**L-shell spectroscopy of Neon and fluorine like copper ions from laser produced plasma**

Channprit Kaur1, 2, S. Chaurasia\*1, 2, Narendra Singh3, John Pasley4, Sunny Aggarwal3, Man Mohan3

*1High Pressure & Synchrotron Radiation Physics Division, BARC, Mumbai 400085, India*

*2Homi Bhabha National Institute, Mumbai-400094, India*

*3Delhi University, Delhi-110007, India*

*4York Plasma Institute, Department of Physics, University of York, York, YO10 5DQ, UK*

*(email: shibu@barc.gov.in)*

The Ne, F and O- like Rydberg resonance lines along with some of the inner shell satellite lines of Copper plasma, in the wavelength range of 7.9-9.5Å, are experimentally observed using a TAP crystal spectrometer. The plasma is produced by the irradiation of a Cu target with a 15 J, 500 ps Nd: Glass laser with a focusable intensity up to 5×1014 W/cm2. The observed lines result from the transitions among 2p-nd, 2p-ns and 2s-nd (n= 4 to 6) levels. Transition wavelengths, transition probabilities and oscillator strengths of these lines are calculated using the Multi-Configuration Dirac-Fock (MCDF) method. In this computation, the contribution of relativistic corrections such as two-body Breit corrections and QED corrections due to vacuum polarization and self-energy have also been considered. FLYCHK simulations are used to analyze the distribution of the various charge states of the Copper ions and to find the temperature and density of plasma. Moreover, the effect of self-absorption of the plasma (opacity), as well as of suprathermal electrons on charge state distribution of ions, are also studied. The synthetic spectrum provides a best-match with the experimental spectrum at a laser intensity of 1.3×1014 W/cm2 for Tc=150eV, Th=1000 eV, f=0.008 and density 4.5×1020cm-3.The temperature and density ranges are also calculated using a radiative hydrodynamic code. The calculated temperature and density range are in agreement with the experimentally determined values. The effect of change in laser intensity on the L-shell spectrum of Cu is studied which indicates the switching between lower (Cu XX) and higher charge states (Cu XXI, Cu XXII) at higher laser intensities.

# INTRODUCTION

The hot dense plasma, produced by high power lasers, has attracted an attention in various fields such as high energy density physics, atomic physics, and as a source of bright X-ray. The soft and hard X-rays produced from laser-produced plasma are useful for diagnosing conditions in dense plasma experiments[1-5](#_ENREF_1). X-rays are produced from plasma via free-free, free-bound and bound-bound transitions occurring in highly charged ions present in the plasma[6](#_ENREF_6).The X-ray line emission from bound-bound transitions from highly charged ions has proven to be an excellent method to obtain a wealth of information about the plasma. K-shell emission can easily be achieved by interacting a moderately intense laser with alow Z target. Extensive studies of K- shell emission spectra from laser-produced plasma have been carried out in the past[7-18](#_ENREF_7). The relative ease with which K-shell spectra can be modelled makes them an obvious target for study. On the other hand, although the L -shell ionization of high Z atoms can be achieved easily, complex models are required to interpret such spectra. L-shell ions are observed in many astrophysical phenomenons and in the laboratory. In astrophysics, the radiation produced by highly charged ions is a key observable from the standpoint of diagnosing the physical properties of non-terrestrial sources. L-shell spectroscopy provides a means to study the average ionization of plasma because the spectrum depends upon the abundance of ions from each charge state which is itself a function of the plasma temperature and density. Mid-Z tracers are used in Inertial Confinement Fusion (ICF) experiments [18](#_ENREF_18) as well as in the production of efficient sources of X-rays[19](#_ENREF_19), in tokomak diagnostics [20](#_ENREF_20), and in astrophysics [21-25](#_ENREF_21). In addition, recent astronomical research[26](#_ENREF_26) shows the presence of Cu ions in dwarf stars.

In the recent past, various efforts have been made to study the L-shell spectroscopy of mid-Z element, such as line identification from L-Shell copper ions[12](#_ENREF_12), [24](#_ENREF_24), [27-33](#_ENREF_27). However, the L-shell spectroscopy of Cu is still not fully understood. In this paper, analysis of the L-shell spectrum of a laser-produced Cu plasma in the wavelength range 7.5-9.5Å is performed. The transition wavelengths, transition probabilities and oscillator strengths of experimentally observed lines are calculated using a Multi-configuration Dirac-Fock (MCDF) method. In this computation, the contribution of relativistic corrections such as two-body Breit corrections and QED corrections due to vacuum polarization and self-energy has also been considered. FLYCHK simulations are used to analyze the distribution of various charge states of Copper and to find the temperature and density of the plasma. Moreover, the effect of opacity on charge state distributions of ions is also studied.

# EXPERIMENTAL PROCEDURE

The experiment was carried out using an Nd: Glass laser with output energy of 15 J per pulse with pulse duration of 500 ps. The laser was focused onto a Cu slab to a spot size of 100 μm using an f/5 lens, yielding intensities up to 4×1014 W/cm2.The experimental chamber was evacuated to a pressure of 4×10-5 mbar. An X-ray crystal spectrometer, made up of a Thallium Acid Phthalate (TAP) crystal placed at 45o with respect to the laser axis, was used for the line-emission studies in the spectral range 7.9-9.5 Å. Two stacked aluminized polycarbonate foils (Alexander Vacuum Research, Inc., trade name: B-10) having 1/e cut-off of 0.9 keV were used to prevent the scattered visible light from entering the plasma chamber. The TAP crystal spectrally resolves X-ray emission from the laser produced plasma. The reflected X-rays were detected by an X-ray CCD camera (Model VISION 4M, from Rigaku innovative) which has a resolution of 25 mÅ. The schematic of the crystal spectrometer along with the experimental setup is shown in Fig. 1. A sample image of a recorded X-ray spectrum is also shown in Fig. 1.

# RESULT AND DISCUSSION

The experimental spectrum consists of transitions corresponding to Oxygen (O-) (Cu XXII), Fluorine (F-) (Cu-XXI), Neon (Ne-) (Cu-XX) ions. The formation of these multiply charged ions takes place via collision of free electrons with atoms. The highest charge state of Cu in the present experiment is O-like with an ionization potential of 1.9 keV followed by F-like (1.8 keV) and Ne-like (1.69 keV). The Ne-like ions are easy to produce due to the low energy requirement and closed shell ground configuration in these ions. The shell configuration of Ne-like ions is 1s22sk2pl (k=1, 2; l=1 to 6). There are in total seven Rydberg series of dipole transitions from 2s22p6-2s22p5*nl’* and 2s22p6-2s2p6*nl’*. A standard notation for Ne-like lines is followed as provided in literature[34](#_ENREF_34) nA: 2s-np3/2, nB: 2s-np1/2, nC: 2p1/2-nd3/2, nD: 2p3/2-nd5/2, nE: 2p3/2-nd3/2, nF: 2p1/2-ns, nG: 2p3/2- ns. The procedure followed for identification of lines and calibration of wavelength is discussed below.

|  |
| --- |
|  |
| FIG. 1. Schematic of the experimental setup |

First identification of the wavelength of the spectral lines is done by ray-tracing, considering the location of the crystal and detector in addition to the dispersion curve of the crystal. Then the dispersion curve of the crystal is coupled with the two most intense line of the Ne-like (2P-4d) 4C and Ne-like (2p-4d) 2D, which results in a spectrum with error of 25 mÅ. Further identification is done using calculations performed with the General Purpose Relativistic Atomic Structure Package (GRASP) code of Grant *et al.*[35](#_ENREF_35), which employs the MCDF method. Configuration interaction has been included for O-like, F-like and Ne-like Cu. Similar calculations have also been performed using the fully relativistic flexible atomic code (FAC). Transition wavelength, oscillator strength, transition probabilities, and line strength are reported for electric dipole (E1), electric quadrupole (E2), magnetic dipole (M1) and magnetic quadrupole (M2) transitions from the ground level. First, for validation of our code, we match our generated data for some very well-known materials such as Aluminum and Silicon K-shell spectrum and data available in the literature. Then we use the same model for our calculation of the spectral lines of Ne-, F-, O- and Na-like ions. Here, we found some new transitions both in measured spectra and calculation, which are not reported earlier and may be useful for plasma parameter estimation.

We compared our calculated results with the available data in the literature. The calculated results are found to be in close agreement with previous results. Furthermore, we predict some new atomic data which may be important for plasma diagnostics.

**For Ne-like Cu**, in our MCDF calculations, we have included 51 configurations, 2s22p6*,* 2s22p53*l (l* = 0–2*)*, 2s2p6 3*l(l* = 0–2*)*, 2s2p6 4*l (l* = 0–3*),* 2s22p54*l (l* = 0–3*)*, 2s22p55*l (l* = 0–3*)*, 2s2p65*l (l* = 0–3*)*, 2s22p56*l(l* = 0–3*)*, 2s2p6 6*l (l* = 0–3*)*, 2s22p57*l (l* = 0–3*)*, 2s2p67*l (l* = 0–3*)*, 2s22p43s2, 2s22p43p2, 2s22p43d2, 2s2p53s2, 2s2p53p2, 2s2p53d2, 2s22p43s3p, 2s22p43p3d, 2s22p43s3d, 2s2p53s3p, 2s2p53p3d and 2s2p53s3d which give rise to 1016 fine structural levels. We mainly focus on one electron excitation by taking 39 configurations of single excitation due to negligible contributions of two or more electron excitations on energy levels and radiative rates.

**For F-like Cu**, in our MCDF calculations, we have included 27 configurations, 2s22p5*,* 2s2p6, 2s22p43*l (l* = 0–2*)*, 2s2p53*l (l* = 0–2*)*, 2p6 3*l (l* = 0–2*),* 2s22p4 4*l (l* = 0–3*),* 2s2p54*l (l* = 0–3*)*, 2s22p45*l (l* = 0–3*)*, 2s2p55*l (l* = 0–3*)* which give rise to 492 fine structural levels.

**For O-like Cu**, in our MCDF calculations, we have included 54 configurations, 2s22p4*,* 2s2p5,2p6,2s22p33*l (l* = 0–2*)*, 2s2p4 3*l (l* = 0–2*)*, 2p5 3*l (l* = 0–2*),* 2s22p3 4*l (l* = 0–3*),* 2s2p44*l (l* = 0–3*)*, 2p54 *l(l* = 0–3*),* 2s22p35*l (l* = 0–4*)*, 2s2p45*l (l* = 0–4*),* 2p5 5*l(l* = 0–4*),*2s22p36*l (l* = 0–4*)*, 2s2p46*l (l* = 0–4*),* 2p5 6*l (l* = 0–4*)* which give rise to 1434 fine structural levels. In our calculations, we have also included the contribution of relativistic corrections such as two-body Breit corrections and QED corrections due to vacuum polarization and self-energy.

**For Na-like Cu**, in our MCDF calculations, we have included 39 configurations, 2s22p63*l (l* = 0–2*)*, 2s22p6 4*l (l* = 0–3*),* 2s22p65*l (l* = 0–4*)*, 2s22p66*l (l* = 0–4*),* 2s22p67*l (l* = 0–4*)*, 2s22p68*l (l* = 0–4*),*2s22p53*l4l’ (l* = 0–2, *l’*= 0–3*)*which give rise to 619 fine structural levels. In our calculations, we have also included the contribution of relativistic corrections such as two-body Breit corrections and QED corrections due to vacuum polarization and self-energy.

The transition wavelength, transition rates, and oscillator strengths are calculated. The details of the calculated lines (wavelength, transition rates and oscillator strength) for Ne-like, O-like, F-like and Na-like Cu transitions are provided in the Table I, Table II, Table III and Table V respectively. For F-like transitions, the configuration detail is given in Table IV. The wavelength and relative transition rates obtained for different charge states, without considering the ion charge-state populations, are shown in Fig. 2(a) overlaid with the experimentally measured spectrum. All the identified lines are labeled in Fig. 2(b).

Most of the transition lines in the wavelength range under consideration belong to the 2p-nd transition giving rise to nC and nD pairs and the transitions from the 2p-ns and 2s-nd (n= 4 to 6) levels. The high energy spectral lines are associated with the transition to a 2p1/2 vacancy in the L-shell and the lower energy spectra result from transitions to fill the 2p3/2 vacancy.

TABLE I. Calculated wavelength, transition rates, oscillator strengths of Ne-like transitions along with corresponding configurations.

|  |  |  |  |  |  |  |  |  |
| --- | --- | --- | --- | --- | --- | --- | --- | --- |
| Notation | Lower level (i) | Upper level (j) | λ (in Å)(MCDF) | Aji(s-1)(MCDF) | fij(MCDF) | Sij(au)(MCDF) | Type | λFLYCHK(Å) |
| 4D | 2s22p61S0 | 2s22p54d 1P1 | 9.2443 | 1.12E+13 | 4.30E-01 | 1.31E-02 | E1 | 9.2673 |
| 4C | 2s22p61S0 | 2s22p54d 3D1 | 9.1144 | 1.06E+13 | 3.97E-01 | 1.19E-02 | E1 | 9.1254 |
|  | 2s22p61S0 | 2s22p55s 1P1 | 8.5723 | 3.89E+11 | 1.29E-02 | 3.63E-04 | E1 |  |
|  | 2s22p61S0 | 2s22p55d 3P1 | 8.4708 | 7.86E+10 | 2.54E-03 | 7.08E-05 | E1 |  |
| 5D | 2s22p61S0 | 2s22p55d 1P1 | 8.4562 | 6.14E+12 | 1.97E-01 | 5.50E-03 | E1 |  |
|  | 2s22p61S0 | 2s22p55s 3P1 | 8.4507 | 1.23E+12 | 3.96E-02 | 1.10E-03 | E1 |  |
| 4A, 4B | 2s22p61S0 | 2s2p64p 3P1 | 8.3899 | 1.12E+12 | 3.55E-02 | 9.80E-04 | E1 | 8.3823 |
|  | 2s22p61S0 | 2s2p64p 1P1 | 8.3763 | 3.14E+12 | 9.91E-02 | 2.73E-03 | E1 |  |
| 5C | 2s22p61S0 | 2s22p55d 3D1 | 8.3419 | 3.98E+12 | 1.24E-01 | 3.42E-03 | E1 | 8.344 |
|  | 2s22p61S0 | 2s22p56s 1P1 | 8.1413 | 2.02E+11 | 6.03E-03 | 1.62E-04 | E1 |  |
|  | 2s22p61S0 | 2s22p56D3P1 | 8.0893 | 3.26E+10 | 9.60E-04 | 2.56E-05 | E1 |  |
| 6D | 2s22p61S0 | 2s22p56D1P1 | 8.0809 | 4.51E+12 | 1.32E-01 | 3.52E-03 | E1 | 8.0844 |
| 6F | 2s22p61S0 | 2s22p56S3P1 | 8.032 | 6.84E+10 | 1.99E-03 | 5.25E-05 | E1 |  |
| 6C | 2s22p61S0 | 2s22p56D3D1 | 7.9761 | 2.56E+12 | 7.32E-02 | 1.92E-03 | E1 | 7.975 |

TABLE II. Calculated wavelength, transition rates, oscillator strengths of O-like transitions along with corresponding configurations.

|  |  |  |  |  |
| --- | --- | --- | --- | --- |
|  Transition | λ (in Å) | Aji(s-1)  | fij | Sij(a.u) |
| Lower level (i) |  Upper level (j) |
| 2s22p43P2 | 2s22p3 (4S)4d 5$D\_{1}^{o}$ | 8.3059 | 4.192E+11 | 2.601E-03 | 3.557E-04 |
| 2s22p43P2 | 2s22p3 (4S)4d 3$D\_{2}^{o}$ | 8.2904 | 2.819E+12 | 2.905E-02 | 3.964E-03 |
| 2s22p43P2 | 2s22p3 (4S)4d 3$D\_{1}^{o}$ | 8.2793 | 3.002E+11 | 1.851E-03 | 2.522E-04 |
| 2s22p43P2 | 2s22p3 (4S)4d 3$D\_{3}^{o}$ | 8.2782 | 1.002E+13 | 1.441E-01 | 1.963E-02 |
| 2s22p43P0 | 2s22p3 (2D)4d 3$D\_{1}^{o}$ | 8.2540 | 8.167E+11 | 2.502E-02 | 6.800E-04 |
| 2s22p43P0 | 2s22p3 (2D)4d 3$P\_{1}^{o}$ | 8.2352 | 1.196E+12 | 3.649E-02 | 9.893E-04 |
| 2s22p43P0 | 2s22p3 (2D)4d 1$P\_{1}^{o}$ | 8.2034 | 3.572E+10 | 1.081E-03 | 2.920E-05 |
| 2s22p43P1 | 2s22p3 (4S)4d3$D\_{3}^{o}$ | 8.2875 | 7.137E+09 | 7.349E-05 | 6.015E-06 |
| 2s22p43P1 | 2s22p3 (2D)4d 3$F\_{2}^{o}$ | 8.2878 | 6.513E+11 | 1.118E-02 | 9.150E-04 |
| 2s22p43P1 | 2s22p3 (2D)4d 1$S\_{0}^{o}$ | 8.2848 | 6.691E+11 | 2.295E-03 | 1.878E-04 |
| 2s22p43P1 | 2s22p3 (2D)4d 1$P\_{1}^{o}$ | 8.2686 | 4.147E+12 | 4.251E-02 | 3.471E-03 |
| 2s22p43P1 | 2s22p3 (2D)4d 3$D\_{2}^{o}$ | 8.2690 | 5.075E+12 | 8.670E-02 | 7.081E-03 |
| 2s22p43P1 | 2s22p3 (2D)4d 3$P\_{0}^{o}$ | 8.2400 | 4.412E+12 | 1.497E-02 | 1.218E-03 |
| 2s22p43P1 | 2s22p3 (2D)4d 3$P\_{2}^{o}$ | 8.2377 | 4.716E+11 | 7.995E-03 | 6.505E-04 |
| 2s22p43P1 | 2s22p3 (2D)4d 3$F\_{2}^{o}$ | 8.2365 | 2.960E+12 | 3.010E-02 | 2.449E-03 |
| 2s22p43P1 | 2s22p3 (2D)4d 3$S\_{1}^{o}$ | 8.2290 | 7.566E+11 | 7.681E-03 | 6.243E-04 |
| 2s22p43P1 | 2s22p3 (2D)4d 1$D\_{2}^{o}$ | 8.2256 | 9.513E+11 | 1.608E-02 | 1.307E-03 |
| 2s22p41D2 | 2s22p3 (2D)4d 3$P\_{2}^{o}$ | 8.2984 | 2.688E+09 | 2.775E-05 | 3.791E-06 |
| 2s22p41D2 | 2s22p3 (2D)4d 1$P\_{1}^{o}$ | 8.2972 | 8.351E+11 | 5.172E-03 | 7.063E-04 |
| 2s22p41D2 | 2s22p3 (2D)4d 3$S\_{1}^{o}$ | 8.2895 | 3.687E+12 | 2.279E-02 | 3.110E-03 |
| 2s22p41D2 | 2s22p3 (2D)4d 1$D\_{2}^{o}$ | 8.2861 | 8.344E+12 | 8.589E-02 | 1.171E-02 |
| 2s22p41D2 | 2s22p3 (2D)4d 1$F\_{3}^{o}$ | 8.2836 | 9.957E+12 | 1.434E-01 | 1.955E-02 |
| 2s22p41D2 | 2s22p3 (2P)4d 3$F\_{2}^{o}$ | 8.2430 | 1.309E+12 | 1.333E-02 | 1.809E-03 |
| 2s22p41D2 | 2s22p3 (2P)4d 3$P\_{2}^{o}$ | 8.2363 | 8.564E+11 | 8.709E-03 | 1.181E-03 |
| 2s22p41D2 | 2s22p3 (2P)4d 3$F\_{3}^{o}$ | 8.2348 | 3.616E+12 | 5.147E-02 | 6.977E-03 |
| 2s22p41D2 | 2s22p3 (2P)4d 3$D\_{1}^{o}$ | 8.2276 | 2.239E+11 | 1.364E-03 | 1.847E-04 |
| 2s22p41S0 | 2s22p3 (2P)4d 3$P\_{1}^{o}$ | 8.3074 | 6.048E+11 | 1.877E-02 | 5.134E-04 |
| 2s22p41S0 | 2s22p3 (2P)4d 1$P\_{1}^{o}$ | 8.2791 | 1.639E+13 | 5.053E-01 | 1.377E-02 |
| 2s2p53$P\_{2}^{o}$ | 2s2p4(2D)4d 3G3 | 8.2069 | 8.830E+10 | 1.248E-03 | 1.686E-04 |
| 2s2p53$P\_{1}^{o}$ | 2s2p4(2D)4d 3F2 | 8.2635 | 5.205E+10 | 8.880E-04 | 7.247E-05 |
| 2s2p53$P\_{1}^{o}$ | 2s2p4(2D)4d 3D1 | 8.2602 | 3.602E+12 | 3.684E-02 | 3.006E-03 |
| 2s2p53$P\_{1}^{o}$ | 2s2p4(2D)4d 3D2 | 8.2551 | 4.838E+12 | 8.238E-02 | 6.716E-03 |
| 2s2p53$P\_{1}^{o}$ | 2s2p4(2D)4d 3P1 | 8.2536 | 2.240E+12 | 2.287E-02 | 1.865E-03 |
| 2s2p53$P\_{1}^{o}$ | 2s2p4(2D)4d 3S1 | 8.2397 | 1.092E+09 | 1.111E-05 | 9.043E-07 |
| 2s2p53$P\_{1}^{o}$ | 2s2p4(2D)4d 3P2 | 8.2378 | 6.660E+10 | 1.129E-03 | 9.188E-05 |
| 2s2p53$P\_{1}^{o}$ | 2s2p4(2D)4d 1D2 | 8.2260 | 7.385E+11 | 1.249E-02 | 1.014E-03 |
| 2s2p53$P\_{1}^{o}$ | 2s2p4(2D)4d 1P1 | 8.2247 | 1.734E+12 | 1.759E-02 | 1.428E-03 |
| 2s2p53$P\_{1}^{o}$ | 2s2p4(2D)4d 1S0 | 8.2210 | 2.454E+12 | 8.288E-03 | 6.729E-04 |
| 2s2p53$P\_{0}^{o}$ | 2s2p4(2D)4d 3P1 | 8.3086 | 8.701E+12 | 2.701E-01 | 7.389E-03 |
| 2s2p53$P\_{0}^{o}$ | 2s2p4(2D)4d 3S1 | 8.2945 | 2.293E+11 | 7.094E-03 | 1.937E-04 |
| 2s2p53$P\_{0}^{o}$ | 2s2p4(2D)4d 1P1 | 8.2793 | 2.470E+11 | 7.613E-03 | 2.075E-04 |
| 2s2p53$P\_{0}^{o}$ | 2s2p4(2S)4d 3D1 | 8.2048 | 3.162E+11 | 9.575E-03 | 2.586E-04 |
| 2s2p51$P\_{1}^{o}$ | 2s2p4(2P)4d 1P1 | 8.3052 | 5.363E+12 | 5.546E-02 | 4.549E-03 |
| 2s2p51$P\_{1}^{o}$ | 2s2p4(2P)4d 1D2 | 8.3046 | 1.026E+13 | 1.767E-01 | 1.450E-02 |
| 2s2p51$P\_{1}^{o}$ | 2s2p4(2P)4d 3D1 | 8.2150 | 4.069E+12 | 4.117E-02 | 3.340E-03 |
| 2s2p51$P\_{1}^{o}$ | 2s2p4(2P)4d 3P2 | 8.2145 | 5.453E+11 | 9.194E-03 | 7.459E-04 |
| 2s2p51$P\_{1}^{o}$ | 2s2p4(2P)4d 3D2 | 8.2094 | 7.559E+12 | 1.273E-01 | 1.032E-02 |

TABLE III. Calculated wavelength, transition rates, oscillator strengths of F-like transitions along with corresponding configurations.

|  |  |  |  |  |  |
| --- | --- | --- | --- | --- | --- |
| i | j | λ (in Å) | Aji(s-1) | fij | Sij(au) |
| 3 | 434 | 8.08E+00 | 1.75E+10 | 1.71E-04 | 9.12E-06 |
| 3 | 424 | 8.08E+00 | 6.09E+07 | 1.19E-06 | 6.33E-08 |
| 2 | 253 | 8.08E+00 | 9.10E+11 | 1.78E-02 | 9.47E-04 |
| 2 | 246 | 8.08E+00 | 9.51E+11 | 9.31E-03 | 4.95E-04 |
| 1 | 209 | 8.09E+00 | 3.27E+12 | 3.20E-02 | 3.41E-03 |
| 3 | 410 | 8.09E+00 | 7.95E+12 | 7.79E-02 | 4.15E-03 |
| 1 | 208 | 8.09E+00 | 1.39E+12 | 2.05E-02 | 2.18E-03 |
| 2 | 242 | 8.09E+00 | 3.71E+11 | 7.28E-03 | 3.88E-04 |
| 2 | 241 | 8.09E+00 | 1.11E+12 | 1.09E-02 | 5.83E-04 |
| 2 | 240 | 8.09E+00 | 3.72E+11 | 7.31E-03 | 3.89E-04 |
| 3 | 407 | 8.09E+00 | 3.28E+12 | 6.44E-02 | 3.43E-03 |
| 1 | 182 | 8.59E+00 | 5.30E+12 | 8.81E-02 | 9.97E-03 |
| 3 | 371 | 8.59E+00 | 7.30E+11 | 1.62E-02 | 9.15E-04 |
| 3 | 370 | 8.60E+00 | 7.27E+11 | 8.05E-03 | 4.56E-04 |
| 1 | 180 | 8.60E+00 | 7.49E+12 | 8.30E-02 | 9.40E-03 |
| 1 | 179 | 8.60E+00 | 9.02E+12 | 5.00E-02 | 5.66E-03 |
| 1 | 178 | 8.60E+00 | 4.42E+11 | 7.35E-03 | 8.32E-04 |
| 1 | 169 | 8.65E+00 | 1.63E+11 | 9.14E-04 | 1.04E-04 |
| 1 | 166 | 8.65E+00 | 5.24E+11 | 5.89E-03 | 6.71E-04 |
| 1 | 165 | 8.66E+00 | 4.38E+12 | 7.39E-02 | 8.43E-03 |
| 1 | 163 | 8.66E+00 | 2.09E+12 | 3.53E-02 | 4.03E-03 |
| 1 | 164 | 8.67E+00 | 4.91E+11 | 8.29E-03 | 9.46E-04 |
| 1 | 162 | 8.67E+00 | 1.04E+12 | 1.18E-02 | 1.34E-03 |
| 1 | 160 | 8.67E+00 | 2.64E+11 | 2.98E-03 | 3.40E-04 |
| 1 | 159 | 8.68E+00 | 2.70E+11 | 1.52E-03 | 1.74E-04 |
| 3 | 351 | 8.70E+00 | 1.98E+10 | 4.50E-04 | 2.58E-05 |
| 3 | 346 | 8.70E+00 | 5.01E+08 | 1.14E-05 | 6.51E-07 |
| 3 | 345 | 8.70E+00 | 1.98E+10 | 2.24E-04 | 1.28E-05 |
| 1 | 155 | 8.70E+00 | 1.93E+12 | 3.28E-02 | 3.76E-03 |
| 2 | 184 | 8.71E+00 | 1.45E+13 | 1.65E-01 | 9.47E-03 |
| 1 | 152 | 8.71E+00 | 7.48E+11 | 8.51E-03 | 9.75E-04 |
| 2 | 183 | 8.71E+00 | 8.77E+12 | 1.99E-01 | 1.14E-02 |
| 2 | 180 | 8.72E+00 | 1.90E+12 | 4.34E-02 | 2.49E-03 |
| 2 | 179 | 8.72E+00 | 6.49E+11 | 7.41E-03 | 4.25E-04 |
| 1 | 140 | 8.76E+00 | 8.29E+12 | 1.43E-01 | 1.65E-02 |
| 3 | 328 | 8.76E+00 | 1.69E+12 | 3.88E-02 | 2.24E-03 |
| 1 | 139 | 8.76E+00 | 7.21E+12 | 8.30E-02 | 9.58E-03 |
| 3 | 320 | 8.76E+00 | 5.94E+07 | 6.84E-07 | 3.95E-08 |
| 3 | 319 | 8.77E+00 | 1.32E+12 | 3.05E-02 | 1.76E-03 |
| 1 | 136 | 8.77E+00 | 3.32E+12 | 1.92E-02 | 2.22E-03 |
| 3 | 318 | 8.77E+00 | 3.69E+12 | 8.52E-02 | 4.92E-03 |
| 3 | 316 | 8.78E+00 | 1.27E+08 | 2.93E-06 | 1.69E-07 |
| 2 | 169 | 8.78E+00 | 8.99E+11 | 1.04E-02 | 6.01E-04 |
| 2 | 166 | 8.78E+00 | 4.48E+12 | 1.04E-01 | 5.99E-03 |
| 1 | 131 | 8.79E+00 | 3.03E+10 | 1.75E-04 | 2.03E-05 |
| 3 | 305 | 8.79E+00 | 4.01E+10 | 4.64E-04 | 2.69E-05 |
| 1 | 130 | 8.79E+00 | 3.16E+09 | 3.66E-05 | 4.23E-06 |
| 1 | 128 | 8.79E+00 | 3.22E+09 | 5.60E-05 | 6.48E-06 |
| 2 | 160 | 8.80E+00 | 2.83E+10 | 6.58E-04 | 3.81E-05 |
| 2 | 162 | 8.80E+00 | 3.46E+11 | 8.04E-03 | 4.66E-04 |
| 3 | 297 | 8.80E+00 | 2.90E+11 | 6.74E-03 | 3.90E-04 |
| 2 | 159 | 8.80E+00 | 3.89E+11 | 4.52E-03 | 2.62E-04 |
| 1 | 126 | 8.81E+00 | 2.77E+10 | 3.23E-04 | 3.75E-05 |
| 1 | 125 | 8.81E+00 | 8.54E+11 | 1.49E-02 | 1.73E-03 |
| 3 | 293 | 8.83E+00 | 5.72E+12 | 1.34E-01 | 7.76E-03 |
| 3 | 289 | 8.84E+00 | 2.35E+12 | 2.75E-02 | 1.60E-03 |
| 2 | 152 | 8.84E+00 | 7.94E+10 | 1.86E-03 | 1.08E-04 |
| 3 | 287 | 8.84E+00 | 1.23E+12 | 2.89E-02 | 1.68E-03 |
| 3 | 285 | 8.84E+00 | 4.73E+11 | 5.54E-03 | 3.22E-04 |
| 3 | 283 | 8.84E+00 | 1.28E+11 | 3.00E-03 | 1.74E-04 |
| 3 | 282 | 8.85E+00 | 1.40E+11 | 3.28E-03 | 1.91E-04 |
| 3 | 277 | 8.85E+00 | 7.27E+10 | 8.54E-04 | 4.98E-05 |
| 3 | 266 | 8.88E+00 | 1.92E+09 | 4.53E-05 | 2.65E-06 |
| 3 | 269 | 8.88E+00 | 5.86E+12 | 6.92E-02 | 4.05E-03 |
| 1 | 121 | 8.88E+00 | 4.84E+11 | 2.86E-03 | 3.34E-04 |
| 3 | 264 | 8.88E+00 | 8.37E+11 | 1.98E-02 | 1.16E-03 |
| 3 | 257 | 8.88E+00 | 3.46E+09 | 8.19E-05 | 4.79E-06 |
| 3 | 247 | 8.88E+00 | 2.43E+12 | 2.88E-02 | 1.68E-03 |
| 1 | 120 | 8.88E+00 | 3.07E+11 | 3.63E-03 | 4.25E-04 |
| 3 | 245 | 8.89E+00 | 3.56E+12 | 4.22E-02 | 2.47E-03 |
| 2 | 139 | 8.89E+00 | 3.80E+11 | 9.01E-03 | 5.28E-04 |
| 3 | 243 | 8.90E+00 | 3.29E+12 | 7.82E-02 | 4.58E-03 |
| 2 | 136 | 8.90E+00 | 2.04E+11 | 2.42E-03 | 1.42E-04 |
| 2 | 131 | 8.92E+00 | 1.24E+11 | 1.48E-03 | 8.68E-05 |
| 2 | 130 | 8.92E+00 | 5.71E+10 | 1.36E-03 | 8.00E-05 |
| 1 | 114 | 8.92E+00 | 3.81E+11 | 2.27E-03 | 2.67E-04 |
| 3 | 226 | 8.93E+00 | 2.23E+10 | 5.32E-04 | 3.13E-05 |
| 3 | 225 | 8.93E+00 | 1.14E+10 | 1.37E-04 | 8.05E-06 |
| 2 | 126 | 8.94E+00 | 6.65E+11 | 1.60E-02 | 9.40E-04 |
| 3 | 221 | 8.95E+00 | 4.53E+10 | 1.09E-03 | 6.41E-05 |
| 3 | 218 | 8.96E+00 | 1.12E+10 | 1.35E-04 | 7.95E-06 |
| 3 | 215 | 8.97E+00 | 1.39E+06 | 3.36E-08 | 1.98E-09 |
| 3 | 213 | 9.00E+00 | 6.07E+11 | 7.36E-03 | 4.36E-04 |
| 1 | 111 | 9.00E+00 | 1.12E+12 | 1.35E-02 | 1.60E-03 |
| 1 | 110 | 9.01E+00 | 6.73E+10 | 1.23E-03 | 1.46E-04 |
| 2 | 121 | 9.01E+00 | 7.52E+11 | 9.15E-03 | 5.43E-04 |

Table IV. Table for configuration for F-like lines

|  |  |  |  |  |  |  |  |  |
| --- | --- | --- | --- | --- | --- | --- | --- | --- |
| Level | Label |  j | Level | Label | j | Level | Label | j |
| 1 | 2s22p5 2P | 3/2 | 179 | 2s22p44d1 4P | 1/2 | 282 | 2s22p45p12P | 3/2 |
| 2 | 2s22p5 2P | 1/2 | 180 | 2s22p44d1 4D | 3/2 | 283 | 2s22p45p12D | 3/2 |
| 3 | 2s12p62S | 1/2 | 182 | 2s22p44d1 4P | 5/2 | 284 | 2s12p54d12D | 5/2 |
| 110 | 2s22p44s1 2D | 5/2 | 183 | 2s22p44d1 4P | 3/2 | 285 | 2s22p45p12P | 1/2 |
| 111 | 2s22p44s1 2D | 3/2 | 184 | 2s22p44d1 4D | 1/2 | 286 | 2s12p54d12F | 7/2 |
| 114 | 2s22p44s1 2S | 1/2 | 208 | 2s12p54p1 4P | 5/2 | 287 | 2s12p54d12P | 3/2 |
| 120 | 2s22p44s1 2P | 3/2 | 209 | 2s12p54p1 2D | 3/2 | 288 | 2s12p54d12F | 5/2 |
| 121 | 2s22p44s1 2P | 1/2 | 210 | 2s22p45s12D | 3/2 | 289 | 2s12p54d12P | 1/2 |
| 125 | 2s22p44s1 4P | 5/2 | 213 | 2s12p54s12P | 1/2 | 293 | 2s12p54d12D | 3/2 |
| 126 | 2s22p44s1 4P | 3/2 | 215 | 2s22p45p12P | 3/2 | 297 | 2s12p54s14P | 3/2 |
| 128 | 2s22p44d1 2D | 5/2 | 218 | 2s22p45p12P | 1/2 | 305 | 2s12p54s14P | 1/2 |
| 130 | 2s22p44d1 2P | 3/2 | 221 | 2s22p45p12D | 3/2 | 316 | 2s22p45p14S | 3/2 |
| 131 | 2s22p44d1 2S | 1/2 | 225 | 2s12p54d12P | 1/2 | 317 | 2s22p45p14D | 5/2 |
| 136 | 2s22p44d1 2P | 1/2 | 226 | 2s12p54d12P | 3/2 | 318 | 2s12p54d14P | 3/2 |
| 139 | 2s22p44d1 2D | 3/2 | 240 | 2s22p45s12P | 3/2 | 319 | 2s22p45p14P | 3/2 |
| 140 | 2s22p44d1 2F | 5/2 | 241 | 2s22p45s12P | 1/2 | 320 | 2s22p45p14P | 1/2 |
| 152 | 2s22p44d1 2D | 3/2 | 242 | 2s22p45d12D | 3/2 | 328 | 2s22p45f12D | 5/2 |
| 155 | 2s22p44d1 2D | 5/2 | 243 | 2s12p54d12D | 3/2 | 345 | 2s22p45f14D | 1/2 |
| 159 | 2s22p44d1 2P | 1/2 | 245 | 2s22p45p12P | 1/2 | 346 | 2s22p45f14F | 3/2 |
| 160 | 2p63d1 2D | 3/2 | 246 | 2s12p54p12P | 1/2 | 351 | 2s22p45f14D | 3/2 |
| 162 | 2s22p44d1 2P | 3/2 | 253 | 2s12p54p14S | 3/2 | 370 | 2s22p45p14D | 1/2 |
| 163 | 2p63d1 2D | 5/2 | 257 | 2s22p45f12D | 3/2 | 371 | 2s22p45p14D | 3/2 |
| 164 | 2s22p44d1 2F | 5/2 | 264 | 2s22p45f12P | 3/2 | 407 | 2s12p55d12D | 3/2 |
| 165 | 2s22p44d1 2D | 5/2 | 266 | 2s22p45p12P | 3/2 | 408 | 2s12p55p12P | 3/2 |
| 169 | 2s22p44s1 4P | 1/2 | 269 | 2s22p45f12P | 1/2 | 410 | 2s12p55d14P | 1/2 |
| 178 | 2s22p44d1 4D | 5/2 | 277 | 2s22p45p12S | 1/2 | 424 | 2s12p55g14F | 3/2 |
|  |  |  |  |  |  | 434 | 2s12p55s12P | 1/2 |

TABLE V. Calculated wavelength, transition rates, oscillator strengths and line strengths of Na-like transitions along with corresponding configurations.

|  |  |  |  |  |
| --- | --- | --- | --- | --- |
| Transitions | λ (in Å) | Aji (in s-1) | fij | Sij (in a.u.) |
| i | j |
| 2s22p6 3s 2S1/2 | 2p53p (1D) 4p 2$D\_{3/2}^{o}$ | 9.338 | 4.4701E+10 | 1.1687E-03 | 7.1853E-05 |
| 2s22p6 3s 2S1/2 | 2p53s (3P) 4d 4$F\_{3/2}^{o}$ | 9.338 | 2.1180E+11 | 5.5376E-03 | 3.4047E-04 |
| 2s22p6 3s 2S1/2 | 2p53s (1P) 4d 2$P\_{3/2}^{o}$ | 9.324 | 7.3068E+10 | 9.5227E-04 | 5.8459E-05 |
| 2s22p6 3s 2S1/2 | 2p53p (1P) 4p 2$S\_{1/2}^{o}$ | 9.327 | 1.1345E+11 | 2.9590E-03 | 1.8171E-04 |
| 2s22p6 3s 2S1/2 | 2p53p (1P) 4p 2$D\_{3/2}^{o}$ | 9.318 | 1.1527E+12 | 3.0006E-02 | 1.8408E-03 |
| 2s22p6 3s 2S1/2 | 2p53p (1P) 4p 2$P\_{3/2}^{o}$ | 9.314 | 5.6739E+12 | 1.4760E-01 | 9.0517E-03 |
| 2s22p6 3s 2S1/2 | 2p53s (3P) 4d 2$P\_{3/2}^{o}$ | 9.311 | 1.5616E+12 | 4.0591E-02 | 2.4884E-03 |
| 2s22p6 3s 2S1/2 | 2p53s (1P) 4d 2$P\_{1/2}^{o}$ | 9.313 | 9.0861E+12 | 1.1813E-01 | 7.2433E-03 |
| 2s22p6 3s 2S1/2 | 2p53p (3S) 4p 2$P\_{1/2}^{o}$ | 9.307 | 7.1495E+10 | 9.2834E-04 | 5.6886E-05 |
| 2s22p6 3s 2S1/2 | 2p53p (1P) 4p 2$P\_{1/2}^{o}$ | 9.301 | 1.2043E+11 | 1.5620E-03 | 9.5661E-05 |
| 2s22p6 3s 2S1/2 | 2p53p (3P) 4p 4$P\_{3/2}^{o}$ | 9.300 | 1.6389E+11 | 4.2500E-03 | 2.6023E-04 |
| 2s22p6 3s 2S1/2 | 2p53p (3P) 4p 4$D\_{7/2}^{o}$ | 9.292 | 7.5641E+10 | 9.7909E-04 | 5.9901E-05 |
| 2s22p6 3s 2S1/2 | 2p53p (3D) 4s 2D3/2 | 9.268 | 3.4377E+10 | 8.8541E-04 | 5.4031E-05 |
| 2s22p6 3s 2S1/2 | 2p53p (3P) 4p 2$P\_{3/2}^{o}$ | 9.255 | 9.5897E+09 | 2.4631E-04 | 1.5010E-05 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53p (3D) 4d 4F3/2 | 9.337 | 5.7154E+12 | 1.4940E-01 | 9.1844E-03 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53p (1P) 4d 2P1/2 | 9.329 | 8.2166E+12 | 1.0719E-01 | 6.5838E-03 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53p (3P) 4d 2P3/2 | 9.309 | 3.3101E+11 | 8.5998E-03 | 5.2707E-04 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53p (1D) 4d 2G7/2 | 9.301 | 7.5888E+10 | 1.9685E-03 | 1.2055E-04 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53p (3P) 4d 4D1/2 | 9.299 | 5.5562E+11 | 7.2022E-03 | 4.4094E-04 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53p (3P) 4d 2D3/2 | 9.290 | 3.5322E+10 | 9.1399E-04 | 5.5905E-05 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53p (3D) 4d 2P1/2 | 9.287 | 1.2769E+11 | 1.6512E-03 | 1.0097E-04 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53p (1D) 4d 2P3/2 | 9.286 | 4.2848E+10 | 1.1079E-03 | 6.7740E-05 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53d (3P) 4p 4D1/2 | 9.278 | 1.7559E+11 | 2.2660E-03 | 1.3843E-04 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53p (3D) 4d 4P1/2 | 9.274 | 6.0317E+08 | 7.7770E-06 | 4.7487E-07 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53d (3P) 4p 4D3/2 | 9.266 | 6.3656E+07 | 1.6386E-06 | 9.9961E-08 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53d (3P) 4p 4P1/2 | 9.262 | 3.1616E+09 | 4.0656E-05 | 2.4792E-06 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53d (3P) 4p 4P3/2 | 9.257 | 2.3225E+10 | 5.9677E-04 | 3.6374E-05 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53d (3P) 4p 2P1/2 | 9.253 | 2.4053E+10 | 3.0871E-04 | 1.8807E-05 |
| 2s22p6 3p 2$P\_{1/2}^{o}$ | 2p53d (3D) 4p 4D3/2 | 9.251 | 2.1206E+10 | 5.4414E-04 | 3.3143E-05 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53p (1D) 4d 2G7/2 | 9.333 | 8.3347E+09 | 1.0883E-04 | 1.3374E-05 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53p (3P) 4d 4D1/2 | 9.330 | 2.7063E+11 | 1.7658E-03 | 2.1695E-04 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53p (1D) 4d 2D5/2 | 9.330 | 3.8666E+06 | 7.5692E-08 | 9.2997E-09 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53p (3P) 4d 2D3/2 | 9.321 | 4.9627E+12 | 6.4639E-02 | 7.9340E-03 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53p (3P) 4d 2D5/2 | 9.320 | 7.7799E+12 | 1.5197E-01 | 1.8652E-02 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53p (3D) 4d 2P1/2 | 9.319 | 5.9580E+10 | 3.8780E-04 | 4.7587E-05 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53p (1D) 4d 2P3/2 | 9.318 | 2.0452E+12 | 2.6619E-02 | 3.2661E-03 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (3P) 4p 4D1/2 | 9.309 | 8.2370E+12 | 5.3508E-02 | 6.5595E-03 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53p (3D) 4d 4P1/2 | 9.305 | 6.5583E+09 | 4.2564E-05 | 5.2155E-06 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (3P) 4p 4D3/2 | 9.297 | 1.5688E+08 | 2.0327E-06 | 2.4884E-07 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (3P) 4p 4P1/2 | 9.293 | 7.3148E+09 | 4.7347E-05 | 5.7938E-06 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (3P) 4p 4P3/2 | 9.288 | 8.7983E+08 | 1.1379E-05 | 1.3918E-06 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (3P) 4p 4D5/2 | 9.284 | 1.1237E+10 | 2.1781E-04 | 2.6629E-05 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (3P) 4p 2P1/2 | 9.284 | 3.5936E+10 | 2.3216E-04 | 2.8382E-05 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (3D) 4p 4D3/2 | 9.282 | 9.1824E+09 | 1.1860E-04 | 1.4496E-05 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (3F) 4p 4F7/2 | 9.279 | 5.7375E+10 | 1.1109E-03 | 1.3575E-04 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (3F) 4p 4F5/2 | 9.274 | 6.1208E+10 | 1.1838E-03 | 1.4457E-04 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (3P) 4p 4P5/2 | 9.271 | 8.2346E+06 | 1.5915E-07 | 1.9429E-08 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (3F) 4p 4F9/2 | 9.269 | 3.1637E+08 | 4.0745E-06 | 4.9730E-07 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (1D) 4p 2D3/2 | 9.266 | 1.6423E+10 | 2.1137E-04 | 2.5789E-05 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (1F) 4p 2F5/2 | 9.262 | 3.5491E+10 | 6.8462E-04 | 8.3499E-05 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (1D) 4p 2F5/2 | 9.258 | 2.3296E+09 | 4.4902E-05 | 5.4743E-06 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53d (3P) 4p 2D3/2 | 9.251 | 6.2299E+10 | 7.9935E-04 | 9.7381E-05 |
| 2s22p6 3p 2$P\_{3/2}^{o}$ | 2p53p (1D) 4f 2$F\_{7/2}^{o}$ | 9.249 | 9.4811E+10 | 6.0792E-04 | 7.4040E-05 |
| 2s22p6 3d 2D3/2 | 2p53d (3P) 4d 2$F\_{7/2}^{o}$ | 9.333 | 2.5352E+10 | 1.6555E-04 | 2.0347E-05 |
| 2s22p6 3d 2D3/2 | 2p53d (1F) 4d 2$P\_{3/2}^{o}$ | 9.330 | 8.4570E+10 | 1.1036E-03 | 1.3559E-04 |
| 2s22p6 3d 2D3/2 | 2p53d (3D) 4d 4$S\_{3/2}^{o}$ | 9.328 | 4.3146E+09 | 5.6281E-05 | 6.9131E-06 |
| 2s22p6 3d 2D3/2 | 2p53d (1F) 4d 2$D\_{5/2}^{o}$ | 9.324 | 4.3145E+09 | 8.4355E-05 | 1.0358E-05 |
| 2s22p6 3d 2D3/2 | 2p53d (3F) 4f 4I11/2 | 9.302 | 9.2469E+10 | 5.9969E-04 | 7.3453E-05 |
| 2s22p6 3d 2D3/2 | 2p53d (1F) 4f 2I13/2 | 9.292 | 4.9919E+12 | 9.6927E-02 | 1.1860E-02 |
| 2s22p6 3d 2D3/2 | 2p53d (1P) 4d 2$D\_{5/2}^{o}$ | 9.284 | 6.8304E+10 | 1.3239E-03 | 1.6186E-04 |
| 2s22p6 3d 2D3/2 | 2p53d (3F) 4f 4H7/2 | 9.280 | 1.8932E+12 | 2.4442E-02 | 2.9868E-03 |
| 2s22p6 3d 2D3/2 | 2p53d (1P) 4d 2$P\_{3/2}^{o}$ | 9.272 | 7.2929E+11 | 9.3999E-03 | 1.1477E-03 |
| 2s22p6 3d 2D3/2 | 2p53d (1D) 4f 2P1/2 | 9.264 | 2.0114E+11 | 1.2939E-03 | 1.5784E-04 |
| 2s22p6 3d 2D5/2 | 2p53d (1D) 4d 2$S\_{1/2}^{o}$ | 9.337 | 7.0642E+12 | 1.2311E-01 | 2.2706E-02 |
| 2s22p6 3d 2D5/2 | 2p53d (3P) 4d 2$F\_{7/2}^{o}$ | 9.335 | 6.4597E+12 | 5.6257E-02 | 1.0373E-02 |
| 2s22p6 3d 2D5/2 | 2p53d (1F) 4d 2$P\_{3/2}^{o}$ | 9.333 | 1.7046E+11 | 1.4839E-03 | 2.7354E-04 |
| 2s22p6 3d 2D5/2 | 2p53d (3D) 4d 4$S\_{3/2}^{o}$ | 9.329 | 1.0654E+09 | 1.3901E-05 | 2.5615E-06 |
| 2s22p6 3d 2D5/2 | 2p53d (1F) 4f 2I13/2 | 9.297 | 2.0000E+11 | 2.5916E-03 | 4.7591E-04 |
| 2s22p6 3d 2D5/2 | 2p53d (1P) 4d 2$D\_{5/2}^{o}$ | 9.289 | 8.7382E+11 | 1.1303E-02 | 2.0738E-03 |
| 2s22p6 3d 2D5/2 | 2p53d (3F) 4f 4H7/2 | 9.285 | 1.9302E+11 | 1.6629E-03 | 3.0498E-04 |
| 2s22p6 3d 2D5/2 | 2p53d (3F) 4f 2D3/2 | 9.284 | 4.4543E+10 | 7.6750E-04 | 1.4075E-04 |
| 2s22p6 3d 2D5/2 | 2p53d (1P) 4d 2$P\_{3/2}^{o}$ | 9.277 | 2.9984E+11 | 2.5791E-03 | 4.7262E-04 |

## Determination of plasma parameters

These highly charged ions are produced via various ionization processes such as photo-ionization, collisional-ionization and auto-ionization. The main recombination processes include radiative recombination, three body recombination and dielectronic recombination. The dominance of collisional or radiative processes decides whether plasma is in complete thermal equilibrium (CTE), local thermal equilibrium (LTE) or collisional radiative (CRE) equilibrium[36](#_ENREF_36). This in turn depends upon the temperature and density ranges present in the plasma. The knowledge of the equilibrium state facilitates the interpretation of spectroscopic data to uncover the charge state distribution and populations of the excited states.

CTE is not possible in lab environments due to the escape of radiation from the plasma. LTE occurs when the collision time between electrons and ions is short relative to that for other processes taking place in plasma. If LTE is present, then the temperature can be determined easily by the Boltzmann plot method [37](#_ENREF_37). The collisional radiative equilibrium is non-LTE in the sense that for this to occur, collisions and radiation both must play dominant roles. The solution of this equilibrium lies in solving the rate equations taking care of rates of all collisional and radiative excitation and de-excitation processes.

|  |
| --- |
| **(a)** |
| **(b)** |
| FIG. 2(a) The experimental spectrum with theoretical predicted lines intensity in the spectral range 7.8 Å-9.4 Å recorded for Ne-, F- and O-like ions of Copper. The relative transition rates are drawn. (b) The labeled experiment spectrum at laser intensity 1.3×1014 W/cm2. |

The validity of applying the LTE equilibrium approximation can be checked by applying the Mc Whirter criterion[38](#_ENREF_38) given by

$n\_{e}\geq 1.6×10^{12}T^{^{1}/\_{2}}(∆E)^{3}$ (1)

where ne is the electron density, T is plasma temperature, and ΔE is largest electronic transition. In our case, for plasma temperatures greater than 50 eV, the minimum density should be greater than 1022 /cm3, for the plasma to be in LTE. The required value is greater than our critical density value. Therefore, we can say that our plasma in not in LTE. We used the CRE based code FLYCHK [39](#_ENREF_39) for the determination of plasma temperature and density and the detailed procedure is provided below. In FLYCHK opacity effects are included by solving radiative transport equations for optically thick plasmas.

In the spectrum shown in Fig. 2 at a laser intensity of 1.3×1014 W/cm2, the identified lines belong to Ne-, F- and O-like copper ions. The intensity of Ne-like lines is higher than that of the F-like lines which indicates that the Ne- like ions are the dominant species. The experimentally observed spectral lines consist of 2p-4d, 2p-5d and 2p-6d transitions. These lines exist in pairs corresponding to resonance lines (nC) with transition 2s22p5nd3/21P1-2s22p6 1S0 and intercombination (IC) (nD) for 2s22p5nd1/23D1-2s22p6 1S0 (n=4 to 6 in our case). The ratio nC/nD is a strong function of the ion abundance which is a function of plasma temperature and density[40](#_ENREF_40). Our experimentally observed spectrum is time integrated but it should be borne in mind that the plasma has a strong temperature and density gradients which are a function of time. The Ne-like spectra of Cu dominates for a range of temperatures and densities.

To get a more accurate temperature and density for the most strongly radiating portion of the plasma generated at 1.4×1014 W/cm2, we compare 4C and 4D lines and the ratio of 4D to the Na-like IS satellite with the synthetic spectra generated by the FLYCHK code. The Na-like satellite lines originated from doubly excited Ne-like ions are observed and are reported by a few authors [41-43](#_ENREF_41) with very high precision. The energy level scheme for Ne-like and Na-like Cu is shown in Fig. 3(a).The 4D resonance line consists of a Na-like satellite line whose identification is provided in ref [42] with accuracy on the order of .001 Å. In ref [42], 4C and 4D lines along with several Na-like satellite transitions are given. Due to the comparatively low resolution of our crystal spectrometer, we are not able to distinguish all IS (inner shell) satellite lines. In Fig. 3(b), the experimentally observed satellite transitions are plotted. These consists of transitions from 3s-4d3/2, 4p3/2, 3p3/2-4d5/2,3/2,4d3/2-3s1/2, 3p1/2, 3d3/2,5/2.This satellite emission is also shown by FLYCHK. To match this experimental spectrum, we have run simulations over a wide range of temperatures and densities. Temperature are varied from 100 eV to 400 eV and the density is varied from 1×1020 cm-3 to 1×1021 cm-3. The effect of varying temperature and density on the ratios (4C/4D and Na-like satellite/4D) is observed by comparing spectra from FLYCHK keeping one parameter (either temperature or density) fixed. The synthetic spectrum is plotted for a temperature range of 150-210 eV (Fig. 4(a)) and for a density ranges of 3×1020- 5×1020 /cm3 (Fig. 4(b)). Fig. 5 (dotted line) shows the best match regarding the ratio 4C/4D and Na-like satellite / Ne-like resonance at a plasma temperature of 170 eV and a density of 6.5×1020 /cm3. Fournier et al. [30] in their work have shown the effect of varying the escape factor and the fraction of suprathermal electrons on the spectral lines. We included the effect of opacity in our simulation as the level population gets altered with the effects of collisions and photon re-absorption in the plasma. The intensity of a line is a function of the relative abundance of ions and their radiative transition rates and can be written as

$I\_{j,i}=n\_{j}A\_{j,i}$ (2)

which is a function of both the density of ions and the radiative transition rates. In the presence of plasma self-absorption radiative decay rates are modified Aeff =ɛAj,i with ɛ being the opacity. In our experiment the plasma is considered to be spherical with a diameter of 100 µm. This influences the distribution of charge states of Cu as shown in Fig. 8(c). Also, the contribution of hot electrons also needs to be taken in to account.

|  |  |
| --- | --- |
| **(a)** | **(b)** |
| FIG. 3 (a) Energy level diagram of resonance lines from Ne-like Cu and their Na-like satellites. (b) Na-like satellite corresponding to a Ne-like resonance line. |

|  |  |
| --- | --- |
| **(a)**(a) | **(b)** |
| FIG. 4. Dependence of the ratio of 4C to 4D lines and Na-like satellite to 4D line on (a) plasma temperature and (b) electron density as calculated by FLYCHK. |

|  |
| --- |
|  |
| FIG. 5. The experimental spectrum (at 1.3×1014 W/cm2 laser intensity) overlaid with a synthetic spectrum at plasmatemperature 170 eV and density 6.5×1020 /cm3. |

So, to accurately determine the ratio of resonance lines which are sensitive to plasma temperature and density, the influence of plasma opacity and the fraction of suprathermal electrons (f) have to be taken into account.This also shifts the ionization balance slightly higher. The effect of opacity in the plasma is already introduced in the simulation. The self-absorption results in a reduction of intensities of the nD lines in comparison to the nC lines. The intensity of our 6C line is more than the 6D resonance line. This indicates a higher opacity in our experimentally produced plasma. But, again, it should be borne in mind that this is a time and space integrated spectrum coming out from different regions of the plasma.

The effect of supra thermal electron temperature and fraction on the spectrum on the spectra is considered as follows. Matching between the synthetic and experimental spectra is performed keeping the thermal electron temperature to around 170 eV, while the supra thermal electron population temperature is varied from 0.5-1.5 keV and the ratio of supra thermal to thermal electron temperature is also varied from 0.005-0.1. An increase in the supra thermal electron temperature results in an alteration of the satellite to resonance intensity ratio and an increase in the fraction of supra thermal electrons raises the abundance of F-like ions and hence corresponding intensity ratios. The increase in f-value also results in a lowering of the intensity of the Na-like satellite line. The F-like lines are not distinguishable in FLYCHK because it follows the super-configuration approach, causing these lines to merge together. After several trials using different combinations of Tc, Th and f the best match of experimental and stimulated spectra at Tc=150eV, Th=1000 eV, f=0.008 and density 4.5×1020 cm-3 is shown in Fig. 6.

|  |
| --- |
|  |
| FIG. 6. Best match of experimental spectrum with synthetic spectrum at Tc=150 eV, Th=1000 eV, f=0.008 and density 4.5×1020 cm-3. |

## Effect of laser intensity on abundance of charge states in plasma

The laser intensity is changed to see the effect on the abundance of ions of different charge states. The spectrum at different laser intensities has been plotted in Fig. 7(a). The intensity of all lines increases with an increase in the intensity of the laser. However, the extent of the change of the intensity of lines from different charge states differs. This is due to the direct dependence of ion charge state populations on the plasma temperature and density. In Fig. 7(b&c), the zoomed part of the highlighted region is drawn. Four pairs of charge states (labeled as I, II, III and IV) are compared at varying laser intensities. The first two pairs indicate the increase in O-like Cu ions in comparison to Ne-like and F-like ions and pairs of lines in sections III and IV show comparison of F-like with Ne-like ions. When the laser intensity is low (1.3×1014 W/cm2), the intensities of Ne-like ion related emission are enhanced in comparison to F-like and O-like emission indicating the abundance of Cu XX ions in comparison to Cu XXI and Cu XXII ions in the plasma. When the laser intensity increases, the intensity ratio of Ne- like ions to F-like ions (shown in circle III and IV) decreases and the intensity of the F-like line at 8.27Å has exceeded the intensity of the Ne Like line at 9.12 Å at a laser intensity of 2.4×1014 W/cm2, whilst it is lower at a laser intensity of 1.3×1014 W/cm2. These observations indicate that with increased laser intensity, the population of F-like ions relative to Ne-like ions increases. At the same time, it can also be seen that, on further increase of laser intensity, O-like ions become increasingly prominent until O-like emission dominates at 4.0×1014 W/cm2. This indicates that the population of Cu XXII ions is more than the population of F-like and Ne-like ions at this intensity. The difference in intensity from O-like ions as compared to Ne-like ions is large in the circled zone II as compared to zone I. This may be due to the fact that in the circled zone I, the Ne-like line is accompanied by an F-like line and by increasing the laser intensity, the population of both F-like and O-like ions increases which results in the intensity of both lines being increased. At 3.2×1014 W/cm2, O-like emission starts to dominate Ne-like and F-like emission.

To validate our measurements, we performed FLYCHK simulations to evaluate the charge state distribution and abundance of the charge state of Cu ions at different plasma temperatures and densities with and without the effect of opacity. In the first case, we fixed the plasma density to 5×1020 cm-3 keeping the opacity off and changing the plasma temperature. The charge state distribution and relative population of the main contributing ions (Ne, F and O-like Cu ions) at different temperature are shown in Fig. 8(a & b). Without opacity, it is observed that the population of Ne-like ions starts dominating at temperatures greater than 150 eV and this continues up to 400 eV. This is against our experimental observation where it has been clearly seen that by increasing the laser intensity (plasma temperature), higher charge states have come to dominate the emission. When the effect of opacity is included, the abundance of F-like ions rises abruptly and their relative population becomes more than the Ne-like Cu above 200 eV as shown in Fig. 8(c & d). Above 210 eV, the population of O-like lines starts dominating and becomes equal to Ne-like at 280 eV. On further increase of temperature, O-like ions are dominating even over F-like ion at 350 eV.

We have measured plasma temperature and density at a laser intensity of 1.3×1014 W/cm2. For the measurement of temperature and density, we have used only Ne-like lines. FLYCHK is not showing the transitions corresponding to F-like and O-like ions which could otherwise be used for further investigation of plasma temperature and density. The presence of O-like lines for the laser intensity 1.3×1014 W/cm2, as per Fig. 8(b & c) indicates the plasma temperature is 250 eV (where O-like ions start to become apparent).

As already mentioned our integrated spectrum comes from plasma that has large temperature and density gradients. So, X-ray emission is taking place from hotter regions as well as from colder regions. So, at a laser intensity of 1.3×1014 W/cm2, our plasma temperature is in the range of 150 to 250 eV. When the laser intensity increases, the plasma temperature increases and with it the populations of higher charge states. Dominance of O-like ions over Ne-like and F-like at a laser intensity of 4×1014 W/cm2 indicates that the plasma temperature is of the order of 350 eV from Fig. 8(d) (where O-like starts dominating). So, for the laser intensity 1.3×1014 W/cm2-4×1014 W/cm2, we can say that the plasma temperature ranges from 150 to 350 eV in the region responsible for the bulk of the emission.

|  |
| --- |
| **(a)** |
| **(b)**(a) | **F:\Channprit\Copper paper\Capture1.JPG**IV**(c)** |
| FIG. 7. (a) Experimental spectrum at various laser intensities (b-c) zoom of part a to show comparison of population of Ne-like, F-like and O-like ions at different laser intensities. |

|  |  |
| --- | --- |
| **(a)** | **(b)** |
| **(c)** | **(d)** |
| **(e)** | **(f)** |
| FIG. 8. Distribution of charge states of copper for different temperaturesat an electron density of5×1020 /cm3as obtained from FLYCHK simulations (a) without opacity (b) with opacity. Distribution of Cu XXII (O-like), Cu-XXI (F-like), Cu-XX (Ne-like) ions at the same electron density as obtained from FLYCHK simulations (c) without opacity (d) with opacity. Distribution of charge states of Cu at 1×1021 /cm3as obtained from FLYCHK simulations (e) without opacity (f) with opacity. |

To better understand our results, simulations were performed using the 1-D Lagrangian radiation hydrodynamics simulation code HYADES[44](#_ENREF_44). These simulations used tabulated SESAME equation of state and a multi-group diffusion approximation for radiation transport based on an average atom LTE ionization model. Electron transport was handled by a flux-limited diffusion approximation. The HYADES results show that there is a region of the plasma that corresponds to the combination of temperature and density conditions that have been estimated above. These simulations further highlight the extremely broad range of plasma density and temperature conditions present, showing temperatures ranging from keV in the low-density corona to below 1 eV in the dense target, with electron densities ranging from below 1019 in the corona[[1]](#footnote-2) to above 1024 in the shock wave that is propagating into the solid target. Summing over these many different conditions, with appropriate weighting for optical depths, would be required to more accurately reproduce the experimental spectra.

# CONCLUSIONS

In summary, we have measured copper spectra from laser irradiated copper targets in the intensity range of 1.3×1014 W/cm2 to 4×1014 W/cm2. The identification of spectral lines and atomic calculations were performed using a Multi-configuration Dirac-Fock (MCDF) method. Lines generated using this simulation are in good agreement with the data available and our experimental results. The average temperature and density of the most strongly radiating portion of the plasma is 150-350 eV and 4.5×1020 /cm3 for laser intensities in the range 1.3×1014 - 4×1014 W/cm2. We have performed simulations using the FLYCHK code and performed hydrodynamic simulations which are in good agreement with our experimental results.

**ACKNOWLEDGEMENTS**

Authors from BARC are thankful to Dr. A. K. Mohanty, Director, Physics Group and Dr. M. N. Deo, Head, LS&FTSS for their consistent support. Authors are thankful to Shri D. S. Munda, Mr. Ritesh Sable and Mr. Krishna Bangre for smooth operation of the laser system and for providing support in the data acquisition during the experiment.

1 A. Benuzzi-Mounaix, F. Dorchies, V. Recoules, F. Festa, O. Peyrusse, A. Levy, A. Ravasio, T. Hall, M. Koenig and N. Amadou, Physical Review Letters **107** (16), 165006 (2011).

2 J. Zhang, Y. Xu, J. Yang, G. Yang, H. Li, Z. Yuan, Y. Zhao, G. Xiong, L. Bao and C. Huang, Physics of Plasmas **18** (11), 113301 (2011).

3 A. Rossall, L. Gartside, S. Chaurasia, S. Tripathi, D. Munda, N. Gupta, L. Dhareshwar, J. Gaffney, S. Rose and G. Tallents, Journal of Physics B: Atomic, Molecular and Optical Physics **43** (15), 155403 (2010).

4 L. Antonelli, S. Atzeni, A. Schiavi, S. D. Baton, E. Brambrink, M. Koenig, C. Rousseaux, M. Richetta, D. Batani, P. Forestier-Colleoni, E. Le Bel, Y. Maheut, T. Nguyen-Bui, X. Ribeyre and J. Trela, Physical Review E **95** (6), 063205 (2017).

5 J. Lindl, Physics of Plasmas **2** (11), 3933-4024 (1995).

6 D. Giulietti and L. A. Gizzi, La Rivista del Nuovo Cimento (1978-1999) **21** (10), 1-93 (1998).

7 C. Kaur, S. Chaurasia, A. Poswal, D. Munda, A. Rossall, M. Deo and S. M. Sharma, Journal of Quantitative Spectroscopy and Radiative Transfer **187**, 20-29 (2017).

8 L. P. Presnyakov, Soviet Physics Uspekhi **19** (5), 387 (1976).

9 C. Biedermann, R. Radtke and K. B. Fournier, Physical Review E **66** (6), 066404 (2002).

10 D. Porquet, J. Dubau and N. Grosso, Space Science Reviews **157** (1-4), 103-134 (2010).

11 L. Presnyakov and A. Urnov, Le Journal de Physique Colloques **40** (C7), C7-279-C277-288 (1979).

12 V. Boiko, A. Y. Faenov and S. Pikuz, Journal of Quantitative Spectroscopy and Radiative Transfer **19** (1), 11-50 (1978).

13 U. Safronova, M. Safronova, R. Bruch and L. Vainshtein, Physica Scripta **51** (4), 471 (1995).

14 F. Bely-Dubau, A. Gabriel and S. Volonté, Monthly Notices of the Royal Astronomical Society **186** (3), 405-419 (1979).

15 S. Glenzer, K. Fournier, C. Decker, B. Hammel, R. Lee, L. Lours, B. MacGowan and A. Osterheld, Physical Review E **62** (2), 2728 (2000).

16 O. Renner, M. Šmíd, D. Khaghani and F. Rosmej, presented at the Journal of Physics: Conference Series, 2016 (unpublished).

17 C. Keane, B. Hammel, D. Kania, J. Kilkenny, R. Lee, A. Osterheld, L. Suter, R. Mancini, C. Hooper Jr and N. Delamater, Physics of Fluids B: Plasma Physics **5** (9), 3328-3336 (1993).

18 C. Keane, B. Hammel, A. Osterheld and D. Kania, Physical Review Letters **72** (19), 3029 (1994).

19 C. Back, J. Grun, C. Decker, L. Suter, J. Davis, O. Landen, R. Wallace, W. Hsing, J. Laming and U. Feldman, Physical Review Letters **87** (27), 275003 (2001).

20 J. Rice, K. Fournier, M. Graf, L. Terry, M. Finkenthal, F. Bombarda, E. Marmar and W. Goldstein, Physical Review A **51** (5), 3551 (1995).

21 G. Brown, P. Beiersdorfer, D. Liedahl, K. Widmann and S. Kahn, The Astrophysical Journal **502** (2), 1015 (1998).

22 N. Brickhouse, A. Dupree, R. Edgar, D. Liedahl, S. Drake, N. White and K. Singh, The Astrophysical Journal **530** (1), 387 (2000).

23 E. Träbert, S. Hansen, P. Beiersdorfer, G. Brown, K. Widmann and H.-K. Chung, Review of Scientific Instruments **79** (10), 10E313 (2008).

24 K. Fournier, C. Constantin, C. Back, L. Suter, H.-K. Chung, M. Miller, D. Froula, G. Gregori, S. Glenzer and E. Dewald, Journal of Quantitative Spectroscopy and Radiative Transfer **99** (1-3), 186-198 (2006).

25 M. Shahzad, G. Tallents, A. Steel, L. Hobbs, D. Hoarty and J. Dunn, Physics of Plasmas **21** (8), 082702 (2014).

26 A. Ecuvillon, G. Israelian, N. Santos, M. Mayor, V. Villar and G. Bihain, Astronomy & Astrophysics **426** (2), 619-630 (2004).

27 J. Larour, L. E. Aranchuk, Y. Danisman, A. Eleyan and M. F. Yilmaz, Physics of Plasmas **23** (3), 033115 (2016).

28 M. Swartz, S. Kastner, E. Rothe and W. Neupert, Journal of Physics B: Atomic and Molecular Physics **4** (12), 1747 (1971).

29 D. Batani, A. Giulietti, L. Palladino, G. J. Tallents and I. C. E. Turcu, presented at the ECO4 (The Hague '91), 1991 (unpublished).

30 K. B. Fournier, A. Y. Faenov, T. A. Pikuz, I. Y. Skobelev, V. S. Belyaev, V. I. Vinogradov, A. S. Kyrilov, A. P. Matafonov, I. Bellucci, S. Martellucci, G. Petrocelli, T. Auguste, S. Hulin, P. Monot and P. D’Oliveira, Physical Review E **67** (1), 016402 (2003).

31 H. Gordon, M. Hobby and N. Peacock, Journal of Physics B: Atomic and Molecular Physics **13** (10), 1985 (1980).

32 K. Fournier, A. Y. Faenov, T. Pikuz, I. Y. Skobelev, F. Flora, S. Bollanti, P. Di Lazzaro, D. Murra, A. Grilli and A. Reale, Journal of Physics B: Atomic, Molecular and Optical Physics **35** (15), 3347 (2002).

33 R. Hutcheon, L. Cooke, M. Key, C. Lewis and G. Bromage, Physica Scripta **21** (1), 89 (1980).

34 J. Parkinson, Astronomy and Astrophysics **24**, 215 (1973).

35 I. P. Grant, B. J. McKenzie, P. H. Norrington, D. F. Mayers and N. C. Pyper, Computer Physics Communications **21** (2), 207-231 (1980).

36 J. Cooper, Reports on Progress in Physics **29** (1), 35 (1966).

37 V. Unnikrishnan, K. Alti, V. Kartha, C. Santhosh, G. Gupta and B. Suri, Pramana **74** (6), 983-993 (2010).

38 T. Fujimoto and R. W. P. McWhirter, Physical Review A **42** (11), 6588-6601 (1990).

39 H. K. Chung, M. H. Chen, W. L. Morgan, Y. Ralchenko and R. W. Lee, High Energy Density Physics **1** (1), 3-12 (2005).

40 G. Brown, P. Beiersdorfer, H. Chen, M. Chen and K. Reed, The Astrophysical Journal Letters **557** (1), L75 (2001).

41 A. Osterheld, J. Nilsen, S. Y. Khakhalin, A. Y. Faenov and S. Pikuz, Physica Scripta **54** (3), 240 (1996).

42 S. Bollanti, P. Di Lazzaro, F. Flora, T. Letardi, L. Palladino, A. Reale, D. Batani, A. Mauri, A. Scafati and A. Grilli, Physica Scripta **51** (3), 326 (1995).

43 R. Bruch, U. Safronova, A. Shlyaptseva, J. Nilsen and D. Schneider, Physica Scripta **57** (3), 334 (1998).

44 p. c. l. c. c. HYADES is a commercial product of Cascade Applied Sciences.

1. Note that since this is a 1-D code the densities will be significantly over-estimated in the extremities of the corona. [↑](#footnote-ref-2)