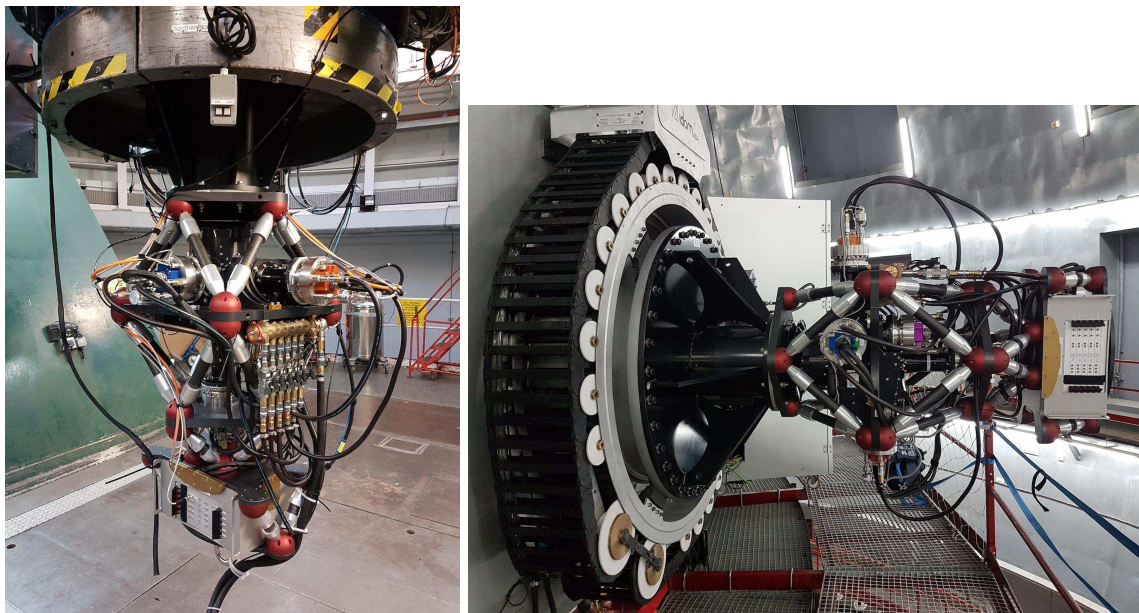


Figure 5. Left: HiPERCAM mounted at the Cassegrain focus of the WHT. Right: HiPERCAM mounted on the Folded Cassegrain focus of the GTC.

pipeline software (see Fig. 7) reads this file and provides a quick-look reduction of the data whilst observing. The pipeline is a fully-featured photometry reduction package and so can also be used off-line to produce publication-ready light curves.

2.5 On-sky performance

We commissioned HiPERCAM on the sky for the first time on the WHT during October 2018. This was a short run of only 1 commissioning night and 4 science nights, and was intended primarily to be a shake-down of the instrument prior to commissioning on the GTC. The instrument was then moved to the GTC, where we were allocated 3 commissioning nights and 10 science nights in February 2018. We suffered from terrible weather during this run, and only observed for the last 3 nights, when we completed all of the commissioning. Three more observing runs followed in April, May and June 2018, totalling 16 nights. A wide range of science was performed during this time, including the observation of black holes, white dwarfs, neutron stars, brown dwarfs, extrasolar planets/asteroids, AGN, FRBs, GRBs, SNe and ultra-diffuse galaxies. We report on the performance of HiPERCAM on the GTC below.

An example star field observed with HiPERCAM on the GTC is shown in Fig. 8. The FWHM of the stars in these images are $u_s = 0.56''$, $g_s = 0.44''$, $r_s = 0.41''$, $i_s = 0.37''$, $z_s = 0.36''$, with no discernible variation with field angle. This indicates that HiPERCAM on the GTC can provide seeing-limited images across the whole field of view in even the very best seeing conditions on La Palma. Since the dichroics operate in a collimated beam and have anti-reflection coatings on their rear surfaces, we do not expect to see any ghosting in our images, and this is indeed the case: a careful inspection of the brightest stars in Fig. 8 reveals no discernible ghosting. The pixel positions of the stars at the corners of the field of view are the same on all 5 CCDs, to within approximately 5 pixels ($75 \mu\text{m}$), indicating that there is no discernible variation of platescale with wavelength and that the CCD heads are well aligned with respect to each other. Twilight-sky flat fields show no discernible vignetting in the corners of the field of view.

The measured photometric zero points of HiPERCAM on the GTC, defined here as the magnitude of a star that would give 1 electron per second above the atmosphere, are: $u_s = 26.92$ (25.76), $g_s = 28.77$ (28.26), $r_s = 28.54$ (28.84), $i_s = 28.30$ (28.49), $z_s = 27.85$ (27.95). The numbers in brackets show the corresponding values for OSIRIS, the common-user, single-channel optical imager at the GTC.⁵ These two sets of zero points

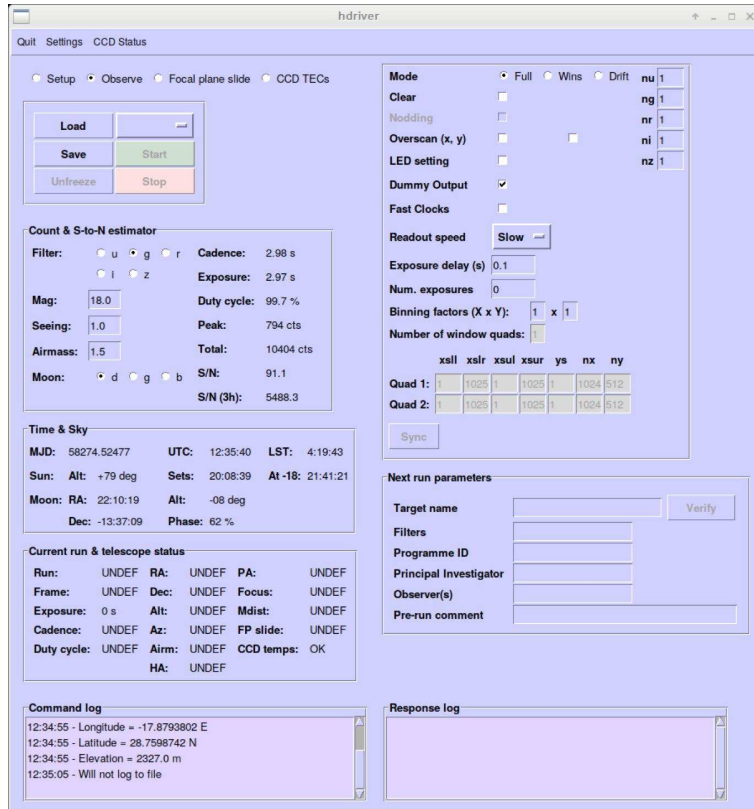


Figure 6. A screenshot of the python-based HiPERCAM GUI, used by astronomers at the telescope to control the instrument. There also exists an engineering GUI for low-level control and telemetry of the CCD controller.

were measured within a few days of each other during May 2018, using observations of SDSS standard stars. It can be seen that unless one needs the full 7.8' field of view of OSIRIS, or just one of the redder bands, it is much more efficient to use HiPERCAM for standard imaging at the GTC, and one would also gain from the huge reduction in dead time between exposures (0.01 s with HiPERCAM, 21 s with OSIRIS).

The 5σ limiting magnitudes of HiPERCAM on the GTC are shown as a function of exposure time in Fig. 9, calculated using the zero points given above. It can be seen that it is possible to achieve a limiting magnitude of $g \sim 16$ in a single 0.001 s exposure, $g \sim 23$ in 1 s, and $g \sim 27$ in 1800 s.

The throughput curves of HiPERCAM, which include all optics and CCDs but not the atmosphere and telescope, are shown in Fig. 10. The throughput peaks at over 60% in g_s , r_s , i_s and 50% in u_s , z_s . This high throughput has been achieved by using high-performance multi-layer coatings on the lenses, dichroics, filters and windows, as well as CCDs optimised for operation in each band.

We measured the accuracy of the GPS absolute timestamping of each CCD frame in HiPERCAM by observing an LED connected to the pulse-per-second (PPS) output of our GPS system. An example observation is shown in Fig. 11, where we plot the light curve of the LED phase-folded on its 1 s period. The shape of the light curve is a convolution of two top-hat functions, one for the LED pulse and the other for the exposure time duration (in this example, 0.003 s). Hence the folded light curve exhibits a ramp, the centre of which should correspond to the start of the GPS second. Fig. 11 shows that this is indeed the case: the offset between the start of the GPS second and centre of the ramp is only $35 \mu\text{s}$ and is set by the measurement accuracy of our experiment. Hence we can say that the absolute timestamping of each CCD frame in HiPERCAM is accurate to better than tens of microseconds.

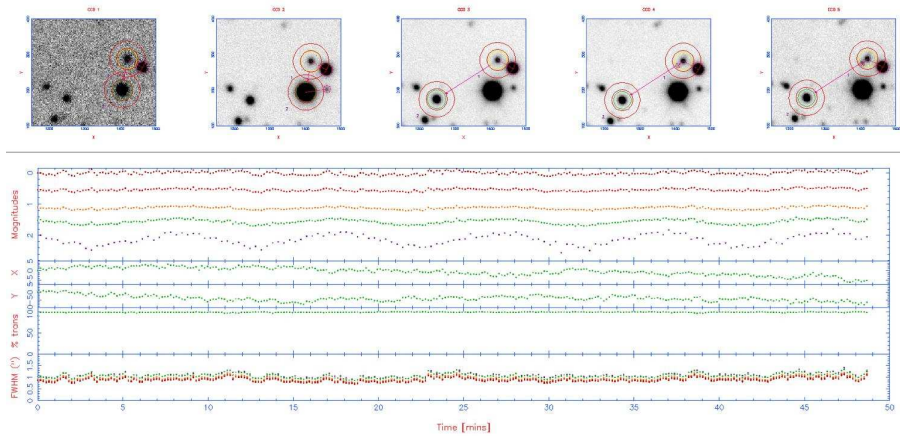


Figure 7. A screenshot of the python-based HiPERCAM data reduction pipeline. The top row shows a zoom-in of the $u_s g_s r_s i_s z_s$ images of the target and comparison stars, with the apertures defining the object and sky regions superimposed. The bottom panel shows the target minus comparison star magnitudes in $u_s g_s r_s i_s z_s$ (top row), the comparison star x, y positions (second and third rows), the sky transparency measured from the comparison star flux (fourth row), and the seeing in $u_s g_s r_s i_s z_s$ measured from the comparison-star FWHM (bottom row).

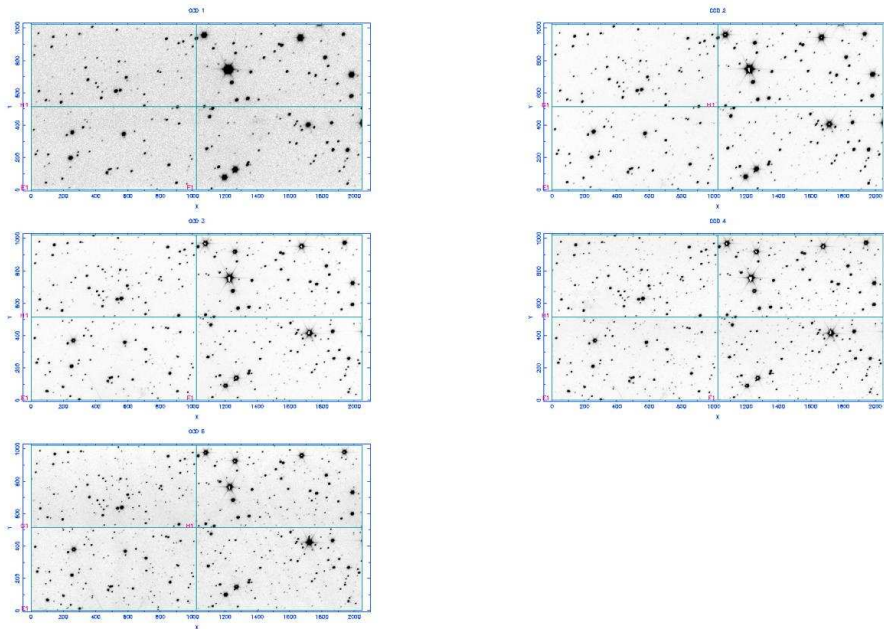


Figure 8. Images of a star field obtained with HiPERCAM on the GTC in u_s and g_s (top row), r_s and i_s (middle row), and z_s (bottom row).

2.6 Future plans

HiPERCAM was removed from the GTC rotator in June 2018 so that an autoguider can be installed. The instrument will be remounted on the telescope from September 2018 to the end of the year for more science runs. In 2019, HiPERCAM will share the Folded Cassegrain focus of the GTC with CanariCam.⁶ The situation from 2020 onwards is less clear at the moment, due to the arrival of MIRADAS at the GTC.⁷

We are planning on making two enhancements to HiPERCAM during the coming year. The first is to modify the CCD preamp boards so that it is possible to switch the bandwidth from the current value of 1.06 MHz^2 to

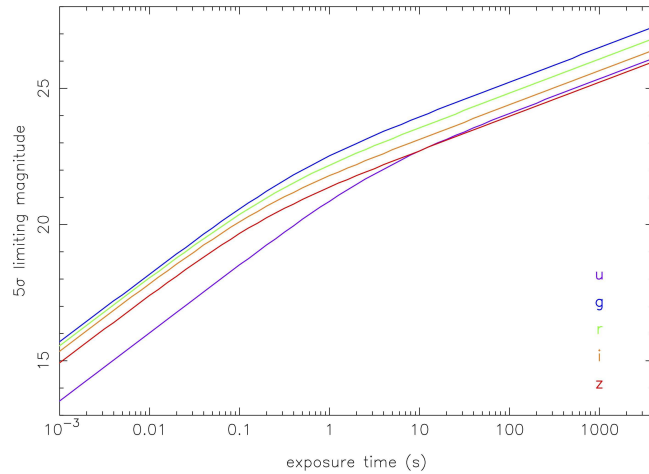


Figure 9. Limiting magnitudes (5σ) of HiPERCAM on the GTC as a function of exposure time. The purple, blue, green, orange and red curves show the results for the $u_s g_s r_s i_s z_s$ filters, respectively. The calculations assume dark moon, observing at the zenith and seeing of $0.8''$.

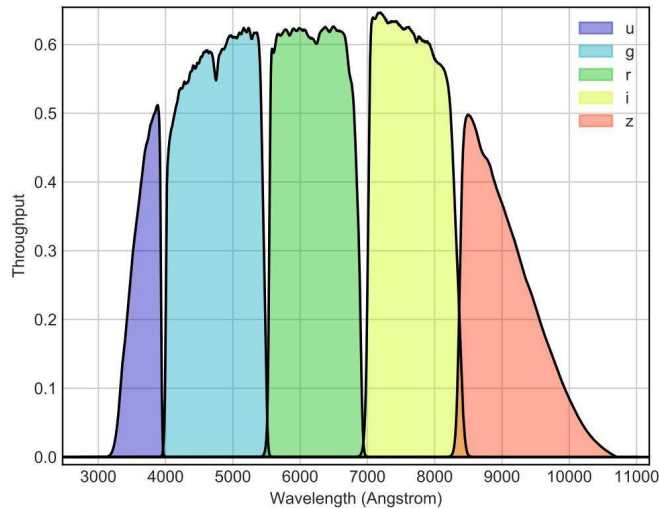


Figure 10. Throughput curves for HiPERCAM in $u_s g_s r_s i_s z_s$. These curves do not include the atmosphere or telescope.

approximately half this value. In this way, we hope to reduce the readout noise from its current value of $\sim 4 e^-$ to $\sim 3 e^-$ in slow readout mode (263 kHz).

The second enhancement is to implement a COMParison star Pick-Off system (COMPO), as shown in Fig. 12. Light from a bright comparison star that falls outside the $3.1'$ diagonal field of view of HiPERCAM, but within the $10'$ diameter field of view of the Folded Cassegrain focus of the GTC, is collected by a pick-off arm. The light is then redirected to a second arm, via a set of relay optics, which injects the starlight onto one of the corners of the HiPERCAM CCDs. In this way it is possible to use much brighter comparison stars than would usually be available for differential photometry. This will be of particular importance when observing bright targets,

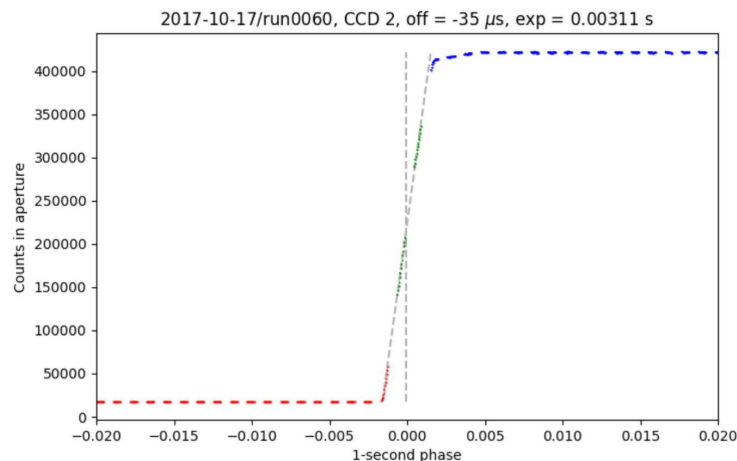


Figure 11. Phase-folded light curve of an LED attached to the PPS output of the HiPERCAM GPS system.

like exoplanet host stars, for which nearby comparison stars of comparable or greater brightness are rare. More quantitatively, without COMPO, there is a 90% probability of finding a comparison star of magnitude $r = 14$ in the field of view, whereas with COMPO one will be able to find comparison stars of magnitude $r = 12$ with the same probability. COMPO will also be of great benefit to most u_s -band observations, as many of the targets observed by HiPERCAM tend to be blue (e.g. white dwarfs) and hence bright in the u_s -band, whereas most comparison stars tend to be red and hence faint in the u_s -band.

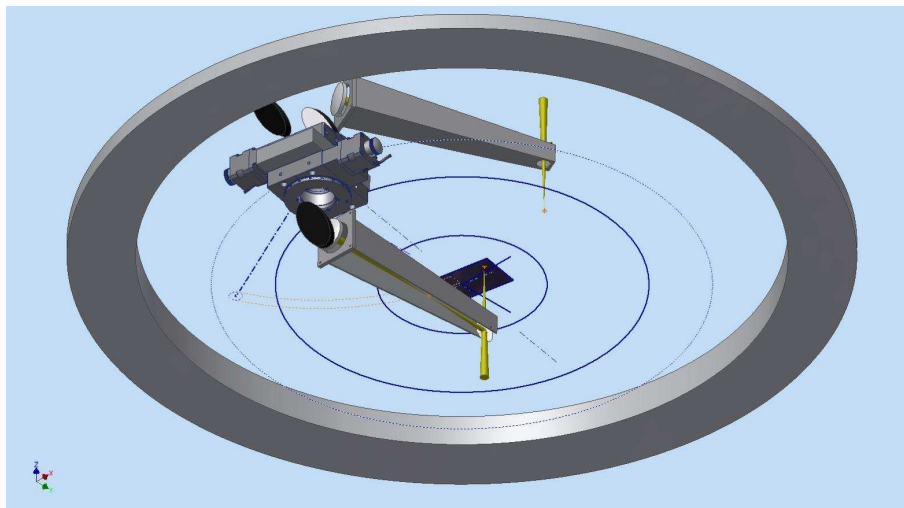


Figure 12. CAD image of COMPO. The upper arm collects light from a star falling outside the HiPERCAM field of view (solid blue rectangle at centre) but inside the 10' diameter view of view at the Folded Cassegrain focus of the GTC (outer blue circle). The lower arm redirects this light to one of the corners of the HiPERCAM field of view, via a set of relay optics.

3. CONCLUSIONS

We have described the as-built design and first-light performance of HiPERCAM on the GTC, as well as some of the enhancements we are planning to implement over the coming year. The on-sky results demonstrate that HiPERCAM is performing to specification. The instrument is now entering its exploitation phase on the GTC, and promises to revolutionise the field of high time-resolution optical astrophysics.

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