Space science applications for conducting polymer particles: synthetic mimics for cosmic dust and micrometeorites

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Over the last decade or so, a range of polypyrrole-based particles have been designed and evaluated for space science applications. This electrically conductive polymer enables such particles to efficiently acquire surface charge, which in turn allows their acceleration up to the hypervelocity regime (\(>1\) km s\(^{-1}\)) using a Van de Graaff accelerator. Either organic latex (e.g. polystyrene or poly(methyl methacrylate)) or various inorganic materials (such as silica, olivine or pyrrhotite) can be coated with polypyrrole; these core–shell particles are useful mimics for understanding the hypervelocity impact ionisation behaviour of micro-meteorites (a.k.a. cosmic dust). Impacts on metal targets at relatively low hypervelocities (\(<10\) km s\(^{-1}\)) generate ionic plasma composed mainly of molecular fragments, whereas higher hypervelocities (\(>10\) km s\(^{-1}\)) generate predominately atomic species, since many more chemical bonds are cleaved if the particles impinge with higher kinetic energy. Such fundamental studies are relevant to the calibration of the cosmic dust analyser (CDA) onboard the Cassini spacecraft, which was designed to determine the chemical composition of Saturn’s dust rings. Inspired by volcanism observed for one of the Jupiter’s moons (Io), polypyrrole-coated sulfur-rich latexes have also been designed to help space scientists understand ionisation spectra originating from sulfur-rich dust particles. Finally, relatively large (20 \(\mu\)m diameter) polypyrrole-coated polystyrene latexes have proven to be useful for understanding the extent of thermal ablation of organic projectiles when fired at ultralow density aerogel targets at up to 6.1 \(\text{km s}^{-1}\) using a Light Gas Gun. In this case, the sacrificial polypyrrole overlayer simply provides a sensitive spectroscopic signature (rather than a conductive overlayer), and the scientific findings have important implications for the detection of organic dust grains during the Stardust space mission.

Background: organic conducting polymers

Most conventional polymers are electrical insulators and many applications make use of such properties (e.g. plastic coatings for cables and wires, printed circuit boards, etc.). Although the subject of various isolated research papers stretching back over a hundred years or more,\(^1\)–\(^5\) organic conducting polymers were only recognised as a distinct class of materials since the Nobel Prize-winning discovery of highly conducting polyacetylene in 1977.\(^6\) Over the last twenty years or so, conducting polymers have become essential components in organic solar cells, polymer-based light emitting diodes and polymer lasers.\(^7\)–\(^9\) The essential pre-requisites for an electrically conductive polymer are an extensively conjugated backbone and mobile charge carriers (i.e. holes or electrons). Prototype conducting polymers such as polycetylene suffer from significant chemical degradation: the highly conjugated backbone is readily attacked by aerial oxygen and/or water, leading to rapid conductivity decay over time scales of hours to days.\(^1\) In contrast, polyheterocycles such as polypyrrole (PPy) or poly(3,4-ethylenedioxy-thiophene) (PEDOT) exhibit much better long-term conductivity stability than polycetylene, whereas polymeric quaternary ammonium salts such as polyaniline (PANI) are completely air-stable over time scales of years.\(^1\)–\(^1\)\(^3\)

In this review article, our focus is on PPy, PANi and PEDOT (see chemical structures in Fig. 1), with particular emphasis being placed on the former material. The chemical synthesis of each of these polymers involves oxidation polymerisation, typically in aqueous acidic solution.\(^1\) In the absence of a suitable stabiliser, the conducting polymer is obtained as an insoluble bulk powder of very limited processability.‡

‡ Indeed, one of the original objectives in developing synthetic routes to colloidal forms of conducting polymers was to improve the intractability of these fascinating materials.
The polypyrrole chains are both lightly cross-linked and highly conjugated. The latter property gives rise to the intensely black colouration of this material and the delocalised positive charge along its backbone (there is typically one ‘hole’ per 3 to 4 polymerised pyrrole rings) leads to solid-state electrical conductivities of approximately 1–10 S cm\(^{-1}\), as judged by four-point probe measurements conducted on pressed pellets at room temperature. This places polypyrrole within the metallic regime (room temperature). This places polypyrrole within the metallic regime (room temperature). This places polypyrrole within the metallic regime (room temperature). This places polypyrrole within the metallic regime (room temperature). This places polypyrrole within the metallic regime (room temperature). This places polypyrrole within the metallic regime (room temperature).

(ii) the ease of coating various colloidal substrates (latexes, inorganic oxides, mineral grains etc.) with a contiguous polypyrrole overlayer from aqueous solution at room temperature. Generally, FeCl\(_3\) is preferred for the polymerisation of pyrrole.\(^{14}\) This oxidant gives a product with a relatively high conductivity and its rate of polymerisation is slow enough to be compatible with a wide range of colloidal formulations. In contrast, (NH\(_4\))\(_2\)S\(_2\)O\(_8\) is a much stronger oxidant that can often introduce carbonyl defects into the polymer backbone, as confirmed by FT-IR spectroscopy studies.\(^{17,18}\) Such over-oxidation leads to lower conductivities by up to an order of magnitude. Moreover, the kinetics of polymerisation of pyrrole is much faster when using (NH\(_4\))\(_2\)S\(_2\)O\(_8\): polymerisations are often complete within a few minutes at room temperature (rather than typically 12–24 h for FeCl\(_3\)-mediated polymerisations).\(^{19,20}\) This is simply too fast for some colloidal formulations, resulting in partial or complete precipitation of the PPy particles. Nevertheless, (NH\(_4\))\(_2\)S\(_2\)O\(_8\) can still be a preferred oxidant in some cases. For example, when coating certain

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delicate substrates that are known to be susceptible to over-oxidation, it may be preferable to use (NH₄)₂S₂O₈ in order to minimise the contact time of the substrate with the acidic oxidising solution (see later). In such cases, excess pyrrole monomer can be employed to ensure that all of the oxidant is consumed and the oxidant can be added last to the reaction mixture so as to minimise any oxidative surface degradation of the mineral grains.

Over the last decade or so, we have examined four classes of conducting polymer-based particles as putative mimics for understanding the behaviour of various types of micro-meteorites (a.k.a. ‘cosmic dust’) in space science experiments. Here the relatively high electrical conductivity of the particles is critical because it allows efficient accumulation of surface charge. This in turn enables the particles to be accelerated up to very high velocities (see later), which correspond to those attained by micro-meteorites within our Solar System. This provides space scientists with the opportunity to conduct fundamental laboratory-based experiments on model dust particles with relatively narrow size distributions under carefully controlled conditions at velocities similar to those encountered in space. The synthetic routes utilised to produce each class of conducting polymer particles are discussed in turn below.

**Sterically-stabilised polypyrrole latexes**

If the polymerisation of pyrrole is conducted in the presence of a suitable water-soluble polymer, colloidal polypyrrole particles can be produced by a process known as aqueous dispersion polymerisation. Initially, the monomer, oxidant and polymeric stabiliser are soluble in the aqueous phase. As the conducting polymer is formed, the polymeric stabiliser adsorbs onto the microscopic precipitating nuclei, preventing their further aggregation by a mechanism known as steric stabilisation. The result is an aqueous dispersion of stericly-stabilised polypyrrole latexes (see Fig. 3). Typically, the polymeric stabiliser is merely physically adsorbed onto the surface of the polypyrrole particles via either hydrogen bonding or electrostatics; although chemically-grafted stabilisers have also been developed, this latter approach has not yet been utilised for space science applications. The biocompatible nature of polypyrrole, coupled with its intense intrinsic pigmentation and strong absorption of near-IR radiation, has led to various biomedical applications being explored for such polypyrrole-based nanoparticles.

One of the most convenient, albeit empirical, methods for adjusting the mean particle size of sterically-stabilised polypyrrole latexes is simply to vary the nature of the steric stabiliser. Thus using poly(vinyl alcohol) typically affords polypyrrole latexes of approximately 100 nm diameter, whereas a poly(2-vinylpyridine)-based stabiliser yields particles with a mean diameter of around 200 nm and utilising a high molecular weight poly(ethylene oxide) produces 300 nm particles. In each case relatively narrow particle size distributions are typically obtained, as judged by charge-velocity analysis studies.

The stabiliser content of such polypyrrole latexes typically depends on both the latex diameter and the stabiliser type. Smaller, higher surface area latexes tend to contain more stabiliser, but stabiliser contents are rarely above 10% by mass. Thus the underlying polypyrrole cores are exposed, which allows sufficient charge to be accumulated to enable efficient electrostatic acceleration using a Van de Graaff instrument (see later).

**Polypyrrole-coated latexes**

Although the mean diameter of sterically-stabilised polypyrrole latexes can be readily varied, it is not possible to prepare micrometer-sized particles by this route. Instead, near-mono-disperse, micrometer-sized sterically-stabilised polystyrene (PS) latexes can be prepared by non-aqueous dispersion polymerisation, followed by the controlled deposition of an
ultrathin polypyrrole overlayer onto these particles from aqueous solution.36–39 (see Fig. 4). It is important to determine the specific surface area of the original polystyrene latex by an appropriate technique such as BET surface area analysis.38 From the latex mass utilised, the total latex surface area can be readily calculated for a given formulation. Thus, given the densities of the latex (ρlatex) and the conducting polymer (ρPPy), the latex mass (Mlatex) and the mean latex radius (Rlatex), this information enables the approximate polypyrrole mass loading (MPPy) required to produce a particular polypyrrole overlayer thickness (x) to be determined using eqn (1).38

\[
x = R_{\text{latex}} \left( \frac{M_{\text{PPy}}}{M_{\text{latex}} + M_{\text{PPy}}} \right)^{1/3} - 1
\]

In practice, this approach is limited to overlayer thicknesses of around 5–30 nm. If much thicker overlayers are targeted, the conducting polymer coating becomes rather inhomogeneous and the deposition process is much less well controlled, with some macroscopic precipitation of the conducting polymer usually observed. However, optimised protocols invariably lead to well-defined ‘core–shell’ particles with electrically insulating cores and electrically conductive shells, with typical polypyrrole loadings of 1–10% by mass depending on the mean latex diameter and the desired shell thickness. This robust protocol was subsequently extended to include somewhat larger polystyrene latex particles of 20 μm.40 Scanning electron microscopy studies indicate that the polypyrrole overlayer is relatively smooth and uniform (see Fig. 5C). Solvent extraction experiments confirm that the overlayer is both robust and contiguous, since a ‘goldfish bowl’ morphology is observed for the insoluble polypyrrole residues after all the underlying polystyrene latex has been removed. This protocol can also be used to coat submicrometer-sized latex particles.41,42 The same approach also works reasonably well for the deposition of PANi and PEDOT, but in these cases the deposited conducting polymer overlayer tends to be somewhat less uniform43–45 (see Fig. 5B and D). Use of less hydrophobic latexes such as poly(methyl methacrylate) (PMMA) particles also tends to produce relatively inhomogeneous conducting polymer coatings46 (see Fig. 13A).

**Polypyrrole/silica nanocomposite particles**

If pyrrole is polymerised in the presence of an ultrafine 20 nm silica sol in aqueous solution under appropriate conditions, then polypyrrole/silica nanocomposite particles are formed with little or no macroscopic precipitation.47–49 (see Fig. 6). The conducting polymer chains adsorb onto the silica nanoparticles and bind them together to form colloiddal stable aggregates. Excess silica sol is readily removed by repeated centrifugation–redispersion cycles and the purified polypyrrole/silica nanocomposite particles can then be characterised by various techniques. Their surface compositions are invariably silica-rich as judged by X-ray photoelectron spectroscopy50 and aqueous electrophoresis,51 which no doubt accounts for their excellent long-term colloid stability. Thermogravimetric analysis readily provides the mean silica
contents of such nanocomposite particles. Alternatively, elemental microanalyses can be used to calculate the organic content of the nanocomposite particles (using polypyrrole bulk powder as a reference material), with the silica content then being obtained by subtraction. It is possible to vary the mean nanocomposite diameter and also the mean silica content, but these two synthesis parameters appear to be inter-dependent.\textsuperscript{17,47} Thus using the FeCl\textsubscript{3} oxidant typically produces larger polypyrrole/silica nanocomposite particles of 250–300 nm diameter and relatively low silica contents (30–40\% by mass) whereas using (NH\textsubscript{4})\textsubscript{2}S\textsubscript{2}O\textsubscript{8} invariably leads to smaller nanocomposite particles of 110–180 nm comprising 50–60\% silica. Both electron microscopy and disc centrifuge photosedimentometry indicate that relatively narrow size distributions are obtained in each case. Given the electrically insulating nature of the silica component, solid-state conductivities for these nanocomposite particles are generally somewhat lower than for the other colloidal forms of polypyrrole, ranging from 10\textsuperscript{–3} to 10\textsuperscript{0} S cm\textsuperscript{–}. Nevertheless, this is sufficient to allow the accumulation of sufficient surface charge to enable acceleration up to hypervelocities (see later). This general approach has also been extended to include ultrafine tin(\textit{iv}) oxide nanoparticles instead of silica sols.\textsuperscript{59} In this case, somewhat higher electrical conductivities can be achieved, particularly if the tin(\textit{iv}) oxide component is doped with antimony in order to render it semi-conductive.\textsuperscript{54}

**Polypyrrole-coated mineral grains**

Various mineral grains (e.g. olivine, pyroxene, pyrrhotite, aluminosilicates etc.) can be conveniently coated with an ultrathin overlayer of polypyrrole by simply using the same formulation developed to coat latex particles (see above) and substituting the mineral grains for the latex particles (see Fig. 7A). Again, knowledge of the ‘sphere-equivalent’ specific surface area and density of the mineral grains of interest is essential for calculating the mean polypyrrole overlayer thickness \textit{via} eqn (1). Generally, the polypyrrole overlayer is not quite as uniform as that achieved for the latex particles, but it is usually sufficient for acceleration up to hypervelocities. The mineral grains are usually obtained by grinding up the corresponding macroscopic material, so they often have rather polydisperse and ill-defined particle morphologies (see Fig. 7B and C). Nevertheless, such projectiles are of considerable interest to space scientists, since their chemical compositions often closely match those of known micro-meteorites and cosmic dust.\textsuperscript{52,58} In the case of certain minerals such as pyrrhotite,\textsuperscript{52} the oxidising conditions used for polypyrrole deposition may be detrimental to their chemical stability. In this case, it is preferable to use (NH\textsubscript{4})\textsubscript{2}S\textsubscript{2}O\textsubscript{8} as the oxidant rather than FeCl\textsubscript{3} since this ensures much shorter reaction times (just a few minutes, rather than many hours). In addition, the oxidant is added last to the reaction mixture in order to minimise chemical degradation of the mineral grains. Selected electron micrographs for various polypyrrole-coated mineral grains are shown in Fig. 7B and C.

Notwithstanding the synthesis of the polypyrrole/silica nanocomposite particles described above, it is relatively difficult to coat large silica particles with an ultrathin contiguous overlayer of polypyrrole.\textsuperscript{58} This is because the highly anionic silica surface is relatively hydrophilic, and polypyrrole does not readily ‘wet’ such substrates, leading to an inhomogeneous, non-uniform coating.\textsuperscript{54} This problem has been recently addressed by surface modification of near-monodisperse silica particles of approximately 1 \textmu m diameter using a commercial organosilane reagent, 3-(methacryloyl)propyl triethoxysilane.\textsuperscript{57} This increases the surface wettability of the silica particles and hence enables much more uniform deposition of the polypyrrole overlayer, as judged by scanning electron microscopy (see Fig. 8). The resulting polypyrrole-coated silica particles can be accelerated up to 7–9 km s\textsuperscript{–}\textsuperscript{1} using a Van de Graaff instrument\textsuperscript{54} (see later) and have been used as model projectiles to generate crater impacts in aluminium foils at various angles of incidence.\textsuperscript{56,60}

**Other conducting polymer colloids**

Although the vast majority of our studies have been conducted using polypyrrole-based projectiles, there are at least two viable alternatives to this conducting polymer. Both polyaniline (PANi) and poly(3,4-ethylenedioxythiophene) (PEDOT) have been prepared...
An alternative approach to conducting polymer deposition: surface metallisation

Recently, surface metallisation of electrically insulating projectiles using colloidal precious metal seeds has been reported. It has enabled silica nanoparticle agglomerate dust particles of 0.1–1.0 μm diameter to be coated with platinum, which is sufficient to enable the projectiles to be accelerated to hypervelocities ranging from 4 to 30 km s⁻¹ in Van de Graaff experiments. More recently, the same approach has been utilised to coat anorthite, mixed-silicate (predominantly orthopyroxene) particles and olivine particles with low-density metallic overlayers and such approaches are actively being developed.

One important advantage of this metallisation approach is that such coatings produce ionic plasmas that contain simple metal cations which can be readily identified by mass spectroscopy. In contrast, conducting polymer coatings can fragment to generate a range of molecular and/or atomic ions, depending on the impact velocity. As discussed above, studies of such ionic plasmas from organic particles usually indicate that their mass spectra are typically dominated by the latex core, rather than the coating material. However, more recent experiments have shown that, at least in the case of impact velocities below approximately 30 km s⁻¹, molecular fragments originating from the polypyrrole coating can contribute to the plasma mass spectra. Depending on the resolution of the mass spectrometer, this can complicate spectra assignment and interpretation. In this context surface metallisation appears to offer a useful alternative approach to deposition of an organic conducting polymer. However, the relatively high density of platinum (~22 g cm⁻³) means that the density of the coated particles will be significantly higher than the precursor particles, even when targeting metallic overlayers of just 5–20 nm. In contrast, the relatively low densities of conducting polymers (~1.50 g cm⁻³) actually results in a slight reduction in particle density for most mineral grains, as well as a far smaller increase in mass. Conducting polymer-coated projectiles are therefore best suited to hypervelocity experiments which are highly sensitive to particle mass and/or density, such as those investigating impact charge, crater or aerogel track morphology.

Hypervelocity experiments

Most of the cosmic dust particles commonly found in our Solar System fall into one of the following four categories: metallic, silicate-rich, carbonaceous or icy. The former class can be
readily mimicked using finely divided metal powders or metal sols. For example, experimental studies often focus on iron particles, since their high intrinsic conductivity enables efficient charging and acceleration up to hypervelocities. However, the other three classes of cosmic dust are primarily electrical insulators and for many years none of these materials were suitable for laboratory-based electrostatic acceleration hypervelocity experiments. The use of electrostatic acceleration is essential when velocities higher than approximately 7 km s⁻¹ or precise control of the particle flux or mass are required. For lower velocities, or the acceleration of multiple grains simultaneously, a light gas gun may be used, which does not require charge-carrying cosmic dust analogues. Advances in the field of conducting polymers have allowed a range of synthetic carbon-rich projectiles to be developed, while conducting polymer coatings have enabled various other mineral grains to be accelerated up to hypervelocities using Van de Graaff accelerators.

The first report of the acceleration of conducting polymer-based particles up to the hypervelocity regime (> 1 km s⁻¹) was by Armes et al. A series of sterically-stabilised polypyrrole particles were evaluated in turn and charge-velocity analysis used to determine their particle size distributions using a small scale 20 kV ‘test bench’ accelerator that allowed velocities of up to 5 km s⁻¹ to be achieved. This approach enabled particles to be accelerated into the hypervelocity regime using the established principle of electrostatic acceleration, confirming the suitability of the particles and their coating for use in the large accelerators, to which access may be limited for purely testing and characterisation purposes.

The relation, \( qV = 0.5mv^2 \), where \( q \) is the particle charge, \( V \) is the applied voltage, \( m \) is the particle mass and \( v \) its velocity, governs the mass-velocity space in which cosmic dust analogues may be electrostatically accelerated. According to this equation, increasing the applied electric field leads to a corresponding increase in the maximum velocity that can be achieved for a given projectile. In practice, a high-field strength Van de Graaff accelerator is required to obtain hypervelocities, with accelerators currently in use utilising potential differences of between 2.0 MV and 3.0 MV. Accurate measurement of the charge carried by individual grains (via induction and charge amplification) and the velocities of the grains (via induction detectors separated by an accurately known distance), combined with the known acceleration voltage allows the particle mass to be determined. If the grain density is also known, then its equivalent spherical radius may be estimated.

For example, a 130 nm diameter polypyrrole particle (whose mass is approximately \( 1.68 \times 10^{-18} \) kg with a charge of \( 6.49 \times 10^{-16} \) C, accelerated through a 2.0 MV applied field can attain a hypervelocity of up to 39.3 km s⁻¹ (approximately 88 000 mph). As far as we are aware, this is the highest velocity ever reported for a synthetic organic projectile.

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Impact ionisation experiments

The enormous energy density of a conducting polymer-based projectile impinging on a metal target at more than a few km s⁻¹ is sufficient to cause extensive bond scission, as well as ionisation (Fig. 10). This leads to the formation of ionic plasma that can be interrogated by time-of-flight mass spectroscopy or charge integrators. The instruments used for mass spectroscopy are typically duplicates of mass spectrometers on-board various spacecraft. As such, their mass resolution is much lower than state-of-the-art analytical mass spectrometers, due to the highly restrictive limitations of size, weight and energy consumption that are necessarily placed on space instrumentation. Nevertheless, space instruments such as the Cosmic Dust Analyser (CDA, Fig. 11) onboard the Cassini spacecraft are capable of discriminating between ions with a mass-dependent mass resolution \((m/\Delta m)\) of approximately 50.

Typically, the impact plasma are separated using an applied electric field (e.g. 330 kV m⁻¹ in CDA), with either the cation or anion component accelerated towards an ion detector, through a low-field (or field-free) region. In recent years, higher resolution mass spectrometers, designed for space-based operation and capable of detecting both cationic and anionic species, have been developed (e.g. the Large Area Mass Analyser, LAMA) (Fig. 11). These spectrometers use a more complicated field geometry (reflectron) within the instrument to remove the effect of the initial ion energies within the impact plasma, as well as providing a longer ion trajectory, thus increasing the mass resolution of the instrument \((m/\Delta m \approx 200)\). In both the drift and reflection cases, the ion detection time \(t_i\) is related to its mass \((m_i)\) and charge \((q_i)\) by \(t_i = a(m_i/q_i)^{0.5} + b\), where \(a'\) is the stretch parameter, which is related to the electric field geometry and strength within the instrument and \(b\) is the shift parameter.

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**Fig. 10** Schematic representation showing the main events during a typical hypervelocity impact experiment. Adapted from Khan et al., Fig. 3.

**Fig. 11** (A) Digital image and (B) schematic of the cosmic dust analyser on board Cassini and (C) schematic of the LAMA spectrometer. Image A adapted from Srama et al., Fig. 1, schematic B adapted from Goldsworthy et al., Fig. 2 and schematic C adapted from Sternovsky et al., Fig. 1.

*It has recently been shown that millimetre-sized ice particles can be fired at speeds up to 6 km s⁻¹ using a two-stage light gas gun. However, it is not currently possible to accelerate micrometre-sized ice particles up to hypervelocities.*
which is related to the instrument recording trigger time. In the case of unknown $a$ and $b$, spectra can be calibrated using two peaks of known masses (assumed to be due to ions of the same energy) and arrival times to solve for $a$ and $b$.

Both the simpler CDA-type mass spectrometers and the more complex LAMA-type instruments utilise the plasma generated during the rapid deceleration of particles travelling at hypervelocities (defined as a velocity higher than the speed of sound in both target and projectile). The energies involved produce an impact cloud, typically composed of macroscopic ejecta, neutral and charged molecules and atoms as well as electrons. The plasma generated during a hypervelocity impact has components due to both the impinging projectile and the target. It is important to examine this impact ionisation phenomenon experimentally as there is a limited understanding of how it works as a function of impact speed and associated shock pressures. The behaviour of most materials, e.g. to what degree they ionise, in what form they will show up in mass spectra etc., thus has to be determined empirically.

At low ($< 5$ km s$^{-1}$) velocities, the charged species produced are typically due to easily ionised atoms, such as alkali metals (e.g. Na, K) or species with a high electron affinity (e.g. CN in the case of organics) (Fig. 12A). Increasing the impact velocity to between 5 km s$^{-1}$ and 15 km s$^{-1}$ results in the generation of additional species, both atomic (e.g. C, O, Si, S) and molecular (e.g. C$_2$H$_3$, C$_2$H$_5$), with lower molecular weights at higher velocities. Between 20 and 30 km s$^{-1}$, molecular organic cations become far less frequent, although molecular organic anions are still present. Above approximately 30 km s$^{-1}$ the impact cloud is dominated by atomic ions. Examples of mass spectra showing the progression from molecular to atomic species in the plasma can be observed for metal-coated particles and organic-rich, polymer-coated particles.

Impact ionisation mass spectrometry experiments using purely organic particles are designed to investigate and simulate mass spectra which would be generated by organic-rich micro-meteorites in space. These particles may come from various sources, including carbonaceous chondrites (i.e. from asteroids within the solar system), organic dust from cometary sources, or the tholin-rich processed surfaces of volatile-rich bodies such as Centaurs or trans-Neptunian objects, and finally poly-aromatic hydrocarbon-rich dust originating from outside the solar system. As well as their compositional similarity to such organic-rich grains, the synthetic polymer particles also possess relatively low densities ($< 1500$ kg m$^{-3}$). This is comparable to that of water ice or “fluffy” aggregate grains, such as those identified during the Stardust Interstellar Preliminary Examination.

Experiments performed using purely organic projectiles and a laboratory model of CDA equipped with a Rh target, prior to the arrival of Cassini in the Saturnian system, were designed to investigate the mass spectra of carbonaceous grains which may have been intercepted during Cassini’s cruise phase. Polypyrrole and PEDOT-coated polystyrene latexes, as well as pure polypyrrole particles, were accelerated over a range of impact velocities, with cation mass spectra being recorded. These spectra show the evolution of the species produced with increasing impact velocity, as well as identifying many important charged molecular fragments, including the stable tropylium cation at a mass of 91 u (see Fig. 12B).

Later experiments using a laboratory model of the LAMA instrument investigated cation mass spectra derived from polypyrrole-coated polystyrene, PANi and poly[bis(4-vinylthiophenyl)sulfide] (PMPV) latexes, as well as anion mass spectra from polypyrrole-coated PMPV latex. These mass spectra were produced by impacts onto an Ag target, whose large area is designed for the detection of cosmic dust in extremely low flux environments. Impacts were performed over an impact velocity range of 3–35 km s$^{-1}$. Large molecular fragments that survived at impact velocities below 10 km s$^{-1}$ were found to be best suited to the identification of the parent organic species, but data collected at higher hypervelocities (10–35 km s$^{-1}$) were also found to be useful.

The chemical structure, as well as the specific composition, of the organic material involved in an impact affects the nature of the molecular fragments produced. This effect was investigated by comparing mass spectra produced by impacts using polypyrrole-coated PMMA latex, an aliphatic-rich projectile, with those produced by impacts using polypyrrole-coated PS latex, an aromatic-rich projectile. Such PMMA latexes are good mimics for the aliphatic organic species found in carbonaceous chondrites. Impacts of such polypyrrole-coated PMMA particles onto a Rh target at 4–8 km s$^{-1}$ resulted in mass spectra

![Fig. 12](image-url) (A) Impact ionisation spectra of PPy-coated PS microparticles onto rhodium at 5 and 9 km s$^{-1}$. (B) Fragmentation scheme for the typical cation fragments observed, resulting from cleavage of aromatic rings from polystyrene. Graph A adapted from Goldsworthy et al., Fig. 10.
exhibiting fewer large molecular species than those found in mass spectra of the polypyrrole-coated PS latex particles at the same speeds. This is in agreement with the known thermal fragility of PMMA compared with that of PS (Fig. 13). In particular, for the aliphatic microparticles there is no distinctive signal at 91 amu, which has been previously attributed to the
tropylium cation, \( \text{C}_7\text{H}_7^+ \). Instead, cationic molecular fragments are observed at 41, 65 and 115 amu, which are traceable to specific bond scissions in the chemical structure of PMMA (Fig. 13c).

Although these studies used solid metallic targets (albeit with some degree of organic contamination, as described by Postberg et al.\(^{97}\)) experiments have also been performed investigating the species produced by sulphur-rich polypyrrole-coated organic projectiles\(^{41}\) impacting porous metal (Ag, Au) targets at velocities of up to \( 30 \text{ km s}^{-1} \).\(^{83}\) The targets comprised highly porous nanostructured surfaces of relatively low density (so-called ‘metal blacks’). After the high-speed impact, latex polymer chains were chemically degraded into molecular fragments. These fragments included both carbon- and sulfur-based species, which were detected as a series of cations and anions by time-of-flight mass spectrometry. Analysis of the mass spectra confirmed that greater chemical degradation occurred at higher velocities until only atomic ions were formed. Furthermore, ‘metal black’ targets led to greater fragmentation than more compact surfaces, which could be a consequence of much smaller impact spots.\(^{83}\) This indicates that the formation of molecular fragments occurs during expansion from the high-pressure shock state.

However, purely organic particles are less likely than those containing some mineral component. Silicates, such as olivines and orthopyroxenes form a major component of both asteroidal and cometary dust within the Solar System,\(^{98,99}\) and that found in the Interstellar Medium and dust-forming regions.\(^{94,100}\) Early impact ionisation mass spectrometry studies\(^{70}\) of polymer–mineral agglomerates used aluminosilicate clay nano- and microparticles, within a matrix of polypyrrole. Speeds as high as \( 50.7 \text{ km s}^{-1} \) were obtained, although the definitive identification of Al or Si in the mass spectra recorded using a CDA laboratory model was complicated by the presence of organic molecular ions.

Subsequent acceleration and analysis of a simpler mineral, pyrrhotite (also coated with polypyrrole), using the higher resolution LAMA instrument showed that ionic species originating from the mineral core could be detected at velocities as low as \( 7 \text{ km s}^{-1} \).\(^{52}\) Although the suspected presence of an oxidised (and possibly hydrated) sulfate layer between the pyrrhotite and polypyrrole unexpectedly complicated the interpretation of mass spectra, organic molecular anions were detected at higher velocities than expected (\( > 20 \text{ km s}^{-1} \)). Several key “fingerprint” molecules were identified as decomposition products of the polypyrrole backbone were also observed at masses of 66, 93 and 105 amu, which were assigned to \( \text{C}_6\text{H}_5\text{N}^+ \), \( \text{C}_5\text{H}_5\text{N}_2^+ \) and \( \text{C}_6\text{H}_3\text{N}_2^+ \) cations, respectively.

Simpler instruments, such as those found on the Ulysses\(^{101}\) and Galileo\(^{102}\) spacecraft, measure only the amplitude and evolution of the overall charge signal produced during a hypervelocity impact. With this information, and suitable laboratory calibration, estimates of particle masses and impact velocities may be made.\(^{103}\) Impacts occurring outside the Chemical Analyser Target (the Rh central target of CDA\(^{73a}\)) produced very similar signals to those produced by Galileo and Ulysses, and a laboratory-based duplicate of the CDA instrument has been used for similar impact charge signal calibration.
The first impact charge signal data originating from the use of conductive polymer particles was reported by Burchell et al. in 1999. A series of experiments with polypyrrole-, PANi- and PEDOT-coated polystyrene latexes were conducted to determine the relationship between impact velocity and the charge produced during an impact \( q = ab^2 \), where \( b = 1.91-2.02 \). This work was later expanded to include further polypyrrole- and PANi-coated polystyrene latexes, sterically-stabilised polypyrrole particles and also silica and tin(IV) oxide nanoparticles embedded within a polypyrrole or PANi matrix.

Ionisation charge yields were calculated for particles impacting Rh, Au or Cu targets, with the purely organic particles producing similar yields to those containing harder mineral grains. The majority of charge yield calibration data prior to these publications was performed using iron dust. These metallic particles produced yields similar to the lower density organic/organic-mineral grains at velocities above 10 km s\(^{-1}\), but significantly underestimated the impact charge produced at lower velocities (by a factor of three at 5 km s\(^{-1}\) and up to a factor of ten at 1 km s\(^{-1}\)). This effect was dependent on target material, with the yield differences for Fe being smaller when using Au and Rh targets, as utilised in the space instrumentation under investigation. Ionisation charge amplitudes were again found to scale exponentially with impact velocity, with ranges of \( b \) (defined above) of 2.6 to 3.6 for speeds below 18 km s\(^{-1}\), with one higher speed sample having \( b = 6.6 \), in comparison with an impact charge per unit mass scaling of \( v^{1.26} \) found for iron dust. Goldsworthy et al. also presented impact charge calibration data using both polypyrrole-coated polystyrene and polypyrrole-coated aluminosilicate clay particles. The charge yield was more strongly dependent on the speed of the impinging particle than that found in the earlier work using polypyrrole-coated mineral grains.

Impact cratering and aerogel tracks

Impact ionisation is the only method available to study cosmic dust compositions in situ. Alternatively, targets passively exposed in space can be returned to Earth for study. In this case they should contain microscopic craters formed as a result of the high-speed impacts. These craters often retain residues from the impinging particles and even fine details of the crater shape (depth, diameter, etc.) can offer useful information regarding the composition, shape and structure of the impactor. Laboratory experiments to simulate such impact events typically use two-stage light gas guns (LGGs, Fig. 14)\(^{14,106–109}\) which can fire at speeds up to 8 or 9 km s\(^{-1}\). These instruments rely on the compression and subsequent rapid expansion of a light gas (e.g. hydrogen or helium) to propel particles up to the hypervelocity regime. In this case a conductive coating is not required for acceleration, although it may still be a useful feature (see later). Cosmic dust analogue grains are usually accelerated en masse, as “buckshot”, producing multiple impacts at a single speed. Although a conducting polymer overlayer is no longer essential for LGG acceleration, it has proved useful to use the same microparticles in LGG work to understand impact cratering. This is discussed further in the following “Cometary dust particles” section. It is also possible to capture particles more or less intact even at impact speeds of a few km s\(^{-1}\), provided that the target material has a sufficiently low density. Again, the microparticles described here have been utilised for LGG calibration experiments to better understand how the high-speed capture of organic particles can alter their chemical nature.

Cassini mission

The Cassini space mission to Saturn carries the Cosmic Dust Analyser discussed above and shown in Fig. 11. Cassini is one of NASA’s most successful multi-year missions. Launched in 1997, it swung past Jupiter in 2000 and arrived at Saturn in 2004, where it became the first spacecraft to enter orbit around Saturn. Given that the Saturnian system has a rich, complex population of dust particles, the CDA has been an essential part of the Cassini mission. This instrument analyses impact plasma when the dust strikes its metal Rh target at a high speed. However, it is not sufficient to merely collect data, the impacts have to be correctly interpreted. This is why the microparticles described herein are so important to space science – by characterising their behaviour in laboratory experiments they have provided us with a much better understanding of the impact ionisation process. Indeed, altering the chemical compositions of these synthetic projectiles has allowed specific hypotheses to be tested regarding the likely ionisation spectra expected for various types of Saturnian dust (i.e., purely organic, mineral grains etc.).

For example, one important question for the Cassini mission concerned the volcanic activity of the Jovian moon Io. Io is the third largest of Jupiter’s moons and its volcanic activity has been known for more than three decades. On its journey to Saturn, Cassini flew past Jupiter and was thus able to investigate the dust ejected from Io. The surface colour of Io and spectroscopic analysis of its atmosphere provides strong evidence for the presence of sulfur on Io and perhaps also within its volcanic plumes. Accordingly, sulfur-containing microparticles were synthesised for hyper-velocity impact experiments. More specifically, this was achieved via dispersion polymerisation of a sulfur-rich divinyl monomer in an ethanol/water mixture, followed by coating with polypyrrole. Although these latexes were somewhat polydisperse in nature, they proved to be suitable mimics for sulfur-rich micrometeorites. They could be accelerated up to \( \sim 35 \text{ km s}^{-1} \) and their impact ionisation on striking a silver target led to both sulfur cations and anions being identified in the resulting ionic plasma using a LAMA detector. Surprisingly, the mass spectra obtained by the CDA near Jupiter were actually dominated by NaCl,
with relatively little sulfur content being detected (despite the strong evidence for sulfur on Io). Nevertheless, the observation of elemental sulfur ions in laboratory-based ionisation spectra confirms that the unexpected absence of sulfur in the spectra obtained by Cassini is a valid result, rather than merely an artefact arising from a poorly understood aspect of the impact ionisation process.

Cometary dust particles

The Stardust mission was commissioned to capture dust grains emanating from comet P81/Wild-2.110 Launched in 1999, this spacecraft flew past this comet at a relative velocity of 6.1 km s\(^{-1}\) in 2004 and returned the captured dust to Earth in 2006.111 In addition, Stardust also collected potential interstellar dust grains impinging over a wide range of impact speeds.112 As expected, a large number of dust grains were captured during the Stardust mission.111 Two types of targets were used: ultralow density aerogel targets designed to provide a ‘soft landing’ for fast-moving grains and aluminium foils, where the impacts produced craters lined with particle residues.110

The use of aerogel targets to collect cosmic dust in space has been reviewed by Burchell et al.112 Stardust utilised a transparent silica aerogel of ultralow density (ranging from 0.005 to 0.050 g cm\(^{-3}\)). Particles tunnel into the aerogel during high-speed impacts and leave long, thin tracks with (semi-)intact dust grains located at their end. Alternatively, bulbous cavities are formed as partial break-up occurs on impact to produce many finer fragments, which line the walls of the cavity or penetrate beneath it as thin tracks. Examples of track types have been given by Hoerz et al.,113 Burchell et al.,114 and Trigo-Rodriguez et al.115 for aerogel-captured cometary dust grains. Tracks from potential interstellar dust grains have been described by Westphal et al.116 Although many inorganic (mineral-based) dust particles were captured within the returned Stardust aerogels, remarkably few cometary organic grains were captured intact. Instead, much of the organic material was volatilised during the high-speed impact into the aerogel target. However, some particles showed evidence of aliphatic organic content.117,118 Many particles also proved to be coated with disordered carbon, as judged by the carbon D and G bands revealed by in situ Raman microscopy studies.117

Accordingly, commercial 20 \(\mu m\) polystyrene latex particles were coated with a polypyrrole overlayer of \(~20\) nm in order to conduct laboratory-based experiments with model organic particles to gain a better understanding of their behaviour during aerogel capture (Fig. 15). In these experiments, the ultrathin polypyrrole coating provided a convenient spectroscopic signature. Thus these projectiles are designed to be exquisitely sensitive to thermal ablation when fired into aerogel targets using an LGG (Fig. 16).109 Such core–shell particles survive aerogel capture intact at 1.07 km s\(^{-1}\), as judged by Raman microscopy studies of individual particles located at the end of ‘carrot tracks’ within the aerogel target (Fig. 17). However, this spectroscopic technique confirmed that surface carbonisation occurred after aerogel capture at approximately 1.95 km s\(^{-1}\), while the observation of D and G Raman bands indicated that the particles were subjected to substantial thermal ablation when captured at or above 3.33 km s\(^{-1}\).109 Indeed, at an impinging velocity of 6.11 km s\(^{-1}\) the mass loss of the captured projectile was estimated to be 84%. Moreover, consideration of the kinetic energies associated with hypervelocity capture under the experimental conditions suggests that these observations are consistent with the mean bond energies required for breaking the C–C, C–H and C=–C bonds found in polystyrene, which constitutes more than 99% of the projectile by mass (Fig. 18). These results confirm the hypothesis that many of the organics returned by Stardust were substantially thermally ablated during their
aerogel capture. Moreover, these findings suggest that relatively low encounter velocities (i.e., 1–2 km s\(^{-1}\)) will be essential for future space missions if organic dust grains originating from comets (or elsewhere) are to be captured intact with minimal thermal ablation.

Conclusions and prospect

Space science offers a fascinating, if rather esoteric, application for conducting polymer-based particles. The ability to systematically vary their particle size (from 100 nm up to 20 \(\mu\)m) and chemical composition makes them ideal mimics for studying a wide range of carbonaceous, sulfur-rich and silicate-based micro-meteorites. In addition, various mineral grains of astronomical interest such as olivine, pyroxene or pyrrhotite can be readily coated with ultrathin conducting polymer overlayers in order to allow their acceleration up to the hypervelocity regime. Compared to alternative approaches such as metallisation, such organic coatings are cost-effective, have much lower densities and, for many particulate substrates, are contiguous in nature. In most cases the conducting polymer coating simply allows the projectiles to acquire sufficient surface charge to enable electrostatic acceleration. However, it has been shown that this coating can also provide a very convenient spectroscopic signature that enables the extent of thermal ablation of aerogel-captured micro-meteorites to be assessed.

In future work, we intend to examine whether polyaromatic hydrocarbon (PAH) particles can be coated with polypyrrole. If successful, such projectiles should serve as interesting mimics for interstellar dust grains.

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Notes and references
