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Some general properties of the renormalized stress-energy tensor for static quantum states on $(n + 1)$ -dimensional spherically symmetric black holes

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We study the renormalized stress-energy tensor (RSET) for static quantum states on $(n + 1)$ -dimensional, static, spherically symmetric black holes. By solving the conservation equations, we are able to write the stress-energy tensor in terms of a single unknown function of the radial co-ordinate, plus two arbitrary constants. Conditions for the stress-energy tensor to be regular at event horizons (including the extremal and “ultra-extremal” cases) are then derived using generalized Kruskal-like co-ordinates. These results should be useful for future calculations of the RSET for static quantum states on spherically symmetric black hole geometries in any number of space-time dimensions.

I. INTRODUCTION

The renormalized stress-energy tensor (RSET) $\langle T_{\mu\nu} \rangle_{\text{ren}}$ is an object of fundamental importance in quantum field theory in curved space-time, since it governs, via the semi-classical Einstein equations

$$G_{\mu\nu} = 8\pi G \langle T_{\mu\nu} \rangle_{\text{ren}}, \quad (1)$$

the back-reaction of the quantum field on the space-time geometry. However, the renormalization process means that detailed calculations of the RSET are notoriously difficult, even in spherically symmetric space-times (see, for example, [1] for some four-dimensional calculations), although calculations in three space-time dimensions are more tractable [2]. Various analytic approximations [3] have been developed for the RSET in various cases, but it is still useful to obtain as much information as possible about the RSET from basic physical principles without a full calculation.

Christensen and Fulling [4] pioneered this approach. From equation (1), since the Einstein tensor $G_{\mu\nu}$ is divergence-free, the same must be true of $\langle T_{\mu\nu} \rangle_{\text{ren}}$. Christensen and Fulling [4] therefore studied solutions of the conservation equations

$$\nabla^\mu \langle T_{\mu\nu} \rangle_{\text{ren}} = 0 \quad (2)$$

on the Schwarzschild black hole geometry, in both two and four space-time dimensions. In four dimensions, the conservation equations, together with elementary symmetry principles, can be solved to give the renormalized stress-energy tensor in terms of one unknown function and two unknown constants. The constants are constrained by the particular choice of vacuum state (Hartle-Hawking [5], Unruh [6] or Boulware [7]), through the regularity properties of the RSET on the event horizon. In two dimensions, the stress-energy tensor is given completely in terms of the trace anomaly [28]. The analysis

of [4] has proved to be powerful for reducing the calculation of the RSET to a single component (typically taken to be T_θ^θ), but also because it revealed certain properties of the RSET (such as its singularity structure on the horizons) without requiring a full computation.

There is now great interest in a wide variety of static black hole space-times, and it is our purpose in this paper to extend the analysis of [4] to more general, $(n + 1)$ -dimensional, static black holes. Our analysis is independent of the asymptotic structure of the geometry at infinity, and therefore includes “topological” black holes which can exist in asymptotically anti-de Sitter space [9] as well as general, spherically symmetric black holes. However, in this paper, for ease of phraseology we will tend to use the phrase “spherically symmetric” even when our results include these topological black holes. A key question in any study of the RSET is whether or not it is regular across a horizon. We examine this for a general horizon (whether non-extremal or extremal), using generalized Kruskal-like co-ordinates [10] which can describe both extremal and non-extremal horizons.

Our analysis in this paper is restricted to static quantum states. This is sufficiently general to cover the Hartle-Hawking [5] and Boulware [7] vacua for all static black hole geometries. These two vacua (particularly the Hartle-Hawking state) are the ones most frequently studied in full calculations of the RSET [1], as they are time-reversal symmetric and therefore possess the greatest symmetries, which makes calculations easier. Our results will certainly be of use for future calculations of the RSET in these two states. However, the Unruh [6] vacuum may not be covered by our analysis in general. For static, spherically symmetric, asymptotically flat black holes with a non-extremal event horizon, it is likely that the Unruh vacuum will be a static state and therefore covered by our approach. However, it is known that for black holes with extremal horizons [11], or with both event and cosmological horizons [12, 13], the equivalent of the Unruh vacuum is not a static state. For non-static states, the conservation equations are considerably more complex and correspondingly less information is accessible from their solution, apart from in two space-time

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dimensions (a similar situation arises if one studies the solutions of the conservation equations on a Kerr black hole [14]). We therefore do not consider this situation further.

The outline of this paper is as follows. In section II we describe our static, black hole metric, and our generalized Kruskal-like co-ordinates (following [10]) which are regular across any horizon. The conservation equations (2) are solved on this background metric in section III. This gives the RSET in terms of a single unknown function of the radial co-ordinate r and two arbitrary constants. We then turn, in section IV, to the regularity of the RSET across an arbitrary horizon, and derive conditions for the RSET to be regular. These conditions are particularly stringent when the horizon is extremal. In section V we focus on the special case of a space-time with distinct event and cosmological horizons, and derive a strong integral constraint which must be satisfied if the RSET is to be regular across both the event and cosmological horizons. Our focus in this paper is the study of $(n+1)$ -dimensional black holes, however in section VI, we restrict attention to the case of two space-time dimensions, where the unknown function of r disappears. For distinct event and cosmological horizons, we show that the RSET can be regular across all horizons only if they have the same temperature, and we are also able to show that no static state can have an RSET which is regular across an extremal horizon. Finally, our conclusions are presented in section VII. Henceforth, the metric has signature $(-, +, +, +)$ and we use units in which $c = G = \hbar = 1$.

II. GENERAL $(n+1)$ -DIMENSIONAL SPHERICALLY SYMMETRIC BLACK HOLES

A. Metric ansatz

We consider an $(n+1)$ -dimensional, static, black hole metric in the following form, after a suitable choice of gauge:

$$ds^2 = -f(r)dt^2 + f^{-1}(r)dr^2 + R^2(r)d\Omega_{n-1}^k, \quad (3)$$

where the metric functions f and R depend on the radial co-ordinate r only. The $(n-1)$ -dimensional metric $d\Omega_{n-1}^k$ (where we have used the notation of [15]) is as follows, for $k = 1$, $k = 0$ and $k = -1$ respectively:

$$\begin{aligned} d\Omega_{n-1}^1 &= dx_2^2 + \sin^2 x_2 dx_3^2 + \dots + \prod_{i=2}^{n-1} \sin^2 x_i dx_n^2; \\ d\Omega_{n-1}^0 &= dx_2^2 + dx_3^2 + dx_4^2 + \dots + dx_n^2; \\ d\Omega_{n-1}^{-1} &= dx_2^2 + \sinh^2 x_2 dx_3^2 + \dots \\ &+ \prod_{i=2}^{n-1} \sinh^2 x_i dx_n^2. \end{aligned} \quad (4)$$

We can only have $k \neq 1$ for asymptotically adS (anti-de Sitter) black holes. The particular value of k describes

the topology of the event horizon, which is spherical for $k = 1$; planar, cylindrical or toroidal (with genus ≥ 1) for $k = 0$; and hyperbolic or toroidal (with genus ≥ 1) for $k = -1$. Although not strictly speaking correct, in the remainder of the paper we shall use the phrase ‘‘spherically symmetric’’ to include these topological black holes.

For most of the exact solutions in which one might be interested (for example, the higher-dimensional Reissner-Nordström(-anti)-de Sitter black holes), it will be the case that $R(r) \equiv r$, which simplifies the analysis. However, our metric (3) is sufficiently general to cover many black hole solutions in dilaton gravity, hairy black holes, and even the Nariai [16] metric (in which case $R \equiv \Lambda > 0$). The general $(n+1)$ -dimensional, spherically symmetric metric is also commonly written in the form

$$ds^2 = -N(R)S^2(R)dt^2 + N(R)^{-1}dR^2 + R^2d\Omega_{n-1}^k. \quad (5)$$

This metric can be transformed into the form (3) by the change of co-ordinates

$$\frac{dr}{dR} = S(R). \quad (6)$$

We will be particularly interested in the regularity of the RSET at a horizon of the black hole geometry. Assuming that $R(r)$ has no zeros, the horizon structure of the black hole is determined by the metric function $f(r)$. We will assume that $f(r)$ has at least one zero, namely a regular black hole event horizon at $r = r_+$, with $f'(r_+) > 0$. In the presence of a positive cosmological constant, there will also be a cosmological event horizon at $r = r_{++}$, where $f'(r_{++}) < 0$. Our analysis also allows the possibility of an inner, Cauchy, horizon at $r = r_-$, with $f'(r_-) < 0$. It is possible for two or more of these horizons to coincide. Therefore, there are several different types of horizon which we need to consider:

1. Regular, non-extremal horizons, of event, cosmological or inner variety (at r_+ , r_{++} or r_- respectively);
2. An extremal horizon, formed by the coincidence of an event and inner horizon (the so-called ‘‘cold’’ black hole [17]);
3. An extremal horizon, formed by the coincidence of an event and cosmological horizon (a ‘‘Nariai’’ black hole [15]);
4. An ‘‘ultra-extremal’’ black hole horizon, formed by the coincidence of all three types of horizon (the ‘‘ultra-cold’’ black hole [17]).

All these possibilities can be illustrated by solutions of Einstein-Maxwell theory [15]. In these specific examples, $R(r) = r$, and the metric function $f(r)$ is given by:

$$f(r) = k - \frac{M}{r^{n-3}} + \frac{q^2}{r^{2(n-3)}} - \frac{\Lambda r^2}{3}, \quad (7)$$

where M , q are related, respectively, to the mass and charge of the black hole [18]. Penrose diagrams for all these possible cases can be found in [15].

B. Kruskal-like co-ordinates

In order to analyze the regularity of the RSET at the horizons of the spacetime, we require Kruskal-like co-ordinates which are regular across each horizon.

We begin by defining the usual ‘‘tortoise’’ co-ordinate r^* by the equation:

$$\frac{dr^*}{dr} = \frac{1}{f(r)}, \quad (8)$$

so that an event horizon $r = r_+$ corresponds to $r_* \rightarrow -\infty$. If there is a cosmological horizon $r = r_{++}$, then it will be the case that $r_* \rightarrow \infty$ there, as is also true as $r \rightarrow \infty$ for asymptotically flat spacetimes. However, if the geometry is asymptotically adS, then r_* tends to a finite constant (which we may as well take to be zero) at infinity.

Starting with the usual null co-ordinates

$$v = t + r^*, \quad w = t - r^*, \quad (9)$$

for all types of non-extremal horizon, we define the standard Kruskal co-ordinates, V , W , in a region in which $f(r) > 0$, by

$$V = e^{\kappa v}, \quad W = -e^{-\kappa w}, \quad (10)$$

where the surface gravity κ is given by

$$\kappa = \frac{1}{2}f'(r_0), \quad (11)$$

with r_0 being the location of the particular horizon under consideration. However, at an extremal horizon, the surface gravity κ (11) vanishes and so the standard Kruskal co-ordinates (10) are constant. In this situation one can use Eddington-Finkelstein co-ordinates in patches across the future and past horizons separately, but we shall instead follow the method of [10] to define new Kruskal-like co-ordinates in this case. In [10], new co-ordinates were defined in the case of coincident event and inner horizons, and we here extend their method to the general situation. In addition, these new co-ordinates can be defined equally well for non-extremal horizons, which will allow us, in section IV, to deal simultaneously with the analysis of the behaviour of the RSET near all types of horizon.

At an extremal horizon, the metric function $f(r)$ will have either a double or a triple zero at $r = r_0$ (the triple zero occurring in the ‘‘ultra-extremal’’ case), and we therefore write $f(r)$ as

$$\begin{aligned} f(r) &= (r - r_0)g_1(r); \\ f(r) &= (r - r_0)^2g_2(r); \\ f(r) &= (r - r_0)^3g_3(r); \end{aligned} \quad (12)$$

for a non-extremal, extremal and ultra-extremal horizon respectively, where $g_1(r)$, $g_2(r)$, $g_3(r)$ are non-zero at $r =$

r_0 . Integrating (8) gives, near the horizon,

$$\begin{aligned} r_* &= a_1 \log(r - r_0) + O(1); \\ r_* &= b_1(r - r_0)^{-1} + b_2 \log(r - r_0) + O(1); \\ r_* &= c_1(r - r_0)^{-2} + c_2(r - r_0)^{-1} \\ &\quad + c_3 \log(r - r_0) + O(1), \end{aligned} \quad (13)$$

for non-extremal, extremal and ultra-extremal horizons, respectively, where the a s, b s and c s are constants given in terms of the g s and their derivatives at $r = r_0$. The ones we need later are:

$$\begin{aligned} a_1 &= g_1(r_0)^{-1}, \\ b_1 &= -g_2(r_0)^{-1}, \\ c_1 &= -\frac{1}{2}g_3(r_0)^{-1}. \end{aligned} \quad (14)$$

Near the horizon, we have $r_* \rightarrow \pm\infty$, with the sign depending on the sign of a_1 , b_1 and c_1 .

Following [10], we now define a function $\psi(\xi)$ as one half of that part of r_* (13) which is singular as $r \rightarrow r_0$, that is, for non-extremal, extremal and ultra-extremal black holes respectively:

$$\begin{aligned} \psi(\xi) &= \frac{1}{2}a_1 \log \xi; \\ \psi(\xi) &= \frac{1}{2}(b_1\xi^{-1} + b_2 \log \xi); \\ \psi(\xi) &= \frac{1}{2}(c_1\xi^{-2} + c_2\xi^{-1} + c_3 \log \xi). \end{aligned} \quad (15)$$

We then define new Kruskal-like co-ordinates \mathcal{V} , \mathcal{W} implicitly by [10]:

$$v = \psi(\mathcal{V}), \quad w = -\psi(\mathcal{W}), \quad (16)$$

where v and w are the null co-ordinates given in equation (9). Note that the definition (16) reduces to the standard Kruskal co-ordinates (10) for a non-extremal horizon. In terms of these new co-ordinates, the metric (3) takes the form

$$ds^2 = -f(r)\psi'(\mathcal{V})\psi'(-\mathcal{W})d\mathcal{V}d\mathcal{W} + R(r)^2d\Omega_{n-1}^k. \quad (17)$$

In order to show that \mathcal{V} and \mathcal{W} are in fact good co-ordinates across the horizons, we need to consider future and past horizons separately. The same argument works for both, so suppose we are considering a horizon \mathcal{H} on which \mathcal{V} is finite and non-zero, so that $\psi'(\mathcal{V})$ is also finite and non-zero there. For many black holes, this will correspond to a future event horizon, where $t \rightarrow \infty$ and $r_* \rightarrow -\infty$, although this will depend on the signs of the constants in (14). The precise causal structure of the event horizon does not affect our construction. Then $t + r_*$ is finite and non-zero on \mathcal{H} , giving, from the definition of ψ ,

$$t - r_* = -2r_* + O(1) = -\psi(r - r_0) + O(1). \quad (18)$$

Therefore we have

$$\begin{aligned}\mathcal{W} &= -\psi^{-1}(-t+r_*) = -\psi^{-1}[\psi(r-r_0) + O(1)] \\ &= -(r-r_0) + O(1).\end{aligned}\quad (19)$$

From the definition of ψ in the non-extremal, extremal and ultra-extremal cases, respectively,

$$\begin{aligned}\psi'(-\mathcal{W}) &= \frac{1}{2}g_1(r_0)(r-r_0)^{-1} + O(1); \\ \psi'(-\mathcal{W}) &= \frac{1}{2}g_2(r_0)^{-1}(r-r_0)^{-2} + O(r-r_0)^{-1}; \\ \psi'(-\mathcal{W}) &= \frac{1}{2}g_3(r_0)^{-1}(r-r_0)^{-3} + O(r-r_0)^{-2}.\end{aligned}\quad (20)$$

In all three cases, then, $f(r)\psi'(-\mathcal{W})$ is finite and non-zero as $r \rightarrow r_0$, so that \mathcal{V} and \mathcal{W} are suitable co-ordinates, because $\psi'(\mathcal{V})$ is also finite and non-zero. Substituting \mathcal{V} for \mathcal{W} in the above argument shows that these are also suitable regular co-ordinates across a horizon where \mathcal{W} is finite and non-zero.

For spacetimes with many distinct horizons, patches of different Kruskal co-ordinates may be required (see, for example, [19]). In the particular case of a black hole with a regular event horizon at $r = r_+$ and cosmological horizon at $r = r_{++}$, one patch of Kruskal co-ordinates, V_+ , W_+ , can be used to cover the region extending from inside the event horizon up to the cosmological horizon, but they will not be regular across the cosmological horizon. We therefore need a second set of Kruskal co-ordinates, V_{++} , W_{++} , which are regular across the cosmological horizon all the way down to the event horizon, but not across the event horizon.

In [13, 20], a co-ordinate system is given which is regular across *both* the event and cosmological horizon in Schwarzschild-de Sitter space-time, and which therefore removes the need to use two sets of Kruskal co-ordinates. We have not used this here because we wish to work as generally as possible, and, although the method of [13, 20] could be used to find globally regular co-ordinates in a more general case, the expressions involved are likely to be algebraically highly complex and make analysis difficult.

III. SOLUTION OF THE CONSERVATION EQUATIONS

Since we are working on a static, spherically symmetric black hole spacetime, we will assume that the RSET (henceforth denoted simply by $T_{\mu\nu}$) is also static and spherically symmetric. The form of the RSET is then:

$$T_{\nu}^{\mu} = \begin{pmatrix} A & -Pf^{-1} & & & \\ Pf & T - A - (n-1)Q & & & \\ & & Q & & \\ & & & \ddots & \\ & & & & Q \end{pmatrix}; \quad (21)$$

with all other entries vanishing, where A, P, Q and T are functions of r only. The RSET will be symmetric under time-reversal symmetry if and only if $P \equiv 0$.

The usual trace anomaly is given by $T_{\alpha}^{\alpha} = T$. For general $(n+1)$ -dimensional spacetimes, the trace anomaly is zero if n is even, and, if n is odd, is given in terms of the appropriate DeWitt-Schwinger coefficient. For example, for a massless, conformally coupled scalar field, the result is [21]:

$$T = \frac{1}{(4\pi)^{(n+1)/2}} \text{Tr } a_{(n+1)/2}. \quad (22)$$

Various DeWitt-Schwinger coefficients have been calculated, giving the trace anomaly in various spacetime dimensions [22]. The trace anomaly is always a geometric scalar, and so is finite everywhere apart from at a curvature singularity. It is independent of the state of the quantum field under consideration, but does depend on the spin of the quantum field.

The conservation equations (2) arising from the x_i co-ordinates are trivial, and the t and r equations give, respectively,

$$\begin{aligned}0 &= \frac{d}{dr} (fPR^{(n-1)}); \\ 0 &= \frac{1}{R^{n-1}} \frac{d}{dr} (R^{n-1}fA) + \frac{(n-1)}{2QR^{2n}} \frac{d}{dr} (Q^2fR^{2n}) \\ &\quad - \frac{1}{2TR^{2(n-1)}} \frac{d}{dr} (T^2fR^{2(n-1)}); \end{aligned}\quad (23)$$

which can be readily integrated to give

$$\begin{aligned}P &= \frac{X}{fR^{(n-1)}}; \\ A &= -(n-1)Q + T + \frac{Z}{fR^{(n-1)}} + J(r_a, r); \end{aligned}\quad (24)$$

where X, Z are integration constants and we define

$$\begin{aligned}J(x, y) &= \frac{1}{2fR^{(n-1)}} \int_x^y [(n-1)Q - T] f'R^{(n-1)} dr \\ &\quad - \frac{(n-1)}{fR^{(n-1)}} \int_x^y QfR^{(n-2)} R' dr, \end{aligned}\quad (25)$$

with r_a being any fixed value of r (to be chosen shortly). The formulae (24) reduce to those in [4] when $n = 1$ or $n = 3$, and the metric (3) is Schwarzschild. Here we find, like [4], that the complete stress-energy tensor is given in terms of two unknown constants X, Z and one unknown function of r , which we can take to be Q . If there are only two spacetime dimensions, then the RSET is determined solely by the trace anomaly and the constants X and Z [4].

IV. BEHAVIOUR OF THE RSET NEAR A HORIZON

We now address the key question of the behaviour of the RSET close to horizons. For all types of horizon at

$r = r_0$, we employ the modified Kruskal-like co-ordinates \mathcal{V} , \mathcal{W} constructed in section II B. The relevant RSET components in these co-ordinates are:

$$\begin{aligned} T_{\mathcal{V}\mathcal{V}} &= \frac{1}{2}f[\psi'(\mathcal{V})]^2 \left\{ \mathcal{F} - \frac{1}{fR^{(n-1)}}[Z - X] \right\}; \\ T_{\mathcal{W}\mathcal{W}} &= \frac{1}{2}f[\psi'(-\mathcal{W})]^2 \left\{ \mathcal{F} - \frac{1}{fR^{(n-1)}}[Z + X] \right\}; \\ T_{\mathcal{V}\mathcal{W}} &= \frac{1}{4}f\psi'(\mathcal{V})\psi'(-\mathcal{W}) \{(n-1)Q - T\}; \end{aligned} \quad (26)$$

where

$$\mathcal{F} = \frac{1}{2}(n-1)Q - \frac{1}{2}T - J(r_0, r) \quad (27)$$

and we have chosen the lower limit in the definition of $J(r_a, r)$ (25) to be the location of the horizon, namely $r_a = r_0$.

Consider firstly a horizon on which \mathcal{V} is finite and non-zero, and \mathcal{W} vanishes from (19). From section II B, on this horizon $\psi'(\mathcal{V})$ is also finite and non-zero, and, from (20), we have

$$\psi'(-\mathcal{W})f = O(1) \quad (28)$$

as $r \rightarrow r_0$. From (26), it is immediately clear that $T_{\mathcal{V}\mathcal{W}} = T_{\mathcal{W}\mathcal{V}}$ is regular on this horizon provided that Q is; which we shall assume to be the case. For the components $T_{\mathcal{V}\mathcal{V}}$ and $T_{\mathcal{W}\mathcal{W}}$ we need to analyze the behaviour of \mathcal{F} (27) as $r \rightarrow r_0$. We assume that $f(r)$ has the form

$$f(r) = f_p(r - r_0)^p + f_{p+1}(r - r_0)^{p+1} + \dots; \quad (29)$$

for $r \sim r_0$, where we are interested particularly in the cases $p = 1$ (non-extremal black hole), $p = 2$ (“cold” [17] black hole or “Nariai” [15] black hole) and $p = 3$ (“ultra-cold” [17] black hole). Performing a Taylor series expansion of all the quantities, we find

$$\mathcal{F}(r) = K_1(r - r_0) + K_2(r - r_0)^2 + O(r - r_0)^3; \quad (30)$$

where

$$\begin{aligned} K_1 &= \frac{1}{2(p+1)} \left[(n-1)\tilde{\mathcal{G}}(r_0)\frac{R'(r_0)}{R(r_0)} + \mathcal{G}'(r_0) \right]; \\ K_2 &= \frac{1}{2(p+1)(p+2)} \left\{ 2(n-1)Q'(r_0)\frac{R'(r_0)}{R(r_0)} \right. \\ &\quad + (p+1)\mathcal{G}''(r_0) - \mathcal{G}'(r_0)\frac{f_{p+1}}{f_p} \\ &\quad + p(n-1)\tilde{\mathcal{G}}'(r_0)\frac{R'(r_0)}{R(r_0)} \\ &\quad + (n-1)\tilde{\mathcal{G}}(r_0) \left[(p+1)\frac{R''(r_0)}{R(r_0)} \right. \\ &\quad \left. \left. - (n+p) \left(\frac{R'(r_0)}{R(r_0)} \right)^2 - \frac{f_{p+1}}{f_p} \frac{R'(r_0)}{R(r_0)} \right] \right\}; \end{aligned} \quad (31)$$

and

$$\begin{aligned} \mathcal{G}(r) &= (n-1)Q(r) - T(r); \\ \tilde{\mathcal{G}}(r) &= (n+1)Q(r) - T(r). \end{aligned} \quad (32)$$

We therefore have that $T_{\mathcal{V}\mathcal{V}}$ is finite at the horizon, while

$$\begin{aligned} T_{\mathcal{W}\mathcal{W}} &= \frac{1}{2} [K_1(r - r_0) + O(r - r_0)^2] f[\psi'(-\mathcal{W})]^2 \\ &\quad - \frac{[\psi'(-\mathcal{W})]^2}{2R^{(n-1)}} [Z + X]. \end{aligned} \quad (33)$$

The second term in (33) is $O(r - r_0)^{-2}$ as $r \rightarrow r_0$ for a non-extremal horizon ($p = 1$), $O(r - r_0)^{-4}$ for a doubly coincident horizon ($p = 2$) and $O(r - r_0)^{-6}$ for the “ultra-cold” black hole with $p = 3$. Therefore, $T_{\mathcal{W}\mathcal{W}}$ will diverge severely unless $X + Z = 0$.

Even if this is the case, $T_{\mathcal{W}\mathcal{W}}$ will still be divergent if $p > 1$ as the first term in (33) is $O(r - r_0)^{-p+1}$. The only way for the RSET to be regular on an extremal horizon is if, as well as imposing $X + Z = 0$, we also have $K_1 = 0$ (for $p = 2$) and, in addition $K_2 = 0$ if $p = 3$. It is clear that these are strong constraints on the RSET, and, in general, it is unlikely that either K_1 or K_2 will vanish, so the RSET will be divergent at an extremal horizon. We will show in section VI B that, in the simpler two-dimensional case, when $n = 1$, these conditions cannot be satisfied at an extremal horizon. It is likely that this result extends to higher dimensions, but a full computation of the unknown function Q is required in this case.

The analysis proceeds similarly for a horizon where \mathcal{W} is finite and non-zero, but \mathcal{V} vanishes. In this case, we require $X - Z = 0$ in order for $T_{\mathcal{V}\mathcal{V}}$ to be regular, and, in addition, $K_1 = 0$ for an extremal horizon, with $K_2 = 0$ as well if the horizon is ultra-extremal.

Most of the extremal black hole geometries shown in [15] have both past and future extremal horizons, and, in order for the RSET to be finite on both, it must be the case that $X = 0 = Z$ and $K_1 = 0$ (with K_2 also zero if $k = 3$). However, the “Nariai”-type black hole (see [15]) is different in that it has a future extremal horizon but no past horizon. In this case the criteria are less stringent, and the RSET will be regular if $X + Z = 0$ and $K_1 = 0$ (note that the horizon in this case is not ultra-extremal).

Typically, horizons where $\mathcal{V} = 0$ will correspond to past horizons and those with $\mathcal{W} = 0$ will correspond to future horizons. Assuming this to be the case, we summarize our results from this section in table I, which gives the conditions on the constants X and Z for the RSET to be regular on the possible combinations of future (\mathcal{H}^+) and past (\mathcal{H}^-) non-extremal horizons. These conditions are necessary and sufficient for regularity on non-extremal horizons; for extremal horizons they are necessary but not sufficient as we also have the additional requirements for the K_i to vanish as outlined above. Table I gives us exactly the same results as in [4]: the only way that the RSET can be regular on both the future and past event horizon is if $X = Z = 0$, so that we have a time-symmetric state which represents the Hartle-Hawking vacuum [5]. For the Boulware vacuum [7] we have $X \neq \pm Z$ and the RSET is divergent on both event horizons. In this case the RSET should be time-reversal

Conditions on X and Z	Regular on	
	\mathcal{H}^-	\mathcal{H}^+
$X = Z = 0$	Yes	Yes
$X = Z \neq 0$	Yes	No
$X = -Z \neq 0$	No	Yes
$X \neq \pm Z$	No	No

TABLE I: Conditions for the RSET to be regular on a non-extremal black hole event horizon.

symmetric so we set $X = 0$ but $Z \neq 0$. Finally, for the Unruh vacuum [6], the RSET is regular on the future but not the past event horizon, so that $X = -Z \neq 0$, and we have a non-zero P (24), which represents the outgoing Hawking flux. The last case, in which the RSET is regular on the past but not the future event horizon, may represent a “future” Unruh state (see, for example, [14]) in which there is an ingoing flux of radiation but none emitted. We do not consider this possibility further.

V. DISTINCT EVENT AND COSMOLOGICAL HORIZONS

We now turn to the case in which the geometry has two distinct, non-extremal, horizons (the prototype being the Reissner-Nordström-de Sitter black hole, with an event and cosmological horizon at r_+ and r_{++} respectively). We consider the region between the horizons and use two sets of Kruskal co-ordinates, V_+ , W_+ and V_{++} , W_{++} , as discussed in section II B. Near the event horizon, we have the same results as in the previous subsection, for the co-ordinate system V_+ , W_+ , leading to table I.

Fixing $r_a = r_+$ as above, we find, near the cosmological horizon,

$$J(r_+, r) = J(r_+, r_{++}) + O(r - r_{++})^2, \quad (34)$$

which gives, near the future cosmological horizon \mathcal{C}^+ ,

$$\begin{aligned} T_{V_{++}}^{W_{++}} &= \left[\tilde{Z} + X \right] \left[C_0 (r - r_{++})^{-2} + C_1 (r - r_{++})^{-1} \right] \\ &\quad + O(r - r_{++})^0; \\ T_{W_{++}}^{V_{++}} &= \left[\tilde{Z} - X \right] \left[D_0 + D_1 (r - r_{++}) \right] + O(r - r_{++})^2, \end{aligned} \quad (35)$$

where

$$\tilde{Z} = Z + J(r_+, r_{++}); \quad (36)$$

whilst, near the past cosmological horizon \mathcal{C}^- :

$$\begin{aligned} T_{V_{++}}^{W_{++}} &= \left[\tilde{Z} + X \right] \left[\tilde{C}_0 + \tilde{C}_1 (r - r_{++}) \right] + O(r - r_{++})^2; \\ T_{W_{++}}^{V_{++}} &= \left[\tilde{Z} - X \right] \left[\tilde{D}_0 (r - r_{++})^{-2} + \tilde{D}_1 (r - r_{++})^{-1} \right] \\ &\quad + O(r - r_{++})^0; \end{aligned} \quad (37)$$

where the C , D , \tilde{C} and \tilde{D} are fixed, non-zero constants. All other components of the RSET are automatically regular across the cosmological horizon. Using (35,37), we

State	Conditions on X , Z and \tilde{Z}	Regular on	
		\mathcal{H}	\mathcal{C}
Boulware	$X = 0 \neq Z$, $\tilde{Z} \neq 0$	No	No
Hartle-Hawking	$X = Z = 0$, $\tilde{Z} \neq 0$	Yes	No
Gibbons-Hawking	$X = \tilde{Z} = 0$, $Z \neq 0$	No	Yes
Regular	$X = Z = \tilde{Z} = 0$	Yes	Yes

TABLE II: Properties of the physically relevant static states on event and cosmological horizons.

can build up a table of conditions similar to table I, but involving \tilde{Z} rather than Z . Combining all the possible behaviours at the event and cosmological horizons, one could build up a long list of different combinations, although most of these will not be physically relevant. Instead, we list in table II the properties of the physically relevant static states (which are all time-reversal symmetric).

Each of the two horizons, the event and cosmological horizon, will have an intrinsic temperature associated to it, given by $\kappa/2\pi$, where κ is the surface gravity of that horizon (11). If we therefore define a zero-temperature, “Boulware” state, then we would expect that it will be divergent on both the event and cosmological horizons. This state will be time-reversal symmetric, so $X = 0$, but Z and \tilde{Z} are unrestricted.

Secondly, we may consider a “Hartle-Hawking” state, which is a finite temperature state at the same temperature as the event horizon, but, in general, will not be at the same temperature as the cosmological horizon. Therefore, in analogy with the Hartle-Hawking state for an asymptotically flat black hole, we expect that this state will be regular on the event horizon. However, due to the temperature difference between the event and cosmological horizons, we expect this state to be divergent on the cosmological horizon. As in the previous section, this means that $X = 0 = Z$ but leaves \tilde{Z} unrestricted.

In a similar way, a “Gibbons-Hawking” [23] state, at the same temperature as the cosmological horizon, will be regular on the cosmological horizon but divergent on the event horizon in general. Therefore, in this case, $X = 0 = \tilde{Z}$ but Z is arbitrary.

Although, in general, the temperatures of the event and cosmological horizons will be different, there are special cases in which they are the same. A good example of this is the “lukewarm” Reissner-Nordström-de Sitter black hole [17, 24]. In this case, it is reasonable to suppose that a state at that temperature might be regular on both the event and cosmological horizons. This places great restrictions on the stress-tensor: it means that $X = 0 = Z$, and, in addition $\tilde{Z} = 0$, or, equivalently, $J(r_+, r_{++}) = 0$, that is,

$$\begin{aligned} 0 &= \frac{1}{2fR^{(n-1)}} \int_{r_+}^{r_{++}} [(n-1)Q - T] f' R^{(n-1)} dr \\ &\quad - \frac{(n-1)}{fR^{(n-1)}} \int_{r_+}^{r_{++}} Q f R^{(n-2)} R' dr. \end{aligned} \quad (38)$$

We will show in section VI that this condition is in fact satisfied for the “lukewarm” black hole in two dimensions, when the unknown function $Q(r)$ is absent. For four or higher dimensions, it is a non-trivial question as to whether a state can be constructed for a “lukewarm” black hole such that (38) holds, to which we plan to return in the future [25]. On the other hand, the Kay/Wald theorem [26] tells us that there is no static state on the Schwarzschild-de Sitter space-time which is regular on both the event and cosmological horizons. In this case, equation (38) does not hold.

We have not mentioned in this section the equivalent of the “Unruh” state for this type of black hole space-time. The reason for this is that the state which represents the emission of Hawking radiation by a black hole formed by gravitational collapse in de Sitter space is not static [12, 13].

We comment that, although our analysis in this section has focussed on the case of an event and cosmological horizon, similar considerations apply to any two distinct horizons, such as an event and inner horizon. In each case, a stringent integral condition similar to (38) must be satisfied if the RSET is to be regular on both horizons. A similar condition has been found for the regularity of the RSET on the inner horizon of a Kerr black hole [27].

VI. TWO-DIMENSIONAL EXAMPLES

In this paper we are primarily concerned with higher-dimensional black holes, but it is informative to study two-dimensional examples as in this case there is no unknown function Q . We consider two particular cases: (i) distinct, non-extremal event and cosmological horizons, and (ii) extremal horizons.

A. Distinct event and cosmological horizons

In this subsection we consider a two-dimensional black hole with two distinct horizons, corresponding to an event horizon and a cosmological horizon. We will show that the RSET can only be regular on both horizons if the temperatures of the two horizons are equal.

From section V, the RSET can only be regular on both the event and cosmological horizons if equation (38) holds. In two dimensions, this reduces to

$$0 = -\frac{1}{2f} \int_{r_+}^{r_{++}} f' T dr. \quad (39)$$

In two dimensions, the trace anomaly T is given by [21]

$$T = \alpha \mathcal{R}, \quad (40)$$

where \mathcal{R} is the Ricci scalar of the two-dimensional metric (3) and α is a constant, independent of the space-time geometry but dependent on the spin of the quantum field

under consideration. For a conformally coupled quantum scalar field, for example,

$$\alpha = \frac{1}{24\pi}. \quad (41)$$

In the two-dimensional case, the Ricci scalar is

$$\mathcal{R} = -f''(r). \quad (42)$$

The integral (39) then becomes

$$\begin{aligned} 0 &= \frac{\alpha}{2f} \int_{r_+}^{r_{++}} f' f'' dr \\ &= \frac{\alpha}{4f} [f'(r_{++})^2 - f'(r_+)^2] \\ &= \frac{\alpha}{f} [\kappa_{++}^2 - \kappa_+^2]; \end{aligned} \quad (43)$$

using (11), where κ_{++} and κ_+ are the surface gravities of the cosmological and event horizons, respectively. Therefore it must be the case that $\kappa_{++} = \kappa_+$ if the RSET is regular on both the horizons. Since the temperatures of the horizons are proportional to the surface gravities, therefore the event and cosmological horizons have the same temperature. This occurs for “lukewarm” Reissner-Nordström-de Sitter black holes [17, 24]. Therefore, for two-dimensional “lukewarm” black holes, the RSET will be regular on both the event and cosmological horizons if $X = 0 = Z$. To extend this result to higher dimensions requires a computation of the function $Q(r)$, to which we will return in the future [25].

B. Extremal horizons

For extremal horizons, if the RSET is to be regular across the event horizon, as well as the conditions on the constants X and Z outlined in table I, we also have conditions on the constants K_i given in equation (31). In two dimensions, the quantity \mathcal{G} in (32) reduces to

$$\mathcal{G} = -T = \alpha f'', \quad (44)$$

and therefore the constants K_i become

$$\begin{aligned} K_1 &= -\frac{\alpha}{2(p+1)} f'''(r_0); \\ K_2 &= -\frac{\alpha}{2(p+1)(p+2)} [(p+1)f''''(r_0) \\ &\quad - \frac{f_{p+1}}{f_p} f'''(r_0)]. \end{aligned} \quad (45)$$

For an extremal horizon, we have $p = 2$, and the condition for regularity is that $K_1 = 0$. However, for an extremal horizon, by definition $f'''(r_0) \neq 0$ and so $K_1 \neq 0$. Therefore the RSET for a static state must diverge on an extremal horizon. On the other hand, for an ultra-extremal horizon, we have $p = 3$ and by definition

$K_1 = 0$. In this case, however, we also have the additional condition that K_2 must vanish if the RSET is to be regular. In the ultra-extremal case we have

$$K_2 = -\frac{\alpha}{10} f''''(r_0) \neq 0, \quad (46)$$

and once again it must be the case that the RSET diverges for a static state.

The question of the regularity of the RSET on extremal horizons has been particularly controversial in the literature (see, for example, [11]). Our results here are in agreement with the general current consensus in the literature [11], namely that, if the RSET is to be regular across an extremal event horizon, the state it describes must be non-static.

VII. CONCLUSIONS

In this paper, we have examined the RSET for static quantum states on $(n+1)$ -dimensional, static, black hole space-times. We have generalized the analysis of [4], solving the conservation equations in this case. The RSET is then given in terms of one unknown function of the

radial co-ordinate r , and two unknown constants X and Z . We have used generalized Kruskal-like co-ordinates, following [10], to study the behaviour of the RSET near a horizon, and have derived conditions for the RSET to be regular there. We hope that our results will be of use for full computations of the RSET on either higher-dimensional black holes, or on black holes with a complicated horizon structure. In particular, for a wide class of black hole space-times, and Hartle-Hawking-like quantum states, the RSET is given by a single unknown function of the radial co-ordinate r , which reduces the amount of computation required.

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