

# From galaxy-scale fueling to nuclear-scale feedback

## The merger-state of radio galaxies 3C 293, 3C 305, and 4C 12.50 (Research Note)

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Received 25 March 2016 / Accepted 12 August 2016

### ABSTRACT

Powerful radio galaxies are often associated with gas-rich galaxy mergers. These mergers may provide the fuel to trigger starburst and active galactic nuclear (AGN) activity. In this Research Note, we study the host galaxies of three seemingly young or re-started radio sources that drive fast outflows of cool neutral hydrogen (HI) gas, namely 3C 293, 3C 305 and 4C 12.50 (PKS 1345+12). Our aim is to link the feedback processes in the central kpc-scale region with new information on the distribution of stars and gas at scales of the galaxy. For this, we use deep optical *V*-band imaging of the host galaxies, complemented with HI emission-line observations to study their gaseous environments. We find prominent optical tidal features in all three radio galaxies, which confirm previous claims that 3C 293, 3C 305, and 4C 12.50 have been involved in a recent galaxy merger or interaction. Our data show the complex morphology of the host galaxies and identify the companion galaxies that are likely involved in the merger or interaction. The radio sources appear to be (re-)triggered at a different stage of the merger; 4C 12.50 is a pre-coalescent and possibly multiple merger, 3C 293 is a post-coalescent merger that is undergoing a minor interaction with a close satellite galaxy, while 3C 305 appears to be shaped by an interaction with a gas-rich companion. For 3C 293 and 3C 305, we do not detect HI beyond the inner ~30–45 kpc region, which shows that the bulk of the cold gas is concentrated within the host galaxy, rather than along the widespread tidal features.

**Key words.** galaxies: active – galaxies: interactions – galaxies: jets – galaxies: starburst – galaxies: evolution – ISM: jets and outflows

## 1. Introduction

Deep optical broadband imaging has revealed that powerful radio galaxies are often associated with gas-rich galaxy mergers or interactions (Ramos Almeida et al. 2012, see also work by Heckman et al. 1986; Smith & Heckman 1989; Roche & Eales 2000; Sabater et al. 2013). Furthermore, powerful radio galaxies often show young stellar populations and contain dust masses that link them to merger activity (Tadhunter et al. 2011, 2014a). These gas-rich mergers and interactions are believed to deposit the cold material that is needed to fuel both starburst activity (e.g., Di Matteo et al. 2007) and the powerful radio sources (e.g., Hardcastle et al. 2007).

Although the majority of early-type galaxies outside clusters contain at least modest amounts of cold gas (few  $\times 10^{6-9} M_{\odot}$ ; Morganti et al. 2006; Oosterloo et al. 2010; Serra et al. 2012), the host galaxies of small radio sources can be particularly rich in HI (Emonts et al. 2007). Moreover, young or recently

re-started radio sources often contain large amounts of HI gas in their central regions (Pihlström et al. 2003; Vermeulen et al. 2003; Gupta & Saikia 2006; Gupta et al. 2006; Chandola et al. 2011, 2013; Geréb et al. 2014). On average, these central regions appear much richer in HI than those occupied by extended radio sources (Gupta et al. 2006; Chandola et al. 2013; Geréb et al. 2014). Moreover, the HI profiles associated with young radio sources more often show blue wings, which suggests that these sources drive gas outflows as they clear their way through the rich ambient medium of the host galaxy (Chandola et al. 2011; Geréb et al. 2015).

Particularly interesting are young or re-started radio sources that have been observed to drive very fast ( $\sim 1000 \text{ km s}^{-1}$ ) outflows of cool neutral and cold molecular gas from the central kpc-scale region (Oosterloo et al. 2000; Morganti et al. 2003, 2005a,b; Dasyra & Combes 2012; Morganti et al. 2013a,b; Mahony et al. 2013). The energy released by these jet-driven outflows approaches values needed to clear the central region of gas and quench star formation (Morganti et al. 2013b; Mahony et al. 2016), certainly when taking into account that part

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**Table 1.** Optical V-band imaging.

Source	$z$	Date	$t_{\text{int}}$ (s)	Airmass
3C 293	0.045	13/03/07	3600 (12 × 300)	1.0
3C 305	0.041	05/04/08	5400 (6 × 900)	1.2
4C 12.50	0.123	06/04/08	5400 (6 × 900)	1.1

**Notes.**  $t_{\text{int}}$  is the on-source integration time (# frames × time per frame).

of the kinetic energy of the jets is put into heating of the cold gas (Guillard et al. 2012; Hardcastle et al. 2012; Tadhunter et al. 2014b; Dasyra et al. 2014; Lanz et al. 2015). Thus, the jet-induced feedback may influence the evolution of these galaxies.

In this paper, we further investigate the link between the nuclear feedback phenomena and large-scale merger processes for three of these seemingly young or re-started radio galaxies with fast outflows of cold gas, namely 3C 293, 3C 305 and 4C 12.50 (PKS 1345+12). The cold outflows in these systems were identified through blueshifted HI seen in absorption against the strong radio continuum (Morganti et al. 2003, 2005a, 2013a; Mahony et al. 2013). Previous optical imaging revealed evidence for tidal debris from a galaxy merger or interaction in these systems (Heckman et al. 1986). We here present more detailed optical imaging, combined with optical spectroscopy and deep HI observations. Our aim is to identify the stage of merger or interaction in these systems. With this, we complement the studies of the feedback processes in the kpc-scale central region with new information on the distribution of stars and gas at scales of the host-galaxy environment.

Throughout this paper we will assume  $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_M = 0.27$  and  $\Omega_\Lambda = 0.73$ .

## 2. Observations

### 2.1. Optical imaging

Deep optical V-band imaging of the three radio galaxies was done in March 2007 and April 2008 and at the Hiltner 2.4 m telescope of the Michigan-Dartmouth-MIT (MDM) observatory, Kitt Peak, Arizona (USA). Imaging was done using the Echelle CCD, with a field-of-view (F.o.V.) of  $9.5' \times 9.5'$ . Observations were taken with a seeing of 1–2". No absolute flux calibration was performed. Table 1 summarizes the observations.

The Image Reduction and Analysis Facility (IRAF; Tody 1993) was used for a standard data reduction (bias subtraction, flat-fielding, frame alignment and cosmic-ray removal). A background gradient, which was most likely introduced by minor shutter issues, had to be subtracted from the final images (see Emonts et al. 2010). Using the Karma software (Gooch 1996), a world coordinate system (accurate to within 1 arcsec) was copied from a Sloan Digital Sky Survey (SDSS) image of the same region.

An optical spectrum with a single 300 s exposure was obtained for the object seen in the vicinity of 3C 293 (Sect. 3.1) using the Intermediate dispersion Spectrograph and Imaging System (ISIS) with R300B/R600R gratings on the *William Herschel* Telescope (WHT) on 29 June 2006. The sole purpose of this spectrum was to identify the object as a companion galaxy and to determine its redshift, hence a standard wavelength calibration was applied, but no absolute flux scale.

### 2.2. Radio HI spectroscopy

Deep 21 cm radio observations were taken with the Westerbork Synthesis Radio Telescope (WSRT) to search for HI emission in 3C 293 and 3C 305. At  $z = 0.12$ , 4C 12.50 is too distant to image HI in emission with the WSRT. 3C 293 was observed on 14, 15, 16, 19 May 2007 for a total of  $4 \times 12$  h, 3C 305 on 29 June and 3 July 2010 for  $2 \times 12$  h. We used the 20 MHz band, with for 3C 293 a double IF with 512 channels ( $\Delta v \sim 8 \text{ km s}^{-1}$ ) and for 3C 305 a single IF with 1024 channels ( $\Delta v \sim 4 \text{ km s}^{-1}$ ). A primary calibrator (3C 286 or CTD 93) was used to calibrate the bandpass every hour to reach the required 1:10 000 spectral dynamic range for detecting weak HI emission in the vicinity of the strong radio continuum cores (3.7 for 3C 293 and 3.0 Jy for 3C 305). The flux scale was set using 3C 286.

The data reduction was done using MIRIAD (Sault et al. 1995). After flagging, a continuum image was made by first fitting a straight line to the signal in the channels that do not contain any HI line, and subsequently Fourier transforming, self-calibrating and cleaning these continuum data. A continuum image from earlier WSRT data of 3C 293 was presented in Emonts et al. (2005). The continuum source of 3C 305 is unresolved. The calibration solutions of the continuum data were then transferred to the original “line+continuum” data, while the continuum model was subtracted from this data set in the uv-plane. Any residual continuum was subtracted from the line data by fitting a straight line to the line-free channels. For 3C 293, any remaining phase and amplitude errors were corrected by a selfcalibration on the strong ( $\sim 200$  mJy) absorption line in the corresponding channels. We then Fourier transformed the line-data using robust +0.5 weighting (Briggs 1995) and cleaned it using a mask (created by smoothing the data) to isolate the regions in which HI emission or absorption was found.

The resulting data of 3C 293 and 3C 305 have a beam-size of  $34.9'' \times 20.5''$  (PA  $-1.0^\circ$ ) and  $20.9'' \times 19.8''$  (PA  $1.1^\circ$ ), respectively. After binning and Hanning smoothing the data, we reached a thermal-noise limited sensitivity of  $0.32 \text{ mJy beam}^{-1}$  per  $8.2 \text{ km s}^{-1}$  channel for 3C 293 and  $0.36 \text{ mJy beam}^{-1}$  per re-binned  $17.8 \text{ km s}^{-1}$  channel for 3C 305.

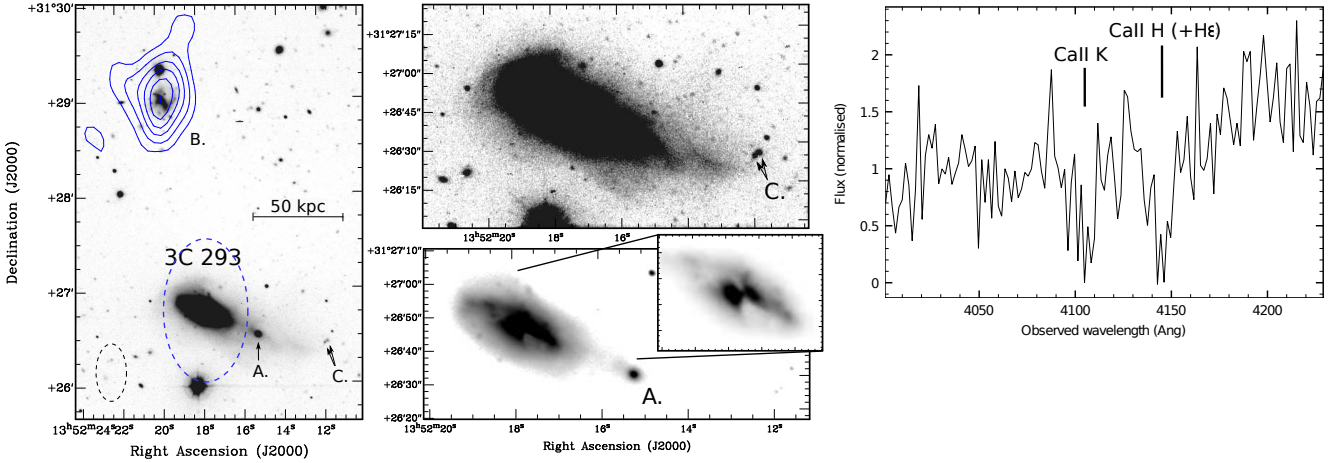
## 3. Results

Figures 1–3 show the V-band imaging of 3C 293, 3C 305 and 4C 12.50. Various features are presented with different scaling. For 3C 293 and 3C 305, the HI results are also included.

### 3.1. 3C 293

Figure 1 shows the V-band image of 3C 293. The high-surface-brightness inner region displays a disc-like morphology, but with prominent dust-lanes (better visualized in the HST imaging by Beswick et al. 2002). The low-surface-brightness envelope shows a distorted morphology, and a faint bridge connecting 3C 293 with the nearby small companion “A” at 13 kpc west-south-west of the nucleus of 3C 293. This small companion shows an elliptical, nearly unresolved, morphology. To confirm this object as a true companion, we obtained an optical spectrum (Fig. 1, right). By fitting a Gaussian profile to the Ca II K and H features, we find that companion A has a velocity of  $13171 \pm 260$ . This corresponds to  $v = -279 \pm 260 \text{ km s}^{-1}$  with respect to the systemic velocity of 3C 293, which is  $v_{\text{sys}} = 13450 \text{ km s}^{-1}$  (Emonts et al. 2005).

The low-surface-brightness envelope of the optical emission around 3C 293 is extended in the direction of this companion A.



**Fig. 1.** Optical *V*-band imaging of 3C 293. *Left+middle*: the various plots show the same image of 3C 293, but with different intensity-scaling and zooming. On the *left* is shown the larger environment of 3C 293, with overlaid in blue the contours of HI 21 cm emission (levels: 0.04, 0.07, 0.10, 0.13, 0.16, 0.19  $\text{Jy km}^{-2} \times \text{km s}^{-1}$ ). The dashed blue ellipse represents the region where the HI absorption dominates (see Morganti et al. 2003). Even though this HI absorption is spatially unresolved, it is so strong that it dominates over any potential HI emission far beyond the FWHM of the synthesized beam (dashed black ellipse on the *bottom-left*). The companion galaxies A and B, and feature C, are described in the text. *Right*: optical spectrum of source A. The Ca II K and H absorption lines confirm that this source is a companion of 3C 293.

About 70 kpc further to the west (beyond companion A) the low-surface-brightness tail shows a “zig-zag” pattern and ends in two knots (indicated as “C” in the top-right plot of Fig. 1). It is interesting to speculate that these two knots could represent star forming regions at the tip of the low-surface-brightness tail.

No HI emission is detected, down to a  $5\sigma$  limit of  $7.0 \times 10^{19} \text{ cm}^{-2}$  across  $100 \text{ km s}^{-1}$ , outside the central  $\sim 45 \text{ kpc}$  region where the HI absorption-signal dominates. Because of the strong HI absorption, any superposed weak HI emission cannot be distinguished within the central region. Part of the HI absorption traces a central disk (Baan & Haschick 1981; Beswick et al. 2002, 2004), which was imaged in CO (Evans et al. 1999; Labiano et al. 2014) and ionized gas (Emonts et al. 2005; Mahony et al. 2016). A fast outflow of both HI and ionized gas was reported by Morganti et al. (2003) and Emonts et al. (2005), and studied in detail against the inner kpc-scale radio source by Mahony et al. (2013, 2016).

Another companion “B” is found at 120 kpc distance north of 3C 293 and detected in HI emission (with  $M_{\text{HI}} = 1.6 \times 10^9 M_{\odot}$  and  $v = 13416 \text{ km s}^{-1}$ ; Emonts 2006). This companion appears to contain a face-on starforming disc with a prominent bar. There is no indication from the optical or HI morphology that this companion is interacting with 3C 293.

### 3.2. 3C 305

Figure 2 shows the *V*-band image of 3C 305. The low-surface-brightness image shows two tidal arms stretching in east-west direction. The major axis of the galaxy seems to turn from  $\sim 145^\circ$  in the low-surface-brightness plot (top-middle) to  $\sim 70^\circ$  in the high-surface-brightness plot (bottom-middle). The inner region also shows a distorted morphology, with arc- or spiral-like isophotes. 3C 305 hosts a single nucleus (Jackson et al. 2003).

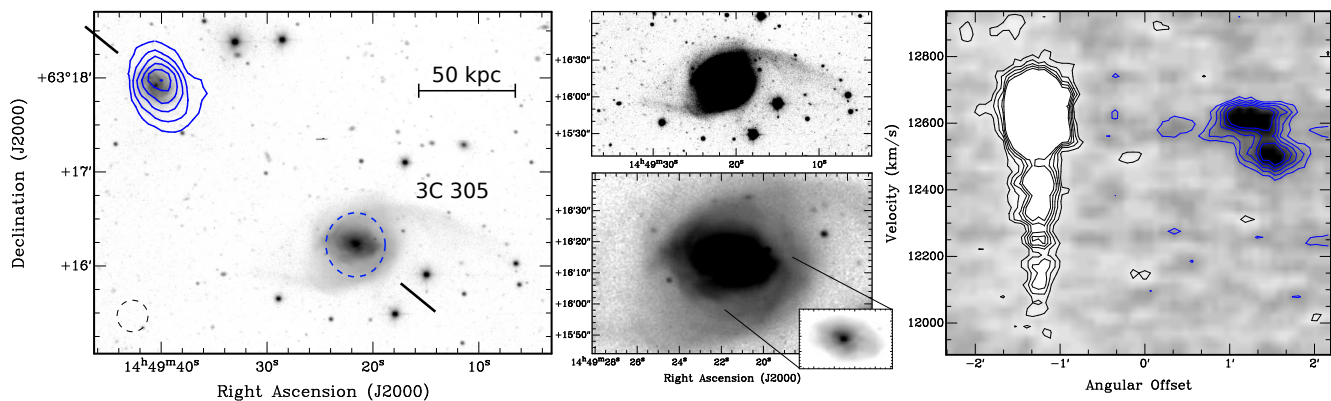
No HI is detected in emission, down to a  $5\sigma$  limit of  $2.1 \times 10^{20} \text{ cm}^{-2}$  across  $100 \text{ km s}^{-1}$ , outside the central  $\sim 30 \text{ kpc}$  region where the HI absorption-signal dominates. As for 3C 293, also in 3C 305 part of this HI absorption likely traces an inner gas disc (Jackson et al. 2003), while another part is a fast gas-outflow detected against the inner radio-jet at kpc-distance from the nucleus (Morganti et al. 2005a).

North-east of 3C 305, at 133 kpc distance, a companion galaxy is detected in HI emission ( $M_{\text{HI}} = 4.2 \times 10^9 M_{\odot}$ ). The HI velocity of this companion is  $v = 12560 \text{ km s}^{-1}$ . This is very close to  $v_{\text{sys}} = 12550 \text{ km s}^{-1}$  for 3C 305 (Morganti et al. 2005a). Part of the HI appears marginally elongated in the direction of 3C 305, as shown by the lowest contours in the position-velocity plot of the HI gas in Fig. 2 (right). If confirmed, it would suggest that an interaction occurred between 3C 305 and this gas-rich companion. It is interesting to note that the axis of the kpc-scale radio source in 3C 305 is aligned not only parallel to the nuclear dust lane (Jackson et al. 2003), but also in the direction of this gas-rich companion. This is reminiscent of similar alignments seen in the nearby radio galaxies NGC 612 and B2 0722+30 (Emonts et al. 2009). However, we cannot rule out that this is a chance-alignment and the radio jet in 3C 305 is merely perpendicular to the large-scale disc (as seen in other radio galaxies by Kaviraj et al. 2015).

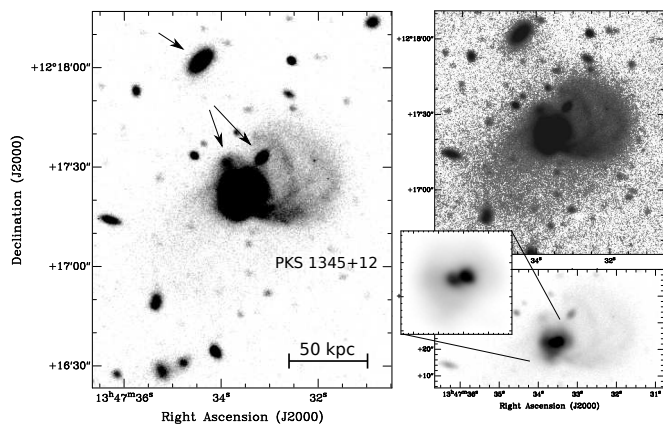
### 3.3. 4C 12.50 (PKS 1345+12)

Figure 3 shows the *V*-band image of 4C 12.50 (PKS 1345+12). It reveals an apparent multiple merger system. The core of 4C 12.50 consists of a double nucleus, with a separation of  $1.6''$  or 3.5 kpc (as previously found by Heckman et al. 1986; and Evans et al. 1999). At the south-western end of the host galaxy, a double tidal-tail stretches  $\sim 60 \text{ kpc}$  north, encircling roughly 30% of the host galaxy. This double tidal-tail stretches in the direction of an apparent bright companion  $\sim 92 \text{ kpc}$  north of 4C 12.50. On the opposite side of 4C 12.50 there is a broad fan of emission that stretches  $\sim 85 \text{ kpc}$  towards the south-east. 4C 12.50 resides in a dense environment. Two apparent close companion galaxies lie  $\sim 20 \text{ kpc}$  north of the host galaxy and may be involved in the merger (see arrows in Fig. 3). However, spectroscopic redshifts are required to confirm their association with 4C 12.50.

At  $z = 0.12$ , 4C 12.50 is too distant to image HI in emission with the WSRT, but the central HI absorption has been studied in great detail (Mirabel 1989; Morganti et al. 2004, 2013a). Against the tip of the southern pc-scale radio jet, VLBI observations revealed a very broad HI absorption feature, which is blueshifted by  $\sim 1000 \text{ km s}^{-1}$  (Morganti et al. 2013a). This broad



**Fig. 2.** Optical V-band imaging of 3C 305. *Left+middle:* the various plots show the same image of 3C 305, but with different intensity-scaling and zooming. On the *left* is shown the larger environment of 3C 305, with overlaid in blue the contours of HI 21 cm emission (levels: 0.03, 0.09, 0.15, 0.21, 0.28, 0.33  $\text{Jy beam}^{-1} \times \text{km s}^{-1}$ ). The dashed blue circle shows the region where the HI absorption dominates (see [Morganti et al. 2005a](#)). Even though this HI absorption is spatially unresolved, it is strong enough to dominate over any potential HI emission beyond the FWHM of the synthesized beam (dashed black circle on the *bottom-left*). *Right:* position-velocity plot of the HI emission (blue contours) and absorption (black contours) along the axis shown in the *left plot*. Contour levels are at  $-3, -2, 2, 3, 4, 5, 6 \times \sigma$ , with  $\sigma = 0.36 \text{ mJy beam}^{-1}$ .



**Fig. 3.** Optical V-band imaging of 4C 12.50 (PKS 1345+12). The various plots show the same image of 4C 12.50, but with different intensity-scaling and zooming. Potential companion galaxies discussed in the text are marked with arrows in the *left plot*.

feature has also been seen in CO by [Dasyra & Combes \(2012\)](#). In addition, a deep HI absorption component occurs against the northern VLBI jet, and could represent a cloud within the central disc that was detected in CO by [Evans et al. \(1999\)](#). Diffuse, faint radio emission is found at larger scales, which is likely a relic source from a past episode of activity ([Stanghellini et al. 2005](#)).

#### 4. Discussion and conclusions

The deep optical imaging that we presented of 3C 293, 3C 305, 4C 12.50 reveals galaxy-morphologies that agree with those previously shown by [Heckman et al. 1986](#) (also [Heckman et al. 1982](#); [van Breugel et al. 1984](#)). However, our data show in more detail a multitude of subtle features. Specifically, combined with our optical spectroscopy and HI imaging, we obtain the following new insights on these three radio galaxies:

- 3C 293: we confirm the presence and redshift of a close companion at  $\sim 13$  kpc distance, beyond which a stellar tail with a “zig-zag” morphology ends in two possible starforming regions.

- 3C 305: the major axis of the host galaxy seems to turn by  $\sim 75^\circ$  from the faint outer envelope with the two tidal arms to the bright inner region. The inner region shows arc- or spiral-like isophotes. There is a tentative indication that a faint HI feature stretches from an HI-rich companion towards 3C 305, along the direction of the radio axis.
- 4C 12.50: a double tidal-tail stretches in the direction of a bright companion galaxy. On the opposite side, a faint but broad fan of emission stretches up to  $\sim 85$  kpc from the galaxy.

Our results also reveal that, for 3C 293 and 3C 305, no HI has been detected along the stellar tidal debris beyond the central region where the HI absorption-signal dominates. However, we note that sensitive HI observations with a spatial resolution that better matches the extended optical features are needed to better study any potential widespread HI gas in emission.

##### 4.1. Merger state and AGN triggering

Our results indicate that all three systems have been involved in a recent galaxy merger or interaction, which is consistent with the presence of young or post-starburst stellar populations across these systems ([Tadhunter et al. 2005](#)). 3C 305 was classified by [Heckman et al. \(1986\)](#) as a post-merger system. However, contrary to the claim by [Heckman et al. \(1986\)](#) that there are no close companion galaxies that could have provoked the observed morphological disturbances, we show that 3C 305 appears to be involved in a galaxy-encounter with a gas-rich companion. This resembles the case of NGC 6872, where two tidal arms were formed by the low-inclination, prograde passage of companion IC 4970 ([Horellou & Koribalski 2007](#); see also [Mihos et al. 1993](#)). If a similar passage created the tidal arms in 3C 305, and assuming that the companion has an average velocity of  $\sim 200 \text{ km s}^{-1}$  with respect to 3C 305, then the time since closest approach must have occurred  $\sim 650$  Myr ago. This time-scale is similar to the 0.5–1 Gyr age of the young stellar population across 3C 305 ([Tadhunter et al. 2005](#)). It could suggest that a galaxy-wide starburst was triggered at the time of closest approach between the two systems, which is consistent with simulations (e.g., [Peirani et al. 2010](#)). Additional observations are needed to investigate this. For 3C 293, the distorted optical morphology and single nucleus ([Floyd et al. 2006](#)) suggest that this

is a post-merger system which is currently experiencing a minor interaction with the close companion 13 kpc towards the west-south-west. 4C 12.50 shows clear evidence from its double nucleus and multiple tidal features that it is currently involved in an ongoing major – possibly multiple – merger event. This is consistent with its high IR luminosity ( $L_{\text{IR}} \sim 2 \times 10^{12} L_{\odot}$ , that is, in the regime of ultra-luminous infrared galaxies; Sanders et al. 1988). These results indicate that 3C 305, 3C 293 and 4C 12.50 are likely at different stages in the merger process.

Interestingly, this also implies that the radio sources in 4C 12.50, 3C 305 and 3C 293 could have been triggered at different stages during the merger. After all, 4C 12.50, 3C 305 and 3C 293 are classified as a Gigahertz Peaked Spectrum (GPS), Compact Steep Spectrum (CSS) and Steep Spectrum Core (SSC) source, respectively, which likely represent young or re-started radio sources ( $\leq 10^6$  yr; Fanti et al. 1995; O’Dea 1998; Shulevski et al. 2012; Collier et al. 2016). This different stage of triggering would be consistent with optical and IR studies, which indicate that a radio source can in general occur at any time during the merger process, either before or after the individual nuclei coalesce (Ramos Almeida et al. 2011; Tadhunter et al. 2011; Dicken et al. 2012). We caution that, while both 3C 293 and 4C 12.50 have radio cores with an estimated age of  $\sim 3 \times 10^4$  yr (Akujor et al. 1996; O’Dea et al. 2000), they also contain extended radio lobes (e.g., Beswick et al. 2004; Stanghellini et al. 2005). Although we assume that these outer lobes represent a previous episode of activity, we cannot rule out that the radio cores are older than they appear from their spectral-index measurements (see Blundell & Rawlings 2000), for example, because they have been frustrated by a dense inter-stellar medium (e.g., O’Dea et al. 1991). If the radio cores are actively feeding the much larger and older radio structures, then the moment of triggering is uncertain. However, we find this scenario less likely, because 3C 293, 3C 305 and 4C 12.50 are among the most powerful steep-spectrum sources at low- $z$  and display among the most extreme jet-ISM interactions known in the low- $z$  Universe. Such features are often associated with young or re-started radio sources that are trying to plough their way through the ambient ISM (Sect. 1; see also Holt et al. 2008, 2011).

Concluding, our results support the notion that galaxy mergers and interactions are the likely mechanisms for accumulating cold gas in the central few kpc of these systems, and possibly also for (re-)triggering the powerful radio sources that violently interact with this dense ISM.

*Acknowledgements.* We thank Jacqueline van Gorkom for her help and useful discussions. We also thank the referee, Stanislav Shabala, for useful feedback that improved this paper. B.E. thanks Columbia University for its hospitality during part of this project. The research leading to these results has received funding from the People Programme (Marie Curie Actions) of the European Union’s Seventh Framework Programme FP7/2007-2013/ under REA grant agreement No. 624351. R.M. gratefully acknowledges support from the European Research Council under the European Union’s Seventh Framework Programme (FP/2007-2013)/ERC Advanced Grant RADIOLIFE-320745. M.V.M. acknowledges support from the Spanish Ministerio de Economía y Competitividad through the grant AYA2012-32295. E.K.M. acknowledges support from the Australian Research Council Centre of Excellence for All-sky Astrophysics (CAASTRO), through project number CE110001020. Data were obtained using the 2.4 m Hiltner Telescope of the Michigan-Dartmouth-MIT (MDM) Observatory, owned and operated by a consortium of the University of Michigan, Dartmouth College, Ohio State University, Columbia University and Ohio University. The Westerbork Synthesis Radio Telescope is operated by ASTRON (Netherlands Institute for Radio Astronomy) with support from the Netherlands Foundation for Scientific Research (NWO). The *William Herschel* Telescope is operated on the island of La Palma by the Isaac Newton Group in the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias.

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